

B-modes and the nature of inflation

This content has been downloaded from IOPscience. Please scroll down to see the full text.

JCAP01(2015)016

(<http://iopscience.iop.org/1475-7516/2015/01/016>)

View [the table of contents for this issue](#), or go to the [journal homepage](#) for more

Download details:

IP Address: 131.169.95.181

This content was downloaded on 18/01/2016 at 09:30

Please note that [terms and conditions apply](#).

B-modes and the nature of inflation

Daniel Baumann,^a Daniel Green^{b,c,d} and Rafael A. Porto^{e,f}

^aD.A.M.T.P., Cambridge University,
Wilberforce Road, Cambridge, CB3 0WA, U.K.

^bCanadian Institute for Theoretical Astrophysics,
60 St. George Street, Toronto, ON M5S 3H8, Canada

^cStanford Institute for Theoretical Physics, Stanford University,
382 Via Pueblo Mall, Stanford, CA 94305, U.S.A.

^dKavli Institute for Particle Astrophysics and Cosmology,
382 Via Pueblo Mall, Stanford, CA 94305, U.S.A.

^eDeutsches Elektronen-Synchrotron DESY, Theory Group,
Notkestrasse 85, D-22603 Hamburg, Germany

^fInstitute for Advanced Study,
Einstein Drive, Princeton, NJ 08540, U.S.A.

E-mail: dbaumann@damtp.cam.ac.uk, drgreen@cita.utoronto.ca, rporto@ias.edu

Received September 15, 2014

Accepted November 26, 2014

Published January 12, 2015

Abstract. Observations of the cosmic microwave background do not yet determine whether inflation was driven by a slowly-rolling scalar field or involved another physical mechanism. In this paper we discuss the prospects of using the power spectra of scalar and tensor modes to probe the nature of inflation. We focus on the leading modification to the slow-roll dynamics, which entails a sound speed c_s for the scalar fluctuations. We derive analytically a lower bound on c_s in terms of a given tensor-to-scalar ratio r , taking into account the difference in the freeze-out times between the scalar and tensor modes. We find that any detection of primordial B-modes with $r > 0.01$ implies a lower bound on c_s that is stronger than the bound derived from the absence of non-Gaussianity in the Planck data. For $r \gtrsim 0.1$, the bound would be tantalizingly close to a critical value for the sound speed, $(c_s)_* = 0.47$ (corresponding to $(f_{\text{NL}}^{\text{equil}})_* = -0.93$), which we show serves as a threshold for non-trivial dynamics beyond slow-roll. We also discuss how an order-one level of equilateral non-Gaussianity is a natural observational target for other extensions of the canonical paradigm.

Keywords: inflation, gravitational waves and CMBR polarization

ArXiv ePrint: [1407.2621](https://arxiv.org/abs/1407.2621)



Contents

1	Introduction	1
2	Implications of a B-mode detection	4
2.1	Spectra of primordial perturbations	4
2.2	Origin of a bound on the sound speed	5
2.3	Degeneracies and second-order corrections	7
3	CMB analysis	9
3.1	Inflationary spectra to second order	9
3.2	Joint analysis of Planck and BICEP2	10
3.3	Robustness of the bound	12
3.4	Prospects of future observations	13
4	A theoretical threshold	15
4.1	Sound waves in the early universe	16
4.2	Energy scales	16
4.3	A critical sound speed	18
5	Physics above threshold	19
5.1	One-scale models & strong coupling	19
5.2	Stable hierarchies	20
5.3	Additional degrees of freedom	21
6	Conclusions and outlook	22
A	Resumming large logarithms	24
A.1	$\varepsilon_2 \ln c_s$	24
A.2	$\varepsilon_3 \ln c_s$	25

1 Introduction

One of the central goals of modern cosmology is to determine the nature of inflation. While measurements of the cosmic microwave background (CMB) and the large-scale structure (LSS) are consistent with the predictions of single-field slow-roll models [1], we should still ask to what degree observations require that inflation occurred in this way.

A systematic way to describe deformations of the canonical framework is the effective field theory (EFT) of inflation [2]. This is a theory of the two massless fields that are guaranteed to be present in any model of inflation: the Goldstone boson of spontaneously broken time translations, π , and the graviton, h_{ij} . In single-field slow-roll inflation the role of the Goldstone boson is played by fluctuations in the inflaton field, which satisfy a relativistic dispersion relation, $\omega = k$, and are only very weakly interacting. Deviations from slow-roll inflation are parameterized by a non-trivial dispersion relation $\omega(k)$, higher-order self-interactions of π or couplings to other fields. A significant advantage of the EFT framework is that it allows us to study scenarios where the accelerated expansion is not necessarily driven by a weakly coupled fundamental scalar field.

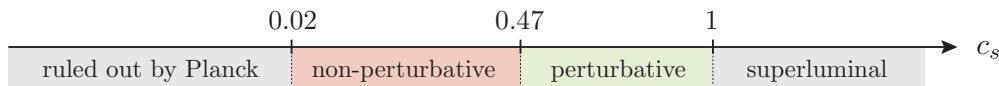


Figure 1. Graphical illustration of the threshold value $(c_s)_* = 0.47$.

A well-motivated possibility is to modify the Goldstone dispersion relation by adding a sound speed, $\omega = c_s k$ [3]. While perturbative higher-derivative corrections to the slow-roll dynamics can only induce small deviations from $c_s = 1$ [4], models with $c_s \ll 1$ are characteristic of non-perturbative physics [5] or non-trivial dynamics [6]. The difference between weakly and strongly coupled inflationary backgrounds is an important qualitative distinction (see figure 1). As we will show, for the theory to be weakly coupled at all relevant energies, the sound speed has to be above the critical value

$$(c_s)_* = 0.47. \quad (1.1)$$

Finding that $c_s > (c_s)_*$, would therefore be a strong indication in favor of the standard scenario, while $c_s < (c_s)_*$ requires physics beyond the slow-roll paradigm.

There are several avenues for probing a sound speed of the inflationary perturbations. Famously, small c_s enhances scalar self-interactions and hence leads to non-Gaussian correlations. The amplitude of the induced equilateral bispectrum is $f_{\text{NL}}^{\text{equil}} \propto c_s^{-2}$, and the absence of significant non-Gaussianity in the Planck data [7] puts a lower bound on the allowed value of the sound speed: $c_s > 0.02$ (95% CL). As we shall see, the critical sound speed (1.1) corresponds to equilateral non-Gaussianity with $(f_{\text{NL}}^{\text{equil}})_* = -0.93$, which is two orders of magnitude below the sensitivity of Planck [7] but two orders of magnitude above the slow-roll expectation. This result is surprisingly robust, as similar thresholds, with $|f_{\text{NL}}^{\text{equil}}| \simeq \mathcal{O}(1)$, exist for other cubic Goldstone interactions even when $c_s = 1$. It is a universal feature that order-one equilateral non-Gaussianity separates qualitatively distinct regions in the parameter space of the EFT of single-field inflation [2] and extensions thereof [8–10]. Probing equilateral non-Gaussianity down to the order-one level is therefore an important and well-motivated observational target [11–15].

The BICEP2 collaboration [16] has recently set a new standard for measurements of B-mode polarization in the CMB by reporting a more than 5σ detection at degree angular scales. While we await confirmation of the primordial origin of the BICEP2 signal, it is a timely issue to address the implications of a detectable tensor-to-scalar ratio for the physics of inflation. In this paper we will show that any detection of primordial tensor modes with $r > 0.01$ puts a lower bound on c_s that is stronger than the Planck-only constraint.

Our bound originates from the fact that a small value of c_s boosts the amplitude of scalar fluctuations and hence suppresses the tensor-to-scalar ratio, which at leading order is given by $r = 16\varepsilon_1 c_s$, where $\varepsilon_1 \equiv -\dot{H}/H^2$ quantifies the deviation of the background from perfect de Sitter space. A large value of r is only compatible with small c_s if the value of ε_1 is much larger than in the canonical expectation in slow-roll models [17]: e.g. $\varepsilon_1 \simeq 0.01$ in $m^2\phi^2$ chaotic inflation. However, for $\varepsilon_1 \gtrsim 0.1$, the expression for r receives an important correction due to the fact that scalar and tensor fluctuations freeze out at different times,

$$r = 16\varepsilon_1 c_s \left(\frac{H_t}{H_s} \right)^2, \quad (1.2)$$

where H_t and H_s denote the Hubble scales at the two freeze-out times. Since $\dot{H} < 0$, when the null energy condition is satisfied, we have $H_t < H_s$ and the tensor-to-scalar ratio receives an additional suppression. This suppression becomes more significant for larger values of ε_1 (faster evolution of H) and smaller values of c_s (larger separation of the freeze-out times). This feature, together with constraints from the shape of the scalar power spectrum, leads to an analytic lower bound on c_s for a given tensor-to-scalar ratio. The ultimate expression for the bound depends on various factors which we will explain in detail throughout the paper. The upshot, however, is clear: a large B-mode signal constrains the sound speed much better than measurements of CMB temperature fluctuations alone. Specifically, for $r > 0.13$ — corresponding to the level reported by BICEP2 and also the benchmark of $m^2\phi^2$ inflation — we find¹

$$c_s > 0.14, \quad (1.3)$$

which would correspond to a bound on the c_s -induced non-Gaussianity of order $|f_{\text{NL}}^{\text{equil}}| \lesssim 10$. Notice that this would be an order of magnitude stronger than the previous Planck-only bound: $f_{\text{NL}}^{\text{equil}} = -42 \pm 75$ [7].

To test our analytical approach we perform a joint likelihood analysis of the CMB data from WMAP, Planck and BICEP2 *without* foreground subtraction.² We find, for constant sound speed, $c_s > 0.25$ (95% CL), which is consistent with (1.3) and supports our analytic arguments. (The small discrepancy is almost entirely due to the preference for a negative running of the spectral index in the CMB data, which favors a larger suppression factor in (1.2). The tendency for negative running disappears if the low- ℓ data is removed, and the results from the numerical and analytic studies approach each other.)

Although a detection of primordial B-modes with $r \gtrsim 0.1$ would lead to a lower bound on c_s remarkably close to the threshold value in (1.1), this does not mean non-Gaussianity may be absent in future observations [12, 17, 21, 22]. In fact, the bispectrum remains an important observable, and there exist at least three distinct possibilities for a detectable signal: first of all, values of c_s and $f_{\text{NL}}^{\text{equil}}$ near the threshold (i.e. $c_s \simeq 0.5$ and $|f_{\text{NL}}^{\text{equil}}| \simeq \mathcal{O}(1)$) arise naturally from strongly coupled backgrounds with a single scale. Second, there are other (technically natural) deformations of slow-roll inflation that are unrelated to the sound speed [2, 23] and therefore allow for potentially observable equilateral non-Gaussianity irrespective of a bound. Finally, the presence of additional light degrees of freedom during inflation may also induce a detectable bispectrum. Some of these scenarios, such as quasi-single-field inflation [24] or non-adiabatic dissipative effects [9, 25], can lead to sizable equilateral non-Gaussianity. These cases may be distinguished from the first two options through correlations between the equilateral shape and certain squeezed limits.

The outline of the paper is as follows: in section 2, we derive an analytic bound on the sound speed for a given value of the tensor-to-scalar ratio. This result follows from the expression in (1.2) and relies on minimal input from the data. In section 3, we present a joint analysis of data from WMAP, Planck and BICEP2 (without foreground subtraction) which confirms our analytic expectation. We discuss the robustness of our results to variations in the inflationary parameters and changes in the cosmological data sets. In section 4, we obtain

¹In order to arrive at (1.3) we have assumed a constant sound speed. This is the case that is most relevant for a direct comparison with the Planck limit $c_s > 0.02$ (95% CL). We will also present bounds on c_s that do not make this assumption yet lead to qualitatively similar results.

²We warn the reader this analysis is performed as a case of study only, and future maps will be required to properly assess the level of foregrounds in the BICEP2 region, e.g. [18–20]. Fortunately, the debate will soon be settled by future CMB polarization measurements.

a critical value for c_s , and a related threshold for $f_{\text{NL}}^{\text{equil}}$, from a unitarity bound in the EFT of inflation. The threshold divides perturbative slow-roll inflation from strongly coupled backgrounds for sound speed models. In section 5, we survey a range of well-motivated scenarios, consistent with current data, which may produce a detectable level of equilateral non-Gaussianity even for c_s close to 1. We conclude with a summary and an outlook on future prospects in section 6.

2 Implications of a B-mode detection

We begin with an analytic discussion of the consequences of a B-mode detection for inflationary models with a sound speed (see also [12, 17, 21, 22]). The main result of this section will be a lower bound on the sound speed as a function of the tensor-to-scalar ratio.

2.1 Spectra of primordial perturbations

Quantum fluctuations of any massless fields get amplified during inflation. Two massless fields that are guaranteed to exist in any model of inflation are the curvature perturbation ζ and the graviton h_{ij} . In comoving gauge, these fields are isotropic and anisotropic perturbations to the spatial metric, respectively,

$$g_{ij}(t, \mathbf{x}) = a^2(t) \left[e^{2\zeta(t, \mathbf{x})} \delta_{ij} + 2h_{ij}(t, \mathbf{x}) \right], \quad (2.1)$$

where h_{ij} is transverse and traceless. The scale factor $a(t)$ is that of an arbitrary quasi-de Sitter background with Hubble expansion rate $H(t) \equiv \partial_t \ln a$. The dynamics of ζ may contain a time-dependent sound speed $c_s(t)$. To describe the evolution of the Hubble parameter and the sound speed, it is convenient to introduce the following *flow parameters*³

$$\varepsilon_{n+1} \equiv \frac{d \ln \varepsilon_n}{dN}, \quad \varepsilon_0 \equiv \frac{H(N_i)}{H(N)}, \quad (2.2)$$

$$\delta_{n+1} \equiv \frac{d \ln \delta_n}{dN}, \quad \delta_0 \equiv \frac{c_s(N)}{c_s(N_i)}, \quad (2.3)$$

where $N \equiv \ln a$ and N_i denotes an arbitrary reference time. Provided fluctuations are produced in the vacuum,⁴ the power spectra of ζ and h_{ij} in models with a sound speed are

$$\Delta_\zeta^2(k) = \frac{1}{8\pi^2} \frac{H^2}{M_{\text{pl}}^2} \frac{1}{\varepsilon_1 c_s} \bigg|_{c_s k = aH}, \quad (2.4)$$

$$\Delta_h^2(k) = \frac{2}{\pi^2} \frac{H^2}{M_{\text{pl}}^2} \bigg|_{k=aH}, \quad (2.5)$$

to leading order in the flow parameters. Notice that while the scalar fluctuations are evaluated when modes cross the sound horizon, $c_s k = aH$, the tensor fluctuations are evaluated at $k = aH$. This fact will play a key role in what follows. As long as the

³In an abuse of terminology, we will sometimes refer the Hubble flow parameters ε_n and the sound flow parameters δ_n as the ‘slow-roll parameters’.

