



Unifying Inflation with the Axion, Dark Matter, Baryogenesis, and the Seesaw Mechanism

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A minimal extension of the standard model (SM) with a single new mass scale and providing a complete and consistent picture of particle physics and cosmology up to the Planck scale is presented. We add to the SM three right-handed SM-singlet neutrinos, a new vectorlike color triplet fermion, and a complex SM-singlet scalar σ that stabilizes the Higgs potential and whose vacuum expectation value at $\sim 10^{11}$ GeV breaks lepton number and a Peccei-Quinn symmetry simultaneously. Primordial inflation is produced by a combination of σ (nonminimally coupled to the scalar curvature) and the SM Higgs boson. Baryogenesis proceeds via thermal leptogenesis. At low energies, the model reduces to the SM, augmented by seesaw-generated neutrino masses, plus the axion, which solves the strong CP problem and accounts for the dark matter in the Universe. The model predicts a minimum value of the tensor-to-scalar ratio $r \simeq 0.004$, running of the scalar spectral index $\alpha \simeq -7 \times 10^{-4}$, the axion mass $m_A \sim 100 \mu\text{eV}$, and cosmic axion background radiation corresponding to an increase of the effective number of relativistic neutrinos of ~ 0.03 . It can be probed decisively by the next generation of cosmic microwave background and axion dark matter experiments.

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Introduction.—The standard model of particle physics (SM) describes with exquisite precision the interactions of all known elementary particles. In spite of intensive searches, no significant deviation from the SM has been detected in collider or other particle physics experiments [1]. However, several long-standing problems indicate that new physics beyond the SM is needed to achieve a complete description of nature. First of all, there is overwhelming evidence, ranging from the cosmic microwave background (CMB) to the shapes of the rotation curves of spiral galaxies, that nearly 26% of the Universe is made of yet unidentified dark matter (DM) [2]. Moreover, the SM cannot generate the primordial inflation needed to solve the horizon and flatness problems of the Universe as well as to explain the statistically isotropic, Gaussian, and nearly scale-invariant fluctuations of the CMB [3]. The SM also lacks enough CP violation to explain why the Universe contains a larger fraction of baryonic matter than of antimatter. Aside from these three problems at the interface between particle physics and cosmology, the SM suffers from a variety of intrinsic naturalness issues. In particular, the neutrino masses are disparagingly smaller than any physics scale in the SM, and, similarly, the strong CP problem states that the θ parameter of quantum chromodynamics (QCD) is constrained from measurements of the neutron electric dipole moment [4,5] to lie below an unexpectedly small value: $|\theta| \lesssim 10^{-10}$.

In this Letter, we show that these problems may be intertwined in a remarkably simple way, with a solution

pointing to a unique new physics scale around 10^{11} GeV. The SM extension we consider consists just of a Kim-Shifman-Vainshtein-Zakharov (KSVZ)-like axion model [6,7] and three right-handed (RH) heavy SM-singlet neutrinos [8]. This extra matter content was recently proposed in Ref. [11], where it was emphasized that, in addition to solving the strong CP problem, providing a good dark matter candidate (the axion), and explaining the origin of the small SM neutrino masses (through an induced seesaw mechanism) and the baryon asymmetry of the Universe (via thermal leptogenesis), it could also stabilize the effective potential of the SM at high energies thanks to a threshold mechanism [12,13]. This extension also leads to successful primordial inflation by using the modulus of the KSVZ SM-singlet scalar field [14]. Adding a cosmological constant to account for the present acceleration of the Universe, this standard model axion seesaw Higgs portal inflation (SMASH) model offers a self-contained description of particle physics from the electroweak scale to the Planck scale and of cosmology from inflation until today. Although some parts of our SMASH model have been considered separately [14–26], a model incorporating all of them simultaneously had not been proposed until now. Remarkably, SMASH can accommodate the constraints from cosmological observations and Higgs stability, successfully reheat the Universe, provide the correct dark matter abundance, and explain the origin of the baryon asymmetry. In this Letter, we present the most important aspects and predictions of SMASH. Further details are given in Ref. [27].

