

The major upgrade of the MAGIC telescopes, Part II: The achieved physics performance using the Crab Nebula observations

J. Aleksić^a, S. Ansoldi^b, L. A. Antonelli^c, P. Antoranz^d, A. Babic^e, P. Bangale^f, M. Barceló^a, J. A. Barrio^g, J. Becerra González^{h,aa}, W. Bednarekⁱ, E. Bernardini^j, B. Biasuzzi^b, A. Biland^k, M. Bitossi^l, O. Blanch^a, S. Bonnefoy^g, G. Bonnoli^c, F. Borracci^f, T. Bretz^{l,ab}, E. Carmona^{m,*}, A. Carosi^c, R. Cecchi^z, P. Colin^{f,*}, E. Colombo^h, J. L. Contreras^g, D. Corti^o, J. Cortina^a, S. Covino^c, P. Da Vela^d, F. Dazzi^f, A. De Angelis^b, G. De Caneva^j, B. De Lotto^b, E. de Oña Wilhelmiⁿ, C. Delgado Mendez^m, A. Dettlaff^f, D. Dominis Prester^e, D. Dorner^l, M. Doró^o, S. Einecke^p, D. Eisenacher^l, D. Elsaesser^l, D. Fidalgo^g, D. Fink^f, M. V. Fonseca^g, L. Font^q, K. Frantzen^p, C. Fruck^f, D. Galindo^r, R. J. García López^h, M. Garczarczyk^j, D. Garrido Terrats^q, M. Gaug^q, G. Giavitto^{aj}, N. Godinović^e, A. González Muñoz^a, S. R. Gozzini^j, W. Haberer^f, D. Hadasch^{n,ac}, Y. Hanabata^s, M. Hayashida^s, J. Herrera^h, D. Hildebrand^k, J. Hose^f, D. Hrupec^e, W. Ideciⁱ, J. M. Illa^a, V. Kadenius^t, H. Kellermann^f, M. L. Knoetig^k, K. Kodani^s, Y. Konno^s, J. Krause^f, H. Kubo^s, J. Kushida^s, A. La Barbera^c, D. Lelas^e, J. L. Lemus^g, N. Lewandowska^l, E. Lindfors^{t,ae}, S. Lombardi^c, F. Longo^b, M. López^g, R. López-Coto^a, A. López-Oramas^a, A. Lorca^g, E. Lorenz^{f,ad}, I. Lozano^g, M. Makariev^u, K. Mallo^t, G. Maneva^u, N. Mankuzhiyil^{b,af}, K. Mannheim^l, L. Maraschi^c, B. Marcote^r, M. Mariotti^o, M. Martínez^a, D. Mazin^f, U. Menzel^f, J. M. Miranda^d, R. Mirzoyan^f, A. Moralejo^a, P. Munar-Adrover^r, D. Nakajima^s, M. Negrello^o, V. Neustroev^t, A. Niedzwieckiⁱ, K. Nilsson^{t,ae}, K. Nishijima^s, K. Noda^f, R. Orito^s, A. Overkemping^p, S. Paiano^o, M. Palatiello^b, D. Paneque^f, R. Paoletti^d, J. M. Paredes^f, X. Paredes-Fortuny^r, M. Persic^{b,ag}, J. Poutanen^t, P. G. Prada Moroni^v, E. Prandini^k, I. Puljak^e, R. Reinthal^l, W. Rhode^p, M. Ribó^r, J. Rico^a, J. Rodríguez García^f, S. Rügamer^l, T. Saito^s, K. Saito^s, K. Satalecka^g, V. Scalzotto^o, V. Scapin^g, C. Schultz^o, J. Schlammer^f, S. Schmid^f, T. Schweizer^f, S. N. Shore^v, A. Sillanpää^t, J. Sitarek^{a,*}, I. Snidaric^e, D. Sobczynskaⁱ, F. Spanier^l, A. Stamerra^c, T. Steinbring^l, J. Storz^l, M. Strzys^f, L. Takalo^t, H. Takami^s, F. Tavecchio^c, L. A. Tejedor^g, P. Temnikov^u, T. Terzić^e, D. Tesaro^h, M. Teshima^f, J. Thiele^p, O. Tibolla^l, D. F. Torres^w, T. Toyama^f, A. Treves^x, P. Vogler^k, H. Wettskind^f, M. Will^h, R. Zanin^r

^aIFAE, Campus UAB, E-08193 Bellaterra, Spain

^bUniversità di Udine, and INFN Trieste, I-33100 Udine, Italy

^cINAF National Institute for Astrophysics, I-00136 Rome, Italy

^dUniversità di Siena, and INFN Pisa, I-53100 Siena, Italy

^eCroatian MAGIC Consortium, Rudjer Boskovic Institute, University of Rijeka and University of Split, HR-10000 Zagreb, Croatia

^fMax-Planck-Institut für Physik, D-80805 München, Germany

^gUniversidad Complutense, E-28040 Madrid, Spain

^hInst. de Astrofísica de Canarias, E-38200 La Laguna, Tenerife, Spain

ⁱUniversity of Łódź, PL-90236 Lodz, Poland

^jDeutsches Elektronen-Synchrotron (DESY), D-15738 Zeuthen, Germany

^kETH Zurich, CH-8093 Zurich, Switzerland

^lUniversität Würzburg, D-97074 Würzburg, Germany

^mCentro de Investigaciones Energéticas, Medioambientales y Tecnológicas, E-28040 Madrid, Spain

ⁿInstitute of Space Sciences, E-08193 Barcelona, Spain

^oUniversità di Padova and INFN, I-35131 Padova, Italy

^pTechnische Universität Dortmund, D-44221 Dortmund, Germany

^qUnitat de Física de les Radiacions, Departament de Física, and CERES-IEEC, Universitat Autònoma de Barcelona, E-08193 Bellaterra, Spain

^rUniversitat de Barcelona, ICC, IEEC-UB, E-08028 Barcelona, Spain

^sJapanese MAGIC Consortium, Division of Physics and Astronomy, Kyoto University, Japan

^tFinnish MAGIC Consortium, Tuorla Observatory, University of Turku and Department of Physics, University of Oulu, Finland

^uInst. for Nucl. Research and Nucl. Energy, BG-1784 Sofia, Bulgaria

^vUniversità di Pisa, and INFN Pisa, I-56126 Pisa, Italy

^wICREA and Institute of Space Sciences, E-08193 Barcelona, Spain

^xUniversità dell'Insubria and INFN Milano Bicocca, Como, I-22100 Como, Italy

^yEuropean Gravitational Observatory, I-56021 S. Stefano a Macerata, Italy

^zUniversità di Siena and INFN Siena, I-53100 Siena, Italy

^{aa}now at: NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA and Department of Physics and Department of Astronomy, University of Maryland, College Park, MD 20742, USA

^{ab}now at Ecole polytechnique fédérale de Lausanne (EPFL), Lausanne, Switzerland

^{ac}Now at Institut für Astro- und Teilchenphysik, Leopold-Franzens- Universität Innsbruck, A-6020 Innsbruck, Austria

^{ad}deceased

^{ae}now at Finnish Centre for Astronomy with ESO (FINCA), Turku, Finland

^{af}now at Astrophysics Science Division, Bhabha Atomic Research Centre, Mumbai 400085, India

^{ag}also at INAF-Trieste

Abstract

MAGIC is a system of two Imaging Atmospheric Cherenkov Telescopes located in the Canary island of La Palma, Spain. During summer 2011 and 2012 it underwent a series of upgrades, involving the exchange of the MAGIC-I camera and its trigger system, as well as the upgrade of the readout system of both telescopes. We use observations of the Crab Nebula taken at low and medium zenith angles to assess the key performance parameters of the MAGIC stereo system. For low zenith observations the trigger threshold of the MAGIC telescopes is ~ 50 GeV. The integral sensitivity for sources with Crab Nebula-like spectrum above 220 GeV is $(0.66 \pm 0.03)\%$ of Crab Nebula flux in 50 h of observations. The angular resolution at those energies is $\lesssim 0.07^\circ$, while the energy resolution is 16%. We also re-evaluate the effect of the systematic uncertainty on the data taken with the MAGIC telescopes after the upgrade. We estimate that the systematic uncertainties can be divided in following components: $< 15\%$ in energy scale, 11-18% in flux normalization and ± 0.15 for the energy spectrum power-law slope.

