



NuSTAR Perspective on High-redshift MeV Blazars

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Abstract

With bolometric luminosities exceeding 10^{48} erg s $^{-1}$, powerful jets, and supermassive black holes at their center, MeV blazars are some of the most extreme sources in the universe. Recently, the *Fermi*-Large Area Telescope detected five new γ -ray emitting MeV blazars beyond redshift $z = 3.1$. With the goal of precisely characterizing the jet properties of these extreme sources, we started a multiwavelength campaign to follow them up with joint *Nuclear Spectroscopic Telescope Array*, *Swift*, and the Southeastern Association for Research in Astronomy's optical telescopes. We observe six high-redshift quasars, four of them belonging to the new γ -ray emitting MeV blazars. Thorough X-ray analysis reveals spectral flattening at soft X-ray for three of these objects. The source NVSS J151002+570243 also shows a peculiar rehardening of the X-ray spectrum at energies $E > 6$ keV. Adopting a one-zone leptonic emission model, this combination of hard X-rays and γ -rays enables us to determine the location of the Inverse Compton peak and to accurately constrain the jet characteristics. In the context of the jet-accretion disk connection, we find that all six sources have jet powers exceeding accretion disk luminosity, seemingly validating this positive correlation even beyond $z > 3$. Our six sources are found to have $10^9 M_\odot$ black holes, further raising the space density of supermassive black holes in the redshift bin $z = [3, 4]$.

Unified Astronomy Thesaurus concepts: [Active galactic nuclei \(16\)](#); [Blazars \(164\)](#); [High energy astrophysics \(739\)](#); [Jets \(870\)](#)

1. Introduction

The hard X-ray band has been crucial to studying some of the most powerful blazars (see e.g., Tavecchio et al. 2000; Massaro et al. 2004a, 2004b, 2006; Donato et al. 2005). More recently, the outstanding sensitivity of the *Nuclear Spectroscopic Telescope Array* (*NuSTAR*, 3–79 keV, Harrison et al. 2013) has enabled us to find and study some of the most distant and luminous ones (e.g., Sbarato et al. 2013; Tagliaferri et al. 2015; Ajello et al. 2016; Paliya et al. 2016; Sbarato et al. 2016; Marcotulli et al. 2017). Harboring highly relativistic jets pointed closely at the observer ($\theta_V < 1/\Gamma$, θ_V being the viewing angle and Γ the bulk Lorentz factor, $\Gamma \sim 10$ –15, Urry & Padovani 1995), this subclass of the Active Galactic Nuclei (AGNs) is home to some of the most energetic particle acceleration and radiation processes known in astrophysics. The boost in flux ascribed to relativistic beaming, arising from the peculiar orientation of the jets, renders them visible at redshifts well beyond $z = 2$ (the farthest blazar detected so far is at $z = 5.47$, Romani et al. 2004), making them extraordinary beacons with which to explore the early universe. Their typical double-hump spectral energy distribution (SED) spans the whole electromagnetic spectrum and is shaped by the nonthermal processes occurring in the jets. Relativistic electrons, spiraling along the magnetic field lines, undergo both synchrotron and inverse Compton (IC) processes. The first produces a peak in the SED located between infrared and X-ray frequencies. The second instead results in a peak located between X- and γ -ray energies. If the electrons interact with a

source of low-energy photons within the jet, this is referred to as Synchrotron Self Compton (SSC, e.g., Ghisellini & Maraschi 1989), whereas if the photons are external to the jet (i.e., the accretion disk, the torus, and/or the broad line region, BLR), it is referred to as External Compton process (EC, e.g., Sikora et al. 1994). Based on their optical spectra, blazars are usually classified either as BL Lacertae objects (BL Lacs) or flat spectrum radio quasars (FSRQs), the first showing weak or no emission lines, the second showing broad (EW > 5 Å) ones. Following Abdo et al. (2010a), these sources can also be classified according to the position of the synchrotron peak (ν_{peak}^S), with low-, intermediate-, and high-synchrotron peak (LSP, ISP, HSP) blazars having, respectively, $\nu_{\text{peak}}^S < 10^{14}$ Hz, $10^{14} < \nu_{\text{peak}}^S < 10^{15}$ Hz, and $\nu_{\text{peak}}^S > 10^{15}$ Hz. FSRQs usually belong to the LSP class and at the high-luminosity end of such subclass are the so-called “MeV blazars,” whose high-energy peak falls in (or close to) the MeV band. With bolometric luminosities exceeding 10^{48} erg s $^{-1}$, these are among the most powerful objects in the universe. In fact, they host powerful relativistic jets (Ghisellini et al. 2014), are usually found at high-redshift ($z > 2$, e.g., Ajello et al. 2009; Ghisellini et al. 2010; Ackermann et al. 2017; Marcotulli et al. 2017), and typically host billion solar mass black holes (e.g., Ghisellini et al. 2010; Paliya et al. 2017a).

In accordance with the *blazar sequence* (Ghisellini et al. 1998), as both source luminosity and distance increase, the two SED peaks drift toward lower energies and the Compton Dominance (CD, the ratio between IC and synchrotron peak

luminosities) increases ($CD > 1$). As a consequence, MeV blazars are bright in the hard X-rays (above 10 keV) and therefore represent perfect targets for *NuSTAR*. Moreover, high-energy emission dominates their energetics (see Madejski & Sikora 2016 for a recent review), hence good coverage of the IC hump is required in order to obtain a precise determination of the peak position. This, in turn, allows us to constrain jet parameters such as viewing angle, power, bulk Lorentz factor, as well as the shape of underlying electron distribution (see, e.g., Celotti & Ghisellini 2008; Ghisellini & Tavecchio 2009; Gao et al. 2011; Chen et al. 2015; Finke 2016; Ghisellini et al. 2019; Paliya et al. 2019b), keys to untangle the physics of these cosmic monsters. Undoubtedly, *NuSTAR* is the best instrument to sample the rising part of their high-energy spectrum, also enabling us to detect possible hard X-ray peculiarities of these sources (e.g., flattening of the spectrum, see Paliya et al. 2016, and variability, e.g., Sbarato et al. 2016). Lacking an all sky MeV instrument (e.g., AMEGO, Rando 2017), complementary to hard X-ray observations are the γ -ray ones obtained with the *Fermi*-Large Area Telescope (*Fermi*-LAT; Atwood et al. 2009). Encompassing a range of energies from 20 MeV to >300 GeV, it allows the sampling of the falling part of the IC hump (e.g., Ackermann et al. 2015) and, together with *NuSTAR*, it is fundamental to place tight constraint on the jet power.

On another important note, the SED peaks' shifts also reveal the underlying thermal emission from the accretion disk at optical-UV energies (e.g., Ghisellini et al. 2010), which can be modeled to extract the black hole mass and the accretion disk luminosity (Shakura & Sunyaev 1973). This fact renders MeV blazars excellent probes for addressing open issues such as the accretion disk-jet connection in jetted quasars (e.g., Paliya et al. 2017a), as well as the understanding of black hole growth and evolution in the early stages of the universe (Volonteri et al. 2011). Indeed, since the detection of a single source with jets aligned to our line of sight implies the existence of $2T^2$ similar sources, at the same z , with same black hole masses but misaligned jets, finding more high-redshift blazars allows us to probe their intrinsic population and constrain the space density of supermassive black holes (e.g., Sbarato et al. 2014).

Even though blazars are the most numerous extragalactic sources in *Fermi*-LAT catalogs (see Abdo et al. 2010b; Nolan et al. 2012; Acero et al. 2015), few are detected at $z > 3$, probably since their peak high-energy emission lies below the *Fermi*-LAT band. Using the improved Pass 8 data set, taking advantage of the higher detection threshold of the *Fermi*-LAT, five new γ -ray emitting MeV blazars have been found to lie at $z > 3$, the farthest located at $z = 4.31$ (Ackermann et al. 2017). Since, on average, high-redshift sources lack good quality multifrequency observations, we started a multiwavelength campaign in order to characterize these newly found γ -ray emitting blazars. In particular, we wanted to obtain good X-ray spectral coverage, and, importantly, sensitive hard X-ray data, in order to derive their X-ray properties and tightly constrain the jet emission.

In this paper, we present the broadband follow up of six sources: NVSS J013126–100931 ($z = 3.51$, hereafter J013126–100931), NVSS J020346+113445 ($z = 3.63$; hereafter J020346+113445), NVSS J064632+445116 ($z = 3.39$; hereafter J064632+445116), NVSS J135406–020603 ($z = 3.71$, hereafter J135406–020603), NVSS J151002+570243 ($z = 4.31$; hereafter J151002+570243), and NVSS J212912–153841 ($z = 3.26$, hereafter

J212912–153841). Four of these objects belong to the newly detected γ -ray blazars, while two are candidate γ -ray sources lying just below the LAT detection threshold. In this work, we publish their first >10 keV detection with *NuSTAR*.¹⁰ Since blazars are extremely variable sources in all energy ranges (e.g., Abdo et al. 2010c; Stroh & Falcone 2013; Marchesini et al. 2016; Paliya et al. 2017b), to accurately constrain the sources' SEDs and to reliably measure the X-ray spectral parameters, we also obtain quasi-simultaneous observations from a variety of facilities. The soft X-ray and UV-optical bands are covered by the *Neil Gehrels Swift Observatory X-ray Telescope* (*Swift*-XRT; Burrows et al. 2005) and Ultraviolet and Optical Telescope (*Swift*-UVOT; Roming et al. 2005). With the aim of studying possible long-term variability properties of these sources in the soft X-ray band, we check if multiple *Swift*-XRT, as well as *Chandra* (Weisskopf et al. 2000) and *XMM-Newton* (Jansen et al. 2001) observations are available. All sources but J151002+570243 and J135406–020603 have multiple *Swift*-XRT observations, while *Chandra* and *XMM-Newton* data are found for J151002+570243 and J212912–153841. We complement UVOT observations with quasi-simultaneous ones from the Southeastern Association for Research in Astronomy (SARA; Keel et al. 2017) facilities, in order to detect the disk emission peak. The γ -ray spectra of the sources are taken from Ackermann et al. (2017). The data reduction pipelines are given in Section 2. In order to derive the important physical parameters of these objects, we perform the SED modeling (described in Section 3.4). Section 3 details the results of the analysis, which are then discussed in Section 4. Throughout the paper, we use cosmological parameters $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.308$, and $\Omega_\Lambda = 0.692$ (Planck Collaboration et al. 2016).

2. Observations and Data Analysis

2.1. Fermi-LAT

Ackermann et al. (2017) recently discovered five new γ -ray emitting high-redshift blazars. In our work, we study four of these *Fermi*-LAT detected sources, namely J064632+445116, J135406–020603, J151002+570243, and J212912–153841, extracting their γ -ray data directly from Ackermann et al. (2017). Two additional sources, J013126–100931 and J020346+113445, were also observed. They are candidate γ -ray emitting blazars that fall below the LAT detection threshold (test statistic, $TS < 25$, Mattox et al. 1996). For these sources we employ the 10 yr LAT sensitivity curve.¹¹

2.2. NuSTAR

NuSTAR observed the six blazars for net exposures of ~ 30 –50 ks during the dates reported in Table 1. We process the data for both instrument Focal Plane Modules (FPMA and FPMB; Harrison et al. 2013) using the *NuSTAR* Data Analysis Software (NUSTARDAS) v1.8.0. The event files are calibrated using the task `nupipeline`, with the response file taken from the latest Calibration Database (CALDB, v.20180419). The generation of source and background spectra, ancillary and response matrix files, is achieved using the `nuproducts` script. For source regions, we select circles with radii of $20''$ centered on the targets; the background events are extracted

¹⁰ These sources were observed as part of our *NuSTAR* proposals 2011, 4301; P.I.: Ajello M., Marcotulli L.