The SMASH model.—We extend the SM with a new complex singlet scalar field σ and a Dirac fermion \mathcal{Q} , which can be split in two Weyl fermions Q and \tilde{Q} in the 3 and $\bar{3}$ representations of $SU(3)_c$ with charges $-1/3$ and $1/3$ under $U(1)_Y$. This ensures that \mathcal{Q} can coannihilate and decay into SM quarks, thereby evading possible overabundance problems [28,29]. We also add three RH fermions N_i . The model is endowed with a new Peccei-Quinn (PQ) global $U(1)$ symmetry [30], which also plays the role of lepton number in our case. Using left-handed Weyl spinors, we denote by q_i , u_i , and d_i the SM quark doublet and the conjugates of the right-handed quarks of each generation $i = 1, 2, 3$, respectively, and by L_i and E_i the corresponding lepton doublet and the conjugate of the right-handed lepton. Denoting the Higgs boson by H , the charges under the PQ symmetry are $q(1/2)$, $u(-1/2)$, $d(-1/2)$, $L(1/2)$, $N(-1/2)$, $E(-1/2)$, $Q(-1/2)$, $\tilde{Q}(-1/2)$, $\sigma(1)$, and $H(0)$. The most general Yukawa couplings involving the new fields are $\mathcal{L} \supset -[F_{ij}L_i\epsilon HN_j + \frac{1}{2}Y_{ij}\sigma N_i N_j + y\tilde{Q}\sigma Q + z_i\sigma Q d_i + \text{H.c.}]$, where ϵ is the two-component antisymmetric symbol. The Yukawa couplings F and Y realize the seesaw mechanism once σ acquires a vacuum expectation value (VEV) $\langle\sigma\rangle = v_\sigma/\sqrt{2}$, giving a neutrino mass matrix of the form $m_\nu = -FY^{-1}F^T v^2/(\sqrt{2}v_\sigma)$, with $v = 246$ GeV. The strong CP problem is solved as in the standard KSVZ scenario, with the role of the axion decay constant f_A played by $v_\sigma = f_A$. Because of non-perturbative QCD effects, the angular part of $\sigma = (\rho + v_\sigma)\exp(iA/f_A)/\sqrt{2}$, the axion field A [31,32], gains a potential with an absolute minimum at $A = 0$. At energies above the QCD scale, the axion-gluon coupling is $\mathcal{L} \supset -(\alpha_s/8\pi)(A/f_A)G\tilde{G}$, solving the strong CP problem when $\langle A \rangle$ relaxes to zero [33]. The latest lattice computation of the axion mass gives $m_A = (57.2 \pm 0.7)(10^{11} \text{ GeV}/f_A) \mu\text{eV}$ [34].

Inflation.—Given the symmetries of SMASH, the most general renormalizable tree-level potential is

$$V(H, \sigma) = \lambda_H \left(H^\dagger H - \frac{v^2}{2} \right)^2 + \lambda_\sigma \left(|\sigma|^2 - \frac{v_\sigma^2}{2} \right)^2 + 2\lambda_{H\sigma} \left(H^\dagger H - \frac{v^2}{2} \right) \left(|\sigma|^2 - \frac{v_\sigma^2}{2} \right). \quad (1)$$

In the unitary gauge, there are two scalar fields that could drive inflation: h , the neutral component of the Higgs doublet $H' = (0, h)/\sqrt{2}$, and the modulus of the new singlet, $\rho^2 = 2|\sigma|^2$. In the context of the SM, it was proposed in Ref. [18] that h could be the inflaton if it is nonminimally coupled to the scalar curvature R through a term $\mathcal{L} \supset -\sqrt{-g}\xi_H H^\dagger H R$ [35], with $\xi_H \sim 10^4$. Such a large value of ξ_H is required by the constraint $\xi_H \sim 10^5 \sqrt{\lambda_H}$ to fit the amplitude of primordial fluctuations, and it implies that perturbative unitarity breaks down at the scale