⁴Scalar and tensor perturbations may also be produced from non-vacuum states. This possibility was studied in detail in [9, 26, 27], where it was shown that non-Gaussianity may be able to discern these types of models.

slow-roll conditions are satisfied, $|\varepsilon_n| \ll 1$ and $|\delta_n| \ll 1$, the induced scalar spectrum has an approximate power-law form

$$\Delta_\zeta^2 = A_s \left(\frac{k}{k_0} \right)^{n_s - 1 + \frac{1}{2}\alpha_s \ln(k/k_0)}, \quad (2.6)$$

with $A_s \simeq 2.2 \times 10^{-9}$ and, to leading order,

$$n_s - 1 = -2\varepsilon_1 - \varepsilon_2 - \delta_1, \quad (2.7)$$

$$\alpha_s = -2\varepsilon_1\varepsilon_2 - \varepsilon_2\varepsilon_3 - \delta_1\delta_2. \quad (2.8)$$

The right-hand sides in (2.7) and (2.8) are to be evaluated when the pivot scale k_0 crosses the sound horizon. Notice that the scalar running α_s starts at second order in the flow parameters. This is a key feature of inflationary models, which will be important in our analysis below. Throughout this section, we will assume $\varepsilon_3 \ll \varepsilon_{1,2}$ and $\delta_2 \ll \delta_1$. This approximation will be relaxed in section 3, where we present a numerical analysis at higher order in the slow-roll expansion.

2.2 Origin of a bound on the sound speed

The fact that tensors and scalars freeze out at different times is important, as it implies that the tensor-to-scalar ratio will depend on the ratio of the Hubble scales at the two freeze-out times:

$$r \equiv \frac{\Delta_h^2}{\Delta_\zeta^2} = 16\varepsilon_1 c_s \left(\frac{H_t}{H_s} \right)^2, \quad (2.9)$$

where $H_t \equiv H(N_t)$ and $H_s \equiv H(N_s)$ denote the Hubble parameters at $k = aH$ and $c_s k = aH$, respectively. For $c_s < 1$, the sound horizon is smaller than the Hubble radius and the scalars freeze out before the tensors do (see figure 2). Since $\dot{H} < 0$ by virtue of the null energy condition, we have $H_t/H_s < 1$ and the tensor-to-scalar ratio in (2.9) receives an additional suppression. In order to quantify this effect, we evolve $H(t)$ in the slow-roll approximation. At next-to-leading order, we have

$$H_t = H_s(1 - \varepsilon_1(N_t - N_s) + \dots) \simeq H_s(1 + \varepsilon_1 \ln c_s + \dots), \quad (2.10)$$

where $\varepsilon_1 \equiv \varepsilon_1(N_s)$ and the ellipses denote terms at higher order in ε_1 and $\varepsilon_{n \geq 2}$. Hence, we get

$$r \simeq 16\varepsilon_1 c_s \left[1 + 2\varepsilon_1 \ln c_s + \dots \right]. \quad (2.11)$$

The leading order expression, $r = 16\varepsilon_1 c_s$, implies that $r \simeq 0.1$ with $c_s \simeq 0.02$ requires $\varepsilon_1 \simeq 0.3$. However, for such a large value of ε_1 and such a small value of c_s , the correction in (2.11) is not small, i.e. $2\varepsilon_1 \ln c_s \simeq -2.4$. To treat this regime of the parameter space, we need to understand the $\ln c_s$ -enhanced contributions in the slow-roll expansion.

By definition, the ratio H_t/H_s is determined by an integral over the time-dependent Hubble flow parameter $\varepsilon_1(N)$:

$$\ln(H_t/H_s) = - \int_{N_s}^{N_t} \varepsilon_1(N) dN \equiv -\bar{\varepsilon}_1 \Delta N, \quad (2.12)$$

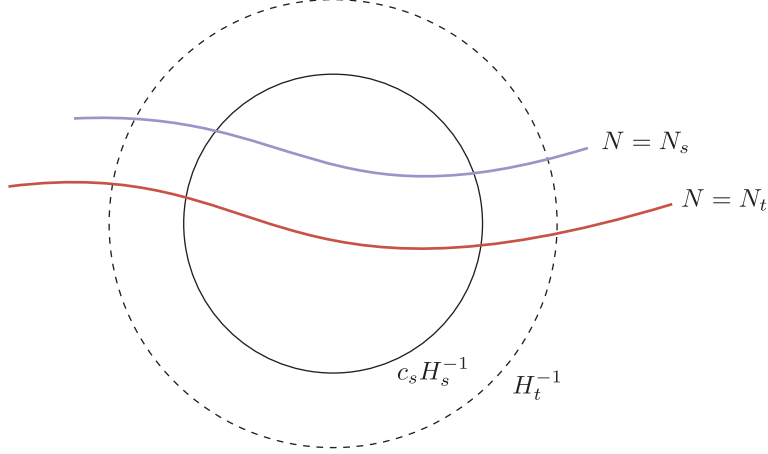


Figure 2. For $c_s < 1$, scalars and tensors freeze out at different times.

where $\Delta N \equiv N_t - N_s$ and $\bar{\varepsilon}_1 \equiv \varepsilon_1(\bar{N})$, for $\bar{N} \in [N_s, N_t]$ as guaranteed by the mean value theorem. At the same time, the horizon crossing conditions $k = aH$ and $c_s k = aH$ imply

$$\frac{H_t}{H_s} = \frac{1}{c_s} \frac{a_s}{a_t} = \frac{1}{c_s} e^{-\Delta N}. \quad (2.13)$$

Combining (2.12) and (2.13), we find

$$\Delta N = \frac{-\ln c_s}{1 - \bar{\varepsilon}_1}, \quad (2.14)$$

which, after substitution into (2.12) and (2.9), leads to

$$r = 16\varepsilon_1 c_s^{\frac{1+\bar{\varepsilon}_1}{1-\bar{\varepsilon}_1}}. \quad (2.15)$$

We see that the tensor-to-scalar ratio is suppressed for $\bar{\varepsilon}_1 \gtrsim \varepsilon_1$, while this effect can be avoided if the dynamics leads to $\bar{\varepsilon}_1 \ll \varepsilon_1$. In order for (2.15) to become a meaningful expression, we need to determine the auxiliary parameter $\bar{\varepsilon}_1$ in terms of the Hubble flow parameters ε_n evaluated at N_s .

To leading order in the slow-roll expansion, we have $\varepsilon_1 \simeq \text{const.}$ and hence $\bar{\varepsilon}_1 \simeq \varepsilon_1$. Since we are not keeping higher orders (unless they are $\ln c_s$ -enhanced), we should expand the exponent in (2.15):

$$r \simeq 16\varepsilon_1 c_s^{1+2\varepsilon_1}. \quad (2.16)$$

This result corresponds to a resummation of the *leading logarithm* of (2.11), i.e. it is valid to all orders in $\varepsilon_1 \ln c_s$, but only holds to leading order in ε_1 . As expected, r in (2.16) receives an additional suppression for large ε_1 and small c_s . In fact, at fixed c_s , the tensor-to-scalar ratio reaches a maximum value at $\varepsilon_1^{\text{max}} = -(2 \ln c_s)^{-1}$, as can be seen in the left panel of figure 3. Turning this around, a detection of r would imply a lower bound on the sound speed, as shown in the right panel of figure 3. We see that $r > 0.13$ requires

$$c_s > 0.1. \quad (2.17)$$

Notice that even values of r an order of magnitude smaller would give a bound on c_s that is as strong as the bound derived from the Planck measurement of the bispectrum.

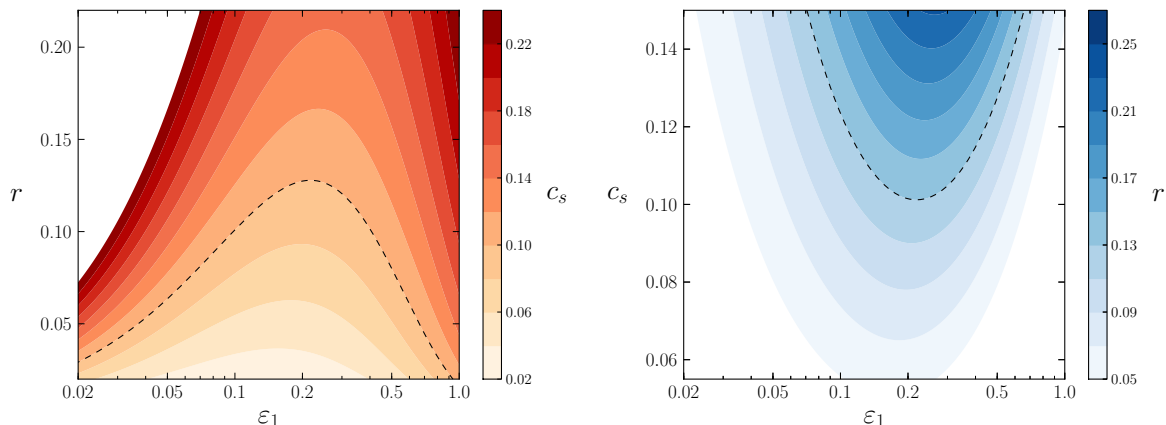


Figure 3. Plots of $c_s(r, \varepsilon_1)$ (left) and $r(c_s, \varepsilon_1)$ (right) as given by (2.16). The dashed contour in the left plot shows the value of r for $c_s = 0.1$. It has a maximum of $r = 0.13$ for $\varepsilon_1 = 0.2$. The dashed contour in the right plot illustrates the lower bound on c_s for $r = 0.13$.

We emphasize that the bound in (2.17) is model-dependent and has assumed that $\varepsilon_1(N)$ does not vary significantly. As we already alluded to, the lower bound on c_s may be relaxed if $\bar{\varepsilon}_1 \ll \varepsilon_1$ in (2.15). This is achieved for large and negative values of ε_2 ,⁵ so that ε_1 is large at sound horizon crossing, but decreases quickly thereafter. However, a non-negligible ε_2 also produces large logarithms. In appendix A, we show how to resum the $\varepsilon_2 \ln c_s$ terms. The result is $\bar{\varepsilon}_1 = \varepsilon_1(1 - c_s^{-\varepsilon_2})/(\varepsilon_2 \ln c_s)$, which implies

$$r = 16\varepsilon_1 c_s^{1+2\varepsilon_1 \cdot (1-c_s^{-\varepsilon_2})/(\varepsilon_2 \ln c_s)}. \quad (2.18)$$

In the left panel of figure 4, we see that $r \simeq 0.13$ is now indeed consistent with small c_s if both ε_1 and $|\varepsilon_2|$ are large. However, having both of these parameters take on large values induces a significant running of the spectral index α_s , cf. (2.8). Currents constraints on the running thus restrict the size of $\varepsilon_1 \varepsilon_2$. Indeed, imposing $|\alpha_s| \lesssim 2 \times 10^{-2}$ [1],⁶ we recover a bound on c_s that is only marginally weaker than the previous bound; see the right panel of figure 4.

2.3 Degeneracies and second-order corrections

Up until now we have only used minimal data input to derive an analytic constraint on c_s : a lower bound on r and an upper bound on $|\alpha_s|$. When we perform a detailed CMB analysis, in the next section, we will get additional constraints from the precise shape of the scalar spectrum. In particular, we have yet to take into account its near scale-invariance, i.e. $n_s \simeq 0.96$ [1]. Moreover, the numerical value of the bound on c_s will be affected by second-order corrections to the primordial spectra (2.4) and (2.5), which so far we have not incorporated. We conclude this section with a brief discussion on how the main feature of our complete analysis — a stronger bound on c_s — can also be understood analytically.

⁵The case $\bar{\varepsilon}_1 > \varepsilon_1$, namely $\varepsilon_2 > 0$, obviously introduces extra suppression in r , which strengthens the bound.

⁶This value follows from the 95% confidence interval, while taking $\alpha_s = 0$ as the central value. We will take into account the full data set in section 3.