¹¹ http://www.slac.stanford.edu/exp/glast/groups/canda/lat_Performance.htm

Table 1
Table of Simultaneous *NuSTAR* and *Swift*-XRT Observations

<i>NuSTAR + Swift</i> -XRT					
Name	$N_{\mathrm{H}}^{\mathrm{a}}$	z	Exp. time (ks) (<i>NuSTAR</i>)	Exp. time (ks) (<i>Swift</i> -XRT)	Date
J013126–100931	3.40	3.51	30	3.4	2018 Jul 21
J020346+113445	5.96	3.63	33	1.7	2016 Nov 8
J064632+445116	10.7	3.39	32	1.8	2016 Nov 28
J135406–020603	3.35	3.71	32	1.9	2017 May 5
J151002+570243	1.57	4.31	53	2.0	2017 Apr 30
J212912–153841	4.85	3.26	37	2.6	2018 Sep 26

Note.

^a Galactic column density in units of $10^{20} \mathrm{cm}^{-2}$ (Kalberla et al. 2005).

from circles of $30''$ radii located in a source-free region on the same frame. The excellent sensitivity of *NuSTAR* provides us with a good signal-to-noise detection of all sources. Therefore, if the source has also good soft X-ray photon statistics (see Sections 2.3–2.4), *NuSTAR* spectra are rebinned to have at least 15 counts per bin. Visual inspection of the light curves shows no detected variability for all six sources.

2.3. Swift

NuSTAR observations are supplemented by quasi-simultaneous *Swift* pointings, each of ~ 2 – 3 ks (see Table 1). The *Swift*-XRT observations are executed in photon counting mode. The data are analyzed with the XRTDAS software package included in HEASoft (v.3.4.1) in combination with the latest calibration database CALDB (v.20180710). The event files are calibrated and cleaned using the standard pipeline task `xrtpipeline`. With the tool `xselect`, we extract source and background regions: the source region was selected as a circle of $10''$ radius, whereas an annular region of inner and outer radii $20''$ and $50''$, respectively, are used for the background. Both regions are centered at the source position. The ancillary response files are generated using `xrtmkarf`, while the response matrix files are already included in the CALDB. In order to study possible long-term soft X-ray variability, if more than one exposure is present in the archival data, we combine all event files and produce light curves using `xselect`. In case no significant variability is found, all event files are considered to extract sources' spectra and all images are added (using the summation task `sum_ima`) in order to increase the signal of the detection.¹² The sources' spectra with high signal-to-noise are rebinned to have at least 15 counts per bin, while for faint sources we use one count per bin.

For *Swift*-UVOT, we follow the standard pipeline procedure of Poole et al. (2008) to extract the final products. We use the tasks `uvotimsum` and `uvotsource` to extract the sources' magnitudes in all filters. We correct the magnitudes for Galactic extinction following the recommendations in Roming et al. (2008) and subsequently derive the fluxes using the zero-points listed in Breeveld et al. (2011). For the fainter sources (J135406–020603 and J151002+570243) we could only derive upper limits.

¹² This was the case for J013126–100931, J020346+113445, J064632+445116, and J212912–153841. In particular, J020346+113445 is very faint, having $\sim 5 \times 10^{-3} \mathrm{cts} \mathrm{s}^{-1}$.

2.4. Chandra and XMM-Newton

In order to study any possible long-term soft X-ray variability properties of the sources, we looked for archival observations with *Chandra* (Weisskopf et al. 2000) and *XMM-Newton* (Jansen et al. 2001), which we find only for J151002+570243 and J212912–153841.

The *Chandra* data are reduced using the CIAO 4.9 software (Fruscione et al. 2006) and the *Chandra* CALDB (v.4.7.8), following standard procedures (`chandra_repro`). The source and background spectra are extracted using, respectively, a circular region of radius $2''$ and an annulus of inner radius $7''$ and outer radius of $15''$, after a visual inspection to avoid contamination from nearby sources. The spectra are obtained using the task `specextract`.

For *XMM* we use the *XMM-Newton* Science Analysis Software (SAS) v16.0.0. The telescope has on board the European Photon Imaging Camera (EPIC), which consists of two CCD arrays, MOS and pn. We use the EPIC/pn, EPIC/MOS1, and EPIC/MOS2 camera observations for our purposes. Standard procedures are followed to reduce the data: event files are generated using `emproc/epproc` (for MOS and PN, respectively); after subtracting bad pixels and bad events from the field, source and background spectra are obtained using `evselect`. The source and background regions are selected from a circle of radius $10''$ and an annulus of inner radius $20''$ and outer radius $40''$, both centered on the source of interest. Both *Chandra* and *XMM* spectra are rebinned to have at least 15 counts per bin.

2.5. SARA

We carry out quasi-simultaneous optical observations for five objects¹³ with the Southern Association for Research in Astronomy (SARA, Keel et al. 2017) consortium's 1.0 m, 0.96 m, and 0.6 m telescopes respectively located at la Roque de los Muchachos Observatory, Canary Islands (SARA-ORM), at Kitt Peak National Observatory, Arizona (SARA-KP) and at Cerro Tololo, Chile (SARA-CT). We obtain our observations using the Bessel *BVRI* filters mounted on SARA-ORM and SARA-KP, and the Sloan Digital Sky Survey *griz* filters mounted on SARA-CT. The data analysis is executed following the standard aperture photometry technique employing IRAF v2.16. The photometric standards utilized for

¹³ Due to bad weather, J013126–100931, J064632+445116, and J212912–153841 were observed within a few days/months to the *NuSTAR* pointing, while J020346+113445 could not be observed due to instrumental issues.

Table 2
Table of SARA Magnitudes

Name	UT Date ^a	SARA			
		B	V	R	I
J064632+445116	2016 Dec 29	...	18.10 ± 0.02	17.92 ± 0.02	19.48 ± 0.03
J135406−020603	2017 May 6	20.33 ± 0.08	19.46 ± 0.05	19.24 ± 0.03	19.32 ± 0.26
J151002+570243	2017 Apr 30	23.41 ± 0.47	20.72 ± 0.06	20.16 ± 0.04	19.17 ± 0.05
		<i>g</i>	<i>r</i>	<i>i</i>	<i>z</i>
J013126−100931	2018 Nov 17	21.43 ± 0.22	21.54 ± 0.46	19.41 ± 0.08	18.88 ± 0.07
J212912−153841	2018 Nov 17	18.26 ± 0.05	16.98 ± 0.01	16.84 ± 0.01	17.26 ± 0.02

Notes.

^a Exposure start time.

^b Corrected for Galactic reddening.

calibration purposes were obtained from Landolt (1992) and Smith et al. (2002). The complete list of available standards is provided in Table 2¹⁴ and Table 8¹⁵ of the respective papers. We selected standards from these tables which were approximately at similar airmass as our targets for each observation. To convert the photometry to the AB system, we apply the standard Vega–AB corrections.¹⁶ The foreground Galactic extinction correction is applied using the calculations from Schlafly & Finkbeiner (2011). The fluxes are extracted using the canonical flux to magnitude conversion (see footnote 15), $F_\nu = 10^{-(m_{AB}+48.6)/2.5}$. The final magnitudes in the AB system are reported in Table 2.

3. Results

3.1. Soft X-Ray Long-term Variability

Blazars are known to be variable sources in all energy bands (e.g., Abdo et al. 2010c; Stroh & Falcone 2013; Marchesini et al. 2016; Rani et al. 2017; Jiang 2018; Mao et al. 2018; Paliya et al. 2019a; Liao et al. 2019). Importantly, for powerful FSRQs in which, according to leptonic models, the EC dominates the high-energy part of the SED, most of the X-ray emission originates from the low-energy electron population (Ghisellini et al. 2007). In this scenario, since the low-energy electron cooling time is expected to be longer with respect to the high-energy electron one ($t_{\text{cool}} \propto \gamma^{-1}$, where γ is the random Lorentz factor of the electrons, see Ghisellini 2013), this radiation is not expected to vary significantly and on long timescales (e.g., Bonning et al. 2009; Sbarato et al. 2016; Prince 2019). Detecting spectral variability in the soft X-ray (0.3–15 keV) band may be indicative of a change of emission region location (see Tagliaferri et al. 2015), and hence magnetic field strength and radiative energy densities, as well as bulk Lorentz factor (Ghisellini et al. 2007). It could also possibly be related to a change of the injection power (see Sbarato et al. 2016), or be indicative of underlying acceleration processes such as shocks or magnetic reconnection (e.g., McHardy 2008; Sikora et al. 2009; De Gouveia Dal Pino et al. 2016; Christie et al. 2019).

Motivated to study such a property for high-redshift blazars, we looked for more sensitive observations in the soft X-ray band. Two of our sources, J151002+570243 and J212912−153841,

have both *XMM-Newton* and *Chandra* data,¹⁷ in addition to XRT data. In order to test the presence/absence of spectral variability, we used XSPEC to separately fit all three observations. To consistently compare the results, we used a simple absorbed power-law model, with absorption (N_{H}) fixed at the Galactic value, and we chose an energy range where all telescopes overlap, from 0.5 to 7 keV. Performing such an exercise leads us to the result that J151002+570243 does not show significant spectral or flux variability, and all parameters are consistent within the errors. On the other hand, J212912−153841 shows both flux and index variability ($>1\sigma$) in all three observations. In the case of J151002+570243, we are able to perform joint fit of *XMM-Newton* and *Chandra* observations with XRT ones (see Section 3.2), assuming that the flux and photon index detected represent the typical status of the source. The same is not applicable to J212912−153841, for which we do not find variability. More detailed study in the soft X-ray band would be required to investigate possible scenarios explaining this property, such as a change in the injected power or emission region location. For sources with multiple XRT observations, light curves show no significant variability (see Section 2.3). Therefore we are able to combine all available XRT data to increase the significance of the detection.

