



BAT AGN Spectroscopic Survey. XVI. General Physical Characteristics of BAT Blazars

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Abstract

The recently released 105-month *Swift*-Burst Alert Telescope (BAT) all-sky hard X-ray survey catalog presents an opportunity to study astrophysical objects detected in the deepest look at the entire hard X-ray (14–195 keV) sky. Here we report the results of a multifrequency study of 146 blazars from this catalog, quadrupling the number compared to past studies, by utilizing recent data from the *Fermi*-Large Area Telescope (LAT), *Swift*-BAT, and archival measurements. In our γ -ray analysis of ~ 10 yr of the LAT data, 101 are found as γ -ray emitters, whereas, 45 remains LAT undetected. We model the broadband spectral energy distributions with a synchrotron-inverse Compton radiative model. On average, BAT detected sources host massive black holes ($M_{\text{bh}} \sim 10^9 M_{\odot}$) and luminous accretion disks ($L_{\text{d}} \sim 10^{46} \text{ erg s}^{-1}$). At high redshifts ($z > 2$), BAT blazars host more powerful jets with luminous accretion disks compared to those detected only with *Fermi*-LAT. We find good agreement in the black hole masses derived from the single-epoch optical spectroscopic measurements and standard accretion disk modeling approaches. Other physical properties of BAT blazars are similar to those known for *Fermi*-LAT detected objects.

Key words: galaxies: active – galaxies: jets

Supporting material: figure set, machine-readable tables

1. Introduction

Active galactic nuclei (AGNs) are among the most energetic objects in the universe and are crucial players in the evolution of galaxies (see, e.g., Kormendy & Ho 2013, for a review). A subset of these interesting objects host relativistic jets and when the jet is closely aligned with the line of sight to the observer, the source is termed as a blazar (Blandford & Rees 1978). Due to their peculiar orientation, the emitted radiation from blazar jets is relativistically amplified. This phenomenon makes blazars visible even at very high redshifts ($z > 2$, e.g., Romani et al. 2004; Ackermann et al. 2017). Blazars are classified as flat spectrum radio quasars (FSRQs) and/or BL Lac objects based on the rest-frame equivalent width (EW) of the optical emission lines (Stickel et al. 1991) with FSRQs exhibiting broad emission lines ($\text{EW} > 5 \text{ \AA}$), while BL Lacs show weak or no emission lines in their optical spectra. This, in turn, suggests a radiatively efficient accretion process that illuminates the broad-line region (BLR) surrounding the central engine of FSRQs (see Ghisellini et al. 2011). The spectral energy distribution (SED) of a blazar is characterized by a double hump structure. The low-energy peak is understood to be due to a synchrotron process and is typically observed between the infrared and the X-ray band, whereas, the high-energy peak is typically explained by the inverse Compton scattering of the low-energy

photons by relativistic electrons in the jet. Alternatively, hadronic models have also been invoked to explain the high-energy SEDs of blazars (e.g., Böttcher et al. 2013). Based on the location of the synchrotron peak frequency ($\nu_{\text{syn}}^{\text{peak}}$), blazars are also classified as low- ($\nu_{\text{syn}}^{\text{peak}} < 10^{14} \text{ Hz}$), intermediate- ($10^{14} < \nu_{\text{syn}}^{\text{peak}} < 10^{15} \text{ Hz}$), and high- ($\nu_{\text{syn}}^{\text{peak}} > 10^{15} \text{ Hz}$) synchrotron peaked, or LSP, ISP, and HSP, respectively (Abdo et al. 2010). Typically FSRQs are LSP/ISP blazars, whereas, BL Lacs are mostly HSP ones.

The sky surveying capabilities of the currently operating hard X-ray and γ -ray instruments, namely the Burst Alert Telescope (BAT; Barthelmy et al. 2005) on board *Neil Gehrels Swift* observatory (Gehrels et al. 2004) and the Large Area Telescope (LAT; Atwood et al. 2009) on board *Fermi* Gamma-Ray Space Telescope, have been allowed to carry out the population studies of blazars and their luminosity-dependent evolution (e.g., Ajello et al. 2009, 2012, 2014). In particular, Ajello et al. (2009) studied 38 blazars detected in the first 36 months of the all-sky survey by *Swift*-BAT and used this sample to derive the 15–55 keV luminosity function (LF) of blazars. Interestingly, it was noted that the evolution of luminous FSRQs peaks at higher redshifts (~ 4) compared to other classes of BAT detected AGNs. Ghisellini et al. (2010) used the LF reported in Ajello et al. (2009) to determine the space density of billion-solar-mass black holes residing in jetted systems and introduced an exponential cut-off in the blazar LF above $z = 4.3$ to remain

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consistent with the number of detected blazars in that redshift bin and also with the space density of massive halos. However, computation of the LF depends strongly on the number of sources detected in a particular redshift bin. For example, Ackermann et al. (2017) reported, for the first time, γ -ray detection of five $z > 3.1$ radio-loud quasars thus confirming their blazar nature. This allowed them to update the space density of $10^9 M_\odot$ black holes residing in radio-loud quasars at $z \approx 4$ and to conclude that radio-loudness may be a crucial ingredient for the rapid black hole growth in the early universe (see also Volonteri et al. 2011). Moreover, by studying high-redshift ($z > 2$) *Swift*-BAT blazars, Ghisellini et al. (2010) found them to host more powerful jets and more luminous accretion disks compared to *Fermi*-LAT detected $z > 2$ sources.

A new catalog of *Swift*-BAT detected sources covering the first 105-month of the mission has recently been released (Oh et al. 2018) along with an extensive multiwavelength observations of BAT detected sources by the BAT AGN Spectroscopic Survey (BASS) collaboration (Koss et al. 2017; Lamperti et al. 2017; Ricci et al. 2017). Motivated by more than quadrupling the number of the sources compared to past studies and also by the larger redshift range it covers,¹⁵ we have carried out a multiwavelength study of all the blazars present in the 105-month *Swift*-BAT catalog. Our primary objectives are to explore the physical properties of *Swift*-BAT blazars by applying a simple leptonic radiative model and also to compare them with that known for the *Fermi*-LAT detected ones, which are yet to be detected by BAT.

In Section 2, we describe the sample adopted in this work and data analysis procedures are explained in Section 3. The adopted leptonic model is elaborated in Section 4 and we discuss the derived SED parameters in Section 5. In Section 6, we compare the accretion-jet connection observed in BAT blazars with *Fermi*-LAT detected sources, and Section 7 is devoted to the luminosity-dependent evolution of blazars. We summarize our findings in Section 8. Throughout this work, we use a flat cosmology with $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_M = 0.308$ (Planck Collaboration et al. 2016).

2. The Sample

Our blazar sample is based on the 1632 sources included in the 105-month *Swift*-BAT catalog. There are 158 sources classified as “beamed AGN,” mainly based on their presence in the BZCAT and Combined Radio All-Sky Targeted Eight-GHz Survey (CRATES) catalogs (Healey et al. 2007; Massaro et al. 2015). We exclude those 32 objects that are reported as “beamed AGNs” (based on the BZCAT catalog) but, however, do not exhibit broadband properties typically observed from the blazar class of AGN. The full list of these sources is provided in Table 1. Additionally, we explore the broadband properties of the whole BAT sample to identify potential blazar candidates that might have been misclassified as Seyferts. This was done by examining their multifrequency SEDs generated using the ASI Data Center (ASDC) SED builder tool¹⁶ and physical properties (e.g., radio detection) reported in the literature (see Griffith & Wright 1993; Gregory et al. 1996). This led to the inclusion of 20 sources (Table 2). Admittedly,

there could be a very few unidentified BAT objects that might be blazars though it may not be possible to pinpoint them due to lack of a confirmed optical/radio counterparts and also because of paucity of the multiwavelength information. The BAT survey completeness and redshift evolution of beamed AGN will be fully examined in a future study (L. Marcotulli et al. 2019, in preparation).

This paper makes use of redshift and black hole mass estimates from BASS, a large effort to collect optical spectra for all the *Swift*-BAT AGN. The Data Release 1 optical spectra were obtained from a large variety of telescopes (Koss et al. 2017) and can all be viewed on the BASS website¹⁷. We also include 15 new Data Release 2 redshifts measurements, from Palomar, Southern Astrophysical Research (SOAR) telescope, and the Very Large Telescope (VLT) that will be published shortly (K. Oh et al. 2019, in preparation) as well as a comparison black hole mass measurement from this catalog. Altogether, our final sample consists of 146 BAT blazars.

