

First Search for Dark Photon Dark Matter with a MADMAX Prototype

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We report the first result from a dark photon dark matter search in the mass range from 78.62 to 83.95 $\mu\text{eV}/c^2$ with a dielectric haloscope prototype for MADMAX (Magnetized Disc and Mirror Axion eXperiment). Putative dark photons would convert to detectable photons within a stack consisting of three sapphire disks and a mirror. The emitted power of this system is received by an antenna and successively digitized using a low-noise receiver. No significant signal attributable to dark photons has been observed above the expected background. Assuming unpolarized dark photon dark matter with a local density of $\rho_\chi = 0.3 \text{ GeV cm}^3$ we exclude a dark photon to photon mixing parameter $\chi > 2.7 \times 10^{-12}$ over the full mass range and $\chi > 1.1 \times 10^{-13}$ at a mass of 80.57 $\mu\text{eV}/c^2$ with a 95% confidence level. This is the first physics result from a MADMAX prototype and exceeds previous constraints on χ in this mass range by up to almost three orders of magnitude.

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Introduction—The nature of dark matter (DM) might be the greatest unsolved mystery of particle physics and cosmology. Many experiments try to directly detect DM particles from the galactic halo in the laboratory. In recent times, very lightweight DM candidates with masses below $\sim 1 \text{ meV}/c^2$ have received increased attention [1]. One

such wavelike DM candidate is the dark photon (DP), also known as hidden photon [2,3]. DPs appear in extensions of the standard model (SM) that postulate an additional $U_\chi(1)$ gauge symmetry. If this symmetry is broken, the DP gains a mass m_χ . Assuming that the SM fields remain uncharged under the new $U_\chi(1)$, DPs would predominantly interact with SM particles via kinetic mixing with the ordinary photon. Its mass and small interaction strength with SM fields makes the DP a suitable DM candidate. Dark photon dark matter (DPDM) production from inflationary perturbations with $m_\chi \sim 100 \mu\text{eV}/c^2$ would easily saturate the observations [4,5] and its discovery could pinpoint the scale of inflation. In the low energy limit, the interaction between DPs and photons can be described by additional source terms in the classical Maxwell's equations.

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In particular, Ampère's law is modified to [6,7]

$$\nabla \times \mathbf{H} - \dot{\mathbf{D}} = \chi \epsilon_0 (\dot{\mathbf{E}}_\chi - c^2 \nabla \times \mathbf{B}_\chi), \quad (1)$$

where $(\mathbf{E}_\chi, \mathbf{B}_\chi)$ are the DP electric and magnetic fields, χ is the kinetic mixing angle, and $(\mathbf{H}, \mathbf{D})$ the magnetic and displacement field of ordinary electrodynamics. For DPDM, the spatial derivative $\nabla \times \mathbf{B}_\chi$ can be neglected as the de Broglie wavelength is much larger than the setup [8,9] whereas the temporal derivative $\dot{\mathbf{E}}_\chi$ acts as an effective current density that oscillates with frequency $\nu_\chi \approx m_\chi c^2/h$. The oscillating DP electric field drives the free charges inside a conductor which leads to emissions of ordinary photons perpendicular to its boundary. The power emitted of a metallic mirror is proportional to its area A and the local DPDM density $\rho_\chi = (\epsilon_0/2)|\mathbf{E}_\chi|^2$ [10],

$$P_0 = \chi^2 c \rho_\chi A \alpha_{\text{pol}}^2. \quad (2)$$

The factor α_{pol} accounts for the DP polarization and is the average fraction of DPs that the experiment is sensitive to. We assume unpolarized DPDM, setting $\alpha_{\text{pol}}^2 = 1/3$ as we are sensitive only to a single polarization [2].