3.2. X-Ray Spectral Analysis

We fit the soft X-ray (0.3–15 keV) and *NuSTAR* FPMA and FPMB (3–79 keV) spectra using XSPEC. We include Galactic absorption (N_{H}), using the Galactic neutral hydrogen column densities from Kalberla et al. (2005). We use the C-statistics (Cash 1979) for sources with low soft X-ray counts (<30 cts), otherwise we employ standard χ^2 statistics. Sources with good X-ray signal are: J013126−100931, J064632+445116, J151002+570243, and J212912−153841. High-redshift blazars have been found to show spectral softening at soft X-rays (Page et al. 2005; Saez et al. 2011; Eitan & Behar 2013), which has been proposed to be linked to either a property of the intrinsic electron population underlying the emission or to the presence of an intervening absorber residing in the intergalactic medium (IGM, see Arcodia et al. 2018), or both. In order to test this scenario, at first these objects are fitted with a simple power-law model along

¹⁴ <http://www.eso.org/sci/observing/tools/standards/Landolt.html>

¹⁵ <https://classic.sdss.org/dr7/algorithms/standardstars/tabc08.dat.txt>

¹⁶ <http://www.astronomy.ohio-state.edu/~martini/usefuldata.html>

¹⁷ The observations of J151002+570243 were taken on 2002 May 11 by *XMM-Newton* and on 2001 June 10 by *Chandra*, while the ones of J212912−153841 were on 2001 May 1 by *XMM-Newton* and on 1999 December 16 by *Chandra*.

with two curved ones, a broken power law and a log-parabola, all with absorption fixed at the Galactic value. Then, if a curved model is favored over a simple power-law one (highlighting the presence of a softening of the soft X-ray part of the spectrum), we also test the possibility of an intrinsic absorber. Therefore, we further fit these objects with both a power-law and a broken power-law model including a cold absorber (`ztbabsin` XSPEC). All X-ray spectral fit parameters are provided in Table 3 and the fits and residuals for each fitted model can be found in the Appendix. The main results are listed below:

1. Fits' residuals and F-test results show that both J013126–100931 and J212912–153841 favor a curved model. Fits including the absorber give a reduced χ^2 close to one and of the same order as the one found for the log-parabola/broken power-law case (see Table 3). Results using the F-test favor a log-parabola fit for J013126–100931, while J212912–153841 is best fitted by a broken power-law model with intrinsic absorber. The radiating electron population is therefore likely in both cases to possess an intrinsic break at low energies, with J212912–153841 also showing the signature of an intervening absorbing medium.
2. The source J151002+570243 represents an exception. Analyzing residuals and F-test results, the broken power-law model plus absorber is favored over the broken power-law (p -value = 8×10^{-5}) or power-law with absorber (p -value = 2×10^{-3}) ones. As can be seen from Figure 1, the spectrum indeed shows a break around 0.8 keV, after which it becomes softer, and a second break around 6 keV, after which it hardens again. While the second break can be explained via detailed modeling of the jet properties (see Section 4), the first one can hint to either the presence of an intrinsic absorber or a break in the electron population itself. To check whether the second option is viable, we further test a double broken power-law model¹⁸ (`b2kn2po` in XSPEC). The fit is again preferred over the broken power-law (p -value = 6×10^{-5}) or power-law with absorber (p -value = 10^{-3}) models. Both the broken power-law plus intrinsic absorber and double broken power-law fits give similar reduced χ^2 values, and the F-test does not favor one over the other.
3. For J020346+113445 and J135406–020603, due to low soft X-ray statistics, a simple power-law model with absorption fixed at the Galactic value is employed.

In order to cross calibrate the instruments, we include a multiplicative constant factor, fixing it equal to 1 for FPMA but leaving it free to vary for FPMB and soft X-ray instruments. For all sources with good X-ray statistics, the constants between instruments are of the order of $\sim 1\%–10\%$ from the fixed FPMA value. Instead, for the ones with poor soft X-ray statistics, the difference for FPMB is in the range $0.6\%–10\%$, while a larger offset in the range $10\%–30\%$ is found for *Swift*-XRT.

3.3. Central Black Hole Mass

Black hole masses are computed following two approaches. The first is through single epoch optical spectroscopy (e.g.,

Shen et al. 2011), assuming a virialized BLR and adopting emission line parameters such as full width at half maximum, luminosity, and continuum emission. The errors associated with this method are of the order ~ 0.4 dex (e.g., Vestergaard & Peterson 2006; Shen et al. 2011). The second is by modeling the accretion disk using a standard Shakura & Sunyaev (1973) model. Blazars' accretion disk emission falls in the IR-UV band and, if measurable, it enables us to derive the black hole mass, even when optical spectra are not available. High-redshift blazars, thanks to the shift of the low-energy SED peak, unveil this emission and are therefore ideal sources for applying this method. With good IR-UV coverage, the uncertainty associated with this approach is of about a factor of 2 and the results have been shown to be in good agreement with the virial method (see, e.g., Castignani et al. 2013; Ghisellini & Tavecchio 2015; Paliya et al. 2017a). The accretion disk SED is assumed to be a multicolor blackbody. The flux density profile of this emission is given by the following (Frank et al. 2002):

$$F_\nu = \nu^3 \frac{4\pi h \cos \theta_v}{c^2 D^2} \int_{R_{\text{in}}}^{R_{\text{out}}} \frac{R \, dR}{e^{h\nu/kT(R)} - 1}, \quad (1)$$

where k , c , and h are the Boltzmann constant, the speed of light, and the Plank constant, respectively, D is the luminosity distance, θ_v is the jet viewing angle and R_{in} and R_{out} are the inner and outer disk radii. We take $R_{\text{in}} = 3R_{\text{Sch}}$ and $R_{\text{out}} = 500R_{\text{Sch}}$ (R_{Sch} is the Schwarzschild radius). The radial temperature profile is given by:

$$T(R) = \frac{3R_{\text{Sch}}L_{\text{disk}}}{16\pi\eta_a\sigma_{\text{SB}}R^3} \left[1 - \left(\frac{3R_{\text{Sch}}}{R} \right)^{1/2} \right]^{1/4}, \quad (2)$$

where σ_{SB} is the Stefan–Boltzmann constant and η_a is the accretion efficiency (we consider $\eta_a = 0.1$). With the above formulation, there are only two free parameters: the black hole mass (M_{BH}) and the accretion rate (\dot{M}_a). The latter is related to the intrinsic accretion luminosity via $L_{\text{disk}} = \eta_a \dot{M}_a c^2$. If the peak of the disk emission is visible, it is possible to obtain L_{disk} from the observations. The only free parameter left, therefore, is M_{BH} , which scales as $M_{\text{BH}} \propto \nu_{\text{peak,disk}}^{-2} L_{\text{disk}}^{1/2}$ (where $\nu_{\text{peak,disk}}$ is the peak of the disk emission, see Calderone et al. 2013).

For our sources, we looked into the literature to find the black hole masses derived through the first method, i.e., optical spectroscopy, then we performed disk modeling to extract them through the second. Both masses are reported in Table 4. They are found to be in reasonable accordance, and match within a factor of 2, apart from J151002+570243, whose difference cannot be accounted for by considering the associated errors. In fact, we note that the SARA magnitudes detected for this source are larger (\sim a factor of 2 higher in flux) than the average archival ones. This could be linked to the fact that we are observing the source in a high state (see, e.g., Kelly et al. 2009; Ruan et al. 2014; Rumbaugh et al. 2018), hence our measurement of the disk luminosity could be higher than the low/average one. However, the peak frequency of the emission is consistent with that found in the archival data, therefore the estimate on the black hole mass is not strongly affected. In the case of J013126–100931 we find that both SARA magnitudes are larger¹⁹ and the peak of the emission is at lower frequencies

¹⁸ A similar test has been performed for J212912–153841, but a double broken power-law does not result in a reliable fit, hence it is excluded from Table 3.

¹⁹ The source J013126–100931 is present in the Catalina Sky Survey (Drake et al. 2009), showing a range of magnitudes compatible with that found by our observations.

Table 3
Table of X-Ray Best-fit Spectral Parameters

J013126–100931									
Model	Best-fit Parameters			Flux ^b	$\chi^2/\text{d.o.f.}$	F-stat ^c	$p\text{-values}^{\text{c}}$		
PL	Γ 1.40 ± 0.05			$(1.08 \pm 0.05) \times 10^{-11}$	192.61/179				
BPL	Γ_1 $1.13_{-0.3}^{+0.1}$	Γ_2 1.52 ± 0.1	E_b (keV) $5.3_{-3.4}^{+2.3}$	$9.50_{-0.6}^{+0.5} \times 10^{-12}$	175.22/177	8.78	2×10^{-4}		
LP ^a	α 1.02 ± 0.17	β 0.23 ± 0.1	E_0 1(scale)	$(9.01 \pm 0.5) \times 10^{-12}$	176.29/178	16.47	7×10^{-5}		
PL+ABS	Γ 1.45 ± 0.06	$N_{\text{H}}(\text{intr.})$ $4.52_{-2.7}^{+3.6} \times 10^{22}$		$1.04_{-0.30}^{+0.04} \times 10^{-11}$	183.46/178	8.77	3×10^{-3}		
BPL+ABS	Γ_1 $1.18_{-0.18}^{+0.16}$	Γ_2 1.56 ± 0.10	E_b (keV) $6.23_{-3.35}^{+1.91}$	$N_{\text{H}}(\text{intr.})$ $0.63_{-0.58}^{+3.39} \times 10^{22}$	$9.45_{-0.72}^{+0.46} \times 10^{-12}$	174.83/176	5.96	6×10^{-4}	
J064632+445116									
Model	Best-fit Parameters			Flux ^b	$\chi^2/\text{d.o.f.}$	F-stat ^c	$p\text{-values}^{\text{c}}$		
PL ^a	Γ 1.63 ± 0.07			$2.27_{-0.17}^{+0.12} \times 10^{-12}$	37.82/48				
BPL	Γ_1 $1.51_{-0.14}^{+0.12}$	Γ_2 $1.76_{-0.09}^{+0.3}$	E_b (keV) $4.11_{-2.41}^{+14.6}$	$2.06_{-0.45}^{+0.17} \times 10^{-12}$	34.54/46	2.18	0.12		
LP	α $1.48_{-0.2}^{+0.1}$	β $0.11_{-0.15}^{+0.17}$	E_0 1(scale)	$2.04_{-0.31}^{+0.22} \times 10^{-12}$	36.63/47	1.52	0.22		
J151002+570243									
Model	Best-fit Parameters			Flux ^b	$\chi^2/\text{d.o.f.}$	F-stat ^c	$p\text{-values}^{\text{c}}$		
PL	Γ 1.47 ± 0.05			$1.95_{-0.09}^{+0.15} \times 10^{-12}$	391.61/276				
BPL	Γ_1 1.41 ± 0.02	Γ_2 0.98 ± 0.20	E_b (keV) $6.25_{-1.10}^{+2.15}$	$(3.36 \pm 0.50) \times 10^{-12}$	381.27/274	3.71	2×10^{-2}		
LP	α 1.44 ± 0.05	β -0.02 ± 0.08	E_0 1 (scale)	$(2.05 \pm 0.30) \times 10^{-12}$	388.32/275	2.32	10^{-1}		
PL+ABS ^a	Γ 1.48 ± 0.04		$N_{\text{H}}(\text{intr.})$ $(0.73 \pm 0.31) \times 10^{22}$	$(1.78 \pm 0.21) \times 10^{-12}$	376.07/275	5.38	8×10^{-4}		
BPL+ABS ^a	Γ_1 1.52 ± 0.05	Γ_2 $0.93_{-0.7}^{+1.17}$	E_b (keV) $6.02_{-1.7}^{+1.1}$	$N_{\text{H}}(\text{intr.})$ $(0.86 \pm 0.32) \times 10^{22}$	$(3.26 \pm 0.5) \times 10^{-12}$	360.27/273	7.91	4×10^{-5}	
2BPL ^a	Γ_1 $1.02_{-0.34}^{+0.27}$	Γ_2 1.49 ± 0.04	Γ_3 0.93 ± 0.20	$E_{b1}(\text{keV})$ $0.85_{-0.10}^{+0.28}$	$E_{b2}(\text{keV})$ $6.09_{-1.50}^{+1.24}$	$(3.41 \pm 0.5) \times 10^{-12}$	355.18/272	6.97	8×10^{-5}
J212912–153841									
Model	Best-fit Parameters			Flux ^b	$\chi^2/\text{d.o.f.}$	F-stat ^c	$p\text{-values}^{\text{c}}$		
PL	Γ 1.38 ± 0.02			$3.53_{-0.07}^{+0.08} \times 10^{-11}$	813.92/631				
BPL	Γ_1 0.91 ± 0.09	Γ_2 1.56 ± 0.03	E_b (keV) $1.75_{-0.17}^{+0.30}$	$3.00_{-0.08}^{+0.06} \times 10^{-11}$	542.37/629	157.462	10^{-56}		
LP	α 1.10 ± 0.04	β 0.25 ± 0.20	E_0 1 (scale)	$2.57_{-0.07}^{+0.08} \times 10^{-11}$	569.69/630	270.085	10^{-50}		
PL+ABS	Γ 1.51 ± 0.02	$N_{\text{H}}(\text{intr.})$ $2.16 \pm 0.33 \times 10^{22}$		$(3.09 \pm 0.09) \times 10^{-11}$	560.98/630	284.06	10^{-53}		
BPL+ABS ^a	Γ_1 $1.22_{-0.13}^{+0.09}$	Γ_2 1.57 ± 0.03	E_b (keV) 2.21 ± 0.35	$N_{\text{H}}(\text{intr.})$ $0.98_{-0.46}^{+0.24} \times 10^{22}$	$2.95_{-0.10}^{+0.04} \times 10^{-11}$	528.13/628	113.278	10^{-58}	
J020346+113445									
Model	Best-fit Parameters			Flux ^b	C-stat/d.o.f.				
PL	Γ 1.59 ± 0.13			$3.26_{-0.33}^{+0.25} \times 10^{-12}$	376.90/456				
J135406–020603									
Model	Best fit parameters			Flux ^b	C-stat/d.o.f.				
PL	Γ 1.32 ± 0.16			$1.99_{-0.32}^{+0.19} \times 10^{-12}$	398.68/477				

Notes.^a The best-fit model extracted from the analysis is highlighted in bold.^b Observed flux in units of $10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the 0.3–79 keV energy band. The errors are at the 90% level of confidence for one parameter of interest and the fluxes are corrected for Galactic absorption.^c The F-statistic and p -values are calculated for the different models with respect to the power-law case.

with respect to archival data. This could lead to an overestimate of the black hole mass, although the value would still be within the errors associated with the modeling approach. Longer term optical monitoring would be needed to study variability of the disk emission.