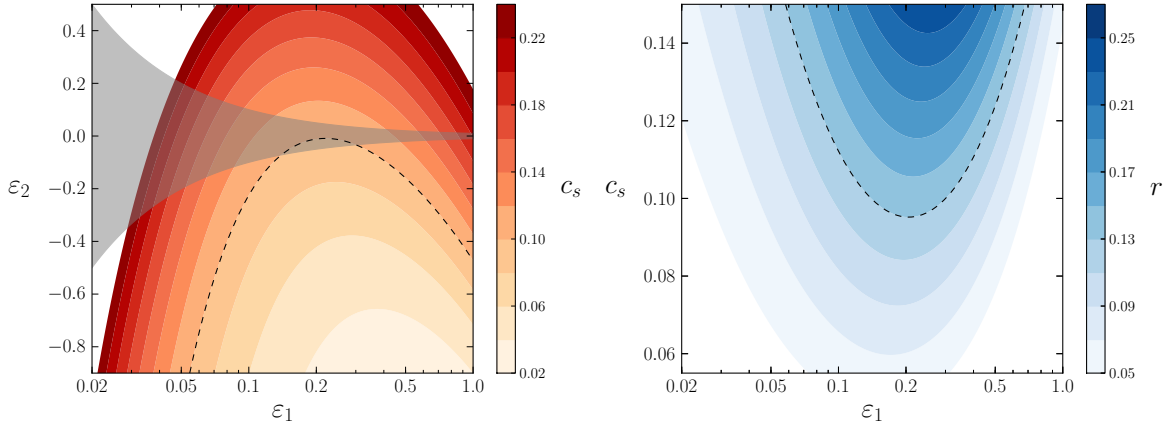


Figure 4. *Left:* plot of $c_s(\varepsilon_1, \varepsilon_2)$, as given by (2.18), for $r = 0.13$. The dashed contour corresponds to the previous bound $c_s = 0.1$. The grey shaded area is the region that is consistent with the constraint from the running: $|\varepsilon_1 \varepsilon_2| < 0.01$. *Right:* plot of $c_s(r, \varepsilon_1, \varepsilon_2)$ as given by (2.18) with $\varepsilon_2 \rightarrow -0.01/\varepsilon_1$ (the most negative value consistent with the bound on α_s).

The most important second-order effect is a correction to the expression (2.18) for the tensor-to-scalar ratio

$$r = 16\varepsilon_1 c_s^{1+2\varepsilon_1 \cdot (1-c_s^{-\varepsilon_2})/(\varepsilon_2 \ln c_s)} \left[1 - \mathcal{C}\varepsilon_2 + (2 - \mathcal{C})\delta_1 \right], \quad (2.19)$$

where $\mathcal{C} \simeq 0.73$. The importance of the extra terms in the square brackets depends on the sizes of the flow parameters ε_2 and δ_1 . Since both parameters appear at leading order in the expression for n_s , cf. (2.7), they are constrained by the near scale-invariance of the spectrum. Moreover, we have seen that the combination of large r and small c_s typically requires relatively large values of ε_1 . Consistency with the measured spectral index then means that either δ_1 or ε_2 (or a combination thereof) has to be chosen to cancel part of the contribution from ε_1 . These two possibilities correspond to two perfect degeneracies that keep n_s within experimental bounds:

- **$\delta_1 \simeq -2\varepsilon_1$.** Allowing for a time dependence in the sound speed gives enough freedom to satisfy the constraints from n_s and α_s simultaneously. In particular, the time variation of $c_s(t)$ can be chosen to nearly cancel the effect of the rather rapid evolution of the Hubble parameter $H(t)$. Setting $\delta_1 \simeq -2\varepsilon_1$ makes any value of ε_1 consistent with the measured value of n_s . Moreover, if $|\varepsilon_2| \lesssim 10^{-2}$ then ε_1 is not constrained by the bound on the running α_s . The analysis is then similar to what led us to figure 3. However, this time we include the second-order correction in r , which along the degeneracy $\delta_1 \simeq -2\varepsilon_1$ becomes

$$r \simeq 16\varepsilon_1 c_s^{1+2\varepsilon_1} \left[1 - 2.54\varepsilon_1 \right]. \quad (2.20)$$

Since $\varepsilon_1 > 0$, the correction in the square brackets suppresses r and we expect a stronger bound on c_s compared to what we found in (2.17). Indeed, $r > 0.13$ now requires

$$c_s > 0.15 \quad \text{for} \quad \delta_1 \simeq -2\varepsilon_1. \quad (2.21)$$

- **$\varepsilon_2 \simeq -2\varepsilon_1$.** Allowing for a varying ε_1 , a potentially large contribution to n_s from ε_1 may be cancelled by choosing $\varepsilon_2 \simeq -2\varepsilon_1$ while keeping c_s constant (i.e. $\delta_1 = 0$).

Along this degeneracy curve any value of ε_1 is consistent with the observed spectral index. However, the running becomes $\alpha_s = -2\varepsilon_1\varepsilon_2 \rightarrow 4\varepsilon_1^2$, so that the constraint $|\alpha_s| < 2 \times 10^{-2}$ implies

$$\varepsilon_1 < 0.07. \quad (2.22)$$

Moreover, the tensor-to-scalar ratio along the degeneracy is given by

$$r \simeq 16\varepsilon_1 c_s^{1 - (1 - c_s^{2\varepsilon_1}) / \ln c_s} \left[1 + 1.46 \varepsilon_1 \right]. \quad (2.23)$$

Combining (2.23) and (2.22), and imposing $r > 0.13$, then leads to

$$c_s > 0.14 \quad \text{for} \quad \varepsilon_2 \simeq -2\varepsilon_1. \quad (2.24)$$

Notice that (2.24) can be read off from figure 4 as the overlap between $\varepsilon_1 \simeq 0.07$ and the grey shaded region.

In the next section, we will present a complete CMB analysis. Remarkably, we will find only small departures from the bound we arrived at analytically.

3 CMB analysis

A complete likelihood analysis can differ from the analytic bounds of the previous section because of additional degeneracies between parameters or stronger constraints arising from the precise form of the scalar power spectrum. We would like to understand to what degree our analytic estimates have accounted for these effects. For purposes of illustration, we will perform a joint likelihood analysis of data from WMAP, Planck, and BICEP2. We appreciate the significant level of uncertainty in the modelling of dust foregrounds in the BICEP2 region of the sky [18–20]. Our analysis will use the BICEP2 likelihood without foreground subtraction, but we caution the reader that quantitative details of our results are subject to change in the event that the BICEP2 likelihood is significantly revised after a better understanding of the foregrounds. Our goal in this section therefore isn't to derive a quantitative bound on the sound speed, but to understand if our analytic approach could have missed an unforeseen degeneracy that would allow the bound on c_s to be evaded completely. We will find that this is not the case and our analytic result is therefore a good reflection of what could be expected of future data analyses. A similar analysis for the case of slow-roll inflation has recently appeared in [28].

3.1 Inflationary spectra to second order

For consistency, we will use results for the scalar and tensor power spectra at second order in the Hubble flow parameters (2.2) and the sound flow parameters (2.3); see [29–33]. We write the power spectra in the standard power-law form

$$\Delta_\zeta^2(k) = A_s \left(\frac{k}{k_0} \right)^{n_s - 1 + \frac{1}{2}\alpha_s \ln(k/k_0)}, \quad (3.1)$$

$$\Delta_h^2(k) = r A_s \left(\frac{k}{k_0} \right)^{n_t + \frac{1}{2}\alpha_t \ln(k/k_0)}, \quad (3.2)$$

where $k_0 = 0.05 \text{ Mpc}^{-1}$ is our chosen pivot scale. At second order, we then have

$$n_s - 1 = -2\varepsilon_1 - \varepsilon_2 - \delta_1 - 2\varepsilon_1^2 - (3 - 2\mathcal{C})\varepsilon_1\varepsilon_2 + \mathcal{C}\varepsilon_2\varepsilon_3 - 3\varepsilon_1\delta_1 - \varepsilon_2\delta_1 - (2 - \mathcal{C})\delta_1\delta_2 - \delta_1^2, \quad (3.3)$$

$$\alpha_s = -2\varepsilon_1\varepsilon_2 - \varepsilon_2\varepsilon_3 - \delta_1\delta_2, \quad (3.4)$$

$$n_t = -2\varepsilon_1 - 2\varepsilon_1^2 - 2(1 - \mathcal{C} - \ln c_s)\varepsilon_1\varepsilon_2, \quad (3.5)$$

$$\alpha_t = -2\varepsilon_1\varepsilon_2, \quad (3.6)$$

$$r = 16\varepsilon_1c_s \left[1 + 2\varepsilon_1 \ln c_s - \mathcal{C}\varepsilon_2 + (2 - \mathcal{C})\delta_1 \right], \quad (3.7)$$

where $\mathcal{C} \equiv 2 - \ln 2 - \gamma$ (with $\gamma = 0.5772$ the Euler-Mascheroni constant). Even the indices n_t and α_t are now expressed in terms of slow-roll parameters evaluated when the pivot scale crosses the sound horizon, $c_s k_0 = aH$. This is responsible for the $\ln c_s$ -terms in the above formulas. In appendix A, we show that the resummation of these logarithms gives

$$r = 16\varepsilon_1c_s^{1+2\varepsilon_1 \cdot (1-c_s^{-\varepsilon_2})/(\varepsilon_2 \ln c_s)} \left[1 - \mathcal{C}\varepsilon_2 + (2 - \mathcal{C})\delta_1 \right]. \quad (3.8)$$

We will use this expression for the tensor-to-scalar ratio in our analysis. This allows us to explore smaller values of c_s while maintaining perturbative control.

3.2 Joint analysis of Planck and BICEP2

This section describes the methodology and the results of a dedicated likelihood analysis using data from WMAP, Planck and BICEP2.

The Planck likelihoods were described in detail in [34]. We use the **CamSpec** likelihood for the Planck temperature power spectrum in the multipole range $50 \leq \ell \leq 2500$ and the **Commander** likelihood for $2 \leq \ell \leq 49$. As in the Planck analysis [35], we use the WMAP polarization data for $\ell \leq 32$. We will also show results with the low- ℓ likelihoods excluded. This helps to identify effects that are driven by the low- ℓ anomalies. We use all nine band-powers of the BICEP2 observations [16, 36]. Pending a resolution in the debate about the importance of dust foregrounds in the BICEP2 region, we use the BICEP2 likelihood without foreground subtraction.

Our theoretical model consists of the standard cosmological parameters of the Λ CDM model, $\theta_{\text{lcdm}} = \{\Omega_b h^2, \Omega_c h^2, \theta_{\text{MC}}, \tau\}$, as well as the inflationary parameters $\theta_i = \{A_s, c_s, \varepsilon_1, \varepsilon_2, \varepsilon_3, \delta_1, \delta_2\}$ characterizing the initial conditions. We also consider various subsets of the inflationary parameters. We modified the Boltzmann code **CAMB** [38] to take the spectra of section 3.1 as input and compute the joint likelihood $P(d|\theta)$, i.e. the probability of the data d given the model parameters θ . We use uniform prior probabilities $P(\theta)$ with prior ranges listed in table 1. The parameter space was sampled by a Markov-Chain Monte Carlo (MCMC) method using the publicly available software package **CosmoMC** [37]. We ran twenty chains until the variation in the means of the chains was small relative to the standard deviation (using $R - 1 < 0.02$ in the Gelman-Rubin [39] criterion). In the plots below we present the posterior probabilities $P(\theta|d)$ for our model parameters. For most of the analysis the cosmological parameters θ_{lcdm} are fixed to their best-fit Planck values [35], except in section 3.3 where we marginalize over them and include baryon acoustic oscillation (BAO) data [40, 41] to break degeneracies. Let us remind the reader that all bounds quoted here are 95% confidence limits unless otherwise stated.

Parameter	Prior	Physical Meaning
$\Omega_b h^2$	[0.005, 0.1]	baryon density today
$\Omega_c h^2$	[0.001, 0.99]	cold dark matter density today
$100\theta_{\text{MC}}$	[0.5, 10.0]	angular sound horizon [37]
τ	[0.01, 0.8]	optical depth to reionization
$\ln(10^{10} A_s)$	[2.7, 4.0]	scalar amplitude
c_s	[0.02, 1.0]	speed of sound
ε_1	[0.001, 0.5]	Hubble flow parameter (HFP)
ε_2	[-0.5, 0.5]	HFP
ε_3	[-0.1, 0.1]	HFP
δ_1	[-0.5, 0.5]	sound flow parameter (SFP)
δ_2	[-0.1, 0.1]	SFP

Table 1. Parameters used in the CosmoMC analysis and their prior ranges.

The current best limit on c_s comes from the absence of primordial non-Gaussianity in the Planck data: $c_s > 0.02$. This limit assumes a constant sound speed. To make a direct comparison⁷ with the Planck analysis, we should therefore take the space of parameters to be $\theta_i = \{c_s, \varepsilon_1, \varepsilon_2\}$, while setting $\varepsilon_3 = \delta_{1,2} = 0$. We will consider this case first, and then discuss what happens when we add additional inflationary parameters. The one-dimensional marginalized posterior probability distribution for c_s is shown in figure 5 and the inferred bound on c_s is

$$c_s > 0.25 \quad \text{for } \delta_1 = 0. \quad (3.9)$$

We remind the reader that this bound was derived from the BICEP2 likelihood *without* foreground subtraction. Including the effects of foregrounds would lower the bound, or might even remove it completely if no primordial tensor signal survives further scrutiny. A complete analysis will have to await the improved likelihoods that will appear with the Planck/BICEP joint analysis.

While the split between $\varepsilon_{1,2}$ and ε_3 is well-motivated by the order at which they appear in the slow-roll expansion, setting $\delta_1 = 0$ is a fairly artificial choice. For example, in (3.3) we see that δ_1 contributes to the spectral index on the same footing as $\varepsilon_{1,2}$. As our *baseline* analysis, we will therefore take $\theta_i = \{c_s, \varepsilon_1, \varepsilon_2, \delta_1\}$. The left panel of figure 5 shows the marginalized posterior probability distribution for the parameters c_s and ε_1 for both the baseline analysis and the case $\delta_1 = 0$. We see that larger values of ε_1 are allowed if we marginalize over finite δ_1 . This feature can be understood in terms of the two degeneracies discussed in section 2.3. For $\varepsilon_1 > 0.01$, one of the degeneracies must be enacted to accommodate the observed value of n_s . When $\delta_1 = 0$, the only option is $\varepsilon_2 \simeq -2\varepsilon_1$, which is constrained by the limit on α_s . On the other hand, if $\delta_1 \simeq -2\varepsilon_1$ the constraint from the running is satisfied provided $|\varepsilon_2| \ll \varepsilon_1$, in which case the bound on ε_1 is relaxed. The one-dimensional posterior for c_s is now given by the dashed curve in figure 5. The bound on c_s is

$$c_s > 0.21 \quad \text{for } \delta_1 \neq 0. \quad (3.10)$$

⁷The Planck limit includes a marginalization over the parameter M_3 discussed in section 5.2. Since this parameter does not contribute to the power spectrum our bound does not need to be adjusted to make this comparison.