Keywords: Gamma-ray astronomy, Cherenkov telescopes, Crab Nebula

1. Introduction

MAGIC (Major Atmospheric Gamma Imaging Cherenkov telescopes) consists of two 17 m diameter Imaging Atmospheric Cherenkov Telescopes (IACT). The telescopes are located at a height of 2200 m a.s.l. on the Roque de los Muchachos Observatory on the Canary Island of La Palma, Spain (28°N, 18°W). They are used for observations of particle showers produced in the atmosphere by very high energy (VHE, $\gtrsim 30$ GeV) γ -rays. Both telescopes are normally operated together in the so-called stereoscopic mode, in which only events seen simultaneously in both telescopes are triggered and analyzed (Aleksić et al., 2012a).

Between summer 2011 and 2012 the telescopes went through a major upgrade, carried out in two stages. In summer 2011 the readout systems of both telescopes were upgraded. The multiplexed FADCs used before in MAGIC-I (Goebel et al., 2007) as well as the Domino Ring Sampler version 2 used in MAGIC-II (DRS2, Tesaro et al., 2009) have been replaced by Domino Ring Sampler version 4 chips (DRS4, Ritt et al., 2010; Sitarek et al., 2013). Besides lower noise, the switch to DRS4 based readout allowed to eliminate the $\sim 10\%$ dead time present in the previous system due to the DRS2 chip. In summer 2012 the second stage of the upgrade followed with an exchange of the camera of the MAGIC-I telescope by a uniformly pixelized one (Mazin et al., 2013). The new MAGIC-I camera is equipped with 1039 photomultipliers (PMTs), as the one of the MAGIC-II telescope. Each of the camera pixels covers a field of view of 0.1° , resulting in the total field of view of $\sim 3.5^\circ$. The upgrade of the camera allowed to increase the area of the trigger region in MAGIC-I by a factor of 1.7 to the value of 4.8° . In the first part of this article (Aleksić et al., 2014a) we described in detail the hardware improvements and the commissioning of the system. In this second part we focus on the performance of the upgraded system based on Crab Nebula observations.

The Crab Nebula is a nearby (~ 2 kpc away) pulsar wind nebula, and the first source detected in VHE γ rays (Weekes et al., 1989). A few years ago, the satellite γ -ray telescopes, AGILE & Fermi-LAT observed flares from the Crab Nebula at GeV energies (Tavani et al., 2011; Abdo et al., 2011). However so far no confirmed variability in the VHE range was found. Therefore, since the Crab Nebula is still considered the brightest steady VHE γ -ray source, it is commonly referred to as the “standard candle” of the VHE γ -ray astronomy, and it is frequently used to evaluate performance of VHE instruments. In this paper we use Crab Nebula data to quantify the improvement in performance of the MAGIC telescopes after the aforementioned upgrade. In Section 2 we describe the different samples of Crab Nebula data used in the analysis. In Section 3 we explain the techniques and methods used for the processing of the MAGIC stereo data. In Section 4 we evaluate the performance parameters of the MAGIC telescopes after the upgrade. In Section 5 we discuss the influence of the upgrade on the systematic

*Corresponding authors: J. Sitarek (jsitarek@ifae.es), E. Carmona (emiliano.carmona@ciemat.es), P. Colin (colin@mppmu.mpg.de)

zenith[°]	offset ξ [°]	time[h]
0-30	0.40	11.1
30-45	0.40	4.0
0-30	0.20	2.3
0-30	0.35	0.9
0-30	0.70	1.9
0-30	1.00	4.2
0-30	1.40	4.6
0-30	1.80	4.1

Table 1: Zenith angle range, wobble offset angles ξ , and effective observation time of the Crab Nebula samples used in this study.

uncertainties of the measurements and quantify them. The final remarks and summary are gathered in Section 6.

2. Data sample

In order to evaluate the performance of the MAGIC telescopes, we use several samples of Crab Nebula data taken in different conditions between October 2013 and January 2014. Notice that as MAGIC is located in the Northern Hemisphere, the Crab Nebula is observable only during the winter season. The data were taken in the standard L1-L3 trigger condition (see Aleksić et al., 2014a). The data selection was based on the rate of background events surviving the stereo reconstruction. All data have been taken in the so-called wobble mode (Fomin et al., 1994), i.e. with the source position offset by a fixed angle, ξ , from the camera center in a given direction. This method allows to estimate the background from other positions in the sky at the same offset ξ . Most of the results are obtained using the data taken at low zenith angles ($< 30^\circ$) and with the standard wobble offset of 0.4° . To evaluate the performance at higher zenith angles we use a medium zenith sample ($30-45^\circ$). In addition, several low zenith samples taken at different offsets are used to study the sensitivity for off-axis observations. All the data samples are summarized in Table 1.

In addition, to analyze the data and to evaluate some of the performance parameters, such as the energy threshold or the energy resolution we used Monte Carlo (MC) simulations.

3. Data analysis

The data have been analyzed using the standard MAGIC tools: MARS (MAGIC Analysis and Reconstruction Software, Zanin et al., 2013). Here we briefly describe all the stages of the standard analysis chain of the MAGIC data.

3.1. Calibration

Each event recorded by the MAGIC telescopes consists of the waveform observed in each of the pixels. The waveforms span 30 ns and are sampled at a frequency of 2Gsamples/s. The telescopes are triggered with a typical rate of 250 – 300 Hz. In the first stage of the analysis the pixel signals are reduced to two numbers: charge and arrival time. The signal is extracted with a simple and robust “sliding window” algorithm (Albert et al.,

2008a), by finding the maximal integral of 6 consecutive time slices (corresponding to 3 ns) within the total readout window. The conversion from integrated readout counts to photoelectrons (phe) is done using the F-Factor (excess noise factor) method (see e.g. Mirzoyan & Lorenz, 1997). On average, one phe generates a signal of the order of ~ 100 integrated readout counts. For typical observation conditions the electronic noise and the light of the night sky with such an extractor result in a noise RMS level of ~ 1 phe and a bias (for very small signals) of ~ 2 phe. The DRS4 readout requires some special calibration procedures, such as the correction of the time inhomogeneity of the domino ring (see Sitarek et al., 2013, for details) which are applied at this stage. A small fraction of channels (typically $< 1\%$) might be malfunctioning, or be affected by bright stars in their field of view. The charge and time information of these pixels are interpolated from their neighbouring ones if at least 3 neighbouring pixels are valid.

3.2. Image cleaning and parametrization

After the upgrade, the camera of each MAGIC telescope has 1039 pixels, however the Cherenkov light of a typical air shower event illuminates only of the order of 10 pixels. Most of the pixel signals are solely induced by the night sky background (NSB) and the electronic noise. In order to remove pixels containing only noise to obtain the image of the shower we perform the so-called sum image cleaning (cf. Rissi, 2009; Aleksić et al., 2011). In the first step, we determine the so-called core pixels. For this, we search for compact groups of 2, 3 or 4 neighbouring pixels (2NN, 3NN, 4NN), with a summed charge above a given threshold. In order to protect against a large signal in a single pixel (e.g. due to an afterpulse) dominating a 3NN or 4NN group the signals are clipped before summation. The signals in those pixels should arrive within a given time window. These time windows were optimized using time resolution of the signal extraction (see Sitarek et al., 2013). For pulses just above the charge thresholds, the coincidence probability of signals from showers falling within the time window is $\approx 80 - 90\%$. The charge thresholds on the sum of 2NN, 3NN and 4NN groups are 2×10.8 phe, 3×7.8 phe, 4×6 phe and the corresponding time windows: 0.5, 0.7 and 1.1 ns. The values of the charge thresholds were optimized to assure that the probability of an event composed of only NSB and electronic noise to survive the cleaning is $\lesssim 6\%$. This translates directly to the maximum fraction of images affected by spurious islands due to noise fluctuations. In the second step, boundary pixels are looked for to reconstruct the rest of the shower image. We loop over all the pixels which have a neighbouring core pixel, and include them into the image if the charge of such boundary pixel lies above 3.5 phe and its signal arrives within 1.5 ns with respect to this core pixel. After the upgrade all the charge and time threshold values are the same in both telescopes.