$\Lambda_U = M_P/\xi_H \ll M_P$ [36,37], where $M_P = 1/\sqrt{8\pi G}$ is the reduced Planck mass. This raises a serious difficulty for Higgs inflation, which requires Planckian values of h and an energy density of the order of Λ_U^2 . Since new physics is expected at or below Λ_U to restore unitarity, the predictivity of Higgs inflation is lost, because the effect of this new physics on inflation is undetermined. This issue affects some completions of the SM such as the νMSM [38,39] and the model proposed in Ref. [23]. Instead, inflation in SMASH is mostly driven by ρ , with a non-minimal coupling $\mathcal{L} \supset -\sqrt{-g}\xi_\sigma \sigma^* \sigma R$, where $\xi_\sigma \lesssim 1$ ensures that the scale of perturbative unitarity breaking is at M_P (provided that also $\xi_H \lesssim 1$). Neglecting ξ_H [40], predictive slow-roll inflation in SMASH can happen along two directions in field space: the ρ direction for $\lambda_{H\sigma} > 0$ and the line $h/\rho = \sqrt{-\lambda_{H\sigma}/\lambda_H}$ for $\lambda_{H\sigma} < 0$. We call them hidden scalar inflation (HSI) and Higgs-hidden scalar inflation (HHSI), respectively. In both cases, inflation can be described in the Einstein frame by a single canonically normalized field χ with potential

$$\tilde{V}(\chi) = \frac{\lambda}{4} \rho(\chi)^4 \left(1 + \xi_\sigma \frac{\rho(\chi)^2}{M_P^2} \right)^{-2}, \quad (2)$$

where λ stands for λ_σ in HSI and for $\tilde{\lambda}_\sigma = \lambda_\sigma - \lambda_{H\sigma}^2/\lambda_H$ in HHSI. The field χ is the solution of $\Omega^2 d\chi/d\rho \approx (b\Omega^2 + 6\xi_\sigma^2 \rho^2/M_P^2)^{1/2}$, $\Omega \approx 1 + \xi_\sigma \rho^2/M_P^2$ being the Weyl transformation into the Einstein frame; and $b = 1$ (for HSI) or $b = 1 + |\lambda_{H\sigma}/\lambda_H|$ (for HHSI). The small value of $|\lambda_{H\sigma}|$ required for stability (see below) typically means that $b \sim 1$ in HHSI, which makes impossible distinguishing in practice between HSI and HHSI from the inflationary potential. However, even a small Higgs component in the inflaton is relevant for reheating, as we will later discuss. The predictions of the potential (2) in the case $\lambda = \lambda_\sigma$ (or $b \rightarrow 1$ in HHSI) for the tensor-to-scalar ratio r vs the scalar spectral index n_s are shown in Fig. 1 for various values of ξ_σ .

In SMASH, the equation of state (EOS) of the Universe after inflation is $w = 1/3$ (like radiation) uninterruptedly until the standard epoch of matter-radiation equality is reached; see the reheating section below. This allows us to compute the number of e -folds of inflation, $N(k)$, for any comoving scale k , matching precisely the predictions for the inflationary spectrum with the observations of the CMB [42]. This determines the thick line in Fig. 1 as the SMASH prediction for $r(n_s)$ and $N(k_0)$ at the fiducial scale $k_0 \equiv 0.002 \text{ Mpc}^{-1}$, which we use through the Letter for all the primordial inflationary parameters. The prediction spans $N \sim (59, 62)$, depending on n_s , and its width (~ 1 e -fold) quantifies the small uncertainty on the transient regime from the end of inflation to radiation domination.

Note that the condition $\xi_\sigma \lesssim 1$ corresponds to $r \gtrsim 0.004$, which is within the planned sensitivities of PIXIE [43],