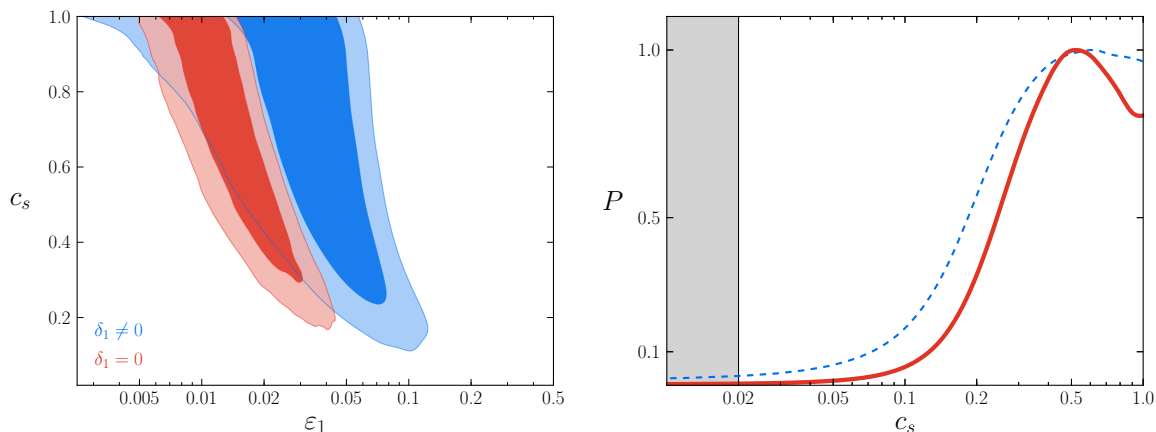


Figure 5. *Left:* 68% and 95% confidence contours of the marginalized posterior probability distribution for the parameters c_s and ε_1 . The red shows that case with $\delta_1 = 0$, while the blue allows for $\delta_1 \neq 0$. *Right:* marginalized posterior probability distribution for c_s , for $\delta_1 = 0$ (solid red line) and $\delta_1 \neq 0$ (dashed blue line). The Planck exclusion, $c_s > 0.02$, is shown in grey.

We note that both (3.9) and (3.10) are somewhat stronger than the corresponding analytic bounds in section 2.3. The goal for the remainder of this section is to understand the reason for these stronger bounds and to determine their robustness to changes in our priors and variations of the data sets. As we shall see, much of the difference is driven by the anomalies in the low- ℓ data.

3.3 Robustness of the bound

As we described in section 2.3, small values of c_s could be made consistent with a large tensor-to-scalar ratio, e.g. $r > 0.13$, if $\varepsilon_2 \ll 0$. However, this would require that the running is positive, $\alpha_s = -2\varepsilon_1\varepsilon_2 > 0$, whereas it is well-known that negative running, $\alpha_s < 0$, improves the fit to the low- ℓ CMB temperature data [35]. When $\delta_1 = 0$, this tension then produces a much stronger constraint on ε_2 , along the $\varepsilon_2 \simeq -2\varepsilon_1 < 0$ degeneracy, than we would expect from the error on α_s alone (see the red contours in figure 6). This constraint is eliminated when we allow for $\delta_1 \neq 0$, and instead there is a clear tendency from the low- ℓ multipoles to prefer $\varepsilon_2 > 0$ (see the blue contours in figure 6).

To test how much our bounds depend on the anomalously low power in the low- ℓ CMB data, we repeat the analysis without the low- ℓ likelihoods. Removing the low- ℓ data eliminates the preference for $\alpha_s < 0$ and hence allows ε_2 to take on a wider range of values. In particular, for $\delta_1 = 0$, excluding the low- ℓ multipoles allows for somewhat larger (and negative) values of ε_2 (and hence smaller values of c_s). This can be seen by comparing the green and red contours in figure 6 or by considering the range of allowed ε_1 with and without the low- ℓ data in figure 7. As expected, the new bound on c_s is weaker than in the case of the full data set,

$$c_s > 0.14 \quad \text{for } \delta_1 = 0, \ell \geq 50. \quad (3.11)$$

For $\delta_1 \neq 0$, removing the low- ℓ data has the effect of allowing large values of ε_1 which lower the bound on c_s , combined with small values of ε_2 to satisfy the constraint from the running (see the orange contours in the left panel of figure 6). In this case the main degeneracy relevant to small values of c_s is $\delta_1 \simeq -2\varepsilon_1$, where the combination of large negative δ_1 and

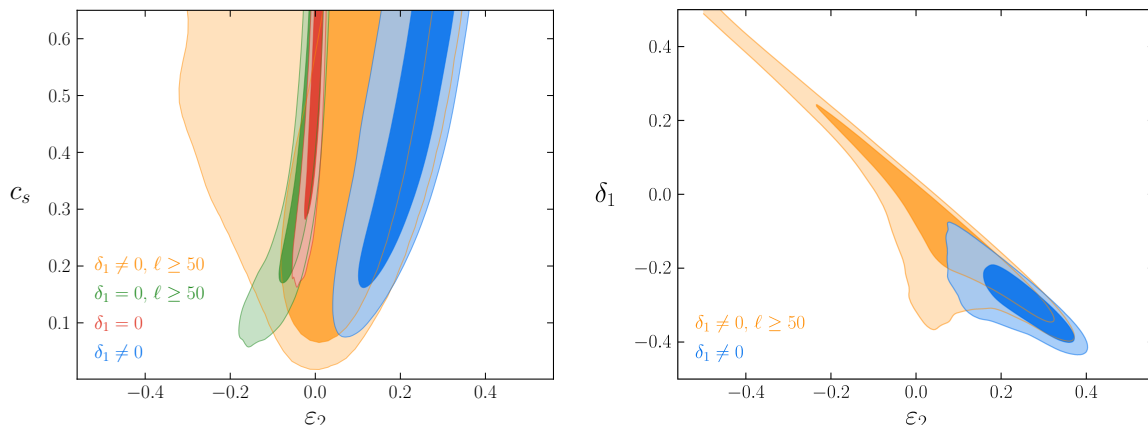


Figure 6. 68% and 95% confidence contours of the marginalized posterior probability distributions. *Left:* contours for the parameters c_s and ε_2 for $\delta_1 = 0$ and $\delta_1 \neq 0$ both with and without the low- ℓ data. We note that the full data set disfavors large negative ε_2 , while it becomes allowed when the low- ℓ data is excluded. The contours for $\delta_1 \neq 0$ also show that the smaller values of c_s push $\varepsilon_2 \rightarrow 0$, as we would expect from the constraint on α_s along the $\delta_1 \simeq -2\varepsilon_1$ degeneracy with large values of ε_1 . A similar feature does not appear when $\delta_1 = 0$, because in that case lower values of c_s are allowed along the $\varepsilon_2 \simeq -2\varepsilon_1$ degeneracy. *Right:* contours for the parameters δ_1 and ε_2 with and without the low- ℓ data. The tendency for positive ε_2 disappears when the low- ℓ data is removed and large negative values of δ_1 are allowed when ε_2 is small.

small ε_2 is allowed (see the orange contours in the right panel of figure 6). As a result, the bound on c_s is also weaker,

$$c_s > 0.15 \quad \text{for } \delta_1 \neq 0, \ell \geq 50. \quad (3.12)$$

The expressions in (3.11) and (3.12) essentially match the corresponding analytic bounds in section 2.3. This strongly suggests that the previous difference between the analytic results and the full likelihood analysis was driven by the low- ℓ data.⁸

Finally, we wish to explore how robust our bounds are to changes in our priors. First, we consider whether there is a significant difference between fixing the cosmological parameters of the Λ CDM model (as we did so far) and marginalizing over them. The posteriors for c_s for both cases are shown in figure 8 and we find little change to our results. Second, we allow for the higher-order flow parameters ε_3 and δ_2 in the initial conditions. From (3.4) we see that both parameters could potentially cancel the contribution to running from $-2\varepsilon_1\varepsilon_2$. However, in order to achieve such a cancelation, ε_3 or δ_2 would need to be large. Allowing for such large values is inconsistent with a hierarchical structure of the slow-roll parameters. Furthermore, third- and higher-order slow-roll corrections which we are neglecting are not necessarily small under such circumstances. If we restrict the range of parameter to be $\{\delta_2, \varepsilon_3\} \in [-0.1, 0.1]$, we find little effect on c_s ; see figure 8 and table 2.

3.4 Prospects of future observations

Our bound on c_s is approaching the critical value, $(c_s)_* = 0.47$, yet experimental improvements are required in order to reach it. The reason our bound is not stronger is due to

⁸Notice that allowing for $\delta_1 \neq 0$ in the analysis without the low- ℓ data does not produce a significant change in the bound on c_s . See also section 2.3.

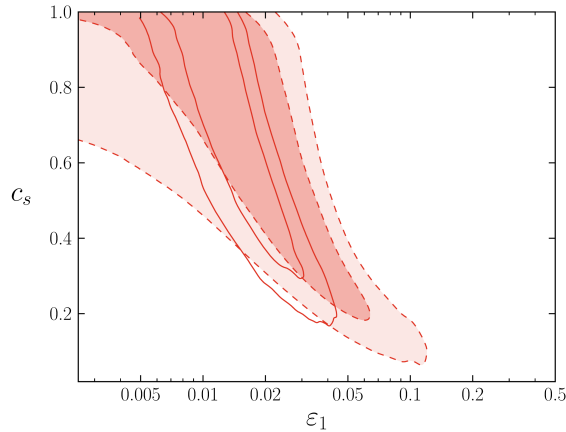


Figure 7. 68% and 95% confidence contours for $\delta_1 = 0$ with and without the low- ℓ data are shown in solid and dashed lines, respectively. The contours without the low- ℓ data are filled in red. We see that removing the low- ℓ data allows for larger values of ε_1 , which favors smaller c_s . This is consistent with the observation that a preference for negative α_s is responsible for the strong limits on ε_1 and c_s shown in figure 5.

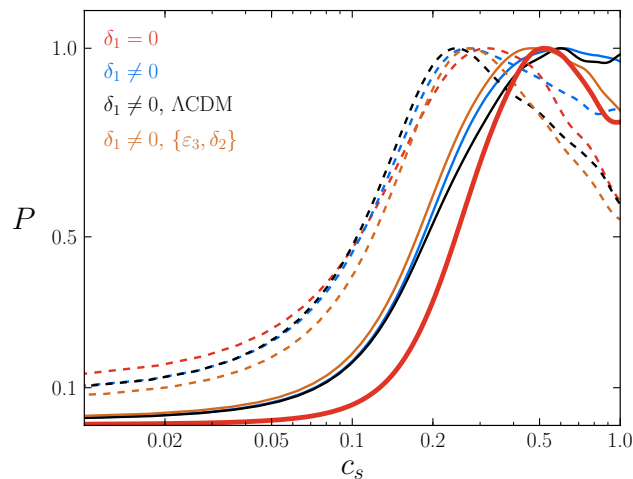


Figure 8. Marginalized posterior probability distributions for c_s for the complete CMB data (solid lines) and without the low- ℓ data (dashed lines). The labels ‘ Λ CDM’ and ‘ $\{\varepsilon_3, \delta_2\}$ ’ denote marginalizations over cosmological parameters and higher-order flow parameters, respectively.

the well-understood degeneracies of section 2.3. To improve the bound, we therefore have to search for observables that break these degeneracies. We will briefly comment on a few possibilities:

- A measurement of the tensor tilt, $n_t = -2\varepsilon_1$, determines ε_1 directly, and therefore breaks all of the important degeneracies that limit the bound on c_s . If $r \simeq 0.13$, then to get to $c_s > 0.47$ at the 95% confidence level, would require $\sigma_{n_t} \sim 0.01$. Unfortunately, to achieve this sensitivity is essentially impossible with CMB measurements alone [42]. Realistic projections for future constraints on n_t may get down to $\sigma_{n_t} \sim 0.1$ [43], which is not sufficient to significantly improve the bound on c_s .

	CMB11	LowLike
$\delta_1 = 0$	$c_s > 0.25$	$c_s > 0.14$
$\delta_1 \neq 0$	$c_s > 0.21$	$c_s > 0.15$
$\delta_1 \neq 0, \Lambda\text{CDM}$	$c_s > 0.21$	$c_s > 0.13$
$\delta_1 \neq 0, \{\varepsilon_3, \delta_2\}$	$c_s > 0.20$	$c_s > 0.13$

Table 2. Bounds on the sound speed with the low- ℓ likelihoods (CMB11) and without (~~LowLike~~).