3.3. Stereo reconstruction

Only events which survive the image cleaning in both telescopes, amounting to about 80% for standard trigger conditions, are retained in the analysis. Afterwards, the events from

both telescopes are paired, and a basic stereo reconstruction is performed. The tentative reconstructed direction of the event is computed from the crossing point from the main axes of the Hillas ellipses (Hillas, 1985). This first stereo reconstruction provides additional event-wise parameters such as *Impact* (defined as the distance of the shower axis to the telescope position, $Impact_1$ and $Impact_2$ are computed with respect to the MAGIC-I and MAGIC-II telescopes respectively) and the height of the shower maximum.

3.4. γ /hadron separation

Most of the events registered by the MAGIC telescopes are cosmic-rays showers, which are mainly of hadronic origin. Even for a bright source such as the Crab Nebula, the fraction of γ -ray events in the raw data is only of the order of 10^{-3} . The rejection of the hadronic background is done on the basis of image shape information and reconstructed direction. The γ /hadron separation is performed with the help of the so-called Random Forest (RF) algorithm (Albert et al., 2008). It allows to combine, in a straight-forward way, the image shape parameters, the timing of the shower, and the stereo parameters into a single classification parameter, *Hadronness*.

The survival probability for γ rays and background events after a cut in *Hadronness* is shown in Fig. 1 for three different energy bins.

A background rejection better than 90% can be achieved, with only a small loss of γ events. The γ /hadron separation performs better at higher energies due to the larger and better defined images. Strong cuts (below 0.1 – 0.2) in *Hadronness* result in a slight mismatch of the γ -ray efficiency obtained from the MC simulations with respect to the one from the data itself. For example, for a requested MC efficiency of 90% there is a small mismatch of $\lesssim 12\%$ between the data and MC simulations efficiency at TeV energies. For lower energies the effect is smaller, $\lesssim 6\%$. For spectral analysis, cuts with a large γ -ray efficiency are used (see Section 4.5). Note that in addition to this *Hadronness* cut, also the direction reconstruction method (described in subsection 3.5) provides additional background suppression.

3.5. Arrival direction reconstruction

The classical method for arrival direction reconstruction uses the crossing point of the main axes of the Hillas ellipses in the individual cameras (Aharonian et al., 1997; Hofmann et al., 1999). In the standard MAGIC analysis the event-wise direction reconstruction of the incoming γ ray is performed with a DISP RF method. This method takes into account image shape and timing information (Aleksić et al., 2010a). For each telescope we compute an estimated distance, DISP, between the image centroid and the source position. As the source position is assumed to be on the line containing the main axis of the Hillas ellipse this results in two possible solutions on either side of the image centroid (see Fig. 2). In general, the ambiguity (the so-called head-tail discrimination) can be solved from the asymmetry along the main axis of the image (Domingo-Santamaria et al., 2005) or from the crossing point

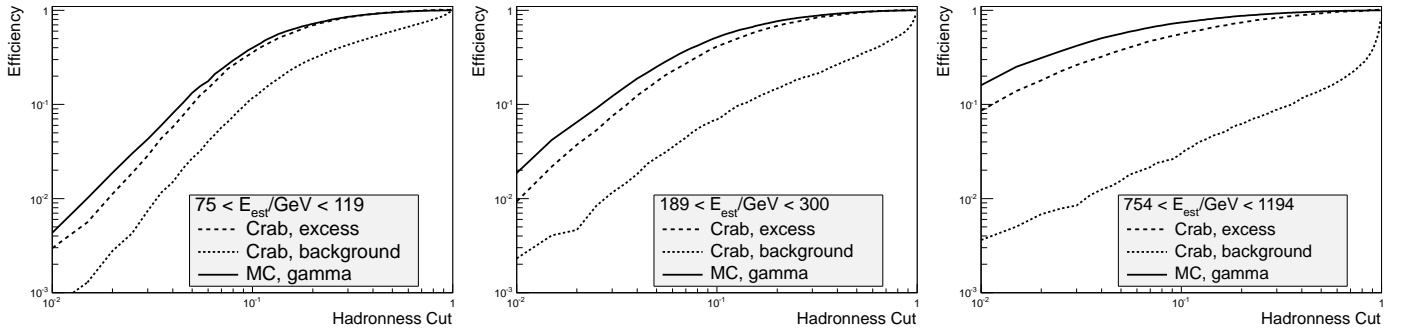


Figure 1: Fraction of events (efficiency) of a given kind: excess γ -rays (dashed), MC simulated γ -ray (solid) and background events (dotted) surviving a *Hadronness* cut. Different panels correspond to different bins in estimated energy: 75-119 GeV (left), 189-300 GeV (center) and 754-1194 GeV (right). Additional cuts of $\theta^2 < 0.03$ and *Size* > 50 phe have been applied beforehand.

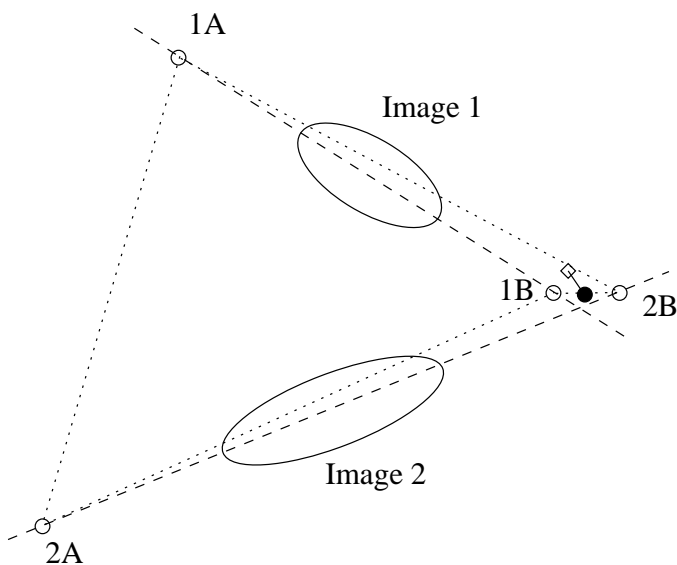


Figure 2: Principle of the Stereo DISP RF method. The main axes of the images are plotted with dashed lines. The two DISP RF reconstructed positions per telescope (1A, 1B, 2A, 2B) are shown with empty circles. The 4 angular distances (1A-2A, 1A-2B, 1B-2A, 1B-2B) are shown with dotted lines. The final reconstructed position (the filled circle) is a weighted average of the two closest '1' and '2' points. The true source position is marked with a diamond.

of the images. However at the lowest energies, where the images consist of few pixels the head-tail discrimination may fail at least in one telescope. The head-tail discrimination based on the crossing point may also fail, in the case of close to parallel events. Therefore we use a more robust method. We compute the 4 distances between the 2 reconstructed positions from each of the telescopes (see dotted lines in Fig. 2). We then select the pair of reconstructed positions which give the smallest distance (in the above example 1B-2B). As the estimation of the DISP parameter is trained with simulated γ rays it often gives non-consistent results for hadronic background events. This provides an extra γ /hadron separation criterion. If none of the four pairs give a similar arrival direction in both telescopes (namely the lowest distance is larger than 0.22°) the event is discarded. After determining the correct pair of points, the reconstructed source position is computed as the average of the positions from

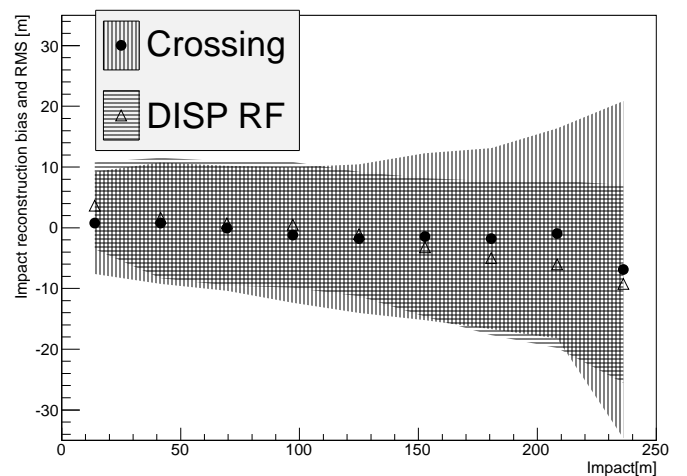


Figure 3: The difference between reconstructed and true *Impact* parameter for events with energy 300-500 GeV. The markers show the bias (mean of Gaussian) in the reconstruction, and the shaded region shows the resolution (RMS of Gaussian) in the reconstruction. Filled circles and vertical lines show the classical reconstruction based on the crossing point of the images. Empty triangles and horizontal lines show the reconstruction based on DISP RF method. Only events with *size* > 50 phe and $\theta^2 < 0.02$ are used.

both telescopes weighted with the number of pixels in each image. The angular distance from this point to the assumed source position is called θ .