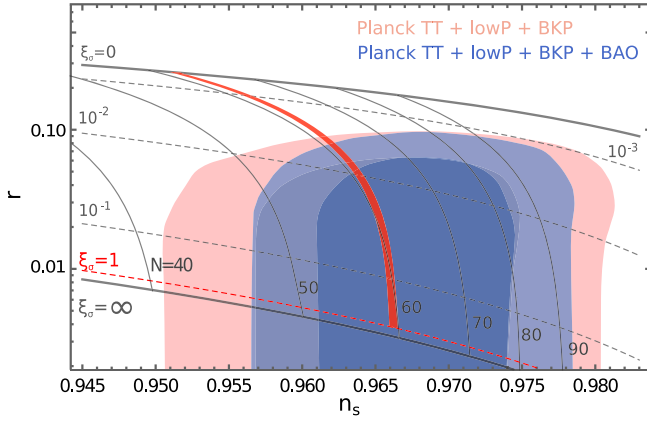


FIG. 1. The tensor-to-scalar ratio r vs the scalar spectral index n_s at $k_0 = 0.002 \text{ Mpc}^{-1}$ for the inflationary potential (2), assuming $|\lambda_{H\sigma}| \ll \lambda_H$. We show lines of constant ξ_σ (dashed) and a constant number of e -folds from the time the scale $k_0 = 0.002 \text{ Mpc}^{-1}$ exits the horizon to the end of inflation (thin solid). In SMASH, the EOS of the Universe is $w = 1/3$ immediately after inflation, which allows us to predict N (thick red line). Colored regions show observational constraints at 68% and 95% C.L. from Ref. [41].

LiteBird [44], CMB-S4 [45], and CoRE+ (which will measure r with an error of $\Delta r \sim 4 \times 10^{-4}$). The joint constraints of the Planck satellite and the BICEP/Keck array [2,46] give $r < 0.07$ at 95% C.L., corresponding in SMASH to $\xi_\sigma \gtrsim 6 \times 10^{-3}$. Taking into account the former constraints, the spectral index n_s at $k = k_0$ lies in the interval (0.962, 0.966), and its running $\alpha = dn_s/d \ln k$ lies in the range $(-7, -6) \times 10^{-4}$, which may be probed, e.g., by future observations of the 21 cm emission line of hydrogen [47]. Since inflation is effectively single-field slow-roll, non-Gaussian features are suppressed by $\sim (1 - n_s)$ [48,49]. These values of the primordial parameters are perfectly compatible with the latest CMB data, and the amount of inflation that is produced solves the horizon and flatness problems. Given the current bounds on r and n_s and the fact that fitting the amplitude of primordial scalar fluctuations requires $\xi_\sigma \sim 10^5 \sqrt{\lambda}$, fully consistent (and predictive) inflation in SMASH occurs if $5 \times 10^{-13} \lesssim \lambda \lesssim 5 \times 10^{-10}$.

Stability.—For the measured central values of the Higgs and top quark masses [1], the Higgs quartic coupling of the SM becomes negative at $h = \Lambda_I \sim 10^{11} \text{ GeV}$ [50]. If no new physics changes this behavior, Higgs inflation is not viable, since it requires a positive potential at Planckian field values. Moreover, this instability is a problem even if another field drives inflation. This is because scalars that are light compared to the Hubble scale \mathcal{H} acquire fluctuations of the order of $\sim \mathcal{H}/2\pi$. These can make the Higgs field move into the instability region of the potential, which would contradict the present electro-weak vacuum [51]. Remarkably, the Higgs portal term $\propto \lambda_{H\sigma}$ in (1) allows stability of the SMASH potential via

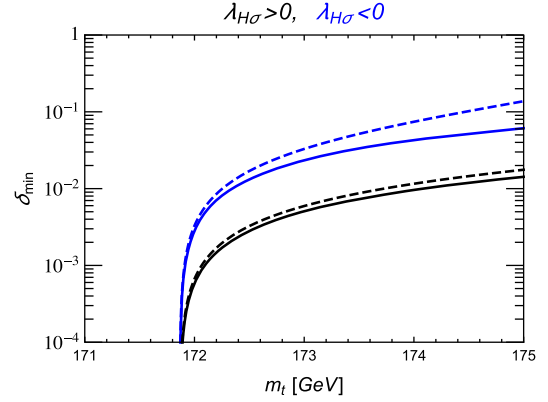


FIG. 2. Minimum value of the threshold correction to the Higgs quartic coupling, $\delta = \lambda_{H\sigma}^2/\lambda_\sigma$, for stable SMASH potentials at RG scales $\mu = m_\rho$ (solid curves) and $\mu = 30M_P$ (dashed curves), for $\lambda_{H\sigma} > 0$ (black curves) and $\lambda_{H\sigma} < 0$ (blue curves).