In this Letter, a search for DPDM using a MADMAX (MAGnetized Disk and Mirror AXion eXperiment) prototype setup is presented. MADMAX is designed to search for axions and axionlike particles in an external static B field in the mass range of 40 to 400 $\mu\text{eV}/c^2$ [11–13]. With no B field applied, the setup is still sensitive to DPs. The proposed mass range is difficult to access using traditional designs like cavity haloscopes because the effective volume where DPs (or equivalently axions) convert into photons naively scales as $\lambda^3 \propto 1/m_\chi^3$. This makes large, and therefore sensitive, conversion volumes hard to achieve. The dielectric haloscope concept of MADMAX removes the dependency of the conversion volume from the wavelength by placing dielectric disks in front of a metallic mirror [14]. This concept has already been successfully employed in searches for DPDM at lower [15] and higher [16] DP masses. The axion or DP-induced photon emission of the disks and mirror, collectively called the booster, can constructively interfere and resonate between disks. The overall increase in expected DP signal power w.r.t. a single mirror is described by the boost factor $\beta^2 = (P_{\text{sig}}/P_0)$, where P_{sig} is the signal power received by the first preamplifier of the receiver system. By controlling the spacing between disks, the boost factor can be tuned in frequency [17], allowing for both broadband and resonant searches. This enables scanning for a broad range of possible DP masses. In this Letter, a single fixed booster configuration at room temperature without scanning capability is used. The expected sensitivity to χ from a DP signal with bandwidth $\Delta\nu_\chi \approx 10^{-6}\nu_\chi$ can be expressed via Dicke's radiometer equation [18] as

$$\begin{aligned} \chi = & 1.0 \times 10^{-13} \left(\frac{640}{\beta^2} \right)^{1/2} \left(\frac{707 \text{ cm}^2}{A} \right)^{1/2} \\ & \times \left(\frac{T_{\text{sys}}}{240 \text{ K}} \right)^{1/2} \left(\frac{11.7 \text{ d}}{\Delta t} \right)^{1/4} \left(\frac{\text{SNR}}{5} \right)^{1/2} \\ & \times \left(\frac{0.3 \text{ GeV cm}^3}{\rho_\chi} \right)^{1/2} \left(\frac{\Delta\nu_\chi}{20 \text{ kHz}} \right)^{1/4}, \end{aligned} \quad (3)$$

where T_{sys} is the system noise temperature, Δt the effective data-taking time, and SNR is the signal to noise ratio of a hypothetical DP signal. Here, DPs are considered to comprise all of the local dark matter density assuming $\rho_\chi = 0.3 \text{ GeV cm}^3$ [19]. Both β^2 and T_{sys} are functions of frequency that depend on the booster configuration.

Experimental setup—A schematic of the setup is shown in Fig. 1. Three sapphire disks with relative permittivity $\epsilon = 9.3(1)$ and thickness 1.00(2) mm as well as an aluminum mirror make up the booster of the dielectric haloscope. The disks and mirror have a diameter of 30 cm each and are held parallel to each other by mechanical spacers in a fixed configuration. The planarity of the three disks has been measured with the methodology presented in [20] and gives an RMS of 50 μm (200 μm min-max). The DP-induced emissions are coupled to a K-band Gaussian beam horn antenna via an off-axis ellipsoidal mirror of focal length $f = 504 \text{ mm}$. The horn antenna's position and orientation are precisely controlled using motorized stages. A low-noise amplifier (LNA) is connected

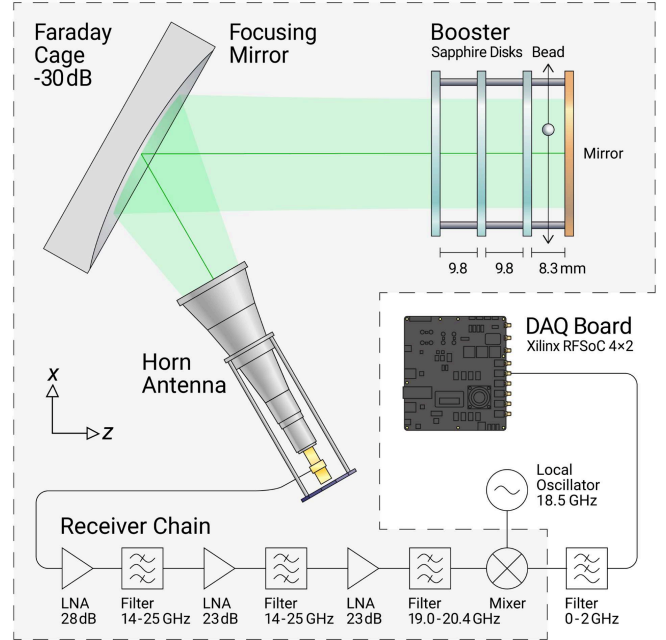


FIG. 1. Experimental setup. It is placed in an RF-isolated Faraday cage indicated by the dashed line. To determine the boost factor, a bead can be inserted into the booster to measure the electric field induced by a reflection measurement. Sketch not to scale.

directly to the horn antenna. Additional LNAs and bandpass filters are connected in series to further amplify the DP signal. It is then mixed down to the intermediate frequency (IF) band using a local oscillator (LO) and mixer. After a final low-pass filter, the signal is digitized with an FPGA-based DAQ using a Xilinx RFSoc4x2, analogous to the system used in [21]. The resulting IF spectrum spans from 0 to 2.4576 GHz with a resolution bandwidth of $\Delta\nu = 9.375$ kHz. The sensitive frequency range is set by the bandwidth of the narrowest band pass filter, which results in an RF range of 19.01 to 20.30 GHz.