3.4. Broadband SED Modeling

We reproduce the broadband SED of all six sources following a simple one-zone leptonic emission model (see Dermer et al. 2009; Ghisellini & Tavecchio 2009) and here we explain it briefly. We assume that radiating electrons are confined in a

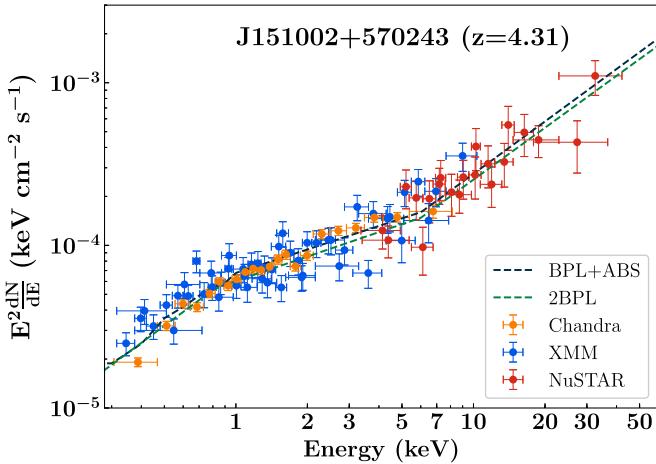


Figure 1. Combined *Chandra*, *XMM*, *Swift*-XRT, and *NuSTAR* spectrum of J151002+570243. The best-fit shapes obtained using XSPEC are a double broken power law (2BPL) as well as a broken power law with excess absorption (BPL+ABS). The low-energy break can be attributed to either an intrinsic property of the underlying electron distribution, or to an intervening absorber in the IGM. The softening (between ~ 0.8 and ~ 6 keV) and rehardening (> 7 keV) of the spectrum can be explained within the one-zone leptonic model (see Section 3.4).

spherical region canvassing the full jet cross section. This region is considered to be located at a distance R_{diss} from the central black hole, moving with bulk Lorentz factor Γ . The jet semi-opening angle is taken to be 0.1 rad. The energy distribution of relativistic electrons is modeled as a broken power law of the following type:

$$N(\gamma) \propto \frac{(\gamma/\gamma_{\text{break}})^{-p}}{(\gamma/\gamma_{\text{break}})^p + (\gamma/\gamma_{\text{break}})^q} \quad (3)$$

where p and q are, respectively, the slopes of the distribution before and after the break energy, γ_{break} . Relativistic electrons, immersed in a randomly oriented but uniform magnetic field (B), emit synchrotron radiation and, in the presence of a photon field, they also undergo IC process. We consider the sources of radiation for IC to be either the same photons emitted via synchrotron or radiation fields external to the jet. For the latter, following Ghisellini & Tavecchio (2009), we take into account (1) the accretion disk emission, (2) the X-ray corona located above the accretion disk, (3) the BLR, and (4) the dusty torus. The corona is assumed to have a cutoff power-law spectral distribution and to reprocess 30% of the disk luminosity. The BLR is considered to be a spherical shell located at a distance $R_{\text{BLR}} = 10^{17} L_{\text{disk},45}^{1/2}$ cm ($L_{\text{disk},45}$ being the accretion disk luminosity in units of 10^{45} erg s $^{-1}$) from the central engine, and reemitting 10% of L_{disk} (Bentz et al. 2006; Kaspi et al. 2007; Tavecchio & Ghisellini 2008). The torus is also assumed to be a spherical shell, located further away at $R_{\text{TORUS}} = 10^{18} L_{\text{disk},45}^{1/2}$ cm and reprocessing 50% of L_{disk} (Lawrence 1991; Nenkova et al. 2002; Cleary et al. 2007). The spectral shapes of both the BLR and the torus are considered to be blackbodies peaking, respectively, at the Ly α frequency and at the characteristic torus temperature ($T_{\text{TORUS}} = 300$ K). Calculating the radiative energy densities of all four components, we are able

Table 4
Table of Black Hole Masses, Derived both from Spectroscopic Approach and SED Modeling

	$M_{\text{BH,SED}}(M_{\odot})$	$M_{\text{BH,spectroscopy}}(M_{\odot})$
J013126-100931	6.0×10^9	...
J020346+113445	3.0×10^9	...
J064632+445116	1.5×10^9	1.3×10^9
J135406-020603	1.0×10^9	8.9×10^8
J151002+570243	6.5×10^9	3.1×10^8
J212912-153841	5.0×10^9	6.3×10^9

to extract the IC spectra (both SSC and EC). In order to calculate the total jet power, we make use of the following assumptions: the kinetic power of the jet is carried only by protons considered to be cold and therefore contributing only to the inertia of the jet; and the number densities of relativistic electrons and protons are equal (e.g., Celotti & Ghisellini 2008). The total jet power is evaluated as the sum of electrons (P_e), protons (P_p), magnetic field (P_m), and radiation (P_r) powers, which are all calculated.

The simultaneous multiwavelength coverage obtained for all our sources is crucial to accurately determine the input parameters to our model (i.e., L_{disk} , R_{diss} , M_{BH} , p , q , Γ , B , low- and high-energy cutoff of the electron distribution (γ_{min} , γ_{max})). Here we list some guidelines behind our choice:

1. Thermal component: optical data of these sources sample the accretion disk emission. It is therefore possible to gauge the peak of this radiation and hence the overall disk luminosity ($L_{\text{disk}} \sim 2\nu_{\text{peak,disk}} L(\nu_{\text{peak,disk}})$). Following the discussion in Section 3.3, optical data also allow us to constrain M_{BH} .
2. Nonthermal component: a blazar's emission region in this one-zone model is required to be compact and located within the BLR or the torus. As a consequence, synchrotron emission from this region is depleted by synchrotron self-absorption below $\sim 10^{12}$ Hz. It follows that the low-energy synchrotron component of the SED is likely to be produced by another region further away along the jet, and therefore cannot be used to constrain the shape of the electron energy distribution. However, good quality X- and γ -rays allow us to exactly determine peak of the IC emission and the shape of the nonthermal continuum, which is directly linked to the shape of the electron distribution. Hence, availability of both *NuSTAR* and *Fermi*-LAT data are key to fix both p and q (see Equation (3)). Furthermore the interplay of Γ , B , and R_{diss} reflects on the level of the high-energy emission. Therefore, the available X-to- γ -ray data allow us to accurately estimate these parameters.

We report the parameters derived from the SED modeling in Table 5. Figure 2 shows the results for all six sources, which are discussed in the next Section.

4. Discussion

High-redshift MeV blazars are among the most extreme sources of the blazar class. Harboring powerful relativistic jets and supermassive black holes at their center, they provide a unique opportunity to study jetted sources in the early universe and are ideal to address open issues such as the accretion disk-jet connection and the growth and evolution of supermassive

Table 5
Summary of the Parameters Used/Derived from the SED Modeling of Six MeV Blazars Shown in Figure 2

Parameter	J013126–100931	J020346+113445	J064632+445116	J135406–020603	J151002+570243	J212912–153841
Slope of the particle distribution below the break energy (p)	1.5	1.8	2.05	0.9	0.5	2.1
Slope of the particle distribution above the break energy (q)	3.8	4.2	3.6	3.9	3.8	4.0
Magnetic field in Gauss (B)	0.5	0.9	1.0	1.2	1.2	0.8
Particle energy density in erg cm $^{-3}$ (U_e)	0.003	0.02	0.004	0.01	0.01	0.009
Bulk Lorentz factor (Γ)	15.8	10.0	14.0	10.0	11.0	10.0
Minimum Lorentz factor (γ_{\min})	12	1	1	1	1.7	4.0
Break Lorentz factor (γ_{break})	74.07	83.85	131.22	59.80	32.42	60.32
Maximum Lorentz factor (γ_{\max})	3e3	3e3	3e3	3e3	5e3	3e3
Dissipation distance (R_{diss}) in parsecs (R_{Sch})	0.43 (750)	0.28 (1000)	0.57 (4000)	0.25 (2700)	0.31 (500)	0.98 (2050)
Size of the BLR in parsec (R_{BLR} , R_{Sch})	0.22 (399)	0.23 (837)	0.38 (2671)	0.23 (2511)	0.27 (438)	0.72 (1514)
Accretion disk luminosity in log scale (L_{disk} , erg s $^{-1}$)	46.70	46.74	47.15	46.74	46.85	47.71
Accretion disk luminosity in Eddington units ($L_{\text{disk}}/L_{\text{Edd}}$)	0.06	0.14	0.74	0.43	0.09	0.79
Jet power in electrons in log scale (P_e , erg s $^{-1}$)	45.19	45.29	45.37	44.93	45.13	45.89
Jet power in magnetic field in log scale (P_B , erg s $^{-1}$)	45.61	45.37	46.36	45.53	45.77	46.33
Radiative jet power in log scale (P_r , erg s $^{-1}$)	45.49	46.11	46.27	46.59	46.37	46.70
Jet power in protons in log scale (P_p , erg s $^{-1}$)	46.90	47.78	47.98	46.89	47.12	48.11
Total jet power in log scale (P_{TOT} , erg s $^{-1}$)	46.93	47.78	47.99	46.91	47.14	48.12

Note. A viewing angle of 3° is adopted for all of them.