The DISP RF method explained above improves not only the reconstruction of the arrival direction but also the estimation of other shower parameters. As an example Fig. 3 shows the difference between the reconstructed and the true *Impact* parameter for MC γ -rays with energies of few hundred of GeV. The *Impact* parameter can be reconstructed with a precision of about 10 m. The reconstruction with the DISP method is clearly superior for larger values of *Impact*, where a good reconstruction of this parameter for events seen outside of the light pool of the shower is important for improving the energy resolution.

3.6. Energy estimation

The event-wise energy estimation is a two-step process. Using simulated γ rays, we build look-up tables relating the event energy to the *Impact* and Cherenkov photons density measured

by each telescope (see Aleksić et al. (2012a) for a detailed explanation). To correct for a zenith dependent bias in the energy reconstruction, an empirical formula is applied. A second correction is applied to account for a small azimuth dependence (due to the geomagnetic field effect). Finally, a third correction improves the energy reconstruction for large images only partially contained in the camera. The final estimated energy, E_{est} , is computed as the average of the energies reconstructed individually for each telescope, weighted by the inverse of their errors.

4. MAGIC performance

In this section we evaluate the main performance parameters of the MAGIC telescopes after the exchange of the readout systems and MAGIC-I camera and compare them with the values from before the upgrade.

4.1. Energy threshold

The energy threshold of an IACT cannot be obtained in a straight-forward way from the data itself. One needs to rely on the Monte Carlo simulations and to make sure that they describe the data appropriately. The energy threshold depends on the trigger settings for a given observations. In particular it depends on an amplitude threshold of individual pixels in the 3NN multiplicity trigger. In order to fine tune the trigger parameters in the MC simulations we follow the method described in Aleksić et al. (2012a). For each event that survived the trigger we search for the 3NN combination with the highest signal. For small showers, just above the threshold, this is the most probable triple which gave a trigger. The pixel in this triple with the lowest signal provides a handle for the trigger threshold. Using many events we build a distribution of such lowest signals in the highest triple. We then compare the position of the peak with a similar distribution produced from MC protons (see Fig.4). We obtain that the individual pixel thresholds are ≈ 3.9 phe for MAGIC-I and ≈ 4.1 phe for MAGIC-II. Those values are within 10-15% consistent with the threshold values obtained directly from the data with an independent method of rate-scans in Aleksić et al. (2014a).

In order to study the energy threshold of the MAGIC telescopes we construct a differential rate plot using MC simulations. A common definition of an energy threshold is a peak energy of such a plot for a hypothetical source with a spectral index of -2.6 . In Fig. 5 we show the differential rate plot in two zenith ranges for events that survived image reconstruction in both telescopes. For low zenith the reconstruction threshold energy is ~ 70 GeV. Note however that the peak is broad and extends far to lower energies. Therefore it is also possible to evaluate the performance of the telescopes and obtain scientific results below such defined threshold.

In Fig. 6 we show the energy threshold of the MAGIC telescopes as the function of zenith distance of observations. The threshold is quite stable for low zenith, (i.e. $< 30^\circ$) observations. It increases rapidly for higher zenith angles, due to larger absorption of the Cherenkov light in the atmosphere and dilution of the photons reaching the ground over a larger light pool.

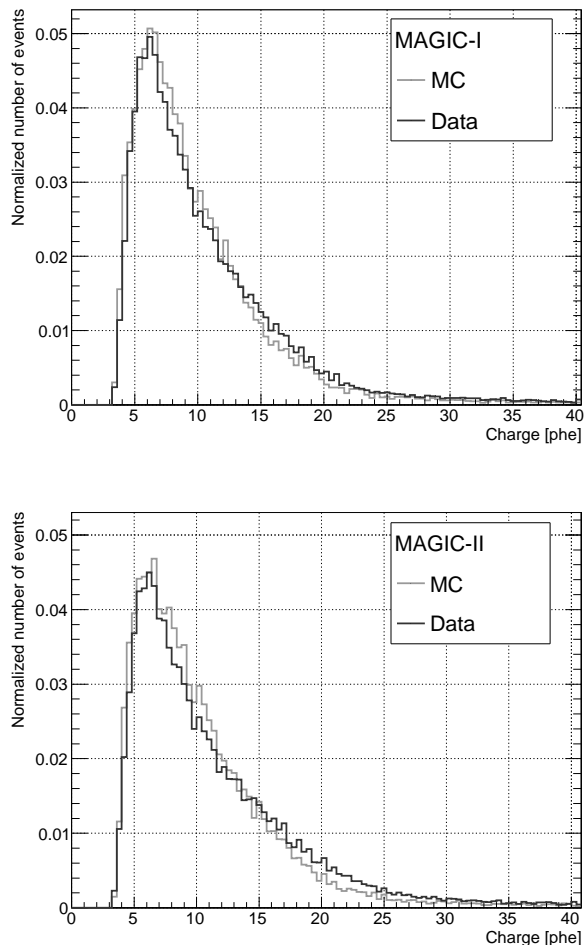


Figure 4: Distribution of the smallest charge in the largest triple of pixels for data (black) and MC protons simulated with a trigger conditions of 3NN above a threshold of 3.9 phe (MAGIC-I) or 4.1 phe (MAGIC-II) (gray). Top panel shows MAGIC-I, bottom MAGIC-II.

The threshold can be evaluated at different stages of the analysis. The trigger threshold computed from all the events that triggered both telescopes is naturally the lowest one, being ~ 50 GeV at low zenith angles. The shower reconstruction procedure involving image cleaning and a typical data quality cut of having at least 50 phe in each telescope raises the threshold to ~ 70 GeV. The events with size lower than this are very small, subjected to high Poissonian fluctuations and therefore harder to reconstruct. Also the separation of γ candidates from the much more abundant hadronic background becomes harder at lower image sizes. Signal extraction cuts (the so-called *Hadronness* cut, and a cut in the angular distance to the nominal source position, θ) increase the threshold further to about 75 GeV at low zenith angles. The value of the energy threshold doubles at zenith angle of 43° . In the investigated zenith angle range the value of the threshold after all cuts can be approximated by an empirical formula: $74 \times \cos(\text{Zenith_Angle})^{-2.3}$ GeV.

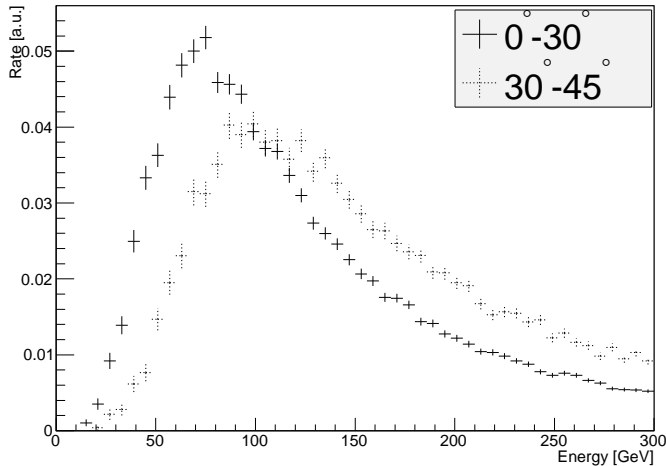


Figure 5: Rate of γ -ray events (in arbitrary units) surviving the image cleaning with at least 50 phe for a source with a spectral index of -2.6 . Solid line: zenith distance below 30° , dotted line: zenith distance between 30° and 45° .