- For a constant sound speed, $\delta_1 = 0$, the main degeneracy that limits the bound on c_s is $\varepsilon_2 \simeq -2\varepsilon_1$. Improving the constraint on $\alpha_s \simeq 4\varepsilon_1^2$ then limits the size of ε_1 , leading to a stronger bound on c_s . In particular, if $r \simeq 0.13$, we need $\varepsilon_1 < 0.017$ to get $c_s > 0.47$. To attain this level of sensitivity requires $\sigma_{\alpha_s} < 4 \times 10^{-4}$. Notice that, for such an experiment, the precise values of n_s and r could in principle exclude $c_s = 1$. Bounding the running at the level $\sigma_{\alpha_s} \sim 10^{-3}$ is within reach of near-term CMB observations [43] and future galaxy surveys [44, 45], although some improvement beyond these projections may be necessary to reach the critical value $(c_s)_* = 0.47$.
- For a varying sound speed, $\delta_1 \neq 0$, any constraints on n_s and α_s can be satisfied if $\delta_1 \simeq -2\varepsilon_1$, provided ε_2 remains small. As a result, the bound on c_s arises only from the analytic properties of $r(\varepsilon_1, \varepsilon_2, c_s)$ along the degeneracy. However, large values of r are only consistent with small c_s if $\delta_1 \simeq -2\varepsilon_1$ is relatively large. For example, to achieve $r \simeq 0.13$ with $c_s \simeq 0.14$ requires $\varepsilon_1 \simeq 0.1$ and $\delta_1 \simeq -0.2$. These values of δ_1 are sufficiently large that they may be detectable. Specifically, the value of c_s relevant for the constraint from r is set at $\ell \sim 10^2$ [16], whereas most of the modes in a given data set are at $\ell \sim \ell_{\max} \gg 10^2$. Using $\delta_1 = \dot{c}_s/(c_s H)$, we have

$$c_s(\ell_{\max}) = c_s(\ell = 10^2) \left(\frac{\ell_{\max}}{10^2} \right)^{\delta_1} = 0.07 \times \frac{c_s(\ell = 10^2)}{0.14} \left(\frac{\ell_{\max}}{3000} \right)^{\delta_1(\ell=10^2)/(-0.2)} . \quad (3.13)$$

The value of $c_s = 0.07$ at $\ell_{\max} \simeq 3000$, translates into $f_{\text{NL}}^{\text{equil}} \simeq -52$, which may be reachable with future LSS surveys and/or high-resolution CMB satellites.⁹ Similarly, reaching the threshold $c_s(\ell = 10^2) = 0.47$, under the same assumptions, requires sensitivity to $f_{\text{NL}}^{\text{equil}} \simeq 5(f_{\text{NL}}^{\text{equil}})_* = -4.4$.

4 A theoretical threshold

Thresholds are an essential part of physics and targets for experimental searches. The LHC was built in part to explore the unitarity threshold derived from the Fermi scale, $\Lambda \simeq 4\pi(\sqrt{2}G_F)^{-1/2} \sim \text{TeV}$, including the possibility (ultimately disfavored by the discovery of a weakly coupled Higgs boson) of strong dynamics, such as ‘technicolor’ or ‘composite Higgs’ models, playing a role in electroweak symmetry breaking. In cosmology, on the other hand, current observational evidence has not yet ruled out inflationary backgrounds driven by strongly coupled dynamics. In this section, we derive a (perturbative) unitarity

⁹Projections for these experiments (e.g. [46]) give $\sigma(f_{\text{NL}}^{\text{equil}}) \sim 10$. Although these forecasts are for constant c_s , most of the information in the surveys is at high ℓ (or k) and therefore insensitive to the variation of c_s . In particular, given that the number of modes scales like ℓ^2 (or k^3), the scale dependence in (3.13) is much weaker than the scale dependence of the signal-to-noise.

threshold for the sound speed, $(c_s)_\star = 0.47$, and the associated equilateral non-Gaussianity, $(f_{\text{NL}}^{\text{equil}})_\star = -0.93$, within the framework of the EFT of inflation [2]. In section 5, we will review well-motivated scenarios that may produce a signal above the order-one threshold on equilateral non-Gaussianity.

4.1 Sound waves in the early universe

The theory of the Goldstone boson $\pi(t, \mathbf{x})$ was introduced in [2, 47]. In the so-called *decoupling limit*, where the mixing with gravity is ignored, π parameterizes the breaking of time translations in a quasi-de Sitter background with a slowly evolving Hubble parameter $H(t)$. The Goldstone boson also characterizes adiabatic density perturbations, which by definition, can be set to zero through the time diffeomorphism $t \rightarrow t - \pi(t, \mathbf{x})$. In this *unitarity gauge*, the fluctuations are described by the curvature perturbation ζ , which at leading order is related to π via $\zeta = -H\pi$. In this framework, standard slow-roll models are described by the Lagrangian: $\mathcal{L}_\pi^{(0)} = M_{\text{pl}}^2 \dot{H}(t) (\partial_\mu \pi)^2$. Different slow-roll models are distinguished only by the function $H(t)$. The first natural deformation of slow-roll inflation can then be parameterized as [2]

$$S_\pi = \int d^4x a^3 \left\{ M_{\text{pl}}^2 \dot{H}(t) (\partial_\mu \pi)^2 + 2M_2^4(t) \left[\dot{\pi}^2 + \dot{\pi}^3 - \frac{\dot{\pi}(\partial_i \pi)^2}{a^2} \right] \right\}, \quad (4.1)$$

where we have only shown terms up to cubic order in π . Turning on finite M_2 induces a sound speed for the Goldstone modes

$$c_s^2 \equiv \frac{M_{\text{pl}}^2 \dot{H}}{M_{\text{pl}}^2 \dot{H} - 2M_2^4}. \quad (4.2)$$

Crucially, a small sound speed (large M_2) simultaneously affects different orders in the Goldstone action. In particular, the contributions to the quadratic, $\dot{\pi}^2$, and cubic, $\dot{\pi}(\partial_i \pi)^2$, Lagrangian are not independent, but are related by a non-linearly realized symmetry [2].¹⁰ This feature allows power spectrum measurements to constrain the strength of certain interactions in the theory.

4.2 Energy scales

Two of the important energy/frequency scales in the theory of the Goldstone bosons are the *freeze-out scale* (H), and the *symmetry breaking scale*¹¹ (f_π); see figure 9. The freeze-out of the Goldstone dynamics happens universally when the physical frequency of a mode (ω) drops below the Hubble scale.¹² (The corresponding momentum depends on the dispersion relation.) The symmetry breaking scale, on the other hand, is model-dependent. For models with a sound speed, one finds [6]

$$f_\pi^4 \equiv 2M_{\text{pl}}^2 |\dot{H}| c_s. \quad (4.3)$$

¹⁰The size of $\dot{\pi}^3$ is not fixed by the value of c_s , and can be adjusted by another term in the EFT; see section 5.2.

¹¹We use f_π to denote the symmetry breaking scale to highlight the analogy with the pion decay constant in chiral perturbation theory.

¹²Here, we assume that the parameters that control the dynamics of the Goldstone boson vary slowly with time or essentially remain constant. When the time variation becomes important, modes do not necessarily freeze-out at $\omega \simeq H$. Nevertheless, the following discussion on the role of the relevant energy scales remains valid [48–51].

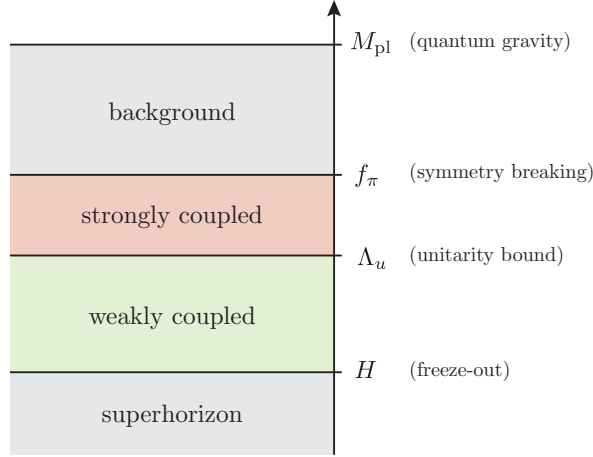


Figure 9. Graphical illustration of the relevant energy scales in the EFT of inflation. Whether Λ_u is above or below f_π is an important qualitative distinction.

In the slow-roll limit, $c_s \rightarrow 1$, this definition reduces to $f_\pi^2 \rightarrow \dot{\phi}$. At energies below f_π a description in terms of the Goldstone boson is appropriate. Assuming vacuum fluctuations, the ratio of H and f_π is fixed by the amplitude of curvature perturbations [52]

$$\Delta_\zeta^2 = \frac{1}{4\pi^2} \left(\frac{H}{f_\pi} \right)^{2\delta_\pi+2}, \quad (4.4)$$

where δ_π is the scaling dimension of π . For sound speed models, we have $\delta_\pi = 1$ and the amplitude of curvature perturbations is enhanced by an inverse power of the sound speed, $\Delta_\zeta^2 \propto c_s^{-1}$. The observed value for (4.4), $\Delta_\zeta^2 = 2.2 \times 10^{-9}$, determines

$$f_\pi \simeq 60 H. \quad (4.5)$$

Armed with a theory of the perturbations, a natural question is at what scale this theory becomes strongly coupled.¹³ To answer this one re-writes (4.1) as [2, 6]

$$S_\pi = \int dt d^3\tilde{x} a^3 \left\{ -\frac{1}{2} (\tilde{\partial}_\mu \pi_c)^2 - \frac{1}{2\Lambda^2} \left[\frac{\dot{\pi}_c (\tilde{\partial}_i \pi_c)^2}{a^2} - c_s^2 \dot{\pi}_c^3 \right] \right\}. \quad (4.7)$$

Here, we have rescaled the spatial coordinates, $\tilde{x}^i = c_s^{-1} x^i$, to reinstall (fake) Lorentz invariance in the quadratic part of the action, defined $\pi_c \equiv f_\pi^2 \pi$, and introduced the scale

$$\Lambda^4 \equiv 2M_{\text{pl}}^2 |\dot{H}| \frac{c_s^5}{(1 - c_s^2)^2}. \quad (4.8)$$

¹³If we assume that a fundamental scalar field ϕ produces the inflationary background, then a sound speed for the fluctuations arises from higher-derivative corrections to the kinetic term [4]

$$\mathcal{L} = -\frac{1}{2} (\partial\phi)^2 + \frac{(\partial\phi)^4}{M^4} + \dots. \quad (4.6)$$

For $\dot{\phi} < M^2$ this is a perturbative correction to the slow-roll dynamics, and as such $c_s \simeq 1$. On the other hand, for $\dot{\phi} > M^2$, the higher-order terms that are hidden in the ellipses in (4.6) become relevant and the background is non-perturbative. The computation of $f_{\text{NL}}^{\text{equil}}$ is not reliable from the term displayed in (4.6) alone, and the relevant scale for the validity of the expansion must be sought within the theory of the perturbations. At the same time, the theory of the Goldstone boson allows for more complicated situations where fluctuations are not necessarily associated with a fundamental scalar field.

We see that the contribution of the operator $\dot{\pi}_c^3$ is suppressed for $c_s \ll 1$ and the dominant non-linearity comes from $\dot{\pi}_c(\tilde{\partial}_i\pi_c)^2$. For this reason, we will, for now, restrict to the effects of $\dot{\pi}_c(\tilde{\partial}_i\pi_c)^2$. We will return to a discussion of the operator $\dot{\pi}_c^3$ in section 5.2.

The high degree of Gaussianity of the observed CMB anisotropies implies that the Goldstone modes are weakly coupled at freeze-out, $\omega \simeq H$. This requires that the scale Λ is above the Hubble scale, so that higher-dimension operators are suppressed by powers of $H/\Lambda < 1$ at horizon crossing. However, extrapolating to higher frequencies the effects of higher-order operators becomes more relevant, until the perturbative description breaks down at the *strong coupling scale* (Λ). Perturbative unitarity is violated at the nearby *unitarity scale*:¹⁴

$$\Lambda_u^4 \equiv \frac{24\pi}{5} (1 - c_s^2) \Lambda^4. \quad (4.10)$$

Whether this happens above or below the symmetry breaking scale is an important qualitative distinction. The theory is close to a weakly coupled slow-roll background if $\Lambda_u > f_\pi$, while for $\Lambda_u < f_\pi$ the Goldstone action involves a completion below the symmetry breaking scale, plausibly¹⁵ signaling strongly coupled dynamics.

4.3 A critical sound speed

The critical value $\Lambda_u = f_\pi$ is an interesting target for future experiments. Using (4.10) and (4.3), this threshold is related to a critical sound speed

$$\frac{24\pi}{5} \left(\frac{2M_{\text{pl}}^2 |\dot{H}| (c_s)_\star^5}{1 - (c_s)_\star^2} \right) = 2M_{\text{pl}}^2 |\dot{H}| (c_s)_\star \quad \Rightarrow \quad (c_s)_\star = 0.47. \quad (4.11)$$

The threshold value $(c_s)_\star$ is still far from the Planck-only bound, $c_s > 0.02$.¹⁶ On the other hand, our analysis in this paper shows that any detection of primordial B-modes with $r > 0.01$ improves this constraint on c_s . Moreover, the bound we obtained using the BICEP2 data (without foreground subtraction), $c_s \geq 0.25$, is already remarkably close to the unitarity threshold.

To relate (4.11) to a threshold for the non-Gaussianity associated with the interaction $\dot{\pi}(\partial_i\pi)^2$ in (4.1), we use

$$\left(f_{\text{NL}}^{\dot{\pi}(\partial_i\pi)^2} \right)_\star = -\frac{85}{324} \frac{1 - (c_s)_\star^2}{(c_s)_\star^2} \simeq -0.93. \quad (4.12)$$

¹⁴This scale is obtained from imposing partial wave unitarity in the quartic interaction: $\frac{1}{8(1-c_s^2)} \frac{1}{\Lambda^4} \frac{(\tilde{\partial}_i\pi_c)^4}{a^4} \subset \mathcal{L}_\pi^{(4)}$. The computation may be found in appendix E of [6] after fixing a minor discrepancy with the numerical coefficient in eq. (E.7), which should read:

$$a_0 = \frac{5}{3} \times \frac{1}{16\pi} \frac{(1 - c_s^2) \omega^4}{2M_{\text{pl}}^2 |\dot{H}| c_s^5} = \frac{5}{48\pi} \frac{\omega^4}{(1 - c_s^2) \Lambda^4}. \quad (4.9)$$

The condition $\omega^4 < \Lambda_u^4$ follows from $a_0 < 1/2$, as required by unitarity.

¹⁵Weakly coupled completions of models with $\Lambda_u < f_\pi$ were studied in [6, 53–57]. These models always involve new physics parametrically below the scale Λ_u and may be observationally distinguishable from their strongly coupled counterparts.

¹⁶The bound on c_s derived from the interaction $\dot{\pi}(\partial_i\pi)^2$ alone is $c_s > 0.04$ [7]. The limit $c_s > 0.02$ includes a marginalization over an additional parameter M_3 (see section 5.2 for the definition of M_3). Since M_3 does not contribute to the two-point function, the limits presented in this paper naturally include such a marginalization and should therefore be compared to the weaker limit.