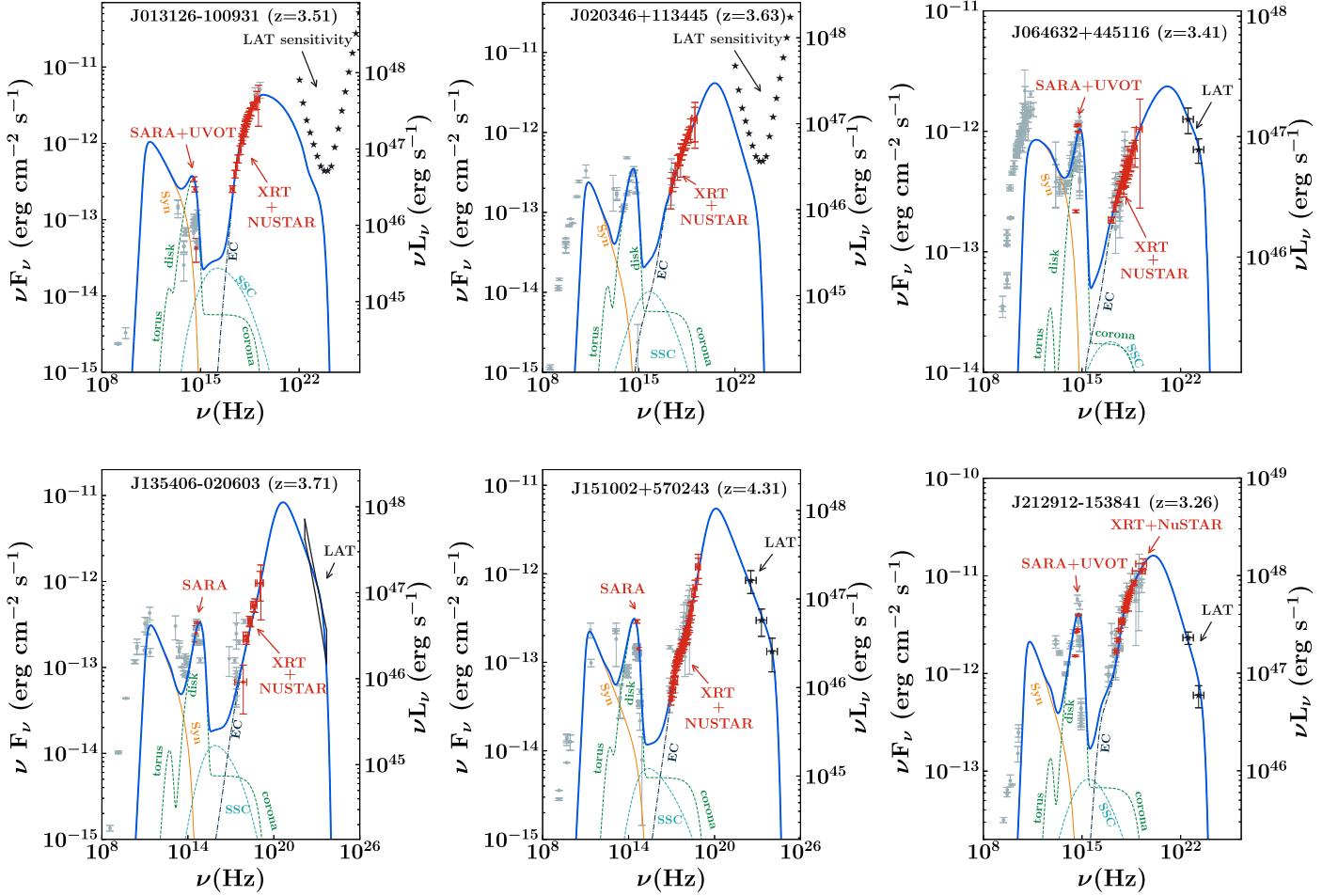


Figure 2. The broadband SED of six quasars using quasi-simultaneous SARA, *Swift*, *NuSTAR*, and *Fermi*-LAT data, modeled via the one-zone leptonic emission model described in the text. Gray circles represent the archival data, the black stars are the *Fermi* data extracted from Ackermann et al. (2017), while the red points show our quasi-simultaneous observations. Since the LAT SED data points for J135406-020603 are not published in Ackermann et al. (2017), we use the parameters listed in their paper to extract the bowtie plot showed here. The lines are the modeling lines illustrating different components considered: the orange line is the synchrotron emission; the green dotted line includes the thermal emission from the torus, the disk, and the corona; the light-blue dotted line represents the SSC component; the dark-blue dotted line is the EC component. The blue solid line is the sum of all the other contributors and delineates our best-fit model (see Section 4 for detailed discussion). For the optical data, when available, we plot the photometric data below the redshifted Ly α frequency since above this limit the emission is significantly affected by absorption from intergalactic hydrogen (e.g., Songaila & Cowie 2010; Croft et al. 2018).

black holes at the dawn of the universe. Since in these sources the high-energy radiation dominates the bolometric luminosity, in order to understand the physics of the jets and the properties of the emitting leptonic population, good coverage at these frequencies is required. Awaiting an all sky MeV instrument (e.g., AMEGO, Rando 2017; e-ASTROGAM, de Angelis et al. 2018) to sample the peak of the IC emission, the most effective way to study the high-energy part of the SED is via a combination of X- and γ -ray data. As a result of the drift of the nonthermal SED peaks to lower frequencies, these high-redshift sources are brighter at X-ray frequencies, displaying a hard X-ray photon index ($\Gamma_X \lesssim 1.5$) and a soft γ -ray one ($\Gamma_\gamma \gtrsim 2.3$). Four of the objects in our study already had γ -ray data available (Ackermann et al. 2017) and, most importantly, we were able to obtain *NuSTAR* observations for all six of them. All our sources show X-ray photon indices < 2 , as can be seen in Table 3, and soft γ -ray ones ($\Gamma_\gamma > 2.5$, Ackermann et al. 2017). We point out that the sources J013126-100931 and J212912-153841 are detected in the BAT 105-months catalog (Oh et al. 2018), with a $\Gamma_{15-150\text{ keV}}$ of $1.81^{+0.53}_{-0.48}$ and

$1.79^{+0.48}_{-0.43}$, respectively. The unprecedented sensitivity of *NuSTAR* allows us to accurately measure the hard X-ray photon index ($\Gamma_X \sim 1.55 \pm 0.1$ and $\Gamma_X \sim 1.56 \pm 0.03$, respectively, see Table 3). Furthermore, *NuSTAR* data are key to sample the rising part of the IC spectrum up to 60–70 keV and, in combination with *Fermi* ones, enable us to better constrain the peak location and consequently the bulk Lorentz factor and the shape of the electron energy distribution underlying this powerful emission. To highlight the importance of *NuSTAR*, in Figures 3–5 we show the zoomed SED for the three targets studied by Ackermann et al. (2017). In these plots, the orange dashed line represents the EC model taken from their paper, while our EC model is shown as the blue solid line. Note that Ackermann et al. (2017) performed the modeling using only X-ray archival observations (shown by orange data points), whereas we benefited from the availability of *NuSTAR* observations (red stars). SED parameters derived in this work and in Ackermann et al. (2017) are also reported in Tables 6–8 for comparison. It is evident how *NuSTAR* data enable us to sample the IC emission closer to the peak than previously

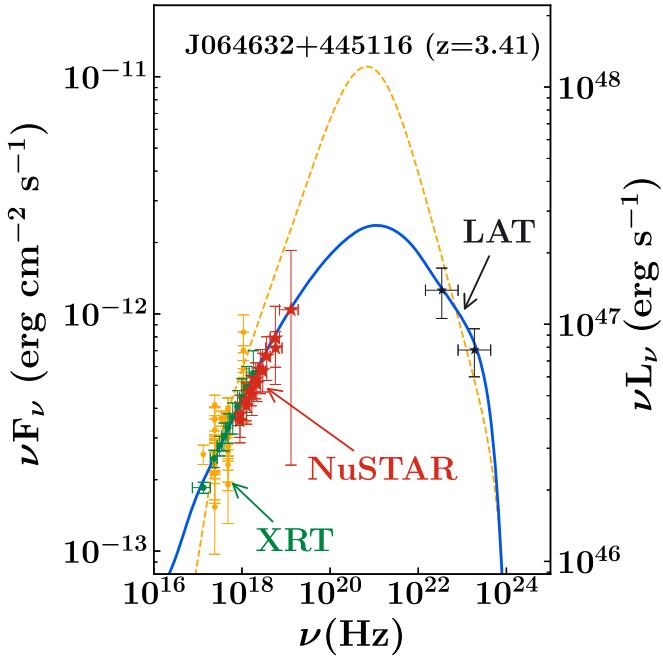


Figure 3. Zoomed high-energy SED of J064632+445116. The orange data points are the archival observations and the orange dashed line is the SED model derived by Ackermann et al. (2017). The blue line is the best-fit model we obtained using combined soft X-ray (green), hard X-ray (red), and γ -ray (black) data. We emphasize that the availability of *NuSTAR* observations allows us to tightly constrain the rising part of the EC peak, enabling us to accurately infer the power of the jet and the underlying particle distribution shape (see Table 6).

Table 6
Table of Comparison for the Parameters Used/Derived from the SED Modeling of J064632+445116 from Ackermann et al. (2017) and This Work

J064632+445116		
Parameter	Ackermann et al. (2017)	This Work
p	1.8	2.05
q	4.4	3.6
B (Gauss)	2.1	1.0
U_e (erg cm $^{-3}$)	0.009	0.004
Γ	12	14
γ_{\min}	1	1
γ_{break}	72	131.22
γ_{\max}	2×10^3	3×10^3
R_{diss} (parsec)	0.25	0.57
R_{BLR} (parsec)	0.37	0.38
P_e (erg s $^{-1}$, in log scale)	44.88	45.37
P_B (erg s $^{-1}$, in log scale)	46.15	46.36
P_r (erg s $^{-1}$, in log scale)	45.89	46.27
P_p (erg s $^{-1}$, in log scale)	47.38	47.98
P_{TOT} (erg s $^{-1}$, in log scale)	47.41	47.99

possible, better constraining the slopes of the radiating leptonic population. For example, looking at J151002+570243 it can be seen how the value of p drastically changes from 1.8 to 0.5, almost a factor of four difference from the previous modeling. These constraints translate into a more precise determination of the IC peak location and luminosity. We point out, as extreme examples, how our modeling predicts the IC peak of J151002+570243, or the IC peak luminosity of J064632+445116, to be

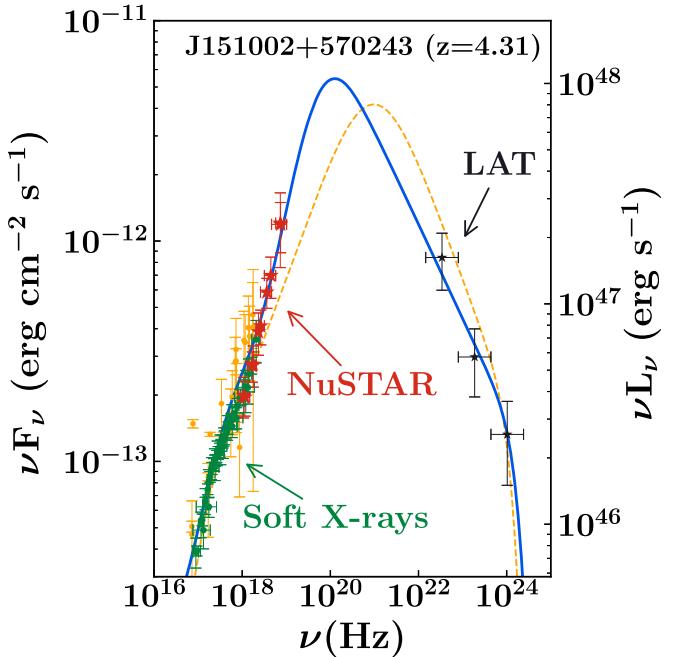


Figure 4. Zoomed high-energy SED of J151002+570243. The orange data points are the archival observations and the orange dashed line is the SED model derived by Ackermann et al. (2017). The blue line is the best-fit model we obtained using combined soft X-ray (green), hard X-ray (red), and γ -ray (black) data.