4.2. Effective collection area

For large arrays of IACTs the collection area well above the energy threshold for low zenith observations is approximately equal to the physical size of the array (Bernlöhr et al., 2013). On the other hand for a single telescope or small arrays such as the MAGIC telescopes, the collection area is mainly determined by the size of the Cherenkov light pool (radius of ~ 120 m). We compute the collection area as the function of the energy E following the standard definition of $A_{\text{eff}}(E) = N(E)/N_0(E) \times \pi r_{\text{max}}^2$. $N_0(E)$ is the number of simulated events, r_{max} is the maximum simulated shower impact and $N(E)$ is the number of events surviving either the trigger condition or a given set of cuts. When computing the collection area in broad bins of energy we use weights to reproduce a given spectral shape. The collection area of the MAGIC telescopes at the trigger level is about 10^5 m² for 300 GeV gamma rays (see Fig. 7). In the TeV range it grows slowly with energy, as some of the large showers can be still caught at large values of *Impact* where the density of the Cherenkov photons on the ground falls rapidly. Around and below the energy threshold the collection area falls rapidly, as only events with a significant upward fluctuation of the light yield can trigger the telescope. At the energy of a few TeV, the trigger collection area after the upgrade is larger by $\sim 30\%$, mostly due to the larger trigger area in the M1 camera. The collection area for observations at higher zenith angles is naturally smaller below ~ 100 GeV due to a higher threshold of the observations. However, at TeV energies it is larger by $\sim 40\%$ due to an increase of the size of the light pool.

In Fig. 7 we also show the collection area after image cleaning, quality and signal extraction cuts optimized for best differential sensitivity (see Section 4.7). The feature of a dip in the collection area after cuts around 300 GeV is caused by a stronger *Hadronness* cut. At energies of a few TeV about 75% of all triggered events are kept.

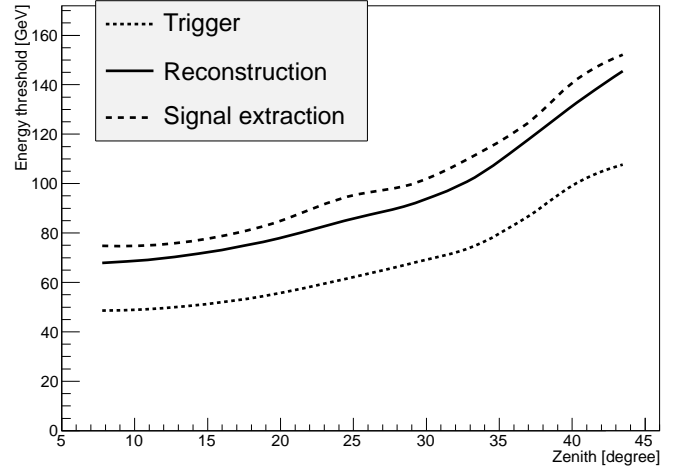


Figure 6: Threshold of the MAGIC telescopes as a function of the zenith angle of the observations. The energy threshold is defined as the peak energy in the differential rate plot for a source with -2.6 spectral index. Dotted curve: threshold at the trigger level. Solid line: only events with images that survived image cleaning in each telescope with at least 50 phe. Dashed line: with additional cuts of $Hadronness < 0.5$ and $\theta^2 < 0.03^\circ^2$ applied.

4.3. Relative light scale between both telescopes

For observations at low zenith angles the density of Cherenkov light photons on the ground produced by a VHE γ -ray shower depends mostly on its energy and its *Impact* parameter. Except for a small dependence on the relative position of the shower axis with respect to the Geomagnetic field, due to the geomagnetic field effect (mostly pronounced at lowest energies, see e.g. Commichau et al., 2008), the density is radially symmetric. Thus, it is possible to compare the light scale of both telescopes by selecting γ -like events from data in which the reconstructed *Impact* parameter is similar in both telescopes (Hofmann, 2003). In the case of hadronic background events, such a correlation is much weaker due to the strong internal fluctuations and poor estimation of the *Impact* parameter. In order to obtain a pure γ -ray sample, we apply rather strict cuts $Hadronness < 0.2$ and $\theta^2 < 0.01$.

The response of MAGIC-II is $(11 \pm 1_{\text{stat}})\%$ larger than that of MAGIC-I for showers observed at similar *Impact* parameter (see Fig. 8). The result is a sum of multiple effects such as differences in the reflectivity of the mirrors, small differences between the two PMT populations, or uncertainty in the *F*-factor used for the calibration of each telescope. The MC simulations are fine tuned to take into account this inter-telescope calibration. Therefore the estimated energy obtained independently from both telescopes is consistent (see Fig. 9).

4.4. Energy resolution

We evaluate the performance of the energy reconstruction with γ -ray MC simulations. The simulations are divided into bins of true energy (5 bins per decade). In each bin we construct a distribution of $(E_{\text{est}} - E_{\text{true}})/E_{\text{true}}$ and fit it with a Gaussian function. The energy resolution is defined as the standard deviation obtained from this fit. The bias of the energy reconstruction method can be computed as the mean value of the dis-

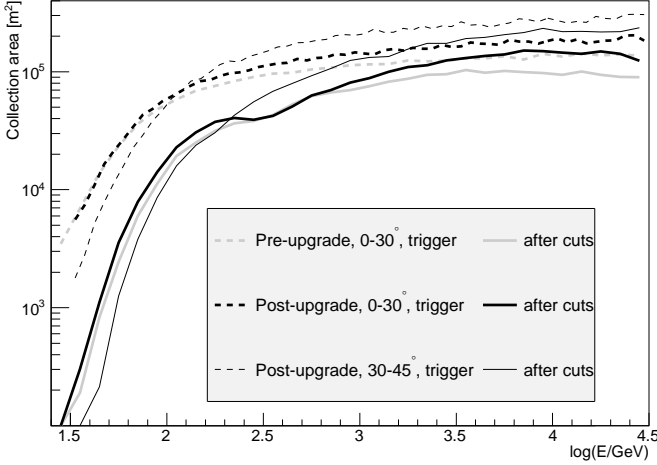


Figure 7: Collection area of the MAGIC telescopes after the upgrade at the trigger level (dashed lines) and after all cuts (solid lines). Thick lines show the collection area for low zenith observations, while thin lines correspond to medium zenith angle. For comparison, the corresponding pre-upgrade collection areas are shown with gray lines.

tribution. The energy resolution and the bias of the MAGIC telescopes as the function of the true energy of the γ rays are shown in Fig. 10.

For low zenith observations in the energy range of a few hundred GeV the energy resolution falls down to 15%. For higher energies it degrades due to an increasing fraction of truncated images, and showers with high *Impact* parameters as well as worse statistics in the training sample. Note that the energy resolution can be easily improved in the multi-TeV range with additional quality cuts (e.g. in the maximum reconstructed *Impact*), however at the price of lowering the collection area. At low energies the energy resolution is degraded, due to worse precision in the image reconstruction (in particular the *Impact* parameters), and higher internal relative fluctuations of the shower. Above a few hundred GeV the absolute value of the bias is below a few per cent. At low energies ($\lesssim 100$ GeV) the estimated energy bias rapidly increases due to the threshold effect. For observations at higher zenith angles the energy resolution is similar. Since an event of the same energy observed at higher zenith will produce a smaller image, the energy resolution at the lowest energies is slightly worse. On the other hand, at multi-TeV energies, where the showers observed at low zenith are often partially truncated at the edge of the camera, and may even saturate some of the pixels, the energy resolution is slightly better for higher zenith observations. As the energy threshold shifts with increasing zenith angle, the energy bias at energies below 100 GeV is much stronger for higher zenith observations.