The expected DP signal power P_{sig} is proportional to the boost factor β^2 . It can be determined directly from measurement using a recently developed method [22] which relates the boost factor, primarily defined for the unknown DP-induced field, to the reflection-induced electric field E_R that is excited by a vector network analyzer (VNA). This field is measured using the nonresonant bead-pull method [23] where small changes in the booster reflection coefficient Γ are related to the electric field at the bead's position. The general procedure to determine β^2 from bead-pull measurements is described in detail in [24,25]. Expressed in terms of measurable quantities, the boost factor is

$$\beta^2 = \frac{P_{\text{sig}}}{P_0} = \frac{4\pi^2\epsilon_0\nu^2 F_{\text{RC}}}{8cP_{\text{in}}A} \left| \int dV E_R \right|^2, \quad (4)$$

where E_R is excited with input power P_{in} at frequency ν and is spatially integrated over the conversion volume. F_{RC} is a dimensionless factor that accounts for an impedance mismatched receiver system. Bead-pull measurements yield E_R relative to P_{in} such that its actual value is not required. We arbitrarily set $P_{\text{in}} = 1$ W for convenience. Figure 2 shows the longitudinal (a) and transverse field distribution (b) inside the booster obtained from bead-pull measurements as well as the derived boost factor as a function of frequency.

The measured field is first integrated in the transverse direction. The absolute value of the integrated electric field is then fitted separately for each frequency by a model that takes the finite size of the bead into account, shown in Fig. 2(a). The shaded band indicates the model uncertainty from geometry and material parameters of both booster and bead. The fit allows us to obtain the deconvoluted field, i.e., the field without the response of the bead, to interpolate between measurements, and to recover the phase. The remaining integration in the longitudinal direction is then performed with the deconvoluted field, yielding β^2 via Eq. (4), shown in Fig. 2(b). On the main resonance at 19.48 GHz, the transverse field is the fundamental Gaussian mode with waist radius $w_0 = 91(6)$ mm that is expected from the optical system and which has a good overlap with the uniform E_x . A mix of higher-order transverse modes can also resonate inside the booster as