3. Data Reduction and Compilation

3.1. Fermi-LAT

The primary objective of this study is to explore the average physical characteristics of BAT blazars rather than any of their specific activity states. With this in mind, we perform an analysis of *Fermi*-LAT data covering ~ 10 yr (2008 August 5–2018 March 3) of *Fermi* operation. The all-sky surveying capability of *Fermi*-LAT ensures that the generated spectra represent an average activity state of the blazar. A standard data reduction procedure is adopted to carry out the analysis.¹⁸ The data reduction is performed using Science Tools (v11r5p3) and the publicly available analysis package *fermiPy* (Wood et al. 2017). We use the latest P8R3 data set (Atwood et al. 2013; Bruel et al. 2018) in the energy range of 0.1–300 GeV and consider only SOURCE class events (evclass=128). We define a region of interest (ROI) of 15° radius centered at the target blazar and apply standard cuts (zenith angle, $z_{\text{max}} < 90^\circ$, and ‘DATA_QUAL>0 and LAT_CONFIG==1’) to select good time intervals. To derive the optimized spectral parameters from the binned-likelihood fitting, we consider all γ -ray sources present in third *Fermi*-LAT detected source catalog (3FGL; Acero et al. 2015) and lying within the ROI and also the isotropic and Galactic diffuse emission models (Acero et al. 2016). Note that the presence of any extended γ -ray source in the ROI is properly taken into account by adopting the publicly available extended source templates.¹⁹ The parameters of all of the considered γ -ray sources are allowed to vary during the likelihood fitting.

The significance of source detection is quantified using maximum likelihood test statistic $TS = 2\Delta \log \mathcal{L}$, where $\mathcal{L}$ denotes the likelihood ratio, between models with and without a point source at the position of interest. Since the time period covered is significantly larger than that used to produce the 3FGL catalog, we adopt the *findsource* tool available in *fermiPy* to generate test statistic (TS) maps and search unmodeled γ -ray sources that are present in the data (with $TS \gtrsim 25$, $\gtrsim 5\sigma$ detection; Mattox et al. 1996) but not in the model. Once found, they are characterized with a power law and are included in the sky model. This procedure is repeated until

¹⁵ The farthest known blazar in 105-month catalog, SWIFT J1430.6+4211, has $z = 4.71$, which is significantly larger than the farthest known sources in previous BAT catalogs ($z = 3.67$; e.g., Ajello et al. 2009).

¹⁶ <https://tools.asdc.asi.it/>

¹⁷ www.bass-survey.com

¹⁸ <http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/>

¹⁹ <https://github.com/fermi-lat/extendedArchive>

Table 1
Swift-BAT Detected Sources that Were Classified as “Beamed AGN” in 105-month Catalog but Excluded in This Work

BAT Name	Counterpart Name	Redshift	Comments
J0048.8+3155	Mrk 348	0.015	Seyfert 2 galaxy (Hernández-García et al. 2015)
J0136.5+3906	B3 0133+388	...	No redshift
J0142.0+3922	B2 0138+39B	0.08	<i>Fermi</i> -LAT undetected HSP BL Lac object
J0156.5–5303	RBS 259	...	No redshift
J0225.8+5946	NVSS J022626+592753	...	No redshift
J0241.3–0816	NGC 1052	0.005	LINER galaxy (Cazzoli et al. 2018)
J0245.2+1047	4C +10.08	0.07	<i>Fermi</i> -LAT undetected HSP BL Lac object
J0319.7+4132	NGC 1275	0.018	Seyfert 1.5 galaxy (Véron-Cetty & Véron 2010)
J0433.0+0521	3C 120	0.033	Radio galaxy (Tombesi et al. 2017)
J0550.7–3215B	2XMM J055054.3–321616	...	No redshift, counterpart doubtful
J0608.9–5507	PKS 0607–549	...	No redshift
J0612.2–4645	PMN J0612–4647	...	No redshift
J0632.1–5404	1RXS J063200.7–540454	0.204	Giant radio quasar (Saripalli et al. 2005)
J0640.3–1286	PMN J0640–1253	...	No redshift
J0709.3–1527	PKS 0706–15	...	No redshift
J0923.2+3850	B2 0920+39	...	No redshift
J1213.2–6020	1RXS J121324.5–601458	...	No redshift
J1304.5–5651	IGR J13045–5630	0.051	Radio and <i>Fermi</i> -LAT undetected, Galactic plane
J1312.1–5631	2MASX J13103701–5626551	...	No redshift
J1325.4–4301	Cen A	0.002	Radio galaxy
J1512.2–1053A	NVSS J151148–105023	...	No redshift
J1557.8–7913	PKS 1549–79	0.150	Radio galaxy (Guillard et al. 2012)
J1655.0–4998	CXOU J165551.9–495732	0.058	Radio and <i>Fermi</i> -LAT undetected, Galactic plane
J1719.7+4900	ARP 102B	0.024	LINER/Radio galaxy (Popović et al. 2014)
J1734.9–2074	NVSS J173459–204533	...	No redshift
J1742.1–6054	PKS 1737–60	0.41	<i>Fermi</i> -LAT undetected HSP BL Lac object
J1941.3–6216	PKS 1936–623	...	No redshift
J2037.2+4151	SSTSL2 J203705.58+415005.3	...	No redshift
J2056.8+4939	RX J2056.6+4940	...	No redshift
J2117.5+5139	2MASX J21174741+5138523	...	No redshift
J2209.4–4711	NGC 7213	0.006	Seyfert 1 galaxy (Halpern & Filippenko 1984)
J2303.1–1837	PKS 2300–18	0.128	Radio galaxy (Robinson et al. 1987)

Note. We exclude three HSP blazars, J0142.0+3922, J0245.2+1047, and J1742.1–6054, which are not detected with *Fermi*-LAT. This is because the lack of the γ -ray spectrum leaves their inverse Compton emission completely unconstrained. LINER is also known as the Low Ionization Nuclear Emission Region.

the TS map stops showing any excess residual. Flux upper limits in the γ -ray spectra are derived at 95% confidence for energy bins with $TS < 9$. Finally, we compute 5σ LAT sensitivity limits by adopting a photon index of 2.4 for each γ -ray undetected BAT FSRQ (median 0.1–300 GeV 5σ flux upper limit $\sim 7 \times 10^{-7}$ MeV cm $^{-2}$ s $^{-1}$). The generated γ -ray spectra are corrected for extragalactic background light absorption following the prescriptions of Domínguez et al. (2011).

Overall, we found 101 γ -ray emitting sources out of 146 BAT blazars and their derived γ -ray spectral parameters are provided in Table 3.

3.2. *Swift*-BAT, NuSTAR, and Archival Measurements

We extract 105-month averaged 14–195 keV spectra for all sources using publicly available spectrum and response files and by fitting a simple power-law model in XSPEC (Arnaud 1996), as done by Oh et al. (2018).

The BAT spectra of 25 FSRQs exhibit a steep falling shape (photon index > 2), in contrast to the soft X-ray (< 10 keV) spectral shape, which predicts a hard, rising (photon index < 2) spectrum. This is likely due to the low signal-to-noise spectrum in BAT where Γ is not well constrained ($\Gamma_{\text{err}} > 0.4$). For such sources, we search for better quality *NuSTAR* observations, when available, and find that in all cases, *NuSTAR* spectra

match well with the soft X-ray one. The data from the *NuSTAR* focal plane modules A and B are reduced following the standard procedure. We first run the tool *nupipeline* to clean and calibrate the event files. The source and background regions of $30''$ and $70''$, respectively, are then selected from the same chip, and the *nuproducts* pipeline is adopted to extract the source and background spectra along with ancillary and response matrix files. The spectra are binned to have at least 20 counts per bin and fitting is performed in XSPEC.

Both *Fermi*-LAT and *Swift*-BAT are all-sky surveying instruments and hence considering their long-time averaged spectra is justified to achieve the objectives proposed in this work. On the other hand, soft X-ray and optical-ultraviolet (UV) instruments, e.g., the X-ray telescope (XRT) and ultraviolet and optical telescope (UVOT) on board *Swift*, operate in pointed modes and observations are often triggered during elevated activity states of blazars. With this in mind, we do not separately analyze XRT and UVOT observations and rather rely on publicly available archival SED measurements from the Space Science Data Center that can be considered as representatives of the average behavior of the target of interest. Note that the BASS collaboration is running a dedicated X-ray spectroscopic follow-up of BAT detected objects, irrespective of their activity states, and the results of a comprehensive an X-ray analysis can be found in Ricci et al. (2017).

Table 2

Swift-BAT Detected Sources that Were Not Classified as Blazars in 105-month Catalog but Exhibit Broadband Properties Similar to Beamed AGNs

BAT Name	Counterpart Name	Redshift	Old Classification
J0144.8–2754	PMN J0145–2733	1.155	Unknown AGN
J0131.5–1007	PMN J0131–1009	3.515	Unknown AGN
J0201.0+0329	[HB89] 0158+031	0.765	Unknown AGN
J1810.0–6554	PMN J1809–6556	0.18	Unknown AGN
J1105.4+0200	PMN J1105+0202	0.106	U3
J1254.9+1165	CGRaBS	0.87	U3
	J1254+1141		
J1153.9+5848	RGB J1153+585	0.202	U3
J0909.0+0358	1RXS J090915.6 +035453	3.2	Sy1; broad- line AGN
J0250.8–3626	6dF J0250552 –361636	1.536	Sy1; broad- line AGN
J1310.9–5553	IGR J13109–5552	0.104	Sy1
J1347.1+7325	RGB J1346+733	0.29	Sy1
J0042.9 +3016B	RX J0042.6+3017	0.141	Sy1
J0506.6–1937	6dF J0506479 –193651	0.094	Sy1
J0547.1–6427	CGRaBS	0.323	Sy1
	J0546–6415		
J1132.9 +1019B	[HB89] 1130+106	0.539	Sy1
J1224.1+4073	NVSS J122359 +404359	0.096	Sy1
J1238.4+5349	RGB J1238+534	0.347	Sy1
J1306.4–7603	PMN J1307–7602	0.183	Sy1
J2033.3–2253	PKS 2030–23	0.131	Sy1
J2033.4+2147	4C+21.55	0.173	Sy1

Note. Last column reports the classification presented in Oh et al. (2018). Source type “U3” were those for which X-ray data did not exist at the time of publication and “Sy1” were those classified as a Seyfert 1 galaxy.