When the data are binned according to estimated energy of individual events (note that, in contrary to MC simulations, in the data only the estimated energy is known) the value of the bias will change depending on the spectral shape of the source. With steeper spectra more events will migrate from lower energies resulting in an overestimation of the energy. Note that this effect does not occur in the case of binning the events according

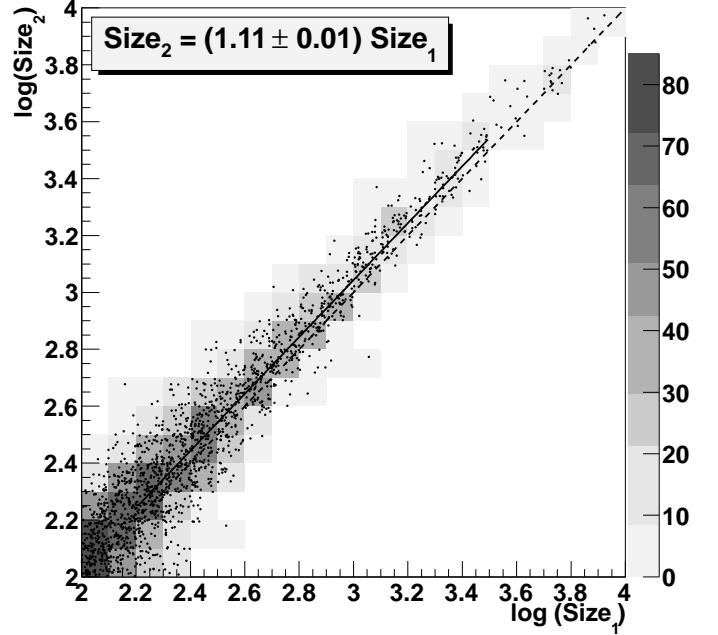


Figure 8: Correlation of $Size_2$ and $Size_1$ for γ -ray events obtained with Crab Nebula data. Only events with $Hadronness < 0.2$, $\theta^2 < 0.01$, $Impact_1 < 150$ m, $Impact_2 < 150$ m and $|Impact_1 - Impact_2| < 10$ m are used. Individual events are marked by black dots, while the gray scale shows the total number of events in a given bin. The black solid line shows the result of the fit. The dashed line corresponds to $Size_1 = Size_2$.

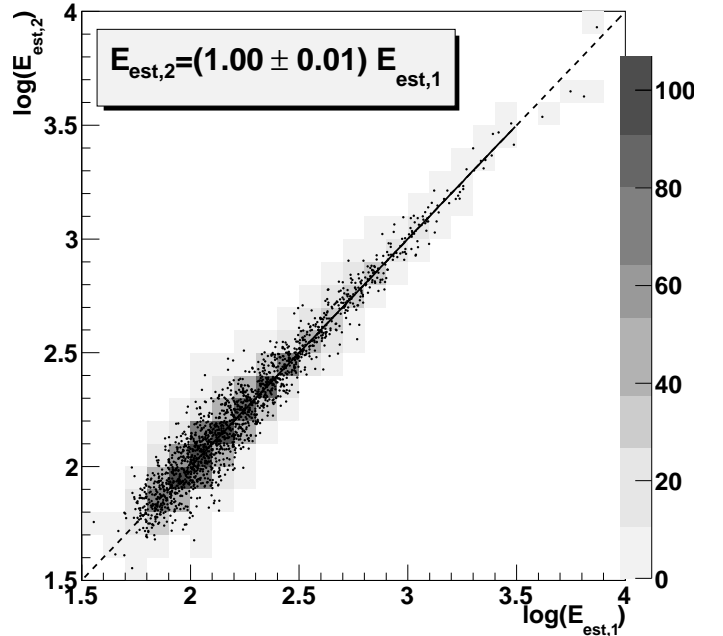


Figure 9: Correlation of the γ -ray estimated energy for events (as in Fig. 8) from the images recorded in each telescope separately. The black solid line shows the result of the fit. The dashed line corresponds to $E_{est,1} = E_{est,2}$.

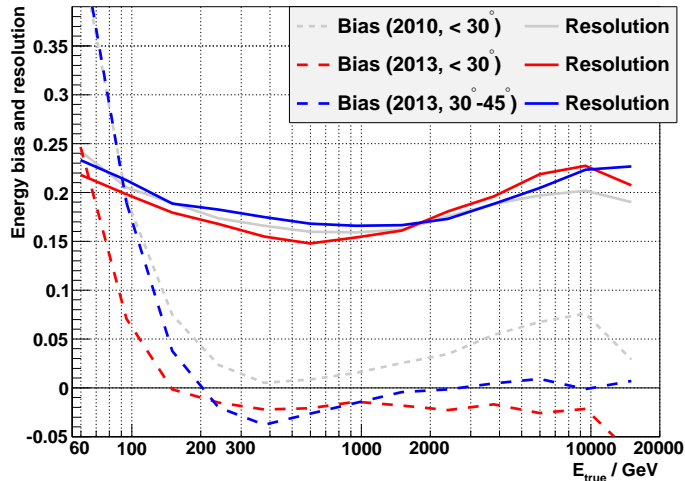


Figure 10: Energy resolution (solid lines) and bias (dashed lines) obtained from the MC simulations of γ -rays. Events are weighted in order to represent a spectrum with a slope of -2.6 . Red: low zenith, blue: medium zenith. For comparison, pre-upgrade values from Aleksić et al. (2012a) are shown in gray lines.

to their true energy (as in Fig. 10). In Fig. 11 we show such a bias as a function of spectral slope for a few values of estimated energy. Note that the bias is corrected in the spectral analysis by means of an unfolding procedure (Albert et al., 2007).

The energy resolution cannot be checked with the data in a straight-forward way and one has to rely on the values obtained from MC simulations. Nevertheless, we can use the fact of having two, nearly independent estimations of the energy, $E_{est,1}$ and $E_{est,2}$ from each of the telescopes to perform a consistency check. We define relative energy difference as $RED = (E_{est,1} - E_{est,2})/E_{est}$. If the $E_{est,1}$ and $E_{est,2}$ estimators were completely independent the energy resolution would be $\approx RMS(RED)/\sqrt{2}$. In Fig. 12 we show a dependence of $RMS(RED)$ on the reconstructed energy. The curve obtained from the data is consistent with the one of MC simulations within a few per cent accuracy. The first point (between 45 and 75 GeV) shows a sudden drop in $RMS(RED)$ compared to the other points, consistently in the data and MC simulations. Note that this point is below the analysis threshold, therefore it is mostly composed of peculiar events in which the shower produces more Cherenkov light than average for this energy. This results in a strong correlation of $E_{est,1}$ and $E_{est,2}$ allowing for a relatively low value of inter-telescope difference in estimated energy, and still a rather poor energy resolution.

4.5. Spectrum of the Crab Nebula

In Fig. 13 we show the spectrum of the Crab Nebula obtained with the total (low + medium zenith) sample. For clarity, the spectrum is presented in the form of spectral energy distribution, i.e. $E^2 dN/dE$. In order to minimize the systematic errors we apply *Hadronness* and θ^2 cuts with high γ -ray efficiency (90% and 75% respectively) for the spectral reconstruction. The spectrum in the energy range 65 GeV – 13.5 TeV can

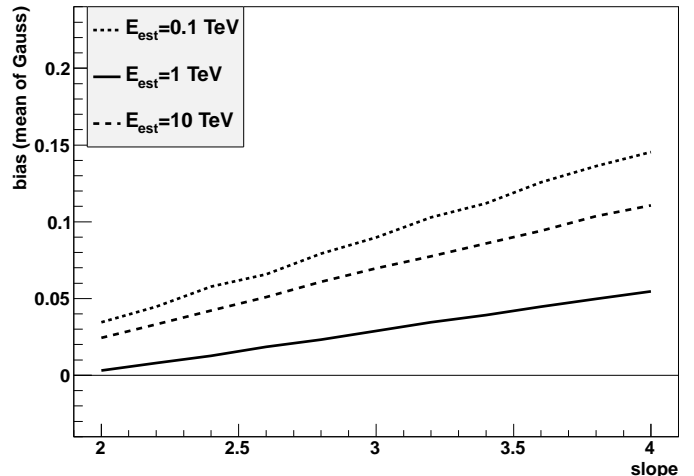


Figure 11: Energy bias as a function of the spectral slope for different estimated energies: 0.1 TeV (dotted line), 1 TeV (solid), 10 TeV (dashed). Zenith angle below 30° .