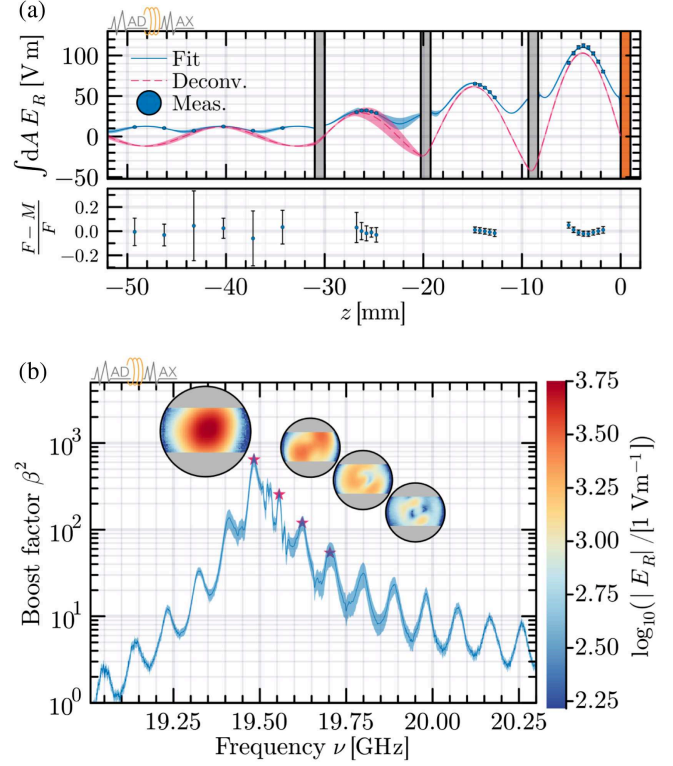


FIG. 2. Results of the boost factor determination using the bead-pull method. (a) Transverse integrated electric field along the optical axis inside the booster at 19.48 GHz. Measurements (circles) fitted with a model (solid line) including the bead. The dashed line shows the field evaluated without the effect of the bead. The shaded band indicates model uncertainties from material and geometry parameters. The relative difference between fit (F) and measurement (M) is shown in the lower panel. (b) Boost factor as a function of frequency, including the corrections for the finite domain and receiver mismatch. The insets show the transverse electric field between the mirror and the first disk at the indicated frequencies (stars), where gray areas indicate regions unprobed due to mechanical constraints.

well as between the antenna and booster, causing the additional smaller peaks, first studied in simulation in [26].

The finite domain of the bead-pull measurements does not cover the full transverse extent of E_R due to mechanical constraints from the booster leaving some of the top and bottom fringe areas of the disks unprobed [indicated by the gray areas in Fig. 2(b)]. We account for these areas in the integral in Eq. (4) by extrapolating the electric field. This increases the naive estimate of β^2 by about 70%. The finite domain correction has been checked against independent measurements with full coverage of the field [25] and contributes an additional 10% to the systematic uncertainty of β^2 .

The receiver chain has an impedance mismatch to the antenna, resulting in an input reflection coefficient of $|\Gamma_{\text{RC}}| \sim 0.25$. Since part of the expected DP signal would be resonating between receiver chain and booster, the boost

factor is further changed by the factor $F_{RC} = [(1 - |\Gamma_{RC}|^2) / (|1 - \Gamma_{RC}\Gamma|^2)]$. This resonance also affects the system noise temperature T_{sys} , which is determined by a Y-factor calibration of the receiver chain [27], establishing its gain and equivalent noise temperature T_e . The overall system noise temperature appears as a standing-wave pattern with a peak-to-peak system temperature of 120 to 332 K over the measurement span of 1.2 GHz, which is dominated by the receiver chain noise of $T_e = 119(28)$ K interfering with itself. This is compatible with the expected noise of the first stage amplifier of around 120 K [28]. Only at the resonance frequency of the booster does its thermal radiation due to physical temperature significantly contribute to the system noise temperature by around 40%. The modulation of T_{sys} is visible in the received power excess in Fig. 3, where a higher T_{sys} leads to larger fluctuations around the baseline.

To properly quantify F_{RC} , the difference in electrical length between the measurements of Γ and Γ_{RC} needs to be known. It is extracted from fitting a 1D model to the system noise temperature, taking Γ_{RC} and Γ as input and simulating the emitted LNA noise as per [29]. The resulting value of 16.016(2) mm matches the length of the adapter used to connect the receiver chain to the antenna. Correcting for the mismatch increases the systematic uncertainty of β^2 by less than 1%. Depending on frequency, F_{RC} ranges between 0.6 and 1.6, and can consequently increase or decrease the boost factor. By tuning the distance between the antenna and booster, it is assured that on resonance frequency the boost factor is increased by $F_{RC} \approx 1.3$ to a maximum value of $\max \beta^2 = 640(110)$. The boost factor shown in Fig. 2(b) includes both corrections from the finite domain of bead-pull measurements and receiver chain mismatch. The shaded band indicates the one standard deviation uncertainty which ranges from 10 to 36%, depending on frequency.

Data-taking run—The data-taking run lasted 16.5 days from December 22, 2023, to January 8, 2024. In this period,

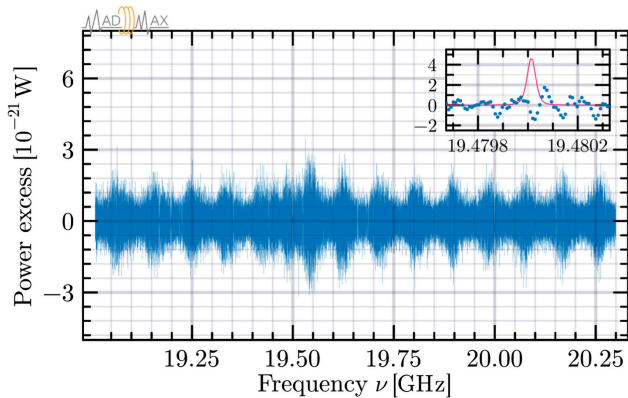


FIG. 3. Observed cross-correlated power excess as a function of frequency. The inset shows a zoomed-in view around 19.48 GHz where the maximum boost factor occurs. A hypothetical DP signal with $\chi = 2 \times 10^{-13}$ (magenta solid line) is superimposed on observations (blue circles).