the threshold-stabilization mechanism of Refs. [12,13], which relies on a nontrivial matching with the SM potential at low energies. The matched Higgs quartic in the SM is $\tilde{\lambda}_H \equiv \lambda_H - \delta$, where the threshold correction is $\delta \equiv \lambda_{H\sigma}^2/\lambda_\sigma$. Even if the running of $\tilde{\lambda}_H$ in the SM makes it negative, the actual Higgs quartic coupling in the UV theory, λ_H , can remain positive provided that δ is large enough. A more detailed analysis [52] shows that, for $\lambda_{H\sigma} > 0$, absolute stability requires [27]

$$\begin{aligned} \tilde{\lambda}_H, \tilde{\lambda}_\sigma &> 0 \quad \text{for } h < \sqrt{2}\Lambda_h, \\ \lambda_H, \lambda_\sigma &> 0 \quad \text{for } h > \sqrt{2}\Lambda_h, \end{aligned} \quad (3)$$

where $\Lambda_h^2 \equiv \lambda_{H\sigma} v_\sigma^2/\lambda_H$ and all the couplings run with the beta functions of SMASH, not the SM. The scale $\sqrt{2}\Lambda_h$ arises as the divide between large and small field values of h , for which v_σ cannot be neglected and the quadratic interactions are relevant, as can be seen from (1). Instead, for $\lambda_{H\sigma} < 0$, the stability condition is just $\tilde{\lambda}_H, \tilde{\lambda}_\sigma > 0$, for all h . The Higgs direction is the one most prone to be destabilized (from top loops), and the potential must remain positive beyond the $h \sim M_P$ values needed for inflation. A one-loop analysis shows that a value of δ above $10^{-3} - 10^{-1}$ (depending on the top mass; see Fig. 2) ensures stability up to $h_{30} \approx 30M_P$ for a Higgs mass of 125.09 GeV. Finally, in SMASH, instabilities could also originate in the direction of ρ due to quantum corrections from N_i and Q, \tilde{Q} . Stability in this direction requires $6y^4 + \sum Y_{ii}^4 \lesssim 16\pi^2 \lambda_\sigma / \log(h_{30}/\sqrt{2}\lambda_\sigma v_\sigma)$ [53].

Reheating.—SMASH provides a complete model of cosmology for which the evolution after inflation can be calculated. The PQ symmetry is spontaneously broken during inflation by the large evolving value of ρ . Slow-roll inflation ends at $\rho_{\text{end}} \sim \mathcal{O}(M_P)$, where the effect of ξ_σ is negligible. Since $\rho_{\text{end}} \gg v_\sigma$, the inflaton starts to undergo Hubble-damped oscillations in a quartic potential.

The first oscillations of the inflaton constitute a phase of so-called preheating [54], during which fluctuations of σ in the direction orthogonal to the inflaton increase exponentially. The postinflationary background can be understood as a homogeneous condensate of particles with energy given by the oscillation frequency $\omega(t) \sim \sqrt{\lambda}\rho_{\text{end}}/a(t)$, where $a(t)$ is the scale factor of the Universe and t denotes cosmic time [55]. In SMASH, λ is the weakest coupling, and thus SM particles coupled to the inflaton have effective masses $\propto \rho(t)$, which are much larger than $\omega(t)$ except when $\rho(t) \sim 0$. Higgs particles and electroweak bosons could, in principle, be produced by parametric resonance [56] at these crossings, but they either have large self-interactions or decay very efficiently into SM fermions. In contrast, the effective mass of σ excitations is $\sim \sqrt{\lambda}\rho(t) \sim \omega(t)$, which allows them to grow by parametric resonance. The growth of fluctuations of a complex inflaton field in a quartic potential was studied analytically in Ref. [56] and numerically in Ref. [57]. Our own numerical simulations [27] corroborate their results. After the first ~ 14 oscillations after inflation, the fluctuations of σ become as large as the inflaton amplitude $\langle |\sigma|^2 \rangle \sim \rho_{\text{end}}^2/a^2$, so the PQ symmetry is nonthermally restored. Only if v_σ were larger than $\sim 10^{-2}M_P$ would the field ρ get trapped around its minimum $\rho = v_\sigma$ before the nonthermal restoration can occur. However, such high values of v_σ are ruled out by CMB axion isocurvature constraints [14,58].

Aside from these common features, reheating progresses differently for HSI and HHSI. The reason is that the small Higgs component of the inflaton in HHSI (which is lacking in HSI) accelerates in that case the production of SM particles. We will now discuss the two cases separately.