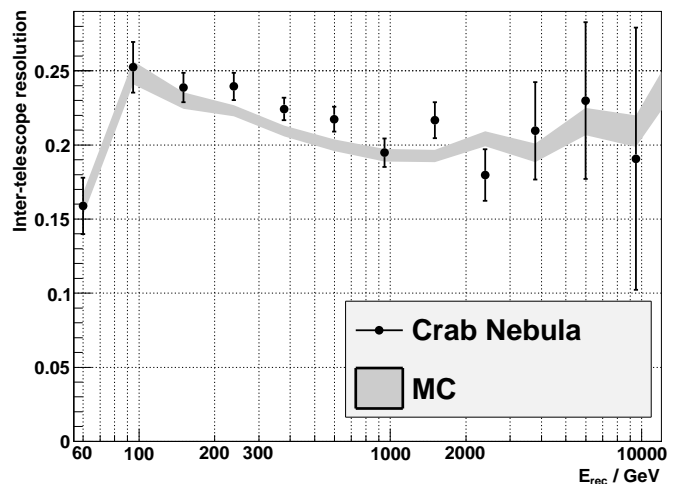


Figure 12: Standard deviation of the distribution of the relative difference between energy estimators from both telescopes as a function of the reconstructed energy for γ -ray MC (gray band) and Crab Nebula observations (black points).

be fitted with a curved power-law:

$$\frac{dN}{dE} = f_0 (E/1 \text{ TeV})^{a+b \log_{10}(E/1 \text{ TeV})} [\text{cm}^{-2} \text{s}^{-1} \text{TeV}^{-1}]. \quad (1)$$

The parameters of the fit are: $f_0 = (3.39 \pm 0.09_{\text{stat}}) \times 10^{-11}$, $a = -2.51 \pm 0.02_{\text{stat}}$, and $b = -0.21 \pm 0.03_{\text{stat}}$. The parameters of the spectral fit were obtained using the robust forward unfolding method which does not require regularization. The forward unfolding requires however an assumption on the spectral shape of the source, and is insensitive to spectral features. Therefore the individual spectral points were computed using the Bertero unfolding method (Albert et al., 2007). The fit parameters obtained from both unfolding methods are consistent.

The spectrum obtained by MAGIC after the upgrade is consistent within $\sim 25\%$ with the previous measurements of the Crab Nebula performed with other IACTs and earlier phases of the MAGIC telescopes.

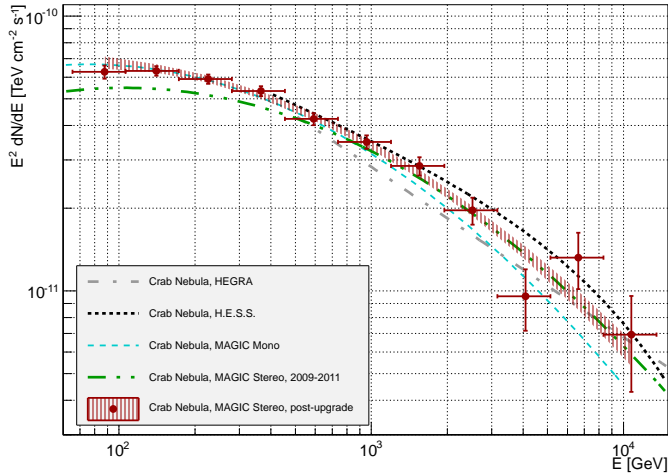


Figure 13: Spectral energy distribution of the Crab Nebula obtained with the MAGIC telescopes after the upgrade (red points and shading) compared with other experiments: MAGIC-I (blue dashed, Albert et al., 2008b), MAGIC Stereo, 2009-2011 (green dot-dot-dashed, Aleksić et al., 2014f), HEGRA (gray dot-dashed, Aharonian et al., 2004), and H.E.S.S. (black dotted, Aharonian et al., 2006). The vertical error bars show statistical errors, while the horizontal ones represent the energy binning.

4.6. Angular resolution

Following the approach in Aleksić et al. (2012a), we investigate the angular resolution of the MAGIC telescopes using two commonly used methods. In the first approach we define the angular resolution Θ_{Gaus} as the standard deviation of a 2-dimensional Gaussian fitted to the distribution of the reconstructed event directions of the γ -ray excess. Such a 2-dimensional Gaussian in the θ_x and θ_y space will correspond to an exponential fitting function for θ^2 distribution. The fit is performed in a narrow range, $\theta^2 < 0.025$, which is a factor ~ 2.5 larger than the typical signal extraction cut applied at medium energies. Therefore it is a good performance quantity for looking for small extensions (comparable with angular resolution) in VHE γ -ray sources. In the second method we compute an angular distance, $\Theta_{0.68}$, around the source, which encircles 68% of the excess events. This method is more sensitive to long tails in the distribution of reconstructed directions. Note that while both numbers assess the angular resolution of the MAGIC telescopes, their absolute values are different, normally $\Theta_{Gaus} < \Theta_{0.68}$. For a purely Gaussian distribution Θ_{Gaus} would correspond to only 39% containment radius of γ -rays originating from a point like source and $\Theta_{0.68} \approx 1.5 \Theta_{Gaus}$.

We use the low and medium zenith samples of the Crab Nebula to investigate the angular resolution. Since the Crab Nebula is a nearby galactic source, it might in principle have an intrinsic size which would artificially spoil the angular resolution measured in this way. However, the extension of the Crab Nebula in VHE γ -rays was constrained to below 0.025° (Aharonian et al., 2000), making it a point-like source for MAGIC.

The angular resolution obtained with both methods is shown in Fig. 14 and summarized in Table 2. At 250 GeV the angular resolution (from a 2D Gaussian fit) is 0.07° . It improves with energy, as larger images are better reconstructed, reach-

E [GeV]	Zenith $< 30^\circ$		Zenith $30 - 45^\circ$	
	$\Theta_{Gaus} [^\circ]$	$\Theta_{0.68} [^\circ]$	$\Theta_{Gaus} [^\circ]$	$\Theta_{0.68} [^\circ]$
95	0.087 ± 0.004	$0.157^{+0.007}_{-0.007}$	0.088 ± 0.013	$0.129^{+0.009}_{-0.021}$
150	0.075 ± 0.002	$0.135^{+0.005}_{-0.005}$	0.078 ± 0.005	$0.148^{+0.017}_{-0.013}$
238	0.067 ± 0.001	$0.108^{+0.004}_{-0.004}$	0.072 ± 0.003	$0.120^{+0.009}_{-0.007}$
378	0.058 ± 0.001	$0.095^{+0.004}_{-0.003}$	0.063 ± 0.003	$0.097^{+0.008}_{-0.006}$
599	0.052 ± 0.001	$0.081^{+0.003}_{-0.003}$	0.054 ± 0.003	$0.083^{+0.007}_{-0.007}$
949	0.046 ± 0.001	$0.073^{+0.004}_{-0.003}$	0.052 ± 0.002	$0.082^{+0.006}_{-0.005}$
1504	0.044 ± 0.001	$0.071^{+0.005}_{-0.003}$	0.046 ± 0.002	$0.077^{+0.007}_{-0.004}$
2383	0.042 ± 0.002	$0.067^{+0.006}_{-0.005}$	0.045 ± 0.003	$0.068^{+0.010}_{-0.006}$
3777	0.042 ± 0.003	$0.065^{+0.011}_{-0.004}$	0.039 ± 0.004	$0.061^{+0.011}_{-0.008}$
5986	0.041 ± 0.004	$0.062^{+0.012}_{-0.011}$	0.038 ± 0.006	$0.059^{+0.031}_{-0.011}$
9487	0.040 ± 0.005	$0.056^{+0.062}_{-0.012}$	0.046 ± 0.009	$0.055^{+0.209}_{-0.005}$

Table 2: Angular resolution Θ_{Gaus} and $\Theta_{0.68}$ of the MAGIC telescopes after the upgrade as a function of the estimated energy E , obtained with the Crab Nebula data sample. Θ_{Gaus} is computed as a sigma of a 2D Gaussian fit. $\Theta_{0.68}$ is the 68% containment radius of the γ -ray excess.

ing a plateau of $\sim 0.04^\circ$ above a few TeV. The angular resolution improved by about 5-10% after the upgrade. The improvement in angular resolution makes slightly more pronounced the small difference between the angular resolution obtained with MC simulations and the Crab Nebula data, also present in the pre-upgrade data. The difference of $\sim 10 - 15\%$ is visible at higher energies and corresponds to an additional 0.02° systematic random component (i.e. added in quadrature) between the MC and the data.