$N_{tot} = 9.484 \times 10^9$ spectra were measured corresponding to an effective data-taking time of $\Delta t = 11.7$ d. The spectra were averaged in batches of $N_{av} = 4.00 \times 10^6$ and saved every 10 minutes, resulting in 2371 files. The setup was located in a shielded laboratory (SHELL) at DESY/University of Hamburg with coordinates: 53.58N/9.89E. The antenna is sensitive only to the zenith-pointing polarization of DPDM. The optical axis of the booster has an azimuth angle of 65° relative to north.

The data-taking routine employed the same LO hopping scheme as described in [21] to smear and suppress radio frequency interference (RFI) in the IF band. Approximately once per second (every 10^4 spectra), the LO frequency is randomly shifted within a window of ~ 9 MHz. The collected power spectra are then realigned in the RF band before averaging. This results in a smearing of any RFI in the IF band, effectively suppressing it without affecting signals in the RF band. Additionally, roughly 50 bins that contain strong RFI signals in the IF band are digitally masked. The frequency hopping also causes these masks to spread in the RF band, rendering their impact on sensitivity negligible.

Communication delays between DAQ and LO are the dominant contribution to the dead time of $\sim 30\%$ and could be improved in the future. In the RF band, RFI signals are further suppressed by ~ 30 db using an additional Faraday cage and RF absorbers. The response of these signals to additional shielding excludes a DP origin. They are removed from the analysis and are listed in Supplemental Material [30]. A plausible source is the surface-monitoring radar of the nearby Hamburg airport.

Analysis—The analysis procedure and nomenclature closely follows HAYSTAC [31]. All saved spectra $P_i(\nu)$ are each filtered with a 6th-order Savitzky-Golay (SG) filter with a window length of $1.04 \text{ MHz} \gg \Delta\nu_\chi$ that would leave a potential DP signal intact. From this we obtain baselines $P_{bl,i}(\nu)$ and processed spectra $p_{proc,i}(\nu) = (P_i/P_{bl,i} - 1)$. For perfect baseline subtraction, the standard deviation σ_i of each processed spectrum, excluding RFI bins, is expected to be $1/\sqrt{N_{av}}$. We observe σ_i to be between 97.9% and 99.2% of this value, indicating the baseline subtraction has a minor impact on our sensitivity. The SG filter reduces the SNR of a potential DP signal by $\eta_{SG} = 0.91$, which is obtained by running the analysis on simulated data with injected synthetic DP signals. During the run, the baseline drifted by no more than 5% in amplitude, which can be attributed to gain variation caused by room temperature changes of the same order. In frequency, spectral features of $P_{bl,i}$ shifted by less than 1 MHz. Assuming the same frequency stability for β^2 , which was determined after the run, translates to a relative systematic uncertainty of $\sim 4\%$ in amplitude.

The processed spectra are combined using weights that take the expected SNR of a hypothetical single-bin DP

TABLE I. Summary of systematic uncertainties. Minimum to maximum is shown for frequency-dependent quantities.

Effect	Uncertainty on χ
Boost factor determination	
Bead-pull measurements	2 to 17%
Bead pull finite domain correction	5%
Receiver chain impedance mismatch	<1%
Subtotal	5 to 18%
Y-factor calibration	4%
Power stability	3%
Frequency stability	2%
Line shape discretization	4%
Total	9 to 19%

signal into account for every bin of each spectrum [31]. The resulting combined spectrum is further cross-correlated with the expected DP line shape [32,33] to account for the fact that the expected DP signal would stretch over ~ 5 bins. We use $\sigma_v = 154.1 \text{ km s}^{-1}$ [34] for the velocity dispersion of the local DM halo and $v_{\text{lab}} = 242.1 \text{ km s}^{-1}$ [35] for the velocity of the laboratory with respect to the local standard of rest. The unknown relative alignment between the frequency bins and DP line shape leads to variation in the discretization of the line shape, affecting the cross-correlation, and is treated as a systematic uncertainty. The observed cross-correlated power excess P_{exc} is shown in Fig. 3, revealing a sensitivity to DP signals of $\sim 2 \times 10^{-21} \text{ W}$. A hypothetical DP signal with $\chi = 2 \times 10^{-13}$ is shown in the inset along with a zoomed-in

view of the observations around 19.48 GHz where the maximum boost factor occurs.