This experimental benchmark is still two orders of magnitude below the limit from the Planck bispectrum measurement [7], $f_{\text{NL}}^{\dot{\pi}(\partial_i\pi)^2} = 8 \pm 146$. In the next section, we will describe the different types of physics involved in non-Gaussianity at or above the threshold given by (4.12).

5 Physics above threshold

It is apparent that there is a large window between the current bounds on non-Gaussianity and the threshold, $|f_{\text{NL}}^{\text{equil}}|_* \simeq \mathcal{O}(1)$. This offers a wonderful opportunity to explore non-canonical models through measurements of non-Gaussianity.¹⁷ In this section, we survey a range of well-motivated theories that can produce a signal *at* or *above* the threshold. (See also [12–14, 17, 21, 22], where combinations of the following scenarios have been reviewed.)

5.1 One-scale models & strong coupling

The standard example of strong coupling realized in nature is QCD. At low energies, chiral symmetry is broken and the dynamics of QCD is described by the associated (pseudo-)Goldstone bosons — the pions — whose interactions are described by the chiral Lagrangian. The order parameter of the symmetry breaking is the vacuum condensate $\langle q\bar{q} \rangle \sim f_\pi^3$, with f_π playing the role of the symmetry breaking scale. Higher-order interactions of the pions are also controlled by f_π , making chiral perturbation theory a ‘one-scale’ theory. (For example, the chiral Lagrangian includes the quartic term $(12\pi^2 f_\pi^4)^{-1} (\partial_\mu \pi)^4$. Cubic interactions are forbidden by parity.)

In cosmology, it may still be the case that interactions in the theory of the Goldstone boson in the EFT of inflation are controlled by a single scale, like for the pion(s) of QCD (without parity). Since inflation breaks time translations the analogy with chiral symmetry breaking would correspond to a (time-dependent) vacuum condensate with $\partial_t \langle q\bar{q} \rangle \sim f_\pi^4$. In such a scenario the cubic interaction $\dot{\pi}_c (\partial_i \pi_c)^2 / f_\pi^2$ (up to an order-one numerical coefficient) would produce an order-one level of equilateral non-Gaussianity. This view is permitted by the data which still allows for a sound speed near the critical value, corresponding to one-scale models with $\Lambda_u \sim f_\pi$. The threshold $(c_s)_* = 0.47$ is therefore a milestone for strongly coupled inflationary dynamics controlled by a single scale.¹⁸ As we have shown in this paper, a confirmed detection of primordial tensor modes by the BICEP2 collaboration and improved bounds on the scalar spectrum by future experiments have the potential to probe these theories.

¹⁷One may worry that a threshold at $|f_{\text{NL}}^{\text{equil}}|_* \simeq \mathcal{O}(1)$ could be experimentally elusive. Indeed, future CMB observations will not be able to reach this threshold [58]. On the other hand, the large number of modes that LSS observations in principle offer may in the future allow us to access smaller values of $f_{\text{NL}}^{\text{equil}}$. This will require a detailed understanding of secondary non-linearities in structure formation [59–63].

¹⁸One may be tempted to take the position that $0.47 \simeq 1$, and no further scrutiny is necessary. However, this attitude ignores well-known examples where the discrepancy between a weak coupling computation and its strong coupling counterpart is a small factor. For example, a computation of the entropy density s in the quark-gluon plasma in terms of an effective ‘quasi-particle’ description fits lattice data qualitatively well, i.e. $s \simeq 0.85s_0$, with s_0 the entropy density of a non-interacting plasma, see e.g. [64]. This small difference, $(s_0 - s)/s_0 \simeq 0.15$, has been taken as evidence of weak coupling at the relevant temperatures. However, a computation in a somewhat related theory ($\mathcal{N} = 4$ SYM) revealed that $s = f(\lambda)s_0$, with $f(\lambda) \rightarrow 3/4$ in the strong coupling limit $\lambda \gg 1$ [65]. The proximity to the lattice computation suggests the quark-gluon plasma may be strongly coupled. In this example the function $f(\lambda)$ changes merely by a factor of $1/4$ when one extrapolates between weak and strong coupling.

5.2 Stable hierarchies

Going beyond sound speed models, the action for the Goldstone boson in the EFT of inflation includes other cubic interactions [2]

$$S_\pi = \int d^4x a^3 \left\{ -M_{\text{pl}}^2 \dot{H} \left[\dot{\pi}^2 - \frac{(\partial_i \pi)^2}{a^2} \right] + 2M_2^4 \left[\dot{\pi}^2 - \frac{\dot{\pi}(\partial_i \pi)^2}{a^2} \right] + \left(2M_2^4 - \frac{4}{3}M_3^4 \right) \dot{\pi}^3 \right. \\ \left. - \frac{H\bar{M}_1^3}{2} \frac{(\partial_i \pi)^2}{a^2} - \frac{\bar{M}_2^2 + \bar{M}_3^2}{2} \frac{(\partial_i^2 \pi)^2}{a^4} + \frac{\bar{M}_1^4}{2} \frac{\partial_i^2 \pi (\partial_j \pi)^2}{a^4} + \dots \right\}. \quad (5.1)$$

As it has been pointed out in the literature, a small sound speed generates a large radiative correction to the parameter M_3 , namely $M_3^4 \sim M_2^4/c_s^2$ [23, 66]. However, the converse is not the case, and various terms in (5.1) may be large without inducing a significant modification of the sound speed. We briefly enumerate these possibilities:

- **M_3 .** It is technically natural to have large M_3 and small M_2 [10, 17, 22]. To see this, let us set $M_2 = \bar{M}_n = 0$ and only turn on finite M_3 :

$$\mathcal{L}_\pi = M_{\text{pl}}^2 \dot{H} (\partial_\mu \pi)^2 - \frac{4}{3} M_3^4 \left(\dot{\pi}^3 - \frac{3}{2} \dot{\pi}^2 (\partial_\mu \pi)^2 + \dots \right). \quad (5.2)$$

We should be concerned that a non-zero value of M_2 will be generated through loops

$$\text{Diagram 1} + \text{Diagram 2} + \dots = 2M_2^4 \left(\dot{\pi}^2 - \dot{\pi} (\partial_\mu \pi)^2 + \dots \right). \quad (5.3)$$

However, running the loops all the way to the cutoff $\Lambda_{\text{UV}} = (M_{\text{pl}}^2 |\dot{H}|)^{3/4}/M_3^2$, we only get $M_2^4 \sim M_{\text{pl}}^2 \dot{H}$, even if $M_3^4 \gg M_{\text{pl}}^2 |\dot{H}|$. A large value of M_3 is therefore consistent with having only small deviations from $c_s = 1$. This allows for large equilateral non-Gaussianity that is not constrained by measurements of the power spectrum.

- **\bar{M}_n .** We can avoid the constraints on c_s by allowing for large values of \bar{M}_2 and \bar{M}_3 , which changes the dispersion relation to be that of ghost inflation [67], namely $\omega = k^2/\rho$. In this case, the relationship between r and ε_1 is broken, and hence the tilt of the power spectrum does not constrain c_s . The phenomenology of this setup was studied in detail by [23], which we briefly summarize below.

The dispersion relation in this model takes the form:

$$\omega^2 = c_s^2 k^2 + \frac{k^4}{\rho^2}, \quad \text{where} \quad \rho^2 \equiv \frac{2M_{\text{pl}}^2 \dot{H}}{c_s^2 (\bar{M}_2^2 + \bar{M}_3^2)} \xrightarrow{c_s \ll 1} \frac{4M_2^4}{\bar{M}_2^2 + \bar{M}_3^2}. \quad (5.4)$$

The c_s^2 -term is negligible at horizon crossing provided that $c_s^2 \ll H/\rho$. Notice that this constraint is not trivially satisfied, because we must require that the modes propagate subluminally at horizon crossing, which implies $2(H/\rho)^{1/2} < 1$. (The group velocity, given by $c_g \equiv d\omega/dk = 2k/\rho$, must be bounded: $c_g \leq 1$.)

The non-linear dispersion enhances operators with a large number of spatial derivatives, and one can determine that [23]

$$f_{\text{NL}}^{\dot{\pi}(\partial_i\pi)^2} = 0.25 \frac{\rho}{H}, \quad (5.5)$$

$$f_{\text{NL}}^{\partial_i^2\pi(\partial_j\pi)^2} = 0.13 \frac{\bar{M}_1^3}{H(\bar{M}_2^2 + \bar{M}_3^2)}, \quad (5.6)$$

for the two leading terms in the cubic action (5.1). It is interesting to note that self-consistency ($c_g < 1$) implies that

$$f_{\text{NL}}^{\dot{\pi}(\partial_i\pi)^2} > 1. \quad (5.7)$$

In other words, for this mechanism to be in operation we necessarily require non-Gaussianity above the threshold.

5.3 Additional degrees of freedom

The EFT of inflation has been extended to include couplings of π to multiple (Goldstone-like) light fields [8]. One of the main signatures of these type of models is non-Gaussianity of the local type. The local shape, however, is highly constrained by the Planck data, $f_{\text{NL}}^{\text{loc}} = 2.7 \pm 5.8$ [7]. There is nonetheless a class of multi-field models, namely *single-clock models*, in which surfaces of constant density are effectively controlled by a single degree of freedom such that the consistency condition (implying a vanishing squeezed limit) is satisfied, thus bypassing the Planck constraint on local non-Gaussianity [68]. On the other hand, equilateral non-Gaussianity may still be generated even when $c_s = 1$. We briefly describe two examples:

- *Dissipative dynamics.* Various mechanisms have been proposed [25, 69] to produce the observed density perturbations through non-vacuum fluctuations in models with dissipative dynamics. Building upon ideas originally developed in the study of black hole absorption [70, 71], an EFT approach to generic classes of dissipative models was put forward in [9, 68]. The main idea is to systematically couple the Goldstone boson to a dissipative sector described by composite operators \mathcal{O} , whose correlation functions are constrained by symmetries.¹⁹ These couplings can induce an effective friction term, $\gamma\dot{\pi}$, in the equations of motion. For strongly dissipative systems, with $\gamma \gg H$, the power spectrum then depends on new parameters and is dominated by non-vacuum fluctuations.

Non-Gaussianity in dissipative models was studied in [9, 68], where it was shown that the non-linearly realized symmetries require the presence of a non-linear term in the dynamics correlated with the leading order dissipation, $\gamma\dot{\pi} \rightarrow -\frac{1}{2}\gamma(\partial_i\pi)^2$, among other contributions. This term induces equilateral non-Gaussianity even when $c_s \simeq 1$, i.e. $f_{\text{NL}}^{\text{equil}} \simeq -\gamma/4H$ [9]. The bound from Planck on the equilateral shape then translates (roughly) into $\gamma \lesssim 10^2 H$. In addition, the bispectrum has another peak at the ‘folded’ configuration: $k_1 = k_2 = \frac{1}{2}k_3$. This contribution is correlated with the size of the equilateral shape. For $\gamma \simeq \mathcal{O}(10)H$, one finds

$$f_{\text{NL}}^{\text{fold}} \simeq -\frac{1}{2}f_{\text{NL}}^{\text{equil}}, \quad (5.8)$$

¹⁹A similar approach was used to study dissipation in fluids [72].

with a negligible squeezed limit (dissipation erases memory quickly) [68]. These features can be considered a smoking-gun for dissipative models, and a probe of the quantum nature of the primordial fluctuations [14].²⁰

- *Quasi-single-field inflation.* Equilateral non-Gaussianity also arises when the Goldstone mode couples to extra scalar fields with masses of order the Hubble scale, as in models of ‘quasi-single-field inflation’ [24]. Since massive fields decay outside the horizon, their interactions are localized at horizon crossing. This effect suppresses the interactions of modes with different wavelengths and produces an approximately equilateral shape for the bispectrum. The squeezed limit of the bispectrum and the collapsed limit of the trispectrum depend on the mass of the field, allowing it to be distinguished from the equilateral shape produced by higher-derivative interactions in single-field inflation, or by non-adiabatic evolution as described above. Models of quasi-single-field inflation arise naturally in inflationary theories with spontaneously broken supersymmetry [10, 73, 74].

Quasi-single-field models can be generalized to include couplings to arbitrary additional sectors parameterized by composite operators \mathcal{O} , as in the case of dissipative dynamics [9]. In [75] the additional sector was taken to be a conformal field theory. For this special case, the predictions are qualitatively similar to those of quasi-single-field inflation. The phenomenology of more general cases is an open problem, but the expectation is that equilateral non-Gaussianity will be common.

6 Conclusions and outlook

Inflationary models are characterized by a small number of important energy scales: the freeze-out scale, H , the scale of time-variation of the background, f_π , and the scale(s) characterizing scalar self-interactions, Λ . Cosmological observables in the scalar sector are only sensitive to the ratios f_π/H and Λ/H . For example, the amplitude of curvature perturbations fixes f_π/H , while non-Gaussianity constrains Λ/H . Measuring tensors, on the other hand, determines H/M_{Pl} and hence normalizes all energy scales relative to the Planck scale. Famously, a large value of f_π in Planck units, then correlates with a super-Planckian field excursion [52, 76] and an enhanced UV sensitivity of the inflationary model [15, 58]. Whether Λ is above or below f_π is an important qualitative distinction. To decide this question, in general, requires precision measurements of the non-Gaussianity corresponding to the interactions associated with the scale Λ . However, in models with a sound speed c_s , the leading cubic interaction is related to the quadratic action by a non-linearly realized symmetry, and both f_π and Λ depend on c_s . In this special case, power spectrum measurements can inform us about the interacting theory by constraining the value of c_s .