Table 7
Table of Comparison for the Parameters Used/Derived from the SED Modeling of J151002+570243 from Ackermann et al. (2017) and This Work

J151002+570243		
Parameter	Ackermann et al. (2017)	This Work
p	1.8	0.5
q	4.1	3.8
B [Gauss]	1.4	1.2
U_e (erg cm $^{-3}$)	0.029	0.01
Γ	11	11
γ_{\min}	1	1.7
γ_{break}	82	32.42
γ_{\max}	3×10^3	5×10^3
R_{diss} (parsec)	0.17	0.31
R_{BLR} (parsec)	0.22	0.27
P_e (erg s $^{-1}$, in log scale)	44.94	45.13
P_B (erg s $^{-1}$, in log scale)	45.38	45.77
P_r (erg s $^{-1}$, in log scale)	45.70	46.37
P_p (erg s $^{-1}$, in log scale)	47.45	47.12
P_{TOT} (erg s $^{-1}$, in log scale)	47.47	47.14

nearly an order of magnitude lower than previously found only using soft X-ray data. The location of the emission regions is now more accurately constrained and found for all three sources to lie outside the BLR. These factors in turn produce a more accurate measure of the jet power and particle energy distribution. The substantial difference between the two EC models reflects the necessity of having hard X-ray coverage.

NuSTAR has also proven to be essential in high-redshift blazar studies since, combined with soft X-ray observations, it has enabled us to detect peculiar X-ray features of these

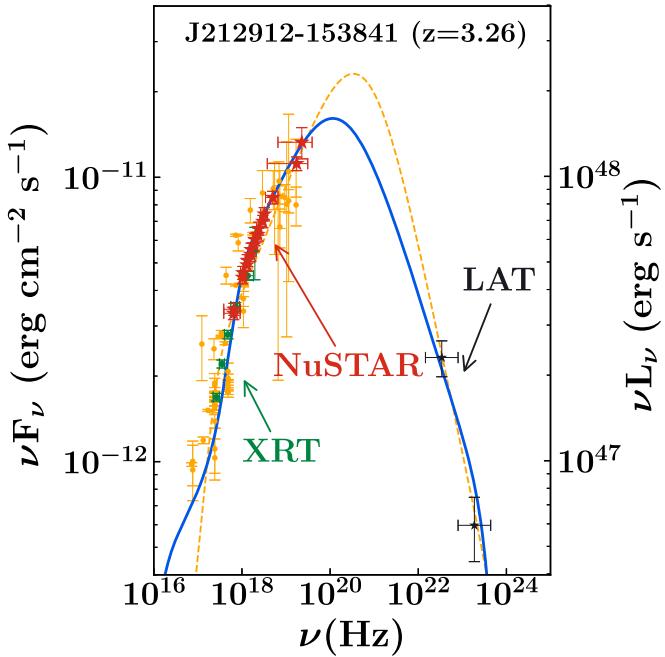


Figure 5. Zoomed high-energy SED of J212912-153841. The orange data points are the archival observations and the orange dashed line is the SED model derived by Ackermann et al. (2017). The blue line is the best-fit model we obtained using combined soft X-ray (green), hard X-ray (red), and γ -ray (black) data.

Table 8

Table of Comparison for the Parameters Used/Derived from the SED Modeling of J212912-153841 from Ackermann et al. (2017) and This Work

J212912-153841		
Parameter	Ackermann et al. (2017)	This Work
p	2.2	2.1
q	4.5	4.0
B (Gauss)	1.3	0.8
U_e (erg cm ⁻³)	0.002	0.009
Γ	14	10
γ_{\min}	1	4
γ_{break}	51	60.32
γ_{\max}	2×10^3	3×10^3
R_{diss} (parsec)	0.59	0.98
R_{BLR} (parsec)	0.79	0.72
P_e (erg s ⁻¹ , in log scale)	45.14	45.89
P_B (erg s ⁻¹ , in log scale)	46.58	46.33
P_r (erg s ⁻¹ , in log scale)	46.31	48.11
P_p (erg s ⁻¹ , in log scale)	47.89	48.12
P_{TOT} (erg s ⁻¹ , in log scale)	47.92	47.99

sources, such as spectral variability (see Tagliaferri et al. 2015; Sbarato et al. 2016) and spectral flattening (see Paliya et al. 2016). The first property is important for various reasons. It could relate to intrinsic changes in the electron distribution (hence injected power) or could be key to unveil acceleration processes happening in the jet, such as magnetic reconnection or shocks (e.g., Spada et al. 2001; Moraïtis & Mastichiadis 2011; Rueda-Becerril et al. 2015; Christie et al. 2019). Furthermore, it has been noticed that often X-ray variability

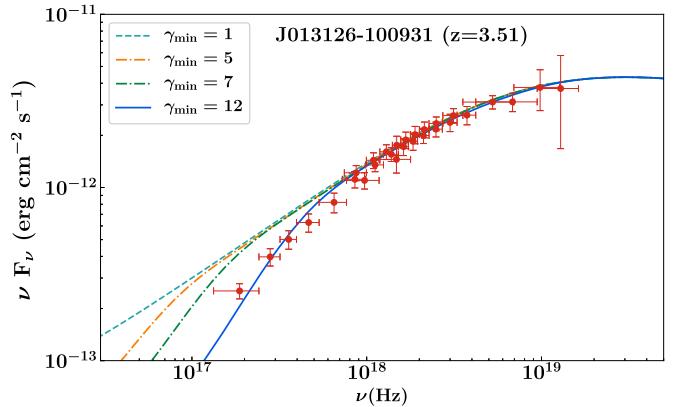


Figure 6. Zoomed SED of J013126-100931, showing the X-ray spectrum. The different lines represent the modeling done with various γ_{\min} values (as labeled). As can be seen, good X-ray coverage allows us to constrain the value of γ_{\min} , which, due to the break in the soft X-ray spectrum, is found to be ~ 12 .

correlates with γ -ray variability, indicating that the assumption of single-electron population models is reliable. However, soft X-ray observations have sometimes not been reflective of this variation. On one hand, sensitivity limits of currently operating instruments could prohibit us from seeing such variability. On the other, the low-energy electron population emitting at these frequencies is not expected to vary significantly (see Section 3.1), while hard X-ray emission could show the signature of such a variation (see discussion in Sbarato et al. 2016). Moreover as these sources become fainter in the γ -rays, observing variability in this high-energy regime becomes a challenge (Li et al. 2018). For all our sources, *NuSTAR* data do not show any variability in the observation period. In the soft X-ray regime, we find both flux and photon index variability only in the case of J212912-153841 during the different observing epochs.

The second property instead finds its cause in different scenarios. It could be related to the intrinsic behavior of the emitting electron population (see Paliya et al. 2016) or could be attributed to an intrinsic absorber along the line of sight, such as the warm-hot IGM (see Arcodia et al. 2018), or a combination of the two. The combination of soft X-ray and *NuSTAR* data enables us to detect spectral curvature in the X-ray spectra of three of our sources (see Section 3.2). Within the available statistics, it appears that the softening below ~ 2 keV in J212912-153841 is a combination of both an intrinsic absorber probably located along the line of sight and a break in the electron distribution. The curvature below ~ 5 keV for J013126-100931 instead is best explained by an intrinsic curvature of the leptonic distribution. Importantly, the shape of the X-ray spectrum constrains the behavior of the electrons underlying this emission. Indeed, if the X-rays are produced via EC rather than SSC, it is possible to constrain the electrons' low-energy cutoff, as can be seen in Figure 6. In fact, for values of $\gamma_{\min} > 1$, our model predicts a significant break which is seen in three of our sources. It follows that joint soft X-ray and *NuSTAR* observations are crucial in determining the minimum energy of the leptonic population underlying the emission, which in turns regulates the jet power.

J151002+570243, the farthest blazar ever detected in γ -rays, is also the most peculiar of our sources in the X-ray regime. Indeed its X-ray spectrum not only favors a softening around

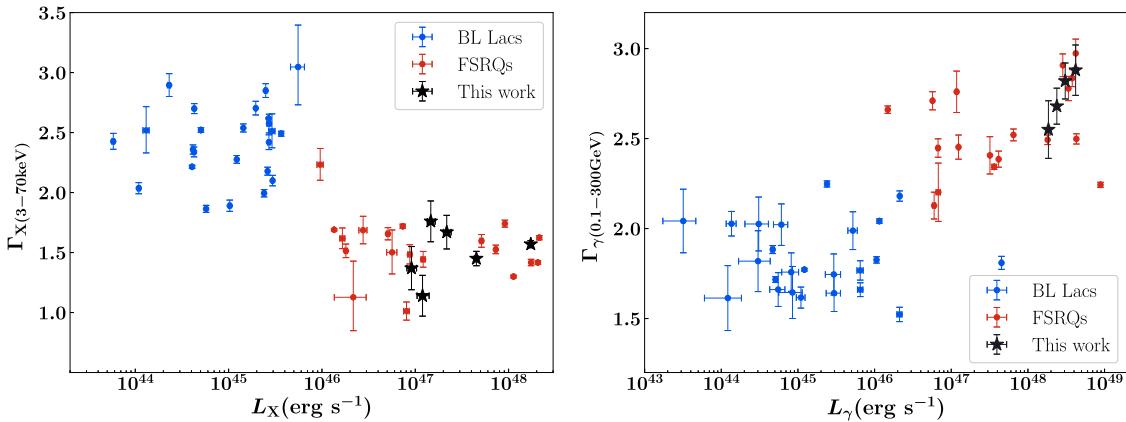


Figure 7. Left panel: *NuSTAR* luminosity vs. X-ray spectral index for all BL Lacs and FSRQs listed in the 3LAC detected by *NuSTAR* (Marcotulli et al. 2020, in preparation). Our sources fall in the lower right corner of the graph, having high X-luminosities and hard Γ_X , and are therefore in agreement with the most powerful blazars. Right panel: *Fermi*-LAT luminosity vs. γ -ray spectral index for the same sources. As can be seen, our high-redshift blazars belong to the most powerful FSRQs, with high γ -luminosities and very soft Γ_γ .

~ 1 keV, but also shows a hardening > 6 keV, in the energy range covered by *NuSTAR*, with an extreme photon index of 0.94. We note that previous studies of this source in the soft X-ray regime found it to have a hard X-ray photon index ($\Gamma_X = 1.55 \pm 0.05$, Wu et al. 2013). This is another instance that reflects the necessity of *NuSTAR* data to constrain the true shape of the X-ray emission, and hence all jet parameters associated with it.

From a modeling point of view, all of our sources have SEDs comparable to the most powerful LSP FSRQ blazars. They show the typical nonthermal humps, as well as the one produced by the accretion disk, unveiled by the SED frequency shift. The low-energy emission is therefore well interpreted by synchrotron process and peaks in the submillimeter regime. Thanks to the optical data obtained with SARA we are able to detect the peak of the disk emission and therefore to estimate disk luminosity and black hole masses from modeling.²⁰ The high-energy part of the SED, covered by *Swift*-XRT, *NuSTAR*, and *Fermi*, instead peaks in (or close to) the MeV energy range and is explained by EC radiation. The X-ray spectra are very hard, which in turn suggests that the EC process is dominant with respect to the SSC (see Ajello et al. 2016 for a detailed discussion). In fact, the SSC is predicted by our modeling to be low for all of the sources, as can be seen in Figure 2. In Figure 7, the six sources here are compared to the *Fermi*-LAT blazars with available *NuSTAR* observations (L. Marcotulli et al. 2020, in preparation). Our sources fall in the same region as the most powerful FSRQs in both plots, showing high X-ray luminosities and hard X-ray indices as well as high γ -ray luminosities and soft γ -ray indices. In all six of them, the IC peak luminosity dominates the synchrotron one ($CD > 1$). In the adopted model, the radiative energy densities of the different AGN components are a function of the distance of the emission region from the central black hole. It follows that we can determine the location of the emission region. For our sources we find it resides outside the BLR and inside the torus (thus in agreement with Costamante et al. 2018). Therefore the torus is the predominant source of radiation for

the EC. Overall, our sources resemble the most extreme MeV blazars.