Table 3

The γ -ray Spectral Parameters of *Fermi*-LAT Detected BAT Blazars

Name (1)	$F_{0.1-300 \text{ GeV}}$ (2)	$\Gamma_{0.1-300 \text{ GeV}/\alpha}$ (3)	$\beta_{0.1-300 \text{ GeV}}$ (4)	TS (5)
J0036.0+5951	38.20 ± 1.92	1.72 ± 0.02	...	3405
J0122.9+3420	3.38 ± 0.80	1.54 ± 0.11	...	164
J0144.8–2754	7.02 ± 0.29	2.59 ± 0.04	...	912
J0213.7+5147	3.59 ± 0.51	1.91 ± 0.09	...	162
J0218.0+7348	10.70 ± 0.41	2.93 ± 0.05	...	841

Note. The column information are as follows: Col. (1): name of the BAT blazar; Col. (2): 0.1–300 GeV energy flux in units of $10^{-6} \text{ MeV cm}^{-2} \text{ s}^{-1}$; Col. (3): 0.1–300 GeV photon index for the power-law model or photon index at pivot energy for log parabola model; Col. (4): the curvature parameter in log parabola model; and Col. (5): the TS derived from the likelihood fitting.

(This table is available in its entirety in machine-readable form.)

4. Modeling the Broadband Emission

4.1. The Model

We adopt a simple one-zone leptonic emission model (e.g., Dermer et al. 2009; Ghisellini & Tavecchio 2009) to reproduce the multifrequency SEDs of BAT blazars. We consider a conical jet with a semi-opening angle of 0.1 rad. The emission region is assumed to be spherical and cover the whole cross-section of the jet. With this assumption, the dissipation distance (R_d) constrains the size of the emitting region. The observer

receives the Doppler-boosted jet emission at an angle of θ_v from the jet. Furthermore, the radiating population of relativistic electrons has a broken power-law distribution of the following type:

$$S(\gamma) = S_0 \frac{(\gamma_b)^{-p}}{(\gamma/\gamma_b)^p + (\gamma/\gamma_b)^q} \quad (1)$$

where S_0 is the normalization constant (cm^{-3}), γ is the Lorentz factor of electrons, and p, q are the spectral indices below and above the break Lorentz factor γ_b , respectively. Note that our modeling technique does not include a time-dependent evolution of the electron spectrum considering particle injection/cooling.

The leptonic population emits synchrotron radiation in presence of uniform and randomly oriented magnetic field B . We also consider various photon fields, both inside and outside of the jet, to calculate the inverse Compton radiation from the emitting population (e.g., Rybicki & Lightman 1979). This includes thermal photons originated from the accretion disk, BLR, and dusty torus (the external Compton or EC process; e.g., Sikora et al. 1994) and also non-thermal synchrotron photons produced by the same electrons (synchrotron self Compton or SSC; Marscher & Gear 1985). Moreover, we assume the BLR and torus are spherical shells with radii $R_{\text{blr}} = 10^{17} \sqrt{L_{\text{d},45}} \text{ cm}$ and $R_{\text{torus}} = 2.5 \times 10^{18} \sqrt{L_{\text{d},45}} \text{ cm}$, respectively, where $L_{\text{d},45}$ is the luminosity of the accretion disk in units of $10^{45} \text{ erg s}^{-1}$ (Ghisellini & Tavecchio 2009). The adopted emission profiles of these components follow a blackbody distribution peaking at Hydrogen Ly α frequency and the characteristic temperature of the torus (T_{torus}), respectively. We assume a geometrically thin and optically thick accretion disk (Shakura & Sunyaev 1973) whose radiative profile is considered to follow a multi-colored blackbody (see, e.g., Frank et al. 2002)

$$F_{\nu, \text{disk}} = \nu^3 \frac{4\pi h \cos \theta_v}{c^2 D_l^2} \int_{R_{\text{d}, \text{in}}}^{R_{\text{d}, \text{out}}} \frac{R dR}{e^{h\nu/kT(R)} - 1} \quad (2)$$

where h is the Planck constant, D_l is the luminosity distance, k is the Boltzmann constant, c is the speed of light, and $R_{\text{d}, \text{in}}$ and $R_{\text{d}, \text{out}}$ are the inner and outer disk radii, assumed as $3R_{\text{Sch}}$ and $500R_{\text{Sch}}$, respectively. R_{Sch} is the Schwarzschild radius. In the above equation, the disk temperature has the following dependence on the radius

$$T(R) = \frac{3R_{\text{Sch}} L_d}{16\pi\eta_{\text{acc}} \sigma_{\text{SB}} R^3} \left[1 - \left(\frac{3R_{\text{Sch}}}{R} \right)^{1/2} \right]^{1/4}, \quad (3)$$

where σ_{SB} is the Stefan–Boltzmann constant and η_{acc} is the accretion efficiency, adopted here as 10%. The thermal emission originated from the X-ray corona is assumed to follow a power law with exponential cut-off ($L_{\text{cor}}(\nu) \propto \nu^{-1} \exp(-h\nu/150 \text{ keV})$, see Ghisellini & Tavecchio 2009).

For more than >50% of sources, the near-infrared-to-ultraviolet (IR-to-UV) SED exhibits a bump which we interpret as due to accretion disk emission. By modeling this bump with a standard Shakura & Sunyaev (1973) disk, we are able to constrain both the accretion disk luminosity (L_d) and the mass of the central black hole (M_{bh}) with the assumption of the accretion efficiency $\eta_a = 0.1$ (see also, Frank et al. 2002; Ghisellini et al. 2010). In order to get independent measurements, we use results from our own ongoing optical

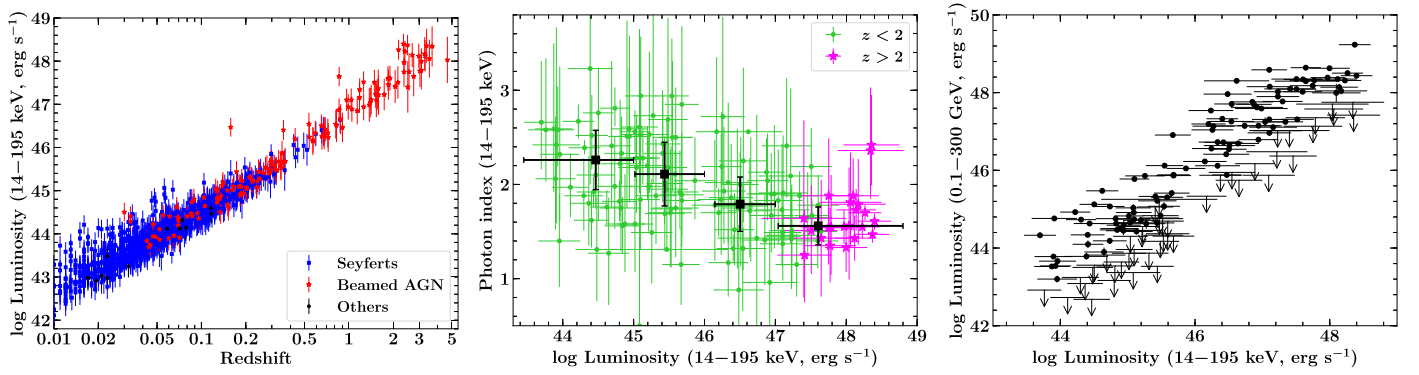


Figure 1. Left: the redshift dependence of 14–195 keV luminosity for various types of BAT detected AGNs. Middle: a plot of the X-ray photon index vs. luminosity. High-redshift blazars are displayed with pink stars. Black squares denote the median values of the plotted quantities in different luminosity bins and the uncertainties are the median absolute deviation. Right: a plot of the hard X-ray vs. γ -ray luminosities. The downward arrows represent γ -ray luminosity upper limits for *Fermi*-LAT undetected sources.

spectroscopic campaign (Koss et al. 2017; J. Mejía-Restrepo et al. 2019, in preparation) and also search for the availability of M_{bh} and L_d derived from single-epoch optical spectroscopy available in the literature (e.g., Shaw et al. 2012). We note that such commonly used single-epoch M_{bh} estimates carry systematic uncertainties of order 0.3–0.5 dex (see, e.g., Shen 2013; Peterson 2014, for reviews of these methods), and that the usage of the broad C IV $\lambda 1549$ emission line may carry additional uncertainties related to the virialized nature of the BLR (e.g., Trakhtenbrot & Netzer 2012; Mejía-Restrepo et al. 2018, and references therein). Furthermore, if optical spectral parameters (e.g., line and continuum luminosities) are reported (see Torrealba et al. 2012), we use them to constrain M_{bh} following empirical relations commonly used for blazars (e.g., Shen et al. 2011; Shaw et al. 2012). We compute the luminosity of the BLR from the emission line luminosities by considering the quasar emission line weights provided by Francis et al. (1991). Assuming that the BLR reprocesses 10% of the disk radiation, we determine L_d .