*Reheating for HSI ($\lambda_{H\sigma} > 0$).—*During preheating, Higgs bosons are nonresonantly produced during inflaton crossings because of the large value of the Higgs self-coupling [59] as well as the fast decay of Higgs bosons into tops and gauge bosons. When the PQ symmetry is nonthermally restored, the induced Higgs mass $\sqrt{\lambda_{H\sigma}}\sqrt{\langle |\sigma|^2 \rangle}$ stabilizes around a large value $\sqrt{\lambda_{H\sigma}}\rho_{\text{end}}/a(t) \gg \omega(t)$, thus blocking Higgs production. Efficient reheating has to wait until the spontaneously symmetry breaking (SSB) of the PQ symmetry, i.e., when $\langle |\sigma|^2 \rangle$ becomes $\sim v_\sigma^2$. We have simulated numerically the phase transition, finding that the energy initially stored in σ fluctuations becomes equipartitioned into axions and ρ particles. The latter can soon decay into Higgs bosons and reheat the SM sector. The corresponding reheating temperature is $T_R \sim v_{11}\lambda_{10}^{3/8}\delta_3^{-1/8}10^7$ GeV, where we introduce SMASH benchmark values $v_{11} = v_\sigma/(10^{11} \text{ GeV})$, $\lambda_{10} = 10^{10}\lambda_\sigma$, and $\delta_3 = \delta/0.03$ [60]. The accompanying axions are relativistic and remain decoupled from such a low-temperature SM thermal bath [61]. They contribute to the late Universe expansion rate as extra (relativistic) neutrino species. We estimate $\Delta N_\nu^{\text{eff}} \sim 0.96(\lambda_{10}/\delta_3 v_{11})^{1/6}$ above the SM value $N_\nu^{\text{eff}}(\text{SM}) = 3.046$

[62]. Current CMB and baryon acoustic oscillation data give $N_\nu^{\text{eff}} = 3.04 \pm 0.18$ at 68% C.L. [2], disfavoring HSI.

*Reheating for HHSI ($\lambda_{H\sigma} < 0$).—*As in HSI, the direct production of Higgs excitations stops when the PQ symmetry is nonthermally restored. However, the Higgs component of the inflaton continues to oscillate around $h \sim 0$ so that W and Z gauge bosons can still be produced during crossings. The fast decay of W, Z into light fermions when h moves away from zero prevents their exponential accumulation but makes the comoving energy in light fermions increase. When light particles thermalize, a population of W, Z bosons is created by the thermal bath during crossings (when their mass is below the temperature) and decays when their mass grows with h . This mechanism enhances the drain of energy from the inflaton to the SM bath. Using Boltzmann equations with thermal and nonthermal sources and accounting for the energy loss of the background fields, we have calculated numerically the reheating temperature, finding $T_R \sim \mathcal{O}(10^{10} \text{ GeV})$ for the values of λ and δ satisfying the requirements for inflation and stability.

The critical temperature for the PQ phase transition is $T_c \simeq 2\sqrt{6}\lambda_\sigma v_\sigma / \sqrt{8(\lambda_\sigma + \lambda_{H\sigma}) + \sum_i Y_{ii}^2 + 6y^2}$ [27]. For SMASH benchmark values $|\lambda_{H\sigma}| \gg \lambda_\sigma$, and requiring the previous stability bound on the Yukawa couplings of the new fermions, $T_c \sim 0.01 v_\sigma < T_R$. Therefore, the PQ symmetry, which had been nonthermally restored by preheating, is also restored thermally at the end of reheating. A few Hubble times later, the temperature drops below T_c and the PQ symmetry becomes spontaneously broken, this time for good. We thus predict a thermal abundance of axions, which decouple at $\min\{T_c, T_A^{\text{dec}}\}$, where $T_A^{\text{dec}} \simeq 2 \times 10^9 \text{ GeV} v_{11}^{2.246}$ [61,63,64]. Considering $g_* = 427/4$ relativistic degrees of freedom at axion decoupling, we get $\Delta N_\nu^{\text{eff}} \simeq 0.03$, which is much smaller than in HSI and in good agreement with current data. This small value of $\Delta N_\nu^{\text{eff}}$ could be probed with future CMB polarization experiments [65,66]. As discussed in Ref. [67], a nondetection of new thermal relics with future CMB probes reaching $\Delta N_\nu^{\text{eff}} \sim 0.01$ will imply that, if such relics exist, they were never in thermal equilibrium with the SM.