It has been found that, for γ -ray detected blazars up to $z = 3$, the accretion disk luminosity is lower than the jet power (Ghisellini et al. 2014). More recent results show that a positive correlation remains for γ -quiet sources (i.e., blazars not detected in γ -rays) beyond $z = 3$ (Paliya et al. 2017a). However, as discussed above, having γ -ray data is crucial to accurately determine the jet power. Therefore, with good γ - and X-ray data in hand, as well as optical ones, we can insert our sources in the context of the disk-jet relation. In Table 5 we see how all six have jet powers exceeding disk luminosities. This further hints to the idea that there has to be an extra reservoir of energy powering the jet rather than the accretion disk alone (e.g., energy could be supplied by the spinning black hole, Blandford & Znajek 1977). However, we need to be cautious and take into account uncertainties in our measurements. For example, in our treatment we do not consider electron-positron pairs, which, if present, would reduce the estimated jet power by a factor n_e/n_p ($n_e = n_{e^+} + n_{e^-}$, see Pjanka et al. 2017). Nevertheless, to avoid the Compton Rocket effect that would otherwise break the jet (see Ghisellini & Tavecchio 2010), the number of pairs could not outnumber by a large amount the number of protons. Moreover, another effect that could reduce the jet power is if we consider it to have a spine-sheath structure (Sikora et al. 2016). An underestimation of the accretion disk luminosity could also lead to a false positive relation. So these caveats have to be kept in mind when analyzing the relation between jet power and disk luminosity in blazars.

An important jet parameter that is obtained via modeling is the bulk Lorentz factor. In fact, assuming that the jet points toward the observer at an angle $\lesssim 1/\Gamma$, by a simple geometrical argument, we can estimate the size of the parent population of these sources. It turns out that the detection of one jetted source implies the existence of $2\Gamma^2$ similar sources, at the same redshift, with similar black hole mass, but with jets pointing away from the observer. From various studies, it has been found that high-redshift blazars host supermassive black holes. Using the $2\Gamma^2$ correction, Sbarato et al. (2015) derived the comoving number density of billion solar mass black holes and found that for radio-quiet blazars (i.e., that host weak or no jets)

²⁰ We note that our sources are assumed to be relatively efficient, even though remaining below the Eddington limit. We therefore impose that $10^{-2}L_{\text{Edd}} < L_{\text{disk}} < L_{\text{Edd}}$.

it peaks at $z \sim 2$, while for radio-loud blazars (i.e., that have strong relativistic jets) it peaks at $z \sim 4$. This fact challenges our understanding of black hole growth and evolution in the early universe since, as pointed out by Ghisellini (2015), simple accretion (even considering highly massive seeds at $z = 20$) cannot account for $M_{\text{BH}} > 10^9 M_{\odot}$ at $z > 4$. On account of the fact that the number density strongly depends on the number of sources detected per redshift bin, finding more such sources is crucial to understand the physics of early black hole accretion. The apparent dichotomy in Sbarrato et al. (2015) between radio-quiet and radio-loud sources further hints to a connection between rapid black hole growth and the presence of powerful jets. From SED modeling, five of our sources have $M_{\text{BH}} > 10^9 M_{\odot}$.²¹ We use the prescription in Ackermann et al. (2017) to estimate the number density (n) of these $M_{\text{BH}} > 10^9 M_{\odot}$ radio-loud blazars in the redshift bin $z = [3, 4]$.²² This number was derived by Sbarrato et al. (2015) to be 50 Gpc^{-3} and Ackermann et al. (2017) increased this estimate by 18 Gpc^{-3} . From our calculation, we find $n \sim 37 \text{ Gpc}^{-3}$ which leads the total space density value to be 87 Gpc^{-3} . This seems to further confirm a connection between the radio-loud phase and rapid accretion. Only detecting more sources would improve our understanding of this rapid and early black hole accretion.

5. Conclusions

In this work, the multiwavelength analysis of six high-redshift ($z > 3$) blazars, four of which were recently discovered to be γ -ray loud by Ackermann et al. (2017), is presented. Employing the excellent capabilities of *NuSTAR*, we are able to precisely study X-ray properties of these sources. In combination with quasi-simultaneous *Swift* and *Fermi*-LAT observations, we accurately constrain their jet properties (i.e., jet power, underlying electron distribution, and location of the emission region). Simultaneous optical monitoring with SARA reveals the peak of the disk emission through which we can determine the accretion disk luminosities and black hole masses. The main results are summarized here:

1. *NuSTAR* spectra are very hard ($\Gamma \lesssim 1.5$) and do not show variability. Combining *NuSTAR* and sensitive soft X-ray observations, we detect spectral flattening in the X-ray spectra of three sources. The curvature in J013126–100931 can be explained by a property of the underlying electron distribution (i.e., it depends on the minimum energy of the electron population). In the case of J212912–153841, it is attributed to a combination of both intrinsic curvature in the leptonic distribution and an intrinsic absorber in the IGM. The peculiar case of J151002+570243 ($z = 4.31$) reveals two breaks in its X-ray spectrum, the first possibly attributed to the minimum energy of the electrons producing this emission and/or an IGM absorber, while the second can be explained in the context of the one-zone leptonic emission model.
2. Using a one-zone leptonic emission model, *NuSTAR* and *Fermi*-LAT spectra allow us to precisely pinpoint the position of the high-energy peak, which is well described by EC emission of the relativistic electrons in the jet. This

enables us to determine jet parameters such as the jet power and the emission region location (for all sources outside the BLR), as well as constrain the shape of the underlying electron distribution.

3. Inserting our sources in the context of accretion disk-jet connection, we find that all sources display jet powers larger than accretion disk luminosity, validating the possibility of an extra source of power to the jet, such as the spinning black hole.
4. The optical data unveil the disk emission and enable us to estimate the disk luminosity and black hole mass. All sources have $M > 10^9 M_{\odot}$, which further raises the space density of supermassive black holes in the redshift bin $z = [3, 4]$ to 87 Mpc^{-3} .

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Part of this work is based on archival data, software, or online services provided by the ASI Data Center (ASDC). This research has made use of the XRT Data Analysis Software (XRTDAS).

Appendix X-Ray Spectral Fits

This Appendix contains the X-ray spectra of the six blazars as well as the residuals of the fits (see Figure 8) in the same order as listed in Table 3. The green dotted line is the best fit found analyzing the residuals as well as performing F-tests (see Section 3.2).

²¹ We consider the masses derived by SED modeling for the sources for which we have obtained optical data.

²² Since four of these sources are the ones discovered by Ackermann et al. (2017), we use the same parameters listed in that paper.

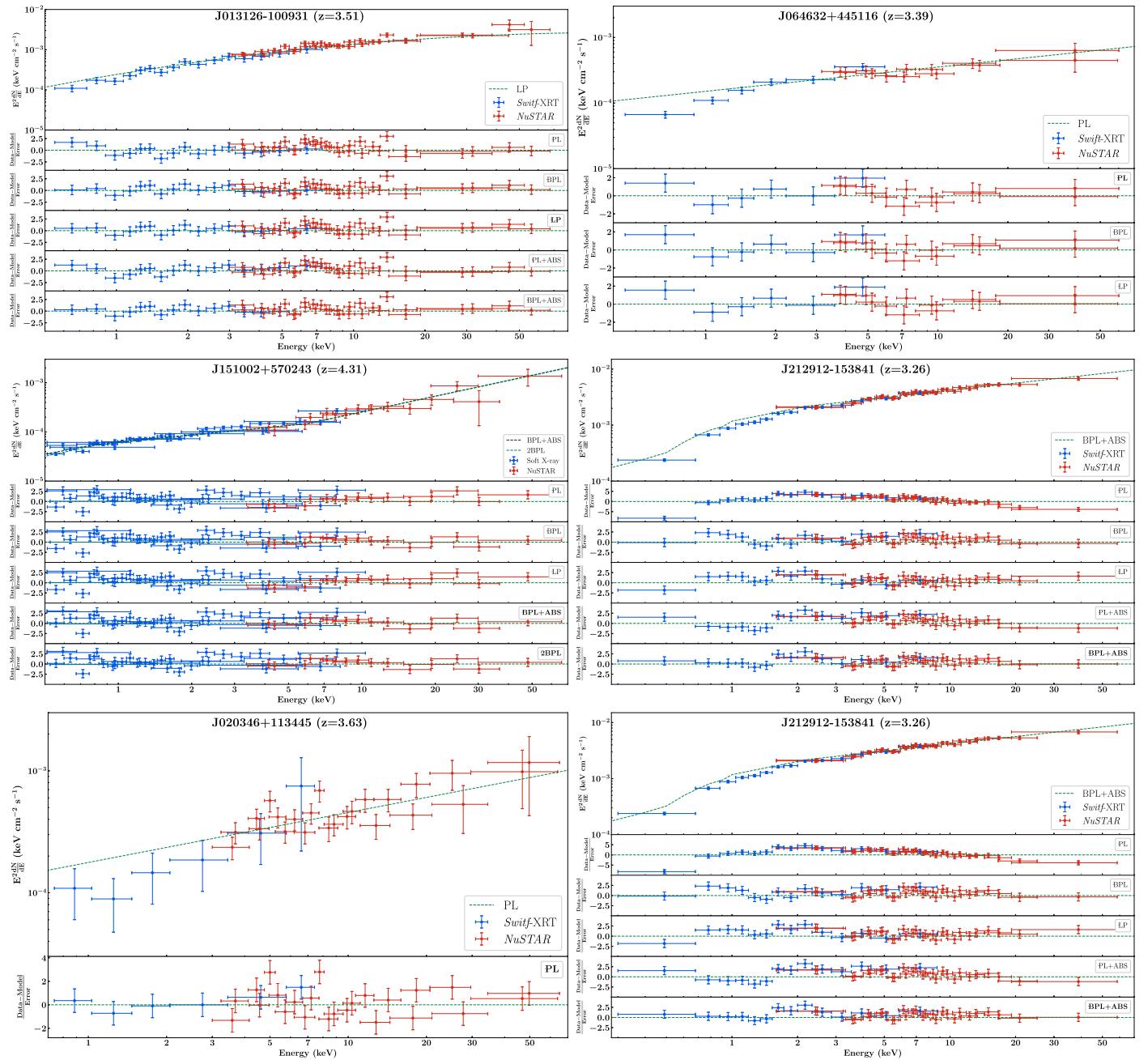


Figure 8. X-ray spectral fits (top panel) and fit residuals for the different fitted models (lower panels) for the six blazars analyzed in this work. The blue and red points are, respectively, the soft and hard X-ray data collected for the sources. The green dashed line of the top panel represents the best-fit model extracted from the X-ray residuals and F-test analysis for the different sources.