4.2. Modeling Constraints

The SED model used in this work does not perform any statistical fitting, and hence the uniqueness of the derived parameters cannot be claimed. However, our modeling efforts are driven by the quality of the multiwavelength data that provide fairly good constraints to the parameters. For example, we determine the M_{bh} and L_d either from the disk modeling or from the optical spectroscopy, and these two parameters regulate the size and radiative energy densities of various external photon fields. The shape of the X-ray spectrum constrains the relative dominance of the EC and SSC mechanisms since different parts of the electron energy distribution are involved. The SSC radiation regulates the size of the emission region and also the magnetic field. A Compton-dominated SED reflects the prevalence of the external photon densities over the magnetic one, and since in our model these are a function of the distance of the emission region from the central black hole, we are able to constrain R_d . In Compton-dominated blazars, the low- and high-energy indices of the broken power-law electron energy distribution are determined from the spectral shapes of the X-ray and γ -ray SEDs, respectively. The optical–UV spectrum further controls the high-energy index, provided it is dominated by the falling synchrotron emission. On the other hand, in blazars whose

synchrotron peak lies at UV–X-ray energies, i.e., HSP sources, we use *Fermi*-LAT spectra to determine the low-energy index. The shape of the BAT spectrum, along with the soft X-ray SED, provides a good constraint to the high-energy index in such objects.

5. Physical Characteristics

In the left panel of Figure 1, we show the 14–195 keV rest-frame luminosity of various types of BAT detected AGNs as a function of their redshifts. As can be seen, blazars dominate the 105-month BAT catalog above $z \approx 0.5$. This is likely a selection effect due to relativistic beaming, which makes blazar jets brighter compared to other classes of astrophysical objects. Considering only blazars, the photon index versus the luminosity distribution is shown in the middle panel, and we highlight $z > 2$ blazars by showing them with pink stars. As can be seen, these are the most powerful ones with luminosities exceeding 10^{47} erg s $^{-1}$. Furthermore, there is a hint of an anti-correlation (the Spearman coefficient of $\rho = -0.29 \pm 0.08$, probability of no correlation or PNC = 0.01) where more luminous objects, mostly FSRQs, tend to have harder BAT spectra. This is because the X-ray to γ -ray emission in FSRQs primarily originates via an EC process that has a hard spectral shape in X-rays (see Figure 2), which becomes more prominent at high redshifts due to the shift of SED peaks to lower frequencies and K -correction. Less luminous HSP blazars generally have a falling synchrotron emission in the BAT energy range and thus exhibit a steep spectrum with a 14–195 keV photon index > 2 . We plot the γ -ray luminosity as a function of the hard X-ray one in the right panel of Figure 1. Interestingly, the *Fermi*-LAT undetected objects share the same range of the 14–195 keV luminosity with γ -ray detected ones. Since all of them are FSRQs, this observation indicates their SED peaks to lie at lower frequencies with respect to *Fermi*-LAT sources.

The modeled SEDs of BAT blazars are shown in Figure 2 and we provide associated parameters in Table 4. Note that there are 28 sources whose broadband SEDs are well represented with synchrotron SSC processes, i.e., without invoking an EC mechanism. All of them are HSP blazars and their names and redshifts are provided in Table 6.

Particle energy distribution. The histograms of the spectral indices of the electron energy distribution are shown in Figure 3 (panels (a) and (b)). The low-energy index (p) distribution peaks at 1.90 and when fitted with a Gaussian

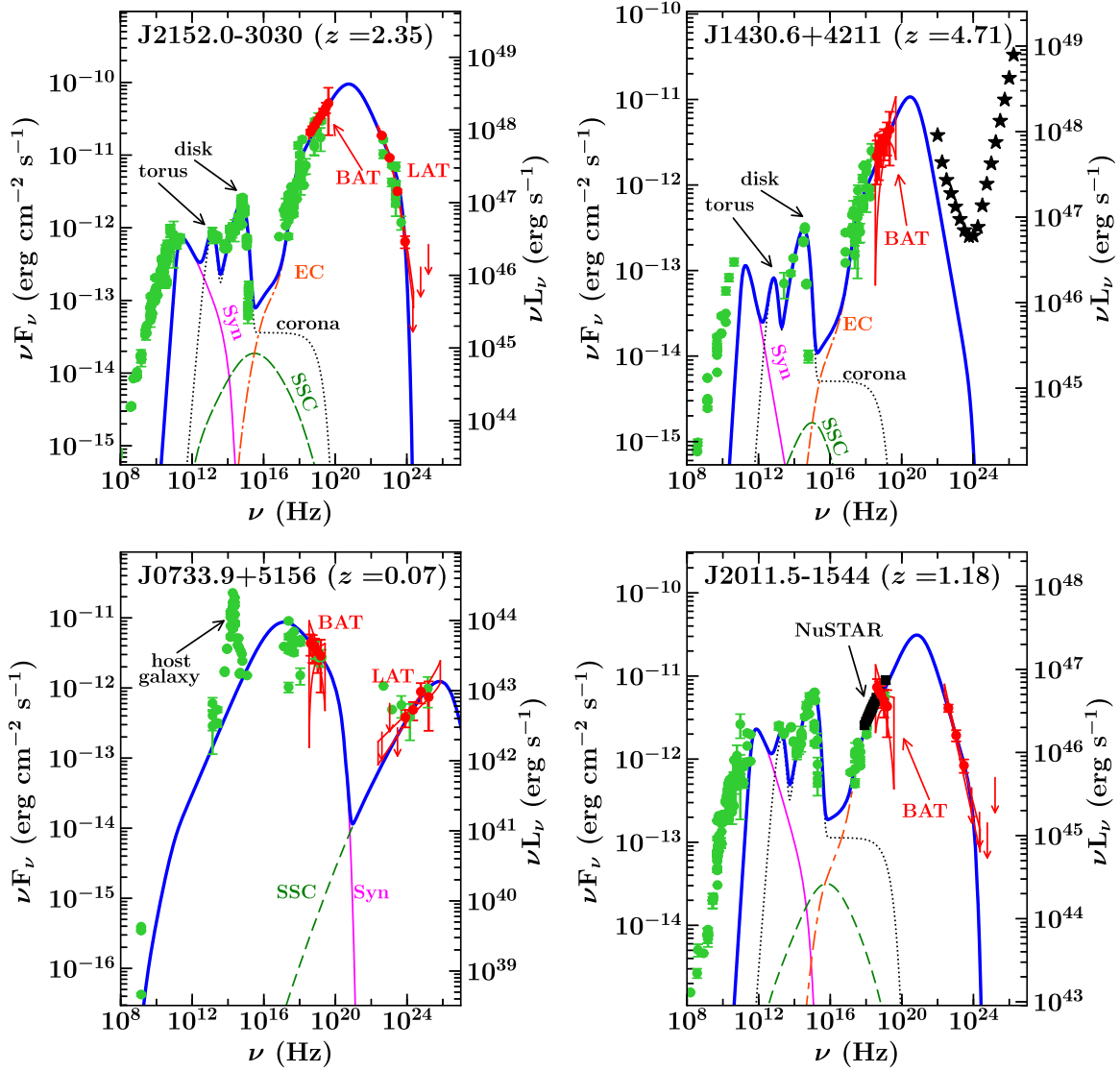


Figure 2. Top: the modeled SED of the *Fermi*-LAT detected (left) and undetected (right) BAT blazars. *Swift*-BAT and *Fermi*-LAT spectral data are shown with red filled circles and bow ties, while archival observations are displayed with lime-green filled circles. Thermal emissions from the torus, accretion disk, and the X-ray corona are represented with the black dotted line. We show the nonthermal synchrotron, SSC, and EC radiations with the pink thin solid, green long dashed and orange dashed-dashed-dotted lines, respectively. The sum of all of the radiations is demonstrated with the blue thick solid line. For the γ -ray undetected sources, we plot 5σ *Fermi*-LAT sensitivity for the duration covered in this work and toward the direction of the source with black stars. Bottom: the left panel illustrates the multiwavelength SED of a BL Lac object. Note the IR-to-UV bump that is likely due to host galaxy emission. For sources whose BAT spectra do not match with the rest of their SED (right panel), we search for the availability of *NuSTAR* data and use it for the modeling. The *NuSTAR* data are shown with black squares. The modeled SED plots for all of the blazars are shown in the figure set.