The largest excess has a local significance of 4.3σ . This is within the expectation of observing thermal noise only as the probability of observing a local excess at least this large is $p \approx 0.21$ for the full dataset. From the nonobservation of any DP signal, we derive a 95% confidence level (CL) upper limit on $P_{\text{exc}} \rightarrow P_{\text{exc}}^{95}$ for groups of 210 neighboring bins using the approach described in [36]. The limit on the kinetic mixing angle χ is obtained by setting $P_{\text{exc}}^{95} = \eta_{\text{SG}} \beta^2 P_0$. Systematic uncertainties, summarized in Table I, are assumed to be independent and propagated to χ . The most conservative value of χ within this uncertainty is adopted for the final limit.

Conclusion—The 95% CL limit on the dark photon kinetic mixing angle χ is shown in Fig. 4. Unpolarized dark photon dark matter ($\rho_\chi = 0.3 \text{ GeV cm}^3$) with $\chi > 2.7 \times 10^{-12}$ can be excluded for masses between 78.62 to 83.95 $\mu\text{eV}/c^2$. Using only three disks, we have improved existing limits by up to almost three orders of magnitude with a peak sensitivity of $\chi = 1.1 \times 10^{-13}$. This demonstrates, for the first time, the feasibility of the MADMAX dielectric haloscope concept and showcases its capability of providing a large conversion volume at high frequencies while retaining resonant enhancement. Further improvements are planned: scaling the conversion volume by adding more dielectric disks, reducing system noise temperature using a cryostat, and implementing a tunable booster via motorized disk control [20,37]. This will vastly improve both the mass range of searches for dark photons as well as increase the sensitivity to their coupling to photons.

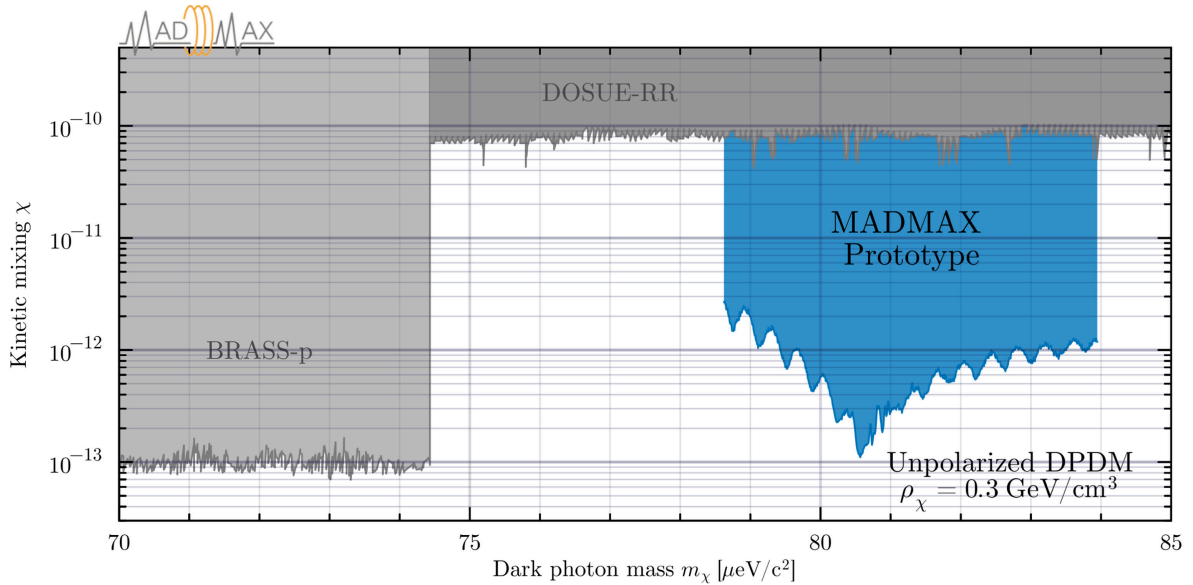


FIG. 4. The 95% CL upper limit on the dark photon kinetic mixing angle χ obtained with the MADMAX prototype as compared to the dish antenna experiments DOSUE-RR [38] and BRASS-p [39], rescaled to a common value of the assumed local dark photon dark matter density of $\rho_\chi = 0.3 \text{ GeV/cm}^3$. Unpolarized dark photons are assumed.