In this paper we have shown that consistency between an observable tensor amplitude and the near scale-invariance of the scalar spectrum enforces a lower bound on the sound speed. We found, analytically, that any observable tensor signal ($r > 0.01$) constrains the sound speed above the bound obtained from the Planck temperature data alone. A joint analysis of the Planck temperature data and the BICEP2 B-mode measurement (without

²⁰ Another way to obtain non-trivial squeezed limits is to assume a state other than the vacuum as an initial condition, i.e. an *excited* state. The most popular choice involves Bogoliubov states. This possibility, however, is severely constrained by the data [49]. In dissipative models, on the other hand, the excited (semi-classical) state is reached dynamically.

foreground subtraction) also supports another related conclusion: a future detection of primordial tensor modes with $r \gtrsim 0.1$ would lead to a bound on c_s which is an order of magnitude stronger than current bounds derived from non-Gaussianity. In such scenario, the resulting bound on c_s would be tantalizingly close to a natural threshold,

$$(c_s)_\star = 0.47, \quad (6.1)$$

namely the value of the sound speed for which the unitarity scale associated with the leading cubic interaction coincides with f_π (see also [12, 13]). This corresponds to a related threshold of equilateral non-Gaussianity with amplitude

$$|f_{\text{NL}}^{\pi(\partial_i\pi)^2}|_\star = 0.93. \quad (6.2)$$

Values of c_s below (6.1) signal non-perturbative physics [5] or non-trivial dynamics [6]. Although finding $c_s > (c_s)_\star$ in future observations would disfavor the simplest deformation of slow-roll inflation, it would not constrain significantly other possible (well motivated) extensions. This is the case because the EFT of inflation allows for stable hierarchies in which large interactions are possible without generating significant radiative corrections to the quadratic action. These scenarios are not constrained by power spectrum measurements and are only probed by the bispectrum and/or higher n -point functions. Similar thresholds of non-Gaussianity, $|f_{\text{NL}}^{\text{equil}}|_\star \sim \mathcal{O}(1)$, exist in these cases. Moreover, the presence of additional light fields may also produce a signal above this threshold. This strongly motivates an experimental effort to improve the bounds on equilateral non-Gaussianity to the order-on level.

Canonical slow-roll models may be thought of as the ‘Higgs mechanism’ of inflation, in the sense that it is a weakly coupled completion of the EFT of inflation involving a fundamental scalar field. However, current observations have not yet ruled out the possibility that the UV completion of the EFT of inflation is characterized by strongly coupled dynamics, such as an analogue of compositeness in the electroweak sector, higher-derivative interactions or additional degrees of freedom. Unlike in particle physics, where higher energies can be explored by building more powerful accelerators, in cosmology we do not have direct access to the unitarity thresholds because observations are always performed at a fixed energy (corresponding to the freeze-out frequency during inflation). This means that sensitivity to physics at (or above) threshold can only be achieved by increasing the precision of the measurements. In this paper we have shown that observations of primordial B-modes together with the study of non-Gaussianity, in particular of the equilateral type, offer a unique window into the physics of the early universe and the nature of inflation. Finding that the primordial perturbations remain Gaussian beyond $|f_{\text{NL}}^{\text{equil}}|_\star \simeq \mathcal{O}(1)$, or that measurements of the power spectra lead to $c_s > (c_s)_\star = 0.47$, would provide strong evidence for canonical slow-roll models. Conversely, a detection of non-Gaussianity above the threshold would open the road to physics beyond the slow-roll paradigm.

Acknowledgments

We thank Valentin Assassi, Carlo Contaldi, Zhiqi Huang, Yi Wang, and Matias Zaldarriaga for helpful discussions. D.B. thanks Yi Wang for his patient assistance with the numerical analysis. D.B. gratefully acknowledges support from a Starting Grant of the European Research Council (ERC STG grant 279617). R.A.P. was partially supported by the German Science Foundation (DFG) within the Collaborative Research Center (SFB) 676 ‘Particles, Strings and the Early Universe’, and by NSF grant AST-0807444 and DOE grant DE-FG02-90ER40542.

A Resumming large logarithms

In inflationary models with a non-trivial sound speed, scalars and tensors do not freeze out simultaneously and the tensor-to-scalar ratio depends non-trivially on the evolution of the Hubble parameter,

$$r = 16\varepsilon_1 c_s^{1+2\bar{\varepsilon}_1} \left[1 - \mathcal{C}\varepsilon_2 + (2 - \mathcal{C})\delta_1 + \dots \right], \quad (\text{A.1})$$

where

$$\bar{\varepsilon}_1 \equiv -\frac{\int \varepsilon_1(N) dN}{\Delta N}, \quad \Delta N \equiv N_t - N_s, \quad (\text{A.2})$$

and the bracket in (A.1) contains higher-order terms in the slow-roll expansion. In general, the auxiliary parameter $\bar{\varepsilon}_1$ will be a function of the Hubble flow parameters ε_n . We will assume a hierarchical structure and truncate the expansion at $\varepsilon_{n \geq 4}(N) \simeq 0$. This implies $\varepsilon_3 \approx \text{const.}$ and

$$\varepsilon_2(N) = \varepsilon_2(N_s) e^{\varepsilon_3(N-N_s)}. \quad (\text{A.3})$$

Our goal in this appendix is to obtain $\bar{\varepsilon}_1$ at leading order in $\varepsilon_{1,2}$ but to all orders in $\varepsilon_2 \ln c_s$. In this way, (A.1) becomes a resummation of the leading logarithms of the form $\bar{\varepsilon}_1 \ln c_s$, with $\bar{\varepsilon}_1$ incorporating all orders in $\varepsilon_2 \ln c_s$.²¹ For completeness, we will also discuss how to include $\varepsilon_3 \ln c_s$ corrections. In what follows, we will suppress explicit reference to N_s , i.e. all the slow-roll parameters are understood to be evaluated at the freeze-out of the scalar modes.

A.1 $\varepsilon_2 \ln c_s$

As we did throughout the main text, we consider $\varepsilon_3 \ll \varepsilon_{1,2}$ but keep all orders in $\varepsilon_{1,2}$. The solution for ε_1 then becomes

$$\varepsilon_1(N) = \varepsilon_1 e^{\varepsilon_2(N-N_s) + \frac{1}{2}\varepsilon_2\varepsilon_3(N-N_s)^2 + \dots}, \quad (\text{A.4})$$

such that (A.2) turns into

$$\bar{\varepsilon}_1 = \varepsilon_1 \left[\frac{e^{\varepsilon_2 \Delta N} - 1}{\varepsilon_2 \Delta N} - e^{\varepsilon_2 \Delta N} \frac{\varepsilon_3 \Delta N}{2(\varepsilon_2 \Delta N)^2} \left\{ 2(1 - e^{-\varepsilon_2 \Delta N}) - 2\varepsilon_2 \Delta N + (\varepsilon_2 \Delta N)^2 \right\} \right]. \quad (\text{A.5})$$

Let us first look at the $\varepsilon_{1,2}$ dependence, as given by the leading term in (A.5). Then, using (2.14), we have

$$\bar{\varepsilon}_1 = \varepsilon_1 \cdot \frac{e^{\varepsilon_2 \Delta N(\bar{\varepsilon}_1)} - 1}{\varepsilon_2 \Delta N(\bar{\varepsilon}_1)} \rightarrow \bar{\varepsilon}_1 = \varepsilon_1 \cdot \frac{(1 - c_s^{\varepsilon_2})}{\varepsilon_1 (1 - c_s^{\varepsilon_2}) - c_s^{\varepsilon_2} \varepsilon_2 \ln c_s}. \quad (\text{A.6})$$

The resummation of the leading logarithms only requires

$$\bar{\varepsilon}_1 \approx \varepsilon_1 \cdot \frac{1}{\varepsilon_2 \ln c_s} (1 - c_s^{-\varepsilon_2}) + \dots. \quad (\text{A.7})$$

Substituting this into (A.1), we get

$$r = 16\varepsilon_1 c_s^{1+2\varepsilon_1 \cdot (1 - c_s^{-\varepsilon_2}) / (\varepsilon_2 \ln c_s)} \left[1 - \mathcal{C}\varepsilon_2 + 2(1 - \mathcal{C})\delta_1 + \dots \right]. \quad (\text{A.8})$$

This is the result that we have implemented in our analysis in section 3.

²¹Notice that for $\bar{\varepsilon}_1 \simeq \varepsilon_1 \approx \text{const.}$ the expression in (A.1) resums all the $\varepsilon_1 \ln c_s$ terms.

A.2 $\varepsilon_3 \ln c_s$

The expression (A.8) is valid to zeroth order in an expansion in ε_3 . To compute the first-order correction we may proceed perturbatively. Moreover, it is instructive to determine when the $\varepsilon_{1,2} \ln c_s$ terms become important. Keeping only the leading logarithms, the expression in (A.5) turns into

$$\bar{\varepsilon}_1 = \varepsilon_1 \left[1 - \frac{1}{2} \varepsilon_2 \ln c_s \left(1 + \frac{1}{3} \varepsilon_3 \ln c_s \right) + \mathcal{O}(\varepsilon_n^2 \ln c_s) \right]. \quad (\text{A.9})$$

For $|\ln c_s| \leq 3.91$ (Planck, 95% CL), the $\ln c_s$ terms produce an enhancement, but only become non-perturbative for somewhat large values of the slow-roll parameters: a perturbative treatment is justified provided $|\varepsilon_2| \ll 0.51$ and $|\varepsilon_3| \ll 0.75$. Values at the boundary of the perturbative regime are highly constrained by the running of the spectral index, $|\alpha_s| \lesssim 2 \times 10^{-2}$. The numerical analysis in section 3.3, using $|\varepsilon_{1,2}| \lesssim 0.5$ and $|\varepsilon_3| \lesssim 0.1$, shows that it is self consistent to work within a slow-roll expansion, and to ignore the ε_3 correction in (A.9) on the boundaries of the 95% confidence contours.

References

- [1] PLANCK collaboration, P.A.R. Ade et al., *Planck 2013 results. XXII. Constraints on inflation*, *Astron. Astrophys.* **571** (2014) A22 [[arXiv:1303.5082](#)] [[INSPIRE](#)].
- [2] C. Cheung, P. Creminelli, A.L. Fitzpatrick, J. Kaplan and L. Senatore, *The Effective Field Theory of Inflation*, *JHEP* **03** (2008) 014 [[arXiv:0709.0293](#)] [[INSPIRE](#)].
- [3] C. Armendariz-Picon, T. Damour and V.F. Mukhanov, *k-inflation*, *Phys. Lett. B* **458** (1999) 209 [[hep-th/9904075](#)] [[INSPIRE](#)].
- [4] P. Creminelli, *On non-Gaussianities in single-field inflation*, *JCAP* **10** (2003) 003 [[astro-ph/0306122](#)] [[INSPIRE](#)].
- [5] E. Silverstein and D. Tong, *Scalar speed limits and cosmology: Acceleration from D-acceleration*, *Phys. Rev. D* **70** (2004) 103505 [[hep-th/0310221](#)] [[INSPIRE](#)].
- [6] D. Baumann and D. Green, *Equilateral Non-Gaussianity and New Physics on the Horizon*, *JCAP* **09** (2011) 014 [[arXiv:1102.5343](#)] [[INSPIRE](#)].
- [7] PLANCK collaboration, P.A.R. Ade et al., *Planck 2013 Results. XXIV. Constraints on primordial non-Gaussianity*, *Astron. Astrophys.* **571** (2014) A24 [[arXiv:1303.5084](#)] [[INSPIRE](#)].
- [8] L. Senatore and M. Zaldarriaga, *The Effective Field Theory of Multifield Inflation*, *JHEP* **04** (2012) 024 [[arXiv:1009.2093](#)] [[INSPIRE](#)].
- [9] D. Lopez Nacir, R.A. Porto, L. Senatore and M. Zaldarriaga, *Dissipative effects in the Effective Field Theory of Inflation*, *JHEP* **01** (2012) 075 [[arXiv:1109.4192](#)] [[INSPIRE](#)].
- [10] D. Baumann and D. Green, *Signatures of Supersymmetry from the Early Universe*, *Phys. Rev. D* **85** (2012) 103520 [[arXiv:1109.0292](#)] [[INSPIRE](#)].
- [11] K.N. Abazajian et al., *Inflation Physics from the Cosmic Microwave Background and Large Scale Structure*, *Astropart. Phys.* **63** (2015) 55 [[arXiv:1309.5381](#)] [[INSPIRE](#)].
- [12] L. Senatore, *Some Theoretical Aspects of B-modes and Inflation*, <http://burkeinstitute.caltech.edu/workshops>.
- [13] D. Green, *Why Search for Primordial Non-Gaussianity?*, <http://www.cita.utoronto.ca/~drgreen/Talks/Snowmass.pdf>.