(The complete figure set (146 images) is available.)

function, the dispersion is $\sigma = 0.20$. On the other hand, the high-energy index (q) has a rather broad range. The distribution peaks at $\langle q \rangle = 4.30$ with $\sigma = 0.49$ and its tail extends up to $q \sim 6$. This is primarily due to the inclusion of γ -ray undetected blazars whose falling part of the inverse Compton spectrum needs to be steep to remain below the *Fermi*-LAT detection threshold. Additionally, the steep index can also be due to a few HSP blazars with falling BAT spectra. The distribution of the break energy, γ_b , which is derived from the synchrotron peak frequency location, peaks at 100 and tails to very large values, $\gtrsim 1000$. The very large γ_b is probably due to inefficient cooling of relativistic electrons in HSP blazars, and overall the distribution has a broad width ($\sigma = 0.55$, Figure 3, panel (c)).

Magnetic field. The magnetic field and Compton dominance²⁰ distributions are shown in panels (d) and (e) of Figure 3. There is a hint of bi-modality that can be understood keeping in mind that a major fraction of BAT blazars studied here are FSRQs and the remaining are mostly HSP blazars with synchrotron dominated SEDs. FSRQs are known to have Compton-dominated ($CD > 1$) SEDs with a larger magnetic field compared to HSP blazars (see, e.g., Tavecchio et al. 2010; Paliya et al. 2017). Therefore, the distributions are skewed toward larger CD and B with $\langle \log CD \rangle = 0.79$ and $\langle B \rangle = 1.10$ Gauss, respectively.

²⁰ Compton dominance, or CD, is defined as the ratio of the inverse Compton to synchrotron peak luminosities.

Table 4
The SED Parameters Associated with the Modeling of the Broadband Emission of BAT Blazars

Name (1)	z (2)	θ_v (3)	M_{bh} (4)	Type (5)	L_d (6)	R_d (7)	R_{blr} (8)	δ (9)	Γ (10)	B (11)	p (12)	q (13)	γ_{min} (14)	γ_b (15)	γ_{max} (16)	U_e (17)	CD (18)
J0010.5+1057	0.09	6.0	8.70	D	45.15	0.215	0.038	9.5	11	0.2	2.2	5.1	4	286	3.0e+03	-2.16	6.2
J0017.1+8134	3.37	4.0	10.04	D	48.00	0.421	1.020	12.2	8	2.2	1.9	4.5	1	41	2.0e+03	-1.68	16.2
J0036.0+5951	0.08	3.0	0.00	N	0.00	0.115	0.000	13.6	8	0.1	2.1	3.7	50	251243	1.5e+06	-2.71	0.4
J0042.9+3016B	0.14	3.0	8.90	A	45.00	0.031	0.032	16.5	11	1.2	1.5	5.3	1	20	1.3e+03	-1.69	247.2
J0057.0+6405	0.29	3.0	8.70	A	44.00	0.120	0.010	18.6	15	1.1	2.1	4.6	5	300	3.0e+03	-2.25	0.6

Note. The column contents are as follows: Col. (1) and (2): name and redshift of the source. Col. (3): viewing angle, in degrees. Col. (4): log scale black hole mass, in units of $M_\odot$. Col. (5): method used to derive M_{bh} and L_d , A: assumed, D: disk fitting, and O: optical spectroscopy. The used suffixes are as follows: B: BASS spectroscopy (Koss et al. 2017; J. Mejia-Restrepo et al. 2019, in preparation), G11: Ghisellini et al. (2011), S12: Shaw et al. (2012), S13: Shaw et al. (2013), WU02: Woo & Urry (2002), T12: Torrealba et al. (2012), W04: Wang et al. (2004), and N: not used. Col. (6): log scale accretion disk luminosity, in erg s^{-1} . Col. (7) and (8): dissipation distance and the size of the BLR, respectively, in parsec. Col. (9) and (10): the Doppler factor and the bulk Lorentz factor, respectively. Col. (11): magnetic field, in Gauss. Col. (12) and (13): spectral indices of the broken power-law electron distribution before and after the break energy (γ_b), respectively. Col. (14)–(16): minimum, break, and maximum Lorentz factors of the emitting electron distribution. Col. (17): log scale particle energy density, in erg cm^{-3} . Col. (18): Compton dominance. Note that M_{bh} and L_d values are quoted as zeros for those blazars whose SEDs are modeled with synchrotron SSC processes only, i.e., without invoking an EC mechanism.

(This table is available in its entirety in machine-readable form.)

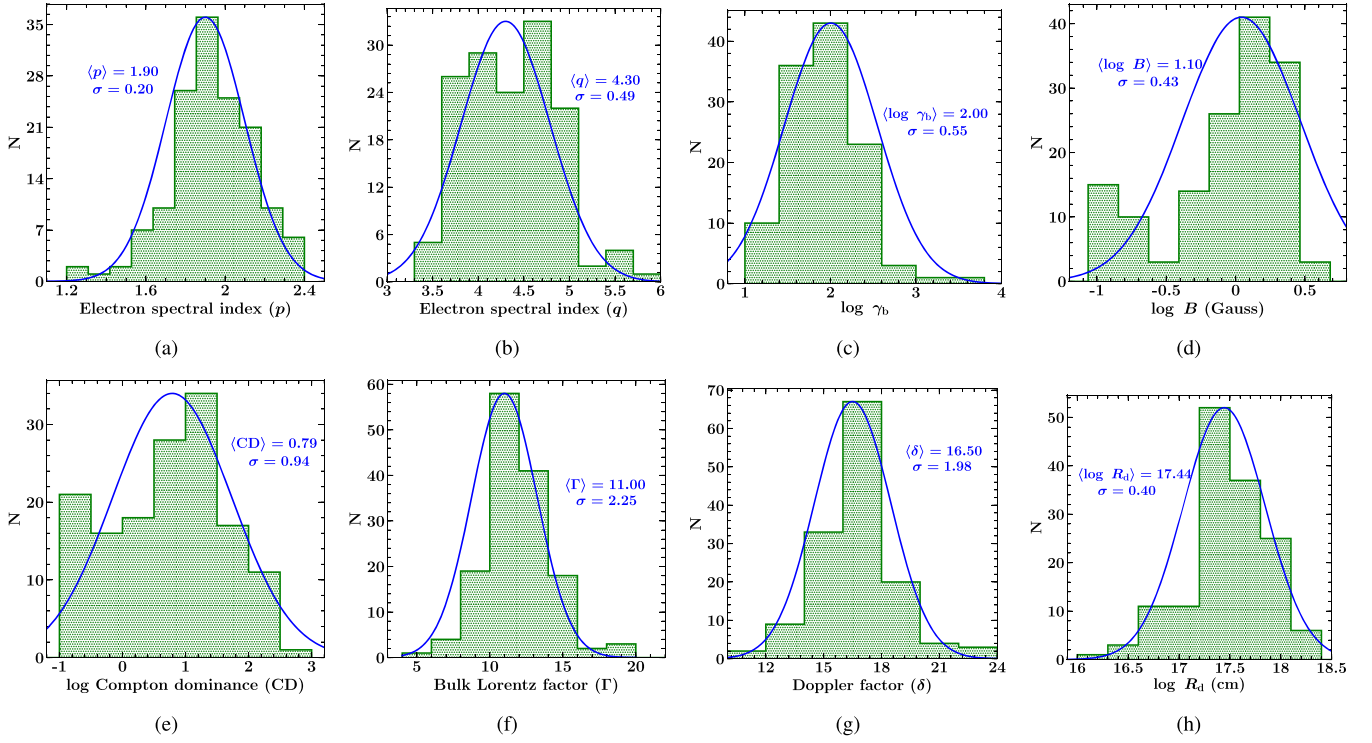


Figure 3. Histograms of modeled SED parameters for the 142 blazars in our sample. We fit the data with a Gaussian distribution, shown as a blue line, and the standard deviations are quoted.

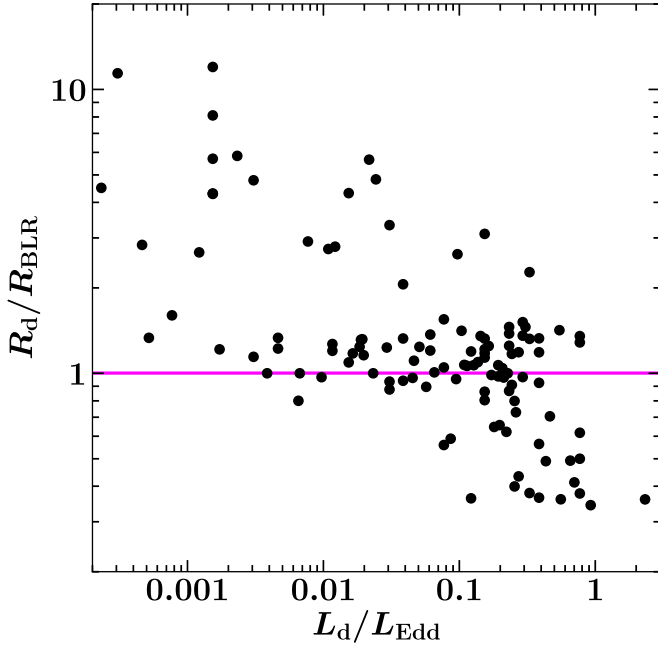


Figure 4. R_d as a function of the accretion disk luminosity, L_d . The horizontal line represents the inner boundary of the BLR.