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Data availability—The data that support the findings of this Letter are openly available [40].

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- [1] J. Billard *et al.*, Direct detection of dark matter—APPEC committee report, *Rep. Prog. Phys.* **85**, 056201 (2022).
 - [2] A. Caputo, A. J. Millar, C. A. J. O’Hare, and E. Vitagliano, Dark photon limits: A handbook, *Phys. Rev. D* **104**, 095029 (2021).
 - [3] M. Fabbrichesi, E. Gabrielli, and G. Lanfranchi, *The Physics of the Dark Photon: A Primer* (Springer International Publishing, 2021), 10.1007/978-3-030-62519-1.
 - [4] P. W. Graham, J. Mardon, and S. Rajendran, Vector dark matter from inflationary fluctuations, *Phys. Rev. D* **93**, 103520 (2016).
 - [5] Y. Nakai, R. Namba, and Z. Wang, Light dark photon dark matter from inflation, *J. High Energy Phys.* **12** (2020) 170.
 - [6] F. P. Huang and H.-S. Lee, Extension of the electrodynamics in the presence of the axion and dark photon, *Int. J. Mod. Phys. A* **34**, 1950012 (2019).
 - [7] G. Marocco, Dark photon limits from magnetic fields and astrophysical plasmas, *arXiv:2110.02875*.
 - [8] A. J. Millar, J. Redondo, and F. D. Steffen, Dielectric haloscopes: Sensitivity to the axion dark matter velocity, *J. Cosmol. Astropart. Phys.* **10** (2017) 006.
 - [9] S. Knirck, A. J. Millar, C. A. O’Hare, J. Redondo, and F. D. Steffen, Directional axion detection, *J. Cosmol. Astropart. Phys.* **11** (2018) 051.
 - [10] D. Horns, J. Jaeckel, A. Lindner, A. Lobanov, J. Redondo, and A. Ringwald, Searching for WISPy cold dark matter with a dish antenna, *J. Cosmol. Astropart. Phys.* **04** (2013) 016.
 - [11] A. Caldwell, G. Dvali, B. Majorovits, A. Millar, G. Raffelt, J. Redondo, O. Reimann, F. Simon, and F. Steffen, Dielectric haloscopes: A new way to detect axion dark matter, *Phys. Rev. Lett.* **118**, 091801 (2017).
 - [12] MADMAX Collaboration, A new experimental approach to probe QCD axion dark matter in the mass range above 40 μeV , *Eur. Phys. J. C* **79**, 186 (2019).
 - [13] A. J. Millar, G. G. Raffelt, J. Redondo, and F. D. Steffen, Dielectric haloscopes to search for axion dark matter: Theoretical foundations, *J. Cosmol. Astropart. Phys.* **01** (2017) 061.
 - [14] J. Jaeckel and J. Redondo, Resonant to broadband searches for cold dark matter consisting of weakly interacting slim particles, *Phys. Rev. D* **88**, 115002 (2013).
 - [15] R. Cervantes *et al.*, ADMX-Orpheus first search for 70 μeV dark photon dark matter: Detailed design, operations, and analysis, *Phys. Rev. D* **106**, 102002 (2022).
 - [16] J. Chiles *et al.*, New constraints on dark photon dark matter with superconducting nanowire detectors in an optical haloscope, *Phys. Rev. Lett.* **128**, 231802 (2022).
 - [17] J. Egge, S. Knirck, B. Majorovits, C. Moore, and O. Reimann, A first proof of principle booster setup for the MADMAX dielectric haloscope, *Eur. Phys. J. C* **80**, 392 (2020).
 - [18] R. H. Dicke, The measurement of thermal radiation at microwave frequencies, *Rev. Sci. Instrum.* **17**, 268 (1946).
 - [19] P. F. de Salas and A. Widmark, Dark matter local density determination: Recent observations and future prospects, *Rep. Prog. Phys.* **84**, 104901 (2021).
 - [20] MADMAX Collaboration, First mechanical realization of a tunable dielectric haloscope for the MADMAX axion search experiment, *J. Instrum.* **19**, T11002 (2024).
 - [21] BREAD Collaboration, First results from a broadband search for dark photon dark matter in the 44 to 52 μeV range with a coaxial dish antenna, *Phys. Rev. Lett.* **132**, 131004 (2024).
 - [22] J. Egge, Axion haloscope signal power from reciprocity, *J. Cosmol. Astropart. Phys.* **04** (2023) 064.
 - [23] C. Steele, A nonresonant perturbation theory, *IEEE Trans. Microwave Theory Tech.* **14**, 70 (1966).
 - [24] J. Egge, M. Ekmedžić, A. Gardikiotis, E. Garutti, S. Heyminck, C. Kasemann, S. Knirck, M. Kramer, C. Krieger *et al.*, Experimental determination of axion signal power of dish antennas and dielectric haloscopes using the reciprocity approach, *J. Cosmol. Astropart. Phys.* **04** (2024) 005.
 - [25] J. Egge, The sensitivity of dielectric haloscopes to dark matter axions and dark photons, Ph.D. thesis, University of Hamburg, 2024. Available at <https://ediss.sub.uni-hamburg.de/handle/ediss/11205>.
 - [26] MADMAX Collaboration, Simulating MADMAX in 3D: Requirements for dielectric axion haloscopes, *J. Cosmol. Astropart. Phys.* **10** (2021) 034.
 - [27] D. Pozar, *Microwave Engineering* (Wiley, New York, 2011).
 - [28] Low noise factory, Datasheet LNF-LNC15_29B, 2 (2023), https://lownoisefactory.com/wp-content/uploads/2023/03/Lnf-lnc15_29b.pdf.
 - [29] H. Haus, W. Atkinson, G. Branch, W. Davenport, W. Fonger, W. Harris, S. Harrison, W. Mcleod, E. Stodola