- [14] R.A. Porto, *Probing the (Quantum) Nature of the Primordial Seed*, <http://online.kitp.ucsb.edu/online/primocosmo13/porto/>.
- [15] D. Baumann and L. McAllister, *Inflation and String Theory*, [arXiv:1404.2601](#) [INSPIRE].
- [16] BICEP2 collaboration, P.A.R. Ade et al., *Detection of B-Mode Polarization at Degree Angular Scales by BICEP2*, *Phys. Rev. Lett.* **112** (2014) 241101 [[arXiv:1403.3985](#)] [INSPIRE].
- [17] P. Creminelli, D. López Nacir, M. Simonović, G. Trevisan and M. Zaldarriaga, ϕ^2 or not ϕ^2 : *Testing the Simplest Inflationary Potential*, *Phys. Rev. Lett.* **112** (2014) 241303 [[arXiv:1404.1065](#)] [INSPIRE].
- [18] PLANCK collaboration, R. Adam et al., *Planck intermediate results. XXX. The angular power spectrum of polarized dust emission at intermediate and high Galactic latitudes*, [arXiv:1409.5738](#) [INSPIRE].
- [19] R. Flauger, J.C. Hill and D.N. Spergel, *Toward an Understanding of Foreground Emission in the BICEP2 Region*, *JCAP* **08** (2014) 039 [[arXiv:1405.7351](#)] [INSPIRE].
- [20] M.J. Mortonson and U. Seljak, *A joint analysis of Planck and BICEP2 B modes including dust polarization uncertainty*, *JCAP* **10** (2014) 035 [[arXiv:1405.5857](#)] [INSPIRE].
- [21] M. Zaldarriaga, *Early Universe Phenomenology after BICEP2*, <http://pirsa.org/14040125/>.
- [22] G. D'Amico and M. Kleban, *Non-Gaussianity after BICEP2*, *Phys. Rev. Lett.* **113** (2014) 081301 [[arXiv:1404.6478](#)] [INSPIRE].
- [23] L. Senatore, K.M. Smith and M. Zaldarriaga, *Non-Gaussianities in Single Field Inflation and their Optimal Limits from the WMAP 5-year Data*, *JCAP* **01** (2010) 028 [[arXiv:0905.3746](#)] [INSPIRE].
- [24] X. Chen and Y. Wang, *Quasi-Single Field Inflation and Non-Gaussianities*, *JCAP* **04** (2010) 027 [[arXiv:0911.3380](#)] [INSPIRE].
- [25] A. Berera, *Warm inflation*, *Phys. Rev. Lett.* **75** (1995) 3218 [[astro-ph/9509049](#)] [INSPIRE].
- [26] L. Senatore, E. Silverstein and M. Zaldarriaga, *New Sources of Gravitational Waves during Inflation*, *JCAP* **08** (2014) 016 [[arXiv:1109.0542](#)] [INSPIRE].
- [27] J.L. Cook and L. Sorbo, *Particle production during inflation and gravitational waves detectable by ground-based interferometers*, *Phys. Rev. D* **85** (2012) 023534 [Erratum *ibid.* **D 86** (2012) 069901] [[arXiv:1109.0022](#)] [INSPIRE].
- [28] J. Martin, C. Ringeval, R. Trotta and V. Vennin, *Compatibility of Planck and BICEP2 in the Light of Inflation*, *Phys. Rev. D* **90** (2014) 063501 [[arXiv:1405.7272](#)] [INSPIRE].
- [29] X. Chen, M.-x. Huang, S. Kachru and G. Shiu, *Observational signatures and non-Gaussianities of general single field inflation*, *JCAP* **01** (2007) 002 [[hep-th/0605045](#)] [INSPIRE].
- [30] L. Lorenz, J. Martin and C. Ringeval, *K-inflationary Power Spectra in the Uniform Approximation*, *Phys. Rev. D* **78** (2008) 083513 [[arXiv:0807.3037](#)] [INSPIRE].
- [31] N. Agarwal and R. Bean, *Cosmological constraints on general, single field inflation*, *Phys. Rev. D* **79** (2009) 023503 [[arXiv:0809.2798](#)] [INSPIRE].
- [32] B.A. Powell, K. Tzirakis and W.H. Kinney, *Tensors, non-Gaussianities and the future of potential reconstruction*, *JCAP* **04** (2009) 019 [[arXiv:0812.1797](#)] [INSPIRE].
- [33] J. Martin, C. Ringeval and V. Vennin, *K-inflationary Power Spectra at Second Order*, *JCAP* **06** (2013) 021 [[arXiv:1303.2120](#)] [INSPIRE].
- [34] PLANCK collaboration, P.A.R. Ade et al., *Planck 2013 results. XV. CMB power spectra and likelihood*, *Astron. Astrophys.* **571** (2014) A15 [[arXiv:1303.5075](#)] [INSPIRE].
- [35] PLANCK collaboration, P.A.R. Ade et al., *Planck 2013 results. XVI. Cosmological parameters*, *Astron. Astrophys.* **571** (2014) A16 [[arXiv:1303.5076](#)] [INSPIRE].

- [36] BICEP2 collaboration, P.A.R. Ade et al., *BICEP2 II: Experiment and Three-Year Data Set*, *Astrophys. J.* **792** (2014) 62 [[arXiv:1403.4302](#)] [[INSPIRE](#)].
- [37] A. Lewis and S. Bridle, *Cosmological parameters from CMB and other data: a Monte Carlo approach*, *Phys. Rev. D* **66** (2002) 103511 [[astro-ph/0205436](#)] [[INSPIRE](#)].
- [38] A. Lewis, A. Challinor and A. Lasenby, *Efficient computation of CMB anisotropies in closed FRW models*, *Astrophys. J.* **538** (2000) 473 [[astro-ph/9911177](#)] [[INSPIRE](#)].
- [39] A. Gelman and D.B. Rubin, *Inference from Iterative Simulation Using Multiple Sequences*, *Statist. Sci.* **7** (1992) 457.
- [40] 2dFGRS collaboration, S. Cole et al., *The 2dF Galaxy Redshift Survey: Power-spectrum analysis of the final dataset and cosmological implications*, *Mon. Not. Roy. Astron. Soc.* **362** (2005) 505 [[astro-ph/0501174](#)] [[INSPIRE](#)].
- [41] SDSS collaboration, D.J. Eisenstein et al., *Detection of the baryon acoustic peak in the large-scale correlation function of SDSS luminous red galaxies*, *Astrophys. J.* **633** (2005) 560 [[astro-ph/0501171](#)] [[INSPIRE](#)].
- [42] S. Dodelson, *How much can we learn about the physics of inflation?*, *Phys. Rev. Lett.* **112** (2014) 191301 [[arXiv:1403.6310](#)] [[INSPIRE](#)].
- [43] W.L.K. Wu et al., *A Guide to Designing Future Ground-based Cosmic Microwave Background Experiments*, *Astrophys. J.* **788** (2014) 138 [[arXiv:1402.4108](#)] [[INSPIRE](#)].
- [44] M. Takada, E. Komatsu and T. Futamase, *Cosmology with high-redshift galaxy survey: neutrino mass and inflation*, *Phys. Rev. D* **73** (2006) 083520 [[astro-ph/0512374](#)] [[INSPIRE](#)].
- [45] P. Adshead, R. Easther, J. Pritchard and A. Loeb, *Inflation and the Scale Dependent Spectral Index: Prospects and Strategies*, *JCAP* **02** (2011) 021 [[arXiv:1007.3748](#)] [[INSPIRE](#)].
- [46] E. Sefusatti, M. Liguori, A.P.S. Yadav, M.G. Jackson and E. Pajer, *Constraining Running Non-Gaussianity*, *JCAP* **12** (2009) 022 [[arXiv:0906.0232](#)] [[INSPIRE](#)].
- [47] P. Creminelli, M.A. Luty, A. Nicolis and L. Senatore, *Starting the Universe: Stable Violation of the Null Energy Condition and Non-standard Cosmologies*, *JHEP* **12** (2006) 080 [[hep-th/0606090](#)] [[INSPIRE](#)].
- [48] S.R. Behbahani, A. Dymarsky, M. Mirbabayi and L. Senatore, *(Small) Resonant non-Gaussianities: Signatures of a Discrete Shift Symmetry in the Effective Field Theory of Inflation*, *JCAP* **12** (2012) 036 [[arXiv:1111.3373](#)] [[INSPIRE](#)].
- [49] R. Flauger, D. Green and R.A. Porto, *On squeezed limits in single-field inflation. Part I*, *JCAP* **08** (2013) 032 [[arXiv:1303.1430](#)] [[INSPIRE](#)].
- [50] P. Adshead and W. Hu, *Bounds on nonadiabatic evolution in single-field inflation*, *Phys. Rev. D* **89** (2014) 083531 [[arXiv:1402.1677](#)] [[INSPIRE](#)].
- [51] D. Cannone, N. Bartolo and S. Matarrese, *Perturbative Unitarity of Inflationary Models with Features*, *Phys. Rev. D* **89** (2014) 127301 [[arXiv:1402.2258](#)] [[INSPIRE](#)].
- [52] D. Baumann and D. Green, *A Field Range Bound for General Single-Field Inflation*, *JCAP* **05** (2012) 017 [[arXiv:1111.3040](#)] [[INSPIRE](#)].
- [53] A.J. Tolley and M. Wyman, *The Gelaton Scenario: Equilateral non-Gaussianity from multi-field dynamics*, *Phys. Rev. D* **81** (2010) 043502 [[arXiv:0910.1853](#)] [[INSPIRE](#)].
- [54] S. Cremonini, Z. Lalak and K. Turzynski, *Strongly Coupled Perturbations in Two-Field Inflationary Models*, *JCAP* **03** (2011) 016 [[arXiv:1010.3021](#)] [[INSPIRE](#)].
- [55] A. Achucarro, J.-O. Gong, S. Hardeman, G.A. Palma and S.P. Patil, *Features of heavy physics in the CMB power spectrum*, *JCAP* **01** (2011) 030 [[arXiv:1010.3693](#)] [[INSPIRE](#)].
- [56] S. Cespedes, V. Atal and G.A. Palma, *On the importance of heavy fields during inflation*, *JCAP* **05** (2012) 008 [[arXiv:1201.4848](#)] [[INSPIRE](#)].

- [57] C.P. Burgess, M.W. Horbatsch and S. Patil, *Inflating in a Trough: Single-Field Effective Theory from Multiple-Field Curved Valleys*, *JHEP* **01** (2013) 133 [[arXiv:1209.5701](#)] [[INSPIRE](#)].
- [58] CMBPOL STUDY TEAM collaboration, D. Baumann et al., *CMBPol Mission Concept Study: Probing Inflation with CMB Polarization*, *AIP Conf. Proc.* **1141** (2009) 10 [[arXiv:0811.3919](#)] [[INSPIRE](#)].
- [59] J.J.M. Carrasco, M.P. Hertzberg and L. Senatore, *The Effective Field Theory of Cosmological Large Scale Structures*, *JHEP* **09** (2012) 082 [[arXiv:1206.2926](#)] [[INSPIRE](#)].
- [60] J.J.M. Carrasco, S. Foreman, D. Green and L. Senatore, *The Effective Field Theory of Large Scale Structures at Two Loops*, *JCAP* **07** (2014) 057 [[arXiv:1310.0464](#)] [[INSPIRE](#)].
- [61] R.A. Porto, L. Senatore and M. Zaldarriaga, *The Lagrangian-space Effective Field Theory of Large Scale Structures*, *JCAP* **05** (2014) 022 [[arXiv:1311.2168](#)] [[INSPIRE](#)].
- [62] R.E. Angulo, S. Foreman, M. Schmittfull and L. Senatore, *The One-Loop Matter Bispectrum in the Effective Field Theory of Large Scale Structures*, [arXiv:1406.4143](#) [[INSPIRE](#)].
- [63] T. Baldauf, L. Mercolli, M. Mirbabayi and E. Pajer, *The Bispectrum in the Effective Field Theory of Large Scale Structure*, [arXiv:1406.4135](#) [[INSPIRE](#)].
- [64] A. Adams, L.D. Carr, T. Schäfer, P. Steinberg and J.E. Thomas, *Strongly Correlated Quantum Fluids: Ultracold Quantum Gases, Quantum Chromodynamic Plasmas and Holographic Duality*, *New J. Phys.* **14** (2012) 115009 [[arXiv:1205.5180](#)] [[INSPIRE](#)].
- [65] S.S. Gubser, I.R. Klebanov and A.A. Tseytlin, *Coupling constant dependence in the thermodynamics of $N = 4$ supersymmetric Yang-Mills theory*, *Nucl. Phys. B* **534** (1998) 202 [[hep-th/9805156](#)] [[INSPIRE](#)].
- [66] L. Senatore and M. Zaldarriaga, *A Naturally Large Four-Point Function in Single Field Inflation*, *JCAP* **01** (2011) 003 [[arXiv:1004.1201](#)] [[INSPIRE](#)].
- [67] N. Arkani-Hamed, P. Creminelli, S. Mukohyama and M. Zaldarriaga, *Ghost inflation*, *JCAP* **04** (2004) 001 [[hep-th/0312100](#)] [[INSPIRE](#)].
- [68] D. Lopez Nacir, R.A. Porto and M. Zaldarriaga, *The consistency condition for the three-point function in dissipative single-clock inflation*, *JCAP* **09** (2012) 004 [[arXiv:1206.7083](#)] [[INSPIRE](#)].
- [69] D. Green, B. Horn, L. Senatore and E. Silverstein, *Trapped Inflation*, *Phys. Rev. D* **80** (2009) 063533 [[arXiv:0902.1006](#)] [[INSPIRE](#)].
- [70] W.D. Goldberger and I.Z. Rothstein, *Dissipative effects in the worldline approach to black hole dynamics*, *Phys. Rev. D* **73** (2006) 104030 [[hep-th/0511133](#)] [[INSPIRE](#)].
- [71] R.A. Porto, *Absorption effects due to spin in the worldline approach to black hole dynamics*, *Phys. Rev. D* **77** (2008) 064026 [[arXiv:0710.5150](#)] [[INSPIRE](#)].
- [72] S. Endlich, A. Nicolis, R.A. Porto and J. Wang, *Dissipation in the effective field theory for hydrodynamics: First order effects*, *Phys. Rev. D* **88** (2013) 105001 [[arXiv:1211.6461](#)] [[INSPIRE](#)].
- [73] V. Assassi, D. Baumann, D. Green and L. McAllister, *Planck-Suppressed Operators*, *JCAP* **01** (2014) 033 [[arXiv:1304.5226](#)] [[INSPIRE](#)].
- [74] N. Craig and D. Green, *Testing Split Supersymmetry with Inflation*, *JHEP* **07** (2014) 102 [[arXiv:1403.7193](#)] [[INSPIRE](#)].
- [75] D. Green, M. Lewandowski, L. Senatore, E. Silverstein and M. Zaldarriaga, *Anomalous Dimensions and Non-Gaussianity*, *JHEP* **10** (2013) 171 [[arXiv:1301.2630](#)] [[INSPIRE](#)].
- [76] D.H. Lyth, *What would we learn by detecting a gravitational wave signal in the cosmic microwave background anisotropy?*, *Phys. Rev. Lett.* **78** (1997) 1861 [[hep-ph/9606387](#)] [[INSPIRE](#)].