Bulk Lorentz factor and Doppler factor. In panels (f) and (g) of Figure 3, we show the histograms of the bulk Lorentz factor (Γ) and Doppler factor (δ). Both of them have narrow distributions peaking at $\langle \Gamma \rangle = 11$ and $\langle \delta \rangle = 16.5$. These values are in agreement with that found from radio studies (e.g., Savolainen et al. 2010), from the SED modeling of a large sample of γ -ray emitting blazars (Paliya et al. 2017), and also from LF studies (Ajello et al. 2012).

Dissipation distance. The distribution of the location of the emission region peaks at $R_d \sim 3 \times 10^{17}$ cm with a majority of sources having $R_d > 10^{17}$ cm (see Figure 3, panel (h)). To understand the emission region location with respect to BLR, we plot R_d/R_{BLR} as a function of L_d in Eddington units in Figure 4. The R_{BLR} is estimated following Ghisellini & Tavecchio (2009), i.e., $R_{\text{BLR}} = 10^{17} (L_d/10^{45} \text{ erg s}^{-1})^{1/2}$ cm. The location is outside BLR, i.e., above the horizontal line, mostly for blazars with less luminous accretion disks, which is typically the case of BL Lacs. Moreover, a majority of sources have R_d comparable to or slightly larger than R_{BLR} , thus indicating the jet environment surrounding the emission region to be transparent enough for γ -rays to avoid absorption via pair production with BLR photons (see also, Böttcher & Els 2016; Paliya et al. 2018 for quantitative discussions). Blazars with most luminous accretion disks, i.e., with the largest BLRs, tend to have emission region well within it. These are the objects with the largest Compton dominance.

Central engine. We have derived M_{bh} in 92 objects using the disk modeling approach, which shows the signature of the accretion disk emission at optical–UV energies. In 20 sources, we use the optical spectroscopic M_{bh} since their optical–UV SED is synchrotron dominated. The multiwavelength SEDs of 28 HSP blazars (Table 6) are well reproduced with a synchrotron SSC model and thus no M_{bh} value was estimated/used. Remaining six sources neither exhibit the big blue bump nor have the spectroscopic M_{bh} . In these objects, we suitably assume a M_{bh} value based on the available observations. The different methods adopted to compute M_{bh} are tabulated in Table 4.

The left and middle panels of Figure 5 represent the M_{bh} and L_d histograms for BAT blazars. The M_{bh} distribution has a rather narrow range ($\sim 10^8$ – $10^{10} M_\odot$) and peaks at $\langle \log M_{\text{bh}} \rangle = 9.0 M_\odot$. Fitting it with a log-normal function

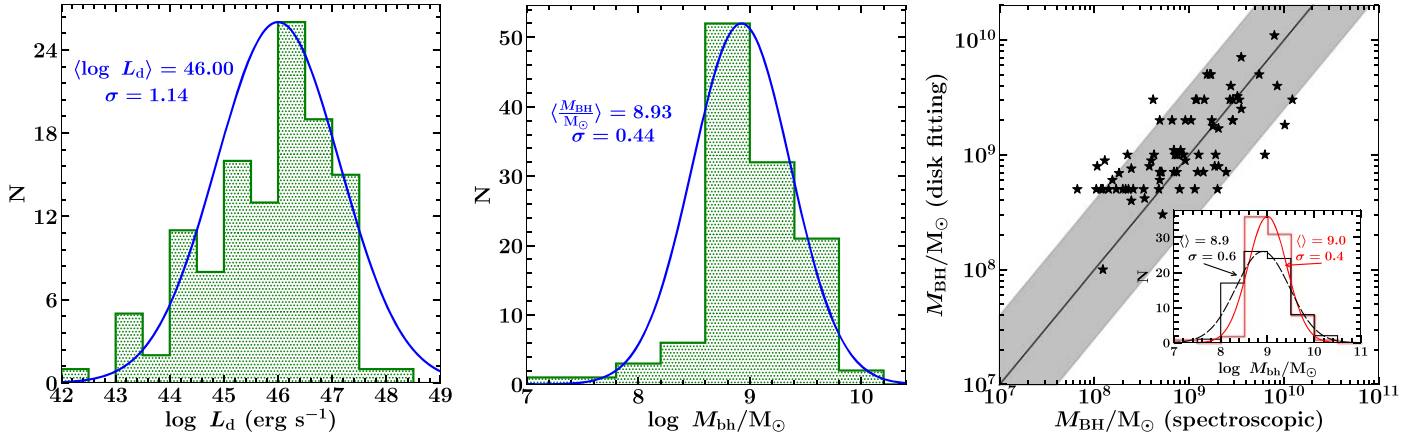


Figure 5. Distributions of L_d (left) and M_{bh} (middle) for BAT blazars studied in this work. In the right panel, we compare M_{bh} values derived from the optical spectroscopy and disk modeling methods. As can be seen, both approaches reasonably agree with each other. The inset shows the histograms of the plotted quantities. The shaded area represent the uncertainty of a factor of 4 associated with the virial estimation (e.g., Shen et al. 2011).

Table 5
Various Jet Powers Derived from the SED Modeling

Name (1)	P_e (2)	P_m (3)	P_r (4)	P_p (5)	P_j (6)
J0010.5+1057	44.84	44.40	45.23	46.97	46.97
J0017.1+8134	45.62	46.59	47.35	48.13	48.14
J0036.0+5951	43.46	42.68	44.26	44.23	44.31
J0042.9+3016B	43.61	44.06	44.25	46.20	46.21
J0057.0+6405	44.51	45.44	45.01	46.51	46.55

Note. The column contents are as follows: Col. (1) name the source; Col. (2)–(6): the electron, magnetic, radiative, proton, and total jet powers, respectively. Note that $P_j = P_e + P_m + P_p$. All jet powers are evaluated for a two-sided jet. (This table is available in its entirety in machine-readable form.)

returns a width of $\sigma = 0.44$. On the other hand, L_d peaks at 10^{46} erg s $^{-1}$ and has a rather broad range with $\sigma = 1.14$.

In our sample, there are a total of 82 blazars that have M_{bh} measurements from both the disk-fitting and optical spectroscopy methods. In the right panel of Figure 5, we compare M_{bh} derived using these two approaches. As can be seen, both methods provide reasonably similar M_{bh} for BAT blazars and thus confirm the findings earlier reported for *Fermi*-LAT detected sources (Paliya et al. 2017).

6. Comparison with *Fermi* Blazars

It may be of a great interest to compare the accretion and jet powers of BAT blazars with that detected by *Fermi*-LAT sources. The powers that the jet carries in the form of the bulk motion of electrons (P_e), “cold” protons (P_p , assuming one proton per electron, i.e., no pairs), radiation (P_r), and magnetic field (P_m) are calculated following Celotti & Ghisellini (2008). The derived jet powers, assuming a two-sided jet, are reported in Table 5. For a comparison, we use the results reported for a large sample of *Fermi* blazars in Ghisellini et al. (2014) and Paliya et al. (2017).

P_r versus L_d . The most robust estimate of the jet power that we can have is the one that the jet produces in the form of radiation. This is because P_r is directly proportional to the observed bolometric luminosity. In the left panel of Figure 6, we show the behavior of P_r as a function of L_d . The low P_r – L_d

Table 6
The List of 28 Blazars that Are Modeled with Synchrotron and SSC Processes, i.e., without Invoking an EC Mechanism

Name	z
J0036.0+5951	0.08
J0122.9+3420	0.27
J0213.7+5147	0.05
J0232.8+2020	0.14
J0244.8–5829	0.26
J0326.0–5633	0.06
J0349.2–1159	0.18
J0353.4–6830	0.09
J0507.7+6732	0.31
J0550.7–3212A	0.07
J0710.3+5908	0.12
J0721.0+7133	0.30
J0733.9+5156	0.07
J0930.1+4987	0.19
J0934.0–1721	0.25
J1031.5+5051	0.36
J1103.5–2329	0.19
J1104.4+3812	0.03
J1136.7+6738	0.13
J1221.3+3012	0.18
J1417.7+2539	0.24
J1428.7+4234	0.13
J1943.5+2120	0.21
J1959.6+6507	0.05
J2009.6–4851	0.07
J2246.7–5208	0.19
J2251.8–3210	0.25
J2359.0–3038	0.17

end is mostly occupied by low-power HSP sources, whereas, luminous BAT FSRQs dominate at the high end. We compute the partial Spearman’s coefficient, ρ (Padovani 1992), and PNC to quantify the strength of the correlation and derive the parameters independent of the common redshift effect. The derived coefficients are $\rho = 0.40 \pm 0.08$ with PNC $< 10^{-10}$ and $\rho = 0.33 \pm 0.18$ with PNC $< 10^{-10}$ for *Fermi*-LAT and *Swift*-BAT blazar populations, respectively. A positive P_r – L_d correlation has been reported previously for *Fermi*-LAT detected sources (e.g., Ghisellini et al. 2014; Paliya et al. 2017), and our