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References

Abdo, A. A., Ackermann, M., Agudo, I., et al. 2010a, *ApJ*, **716**, 30
 Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010b, *ApJS*, **188**, 405
 Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010c, *ApJ*, **722**, 520
 Acero, F., Ackermann, M., Ajello, M., et al. 2015, *ApJS*, **218**, 23
 Ackermann, M., Ajello, M., Atwood, W. B., et al. 2015, *ApJ*, **810**, 14
 Ackermann, M., Ajello, M., Baldini, L., et al. 2017, *ApJL*, **837**, L5
 Ajello, M., Costamante, L., Sambruna, R. M., et al. 2009, *ApJ*, **699**, 603
 Ajello, M., Ghisellini, G., Paliya, V. S., et al. 2016, *ApJ*, **826**, 76
 Arcodia, R., Campana, S., Salvaterra, R., & Ghisellini, G. 2018, *A&A*, **616**, A170
 Atwood, W. B., Abdo, A. A., Ackermann, M., et al. 2009, *ApJ*, **697**, 1071
 Bentz, M. C., Peterson, B. M., Pogge, R. W., Vestergaard, M., & Onken, C. A. 2006, *ApJ*, **644**, 133
 Blandford, R. D., & Znajek, R. L. 1977, *MNRAS*, **179**, 433
 Bonning, E. W., Bailyn, C., Urry, C. M., et al. 2009, *ApJL*, **697**, L81
 Breeveld, A. A., Landsman, W., Holland, S. T., et al. 2011, in AIP Conf. Ser. 1358, Gamma Ray Bursts, ed. J. E. McEnery, J. L. Racusin, & N. Gehrels (Melville, NY: AIP), 373
 Burrows, D. N., Hill, J. E., Nousek, J. A., et al. 2005, *SSRv*, **120**, 165

Calderone, G., Ghisellini, G., Colpi, M., & Dotti, M. 2013, *MNRAS*, **431**, 210

Cash, W. 1979, *ApJ*, **228**, 939

Castignani, G., Haardt, F., Lapi, A., et al. 2013, *A&A*, **560**, A28

Celotti, A., & Ghisellini, G. 2008, *MNRAS*, **385**, 283

Chen, Y.-Y., Zhang, X., Xiong, D., & Yu, X. 2015, *AJ*, **150**, 8

Christie, I. M., Petropoulou, M., Sironi, L., & Giannios, D. 2019, *MNRAS*, **482**, 65

Cleary, K., Lawrence, C. R., Marshall, J. A., Hao, L., & Meier, D. 2007, *ApJ*, **660**, 117

Costamante, L., Cutini, S., Tosti, G., Antolini, E., & Tramacere, A. 2018, *MNRAS*, **477**, 4749

Croft, R. A. C., Miralda-Escudé, J., Zheng, Z., Blomqvist, M., & Pieri, M. 2018, *MNRAS*, **481**, 1320

de Angelis, A., Tatischeff, V., Grenier, I. A., et al. 2018, *JHEAp*, **19**, 1

De Gouveia Dal Pino, E., del Valle, M. V., Kadowaki, L., et al. 2016, arXiv:1608.03173

Dermer, C. D., Finke, J. D., Krug, H., & Böttcher, M. 2009, *ApJ*, **692**, 32

Donato, D., Sambruna, R. M., & Gliozzi, M. 2005, *A&A*, **433**, 1163

Drake, A. J., Djorgovski, S. G., Mahabal, A., et al. 2009, *ApJ*, **696**, 870

Eitan, A., & Behar, E. 2013, *ApJ*, **774**, 29

Finke, J. D. 2016, *ApJ*, **830**, 94

Frank, J., King, A., & Raine, D. J. 2002, in *Accretion Power in Astrophysics*, ed. J. Frank, A. King, & D. Raine (Cambridge: Cambridge Univ. Press), 398

Fruscione, A., McDowell, J. C., Allen, G. E., et al. 2006, *Proc. SPIE*, **6270**, 62701V

Gao, X.-Y., Wang, J.-C., & Zhou, M. 2011, *RAA*, **11**, 902

Ghisellini, G. (ed.) 2013, *Lecture Notes in Physics, Radiative Processes in High Energy Astrophysics*, Vol. 873 (Berlin: Springer)

Ghisellini, G. 2015, *JHEAp*, **7**, 163

Ghisellini, G., Celotti, A., Fossati, G., Maraschi, L., & Comastri, A. 1998, *MNRAS*, **301**, 451

Ghisellini, G., Della Ceca, R., Volonteri, M., et al. 2010, *MNRAS*, **405**, 387

Ghisellini, G., Foschini, L., Tavecchio, F., & Pian, E. 2007, *MNRAS*, **382**, L82

Ghisellini, G., & Maraschi, L. 1989, *ApJ*, **340**, 181

Ghisellini, G., Perri, M., Costamante, L., et al. 2019, *A&A*, **627**, A72

Ghisellini, G., & Tavecchio, F. 2009, *MNRAS*, **397**, 985

Ghisellini, G., & Tavecchio, F. 2010, *MNRAS*, **409**, L79

Ghisellini, G., & Tavecchio, F. 2015, *MNRAS*, **448**, 1060

Ghisellini, G., Tavecchio, F., Maraschi, L., Celotti, A., & Sbarato, T. 2014, *Natur*, **515**, 376

Harrison, F. A., Craig, W. W., Christensen, F. E., et al. 2013, *ApJ*, **770**, 103

Jansen, F., Lumb, D., Altieri, B., et al. 2001, *A&A*, **365**, L1

Jiang, N. 2018, *RNAAS*, **2**, 134

Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, *A&A*, **440**, 775

Kaspi, S., Brandt, W. N., Maoz, D., et al. 2007, *ApJ*, **659**, 997

Keel, W. C., Oswalt, T., Mack, P., et al. 2017, *PASP*, **129**, 015002

Kelly, B. C., Bechtold, J., & Siemiginowska, A. 2009, *ApJ*, **698**, 895

Landolt, A. U. 1992, *AJ*, **104**, 340

Lawrence, A. 1991, *MNRAS*, **252**, 586

Li, S., Xia, Z.-Q., Liang, Y.-F., Liao, N.-H., & Fan, Y.-Z. 2018, *ApJ*, **853**, 159

Liao, N.-H., Dou, L.-M., Jiang, N., et al. 2019, *ApJL*, **879**, L9

Madejski, G., & Sikora, M. 2016, *ARA&A*, **54**, 725

Mao, L., Zhang, X., & Yi, T. 2018, *Ap&SS*, **363**, 167

Marchesini, E. J., Andruhov, I., Cellone, S. A., et al. 2016, *A&A*, **591**, A21

Marcotulli, L., Paliya, V. S., Ajello, M., et al. 2017, *ApJ*, **839**, 96

Massaro, E., Perri, M., Giommi, P., & Neschi, R. 2004a, *A&A*, **413**, 489

Massaro, E., Perri, M., Giommi, P., Neschi, R., & Verrecchia, F. 2004b, *A&A*, **422**, 103

Massaro, E., Tramacere, A., Perri, M., Giommi, P., & Tosti, G. 2006, *A&A*, **448**, 861

Mattox, J. R., Bertsch, D. L., Chiang, J., et al. 1996, *ApJ*, **461**, 396

McHardy, I. 2008, in *Blazar Variability across the Electromagnetic Spectrum (BLAZARS 2008)*, 14 (Trieste: SISSA), 14

Moraitis, K., & Mastichiadis, A. 2011, *A&A*, **525**, A40

Nenkova, M., Ivezic, Z., & Elitzur, M. 2002, *ApJL*, **570**, L9

Nolan, P. L., Abdo, A. A., Ackermann, M., et al. 2012, *ApJS*, **199**, 31

Oh, K., Koss, M., Markwardt, C. B., et al. 2018, *ApJS*, **235**, 4

Page, K. L., Reeves, J. N., O'Brien, P. T., & Turner, M. J. L. 2005, *MNRAS*, **364**, 195

Paliya, V. S., Ajello, M., Ojha, R., et al. 2019a, *ApJ*, **871**, 211

Paliya, V. S., Koss, M., Trakhtenbrot, B., et al. 2019b, *ApJ*, **881**, 12

Paliya, V. S., Marcotulli, L., Ajello, M., et al. 2017a, *ApJ*, **851**, 33

Paliya, V. S., Parker, M. L., Fabian, A. C., & Stalin, C. S. 2016, *ApJ*, **825**, 74

Paliya, V. S., Stalin, C. S., Ajello, M., & Kaur, A. 2017b, *ApJ*, **844**, 32

Pjanka, P., Zdziarski, A. A., & Sikora, M. 2017, *MNRAS*, **465**, 3506

Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, **594**, A13

Poole, T. S., Breeveld, A. A., Page, M. J., et al. 2008, *MNRAS*, **383**, 627

Prince, R. 2019, *ApJ*, **871**, 101

Rando, R. 2017, *JInst*, **12**, C11024

Rani, P., Stalin, C. S., & Rakshit, S. 2017, *MNRAS*, **466**, 3309

Roman, R. W., Sowards-Emmerd, D., Greenhill, L., & Michelson, P. 2004, *ApJL*, **610**, L9

Roming, P. W. A., Kennedy, T. E., Mason, K. O., et al. 2005, *SSRv*, **120**, 95

Roming, P. W. A., Koch, T. S., Oates, S. R., Porterfield, B. L., & vander Berk, D. E. 2008, in *AIP Conf. Ser.* 1065, ed. Y.-F. Huang, Z.-G. Dai, & B. Zhang (Melville, NY: AIP), 81

Ruan, J. J., Anderson, S. F., Dexter, J., & Agol, E. 2014, *ApJ*, **783**, 105

Rueda-Becerril, J. M., Mimica, P., & Aloy, M. A. 2015, arXiv:1502.07882

Rumbaugh, N., Shen, Y., Morganson, E., et al. 2018, *ApJ*, **854**, 160

Saez, C., Brandt, W. N., Shemmer, O., et al. 2011, *ApJ*, **738**, 53

Sbarato, T., Ghisellini, G., Tagliaferri, G., et al. 2014, *MNRAS*, **446**, 2283

Sbarato, T., Ghisellini, G., Tagliaferri, G., et al. 2015, *MNRAS*, **446**, 2483

Sbarato, T., Ghisellini, G., Tagliaferri, G., et al. 2016, *MNRAS*, **462**, 1542

Sbarato, T., Tagliaferri, G., Ghisellini, G., et al. 2013, *ApJ*, **777**, 147

Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, **737**, 103

Shakura, N. I., & Sunyaev, R. A. 1973, *A&A*, **24**, 337

Shen, Y., Richards, G. T., Strauss, M. A., et al. 2011, *ApJS*, **194**, 45

Sikora, M., Begelman, M. C., & Rees, M. J. 1994, *ApJ*, **421**, 153

Sikora, M., Rutkowski, M., & Begelman, M. C. 2016, *MNRAS*, **457**, 1352

Sikora, M., Stawarz, Ł., Moderski, R., Nalewajko, K., & Madejski, G. M. 2009, *ApJ*, **704**, 38

Smith, J. A., Tucker, D. L., Kent, S., et al. 2002, *AJ*, **123**, 2121

Songaila, A., & Cowie, L. L. 2010, *ApJ*, **721**, 1448

Spada, M., Ghisellini, G., Lazzati, D., & Celotti, A. 2001, *MNRAS*, **325**, 1559

Stroh, M. C., & Falcone, A. D. 2013, *ApJS*, **207**, 28

Tagliaferri, G., Ghisellini, G., Perri, M., et al. 2015, *ApJ*, **807**, 167

Tavecchio, F., & Ghisellini, G. 2008, *MNRAS*, **386**, 945

Tavecchio, F., Maraschi, L., Ghisellini, G., et al. 2000, *ApJ*, **543**, 535

Urry, C. M., & Padovani, P. 1995, *PASP*, **107**, 803

Vestergaard, M., & Peterson, B. M. 2006, *ApJ*, **641**, 689

Volonteri, M., Haardt, F., Ghisellini, G., & Della Ceca, R. 2011, *MNRAS*, **416**, 216

Weisskopf, M. C., Tananbaum, H. D., Van Speybroeck, L. P., & O'Dell, S. L. 2000, *Proc. SPIE*, **4012**, 2

Wu, J., Brandt, W. N., Miller, B. P., et al. 2013, *ApJ*, **763**, 109