Finally, we remark that the EOS of the Universe is $w = 1/3$ in both the period of inflaton oscillations in a quartic potential [68] and the nonthermally PQ restored phase, because the evolution is conformal in a quartic potential. This is so for both HHSI and HSI. However, in HSI, there is a small period of matter domination before the ρ particles decay to reheat the SM, whose effects on N are within the uncertainties.

Dark matter.—At the spontaneous breaking of the PQ symmetry, a network of cosmic strings is formed in both HHSI and HSI. In the first case, this happens by the standard Kibble mechanism in thermal equilibrium [69], and in the second, nonthermally [57]. The evolution of the network leads to a population of low-momentum axions

that together with those arising from the realignment mechanism [70–72] constitute the dark matter in SMASH. Requiring that all the DM is made of axions demands

$$3 \times 10^{10} \text{ GeV} \lesssim v_\sigma \lesssim 1.2 \times 10^{11} \text{ GeV}, \quad (4)$$

which translates into the mass window

$$50 \mu\text{eV} \lesssim m_A \lesssim 200 \mu\text{eV}, \quad (5)$$

where we have updated the results of Ref. [9] with the latest axion mass data [34]. The main uncertainty arises from the string contribution [9,73], which we estimate as 3–4 times larger than the misalignment one; the uncertainty is expected to be diminished in the near future [74,75]. The SMASH axion mass window (5) will be probed in the upcoming decade by direct detection experiments such as MADMAX [76,77] and ORPHEUS [78]. A sizable part of the DM in this scenario may be in the form of axion miniclusters [79], which offer interesting astrophysical signatures [80,81].

Baryogenesis.—The origin of the baryon asymmetry of the Universe is explained in SMASH from thermal leptogenesis [82]. This requires the massive RH neutrinos N_i acquiring equilibrium abundances and then decaying when their production rates become Boltzmann suppressed. As we have seen, in HHSI, $T_R > T_c$ for stable models in the DM window (5). The RH neutrinos become massive after the PQ SSB, and those with masses $M_i < T_c$ retain an equilibrium abundance. The stability bound on the Yukawa couplings Y_{ii} enforces $T_c > M_1$, so that at least the lightest RH neutrino stays in equilibrium. Moreover, the annihilations of the RH neutrinos tend to be suppressed with respect to their decays. This allows for vanilla leptogenesis from the decays of a single RH neutrino, which demands $M_1 \gtrsim 5 \times 10^8 \text{ GeV}$ [83,84]. However, for v_σ as in (4), this is just borderline compatible with stability. Nevertheless, leptogenesis can occur with a mild resonant enhancement [85] for a less hierarchical RH neutrino spectrum, which relaxes the stability bound and ensures that all the RH neutrinos remain in equilibrium after the PQ SSB.

Future perspectives.—SMASH provides very clear predictions, which will be tested by the next generation of CMB, large-scale structure, and axion DM experiments. The model predicts a correlation between r , n_s , and a small negative value of α , as well as tiny non-Gaussianities. It also implies the existence of a cosmic background of relativistic axions which may be detected with future CMB polarization experiments. In SMASH, the totality of the DM in the Universe is made of cold axions with a mass in the range (5), which will be explored in the next decade. If all these features are met simultaneously, it will be a very compelling hint in favor of SMASH. If only one is not, the model will be ruled out. We recall that the

cosmological predictions of SMASH are reliable, as opposed to those of incomplete models such as Higgs inflation, which suffers from an early breaking of perturbative unitarity.

SMASH provides an explanation for five of the most pressing problems in particle physics and cosmology: inflation, DM, baryogenesis, the strong CP problem, and the smallness of neutrino masses, some of which are naturalness issues. However, the model does not solve the hierarchy problem nor the cosmological constant problem. It would be interesting to explore if, e.g., some relaxation mechanism along the lines of Refs. [86–89] could be embedded in SMASH to solve also these problems while maintaining its minimality.

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