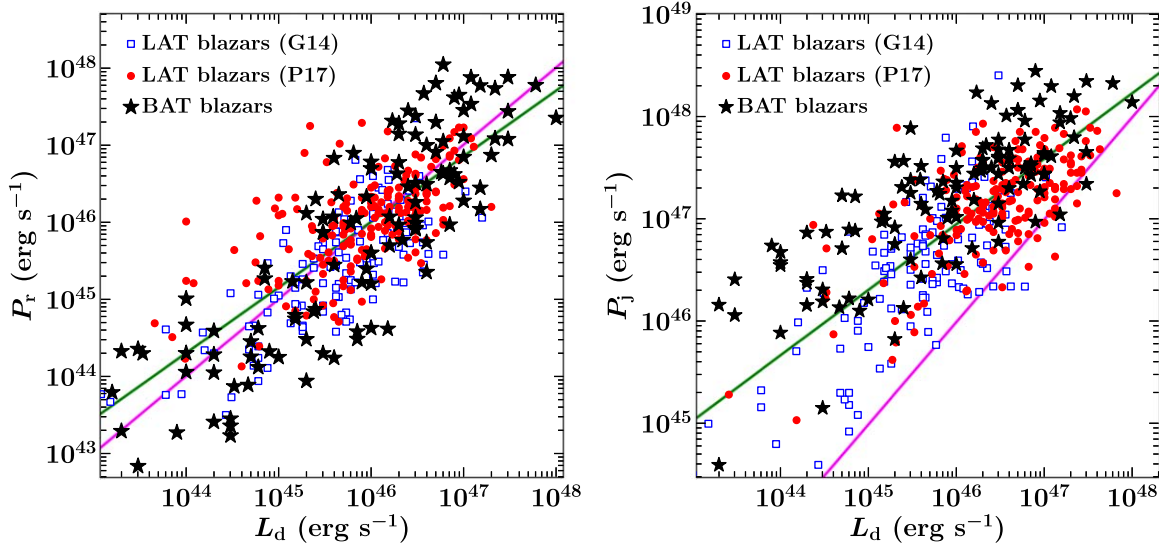


Figure 6. Left: the power that jet produces in the form of radiation (P_r) as a function of the luminosity of the accretion disk. Right: the total jet power ($P_j = P_p + P_e + P_m$) as a function of the L_d . In both the plots, *Fermi*-LAT blazars are displayed with red circles (Paliya et al. 2017, P17) and empty blue squares (Ghisellini et al. 2014, G14), whereas, BAT blazars are shown with black stars. The pink and green solid lines refer to the equality and best-fit of the plotted quantities, respectively.

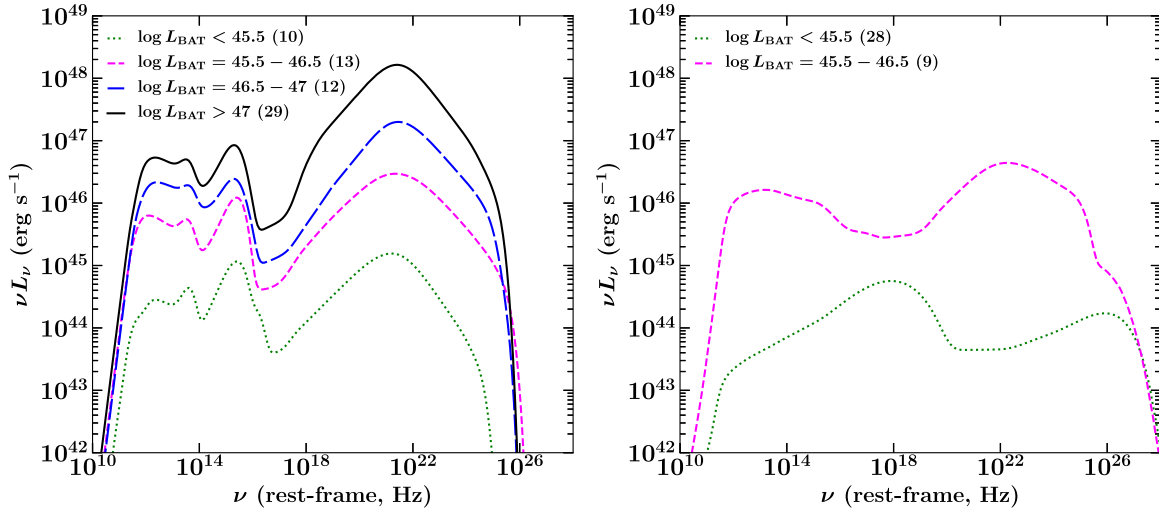


Figure 7. Left: averaged SEDs of BAT FSRQs in four different L_{BAT} bins, as labeled. Right: same as left but for BL Lacs. Note that there are no BL Lac objects detected with *Swift*-BAT with $L_{\text{BAT}} > 10^{47} \text{ erg s}^{-1}$ and there is only one (J0428.6–3788 or PKS 0426–380, $z = 1.11$) in the $\log L_{\text{BAT}} = 46.5\text{--}47$ bin, which is hence not shown. The number shown in brackets refers to the number of sources in the respective luminosity bin. See the text for details.

findings establishes the fact that *Swift*-BAT blazars also follow the same positive trend.

P_j versus L_d . In the middle panel of Figure 6, we show the total jet power ($P_j = P_p + P_e + P_m$) versus L_d . A positive correlation is observed, which is statistically confirmed with $\rho = 0.53 \pm 0.08$ with $\text{PNC} < 10^{-10}$ and $\rho = 0.69 \pm 0.05$ with $\text{PNC} < 10^{-10}$ for *Fermi*-LAT and *Swift*-BAT blazars, respectively. Similar to LAT blazars, BAT detected ones to host jets with $P_j > L_d$. Interestingly, at the high end of the jet power ($P_j > 10^{48} \text{ erg s}^{-1}$), BAT blazars seems to have more extreme jets compared to LAT blazars. This is probably a selection effect due to higher detection threshold of *Swift*-BAT, which leads to the identification of the most luminous sources. In a more physical scenario, the observational shift of the SED peaks to lower frequencies with an increase in the bolometric luminosity (e.g., Fossati et al. 1998) can also explain this

observation. BAT blazars that likely have inverse Compton peaks located at lower energies with respect to LAT detected ones, host more powerful jets, and luminous accretion disks.

7. Luminosity-dependent Evolution

It is instructive to determine the average SEDs of BAT blazars with the motivation to explore their luminosity-dependent evolution (Ajello et al. 2012). We divide the sources in different BAT luminosity (L_{BAT}) bins, and also based on their SED appearances, to identify FSRQs/BL Lacs. The latter is necessary since the 105-month BAT catalog does not provide FSRQ/BL Lac classification. Moreover, we consider only those sources that have both *Fermi*-LAT and *Swift*-BAT detections. We take the arithmetic average of the broadband SEDs in each L_{BAT} bin and show in Figure 7. Note that the adopted choice of L_{BAT} bins are mainly driven to distinguish

BAT blazars of various powers and does not carry any other physical meaning.

In the left panel, we show averaged SEDs for FSRQs in four L_{BAT} bins and note the increasing dominance of the inverse Compton peak (i.e., more Compton-dominated SED) with the increasing L_{BAT} and L_d . Also, there is an indication for the shift of the inverse Compton peak to lower frequencies as their power increases, especially in the highest L_{BAT} bin. In the right-hand panel of Figure 7, we plot the SEDs for BL Lacs. There are no BL Lacs in our sample with $L_{\text{BAT}} > 10^{47}$ erg s⁻¹ and only one (J0428.6–3788 or PKS 0426–380, $z = 1.11$) has log L_{BAT} in the bin 46.5–47, which is hence not plotted. The lowest L_{BAT} bin is dominated by HSP BL Lacs, and a pronounced shift of the SED peaks to lower frequencies can be observed in the higher luminosity bin. These results are aligned with our current understanding about the blazar sequence where the most powerful objects are LSP FSRQs with Compton-dominated SEDs, and the luminosity-dependent SED evolution is more pronounced in BL Lacs (see Ghisellini et al. 2017, for latest results).

8. Summary

In this work, we study 146 blazars detected in the first 105-month all-sky survey of *Swift*-BAT. We analyze the *Fermi*-LAT data and 14–195 keV BAT spectra and supplement them with multiwavelength archival spectral measurements. We determine the physical properties of the jet and central engine by modeling the broadband SED with a simple leptonic emission model. Our main findings are summarized below.