- et al.*, Representation of noise in linear twoports, *Proc. Inst. Radio Eng.* **48**, 69 (1960).
- [30] See Supplemental Material at <http://link.aps.org/supplemental/10.1103/PhysRevLett.134.151004> for additional information on the setup and analysis.
- [31] B. M. Brubaker, L. Zhong, S. K. Lamoreaux, K. W. Lehnert, and K. A. van Bibber, HAYSTAC axion search analysis procedure, *Phys. Rev. D* **96**, 123008 (2017).
- [32] J. Diehl, J. Knollmüller, and O. Schulz, Bias-free estimation of signals on top of unknown backgrounds, *Nucl. Instrum. Methods Phys. Res., Sect. A* **1063**, 169259 (2024).
- [33] C. A. J. O'Hare and A. M. Green, Axion astronomy with microwave cavity experiments, *Phys. Rev. D* **95**, 063017 (2017).
- [34] J. Bovy, C. A. Prieto, T. C. Beers, D. Bizyaev, L. N. da Costa, K. Cunha, G. L. Ebelke, D. J. Eisenstein, P. M. Frinchaboy *et al.*, The Milky Way's circular-velocity curve between 4 and 14 kpc from apogee data, *Astrophys. J.* **759**, 131 (2012).
- [35] K. Malhan, R. A. Ibata, and N. F. Martin, Measuring the sun's velocity using gaia EDR3 observations of stellar streams, [arXiv:2012.05271](https://arxiv.org/abs/2012.05271).
- [36] C. Bartram *et al.* (ADMX Collaboration), Axion dark matter experiment: Run 1B analysis details, *Phys. Rev. D* **103**, 032002 (2021).
- [37] E. Garutti, H. Janssen, D. Kreikemeyer-Lorenzo, C. Krieger, A. Lindner, B. Majorovits, J. Schaffran, and B. van Bree, Qualification of piezo-electric actuators for the MADMAX booster system at cryogenic temperatures and high magnetic fields, *J. Instrum.* **18**, P08011 (2023).
- [38] DOSUE-RR Collaboration, Search for dark photon dark matter in the mass range 74–110 μeV with a cryogenic millimeter-wave receiver, *Phys. Rev. Lett.* **130**, 071805 (2023).
- [39] F. Bajjali, S. Dornbusch, M. Ekmedžić, D. Horns, C. Kasemann, A. Lobanov, A. Mkrтчyan, L. H. Nguyen, M. Tluczykont *et al.*, First results from BRASS-p broadband searches for hidden photon dark matter, *J. Cosmol. Astropart. Phys.* **08** (2023) 077.
- [40] MADMAX Collaboration, First search for dark photon dark matter with a MADMAX prototype: Datasets (2025), [10.5281/zenodo.15052442](https://doi.org/10.5281/zenodo.15052442).