1. We nearly quadruple the number of *Swift*-BAT detected blazars and also the number of *Fermi*-LAT undetected ones compared to earlier studies (146 versus 38 and 45 versus 12, respectively; Ajello et al. 2009).
2. In the photon index versus 14–195 keV luminosity plane, we find tentative evidences ($\rho = -0.29 \pm 0.08$, PNC = 0.01) of inverse correlation with more luminous blazars that tend to have flatter spectra.
3. The overall physical properties of BAT blazars resemble well with that known for the LAT blazar population studied in Paliya et al. (2017).
4. The mass of the central black hole derived from the single-epoch optical spectroscopy matches reasonably well with that derived by modeling the optical–UV bump with a standard accretion disk.
5. BAT blazars exhibit a positive correlation between the power that the jet produces in the form of radiation and L_d . Blazars with most luminous accretion disks host the most radiatively powerful jets.
6. Comparing the averaged SEDs of BAT blazars binned in different 14–195 keV luminosity bins, we find that more luminous sources have more Compton-dominated SEDs. There is also a hint of the shift of SED peaks to lower frequencies as the power of the blazars increases and it is found to be more pronounced in low-luminosity BL Lac objects.

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Software: fermiPy (Wood et al. 2017).

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References

- Abdo, A. A., Ackermann, M., Agudo, I., et al. 2010, *ApJ*, 716, 30
Acero, F., Ackermann, M., Ajello, M., et al. 2015, *ApJS*, 218, 23
Acero, F., Ackermann, M., Ajello, M., et al. 2016, *ApJS*, 223, 26
Ackermann, M., Ajello, M., Baldini, L., et al. 2017, *ApJL*, 837, L5
Ajello, M., Costamante, L., Sambruna, R. M., et al. 2009, *ApJ*, 699, 603
Ajello, M., Romani, R. W., Gasparini, D., et al. 2014, *ApJ*, 780, 73
Ajello, M., Shaw, M. S., Romani, R. W., et al. 2012, *ApJ*, 751, 108

- Arnaud, K. A. 1996, in ASP Conf. Ser. 101, *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby & J. Barnes (San Francisco, CA: ASP), 17
- Atwood, W., Albert, A., Baldini, L., et al. 2013, arXiv:1303.3514
- Atwood, W. B., Abdo, A. A., Ackermann, M., et al. 2009, *ApJ*, 697, 1071
- Barthelmy, S. D., Barbier, L. M., Cummings, J. R., et al. 2005, *SSRv*, 120, 143
- Blandford, R. D., & Rees, M. J. 1978, *PhyS*, 17, 265
- Böttcher, M., & Els, P. 2016, *ApJ*, 821, 102
- Böttcher, M., Reimer, A., Sweeney, K., & Prakash, A. 2013, *ApJ*, 768, 54
- Bruel, P., Burnett, T. H., Digel, S. W., et al. 2018, arXiv:1810.11394
- Cazzoli, S., Márquez, I., Masegosa, J., et al. 2018, *MNRAS*, 480, 1106
- Celotti, A., & Ghisellini, G. 2008, *MNRAS*, 385, 283
- Dermer, C. D., Finke, J. D., Krug, H., & Böttcher, M. 2009, *ApJ*, 692, 32
- Domínguez, A., Primack, J. R., Rosario, D. J., et al. 2011, *MNRAS*, 410, 2556
- Fossati, G., Maraschi, L., Celotti, A., Comastri, A., & Ghisellini, G. 1998, *MNRAS*, 299, 433
- Francis, P. J., Hewett, P. C., Foltz, C. B., et al. 1991, *ApJ*, 373, 465
- Frank, J., King, A., & Raine, D. J. (ed.) 2002, *Accretion Power in Astrophysics* (Cambridge: Cambridge Univ. Press), 398
- Gehrels, N., Chincarini, G., Giommi, P., et al. 2004, *ApJ*, 611, 1005
- Ghisellini, G., Della Ceca, R., Volonteri, M., et al. 2010, *MNRAS*, 405, 387
- Ghisellini, G., Righi, C., Costamante, L., & Tavecchio, F. 2017, *MNRAS*, 469, 255
- Ghisellini, G., & Tavecchio, F. 2009, *MNRAS*, 397, 985
- Ghisellini, G., Tavecchio, F., Foschini, L., & Ghirlanda, G. 2011, *MNRAS*, 414, 2674
- Ghisellini, G., Tavecchio, F., Maraschi, L., Celotti, A., & Sbarbato, T. 2014, *Natur*, 515, 376
- Gregory, P. C., Scott, W. K., Douglas, K., & Condon, J. J. 1996, *ApJS*, 103, 427
- Griffith, M. R., & Wright, A. E. 1993, *AJ*, 105, 1666
- Guillard, P., Ogle, P. M., Emonts, B. H. C., et al. 2012, *ApJ*, 747, 95
- Halpern, J. P., & Filippenko, A. V. 1984, *ApJ*, 285, 475
- Healey, S. E., Romani, R. W., Taylor, G. B., et al. 2007, *ApJS*, 171, 61
- Hernández-García, L., Masegosa, J., González-Martín, O., & Márquez, I. 2015, *A&A*, 579, A90
- Kormendy, J., & Ho, L. C. 2013, *ARA&A*, 51, 511
- Koss, M., Trakhtenbrot, B., Ricci, C., et al. 2017, *ApJ*, 850, 74
- Lamperti, I., Koss, M., Trakhtenbrot, B., et al. 2017, *MNRAS*, 467, 540
- Marscher, A. P., & Gear, W. K. 1985, *ApJ*, 298, 114
- Massaro, E., Maselli, A., Leto, C., et al. 2015, *Ap&SS*, 357, 75
- Mattox, J. R., Bertsch, D. L., Chiang, J., et al. 1996, *ApJ*, 461, 396
- Mejía-Restrepo, J. E., Trakhtenbrot, B., Lira, P., & Netzer, H. 2018, *MNRAS*, 478, 1929
- Oh, K., Koss, M., Markwardt, C. B., et al. 2018, *ApJS*, 235, 4
- Padovani, P. 1992, *A&A*, 256, 399
- Paliya, V. S., Marcotulli, L., Ajello, M., et al. 2017, *ApJ*, 851, 33
- Paliya, V. S., Zhang, H., Böttcher, M., et al. 2018, *ApJ*, 863, 98
- Peterson, B. M. 2014, *SSRv*, 183, 253
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, 594, A13
- Popović, L. Č., Shapovalova, A. I., Ilić, D., et al. 2014, *A&A*, 572, A66
- Ricci, C., Trakhtenbrot, B., Koss, M. J., et al. 2017, *ApJS*, 233, 17
- Robinson, A., Binette, L., Fosbury, R. A. E., & Tadhunter, C. N. 1987, *MNRAS*, 227, 97
- Romani, R. W., Sowards-Emmerd, D., Greenhill, L., & Michelson, P. 2004, *ApJL*, 610, L9
- Rybicki, G. B., & Lightman, A. P. 1979, *Radiative Processes in Astrophysics* (New York: Wiley-Interscience), 393
- Saripalli, L., Hunstead, R. W., Subrahmanyam, R., & Boyce, E. 2005, *AJ*, 130, 896
- Savolainen, T., Homan, D. C., Hovatta, T., et al. 2010, *A&A*, 512, A24
- Shakura, N. I., & Sunyaev, R. A. 1973, *A&A*, 24, 337
- Shaw, M. S., Romani, R. W., Cotter, G., et al. 2012, *ApJ*, 748, 49
- Shaw, M. S., Romani, R. W., Cotter, G., et al. 2013, *ApJ*, 764, 135
- Shen, Y. 2013, *BASI*, 41, 61
- Shen, Y., Richards, G. T., Strauss, M. A., et al. 2011, *ApJS*, 194, 45
- Sikora, M., Begelman, M. C., & Rees, M. J. 1994, *ApJ*, 421, 153
- Stickel, M., Padovani, P., Urry, C. M., Fried, J. W., & Kuehr, H. 1991, *ApJ*, 374, 431
- Tavecchio, F., Ghisellini, G., Ghirlanda, G., Foschini, L., & Maraschi, L. 2010, *MNRAS*, 401, 1570
- Tombesi, F., Mushotzky, R. F., Reynolds, C. S., et al. 2017, *ApJ*, 838, 16
- Torrealba, J., Chavushyan, V., Cruz-González, I., et al. 2012, *RMxAA*, 48, 9
- Trakhtenbrot, B., & Netzer, H. 2012, *MNRAS*, 427, 3081
- Véron-Cetty, M.-P., & Véron, P. 2010, *A&A*, 518, A10
- Volonteri, M., Haardt, F., Ghisellini, G., & Della Ceca, R. 2011, *MNRAS*, 416, 216
- Wang, J.-M., Luo, B., & Ho, L. C. 2004, *ApJL*, 615, L9
- Woo, J.-H., & Urry, C. M. 2002, *ApJ*, 579, 530
- Wood, M., Caputo, R., Charles, E., et al. 2017, arXiv:1707.09551