The distributions of angular distances between the true and reconstructed source position can be reasonably well fitted with a single Gaussian for $\theta^2 < 0.025$. Nevertheless, a proper description of the tail in the θ^2 distribution requires a more complicated function (see Fig. 15). One possibility is to use a combination of two two-dimensional Gaussian distributions as in Aleksić et al. (2012a). For example, the two-dimensional double Gaussian fit to the distribution shown in Fig. 15 for the Crab Nebula data yields $\chi^2/N_{\text{dof}} = 1.48/6$ corresponding to a probability of 96.1%.

4.7. Sensitivity

In order to provide a fast reference and comparison with other experiments we calculate the sensitivity of the MAGIC telescopes following the two commonly used definitions. For a weak source, the significance of an excess of N_{excess} events over a perfectly-well known background of N_{bkg} events can be computed with the simplified formula $N_{\text{excess}}/\sqrt{N_{\text{bkg}}}$. Therefore, one defines the sensitivity $S_{N_{\text{ex}}/\sqrt{N_{\text{bkg}}}}$ as the flux of a source giving $N_{\text{excess}}/\sqrt{N_{\text{bkg}}} = 5$ after 50 h of effective observation time. The sensitivity can also be calculated using the Li & Ma (1983), eq. 17 formula, which is the standard method in the VHE γ -ray astronomy for the calculation of the significances. Note that the sensitivity computed according to the Li & Ma formula will depend on the number of OFF positions used for background estimation.

For a more realistic estimation of the sensitivity (in both methods), we apply conditions $N_{\text{excess}} > 10$ and $N_{\text{excess}} >$

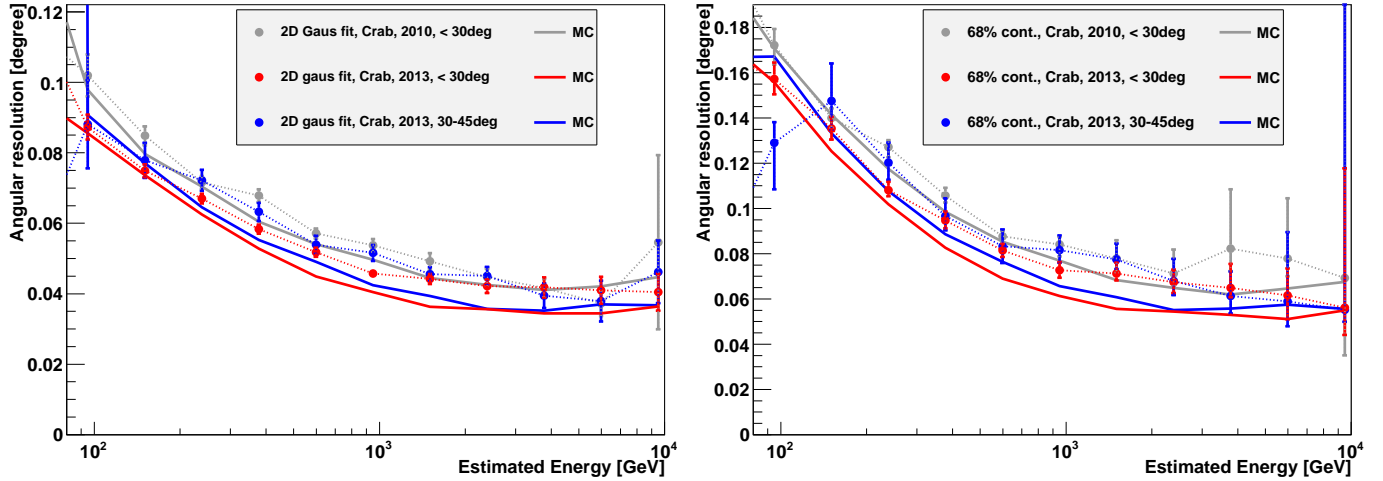


Figure 14: Angular resolution of the MAGIC telescopes after the upgrade as a function of the estimated energy obtained with the Crab Nebula data sample (points) and MC simulations (solid lines). Left panel: 2D Gaussian fit, right panel: 68% containment radius. Red points: low zenith sample, blue points: medium zenith sample. For comparison the low zenith pre-upgrade angular resolution is shown as gray points Aleksić et al. (2012a).

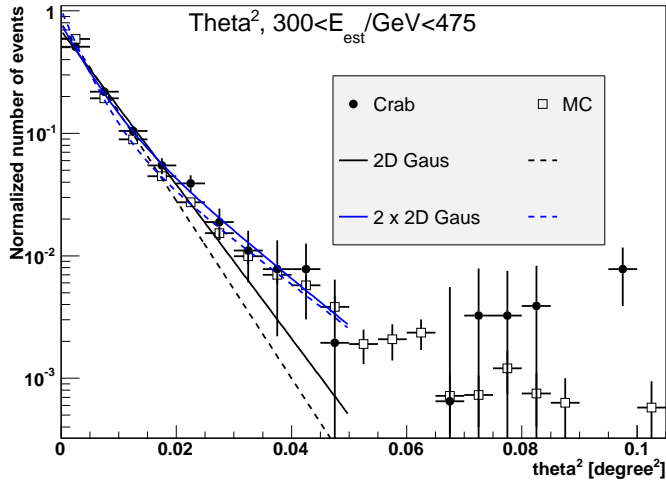


Figure 15: θ^2 distribution of excess events for the Crab Nebula (filled circles, solid lines) and MC (empty squares, dashed lines) samples in the energy range of 300 – 475 GeV. The distributions are fitted with a single or a double two dimensional Gaussian (black and blue lines respectively).

$0.05N_{\text{bkg}}$. The first condition assures that the Poissonian statistics of the number of events can be approximated by a Gaussian distribution. The second condition protects against small systematic discrepancies between the ON and OFF distributions, which may mimic a statistically significant signal if the residual background rate is large.

The integral sensitivity of the different phases of the MAGIC experiment for a source with a Crab Nebula-like spectrum are shown in Fig. 16. The sensitivity values both in Crab Nebula Units (C.U.) and in absolute units are summarized in Table 3 for low zenith and in Table 4 for medium zenith angles. We used here the $N_{\text{excess}}/\sqrt{N_{\text{bkg}}} = 5$ definition, recomputing the original MAGIC-I mono sensitivities to include also the $N_{\text{excess}} > 10$ and $N_{\text{excess}} > 0.05N_{\text{bkg}}$ conditions ¹.

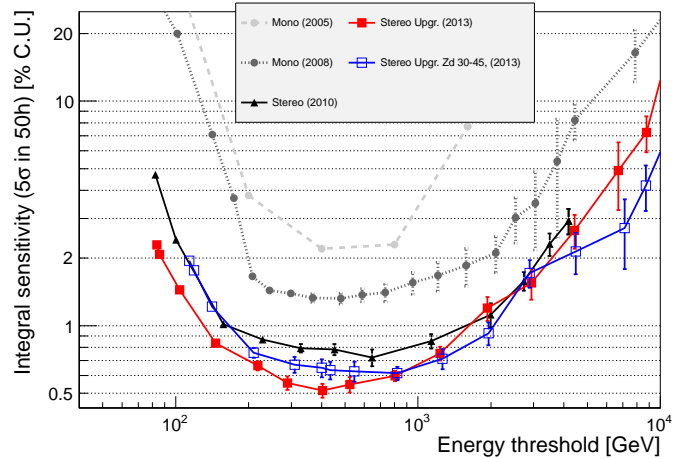


Figure 16: Evolution of integral sensitivity of the MAGIC telescopes, i.e. the integrated flux of a source above a given energy for which $N_{\text{excess}}/\sqrt{N_{\text{bkg}}} = 5$ after 50 h of effective observation time, requiring $N_{\text{excess}} > 10$ and $N_{\text{excess}} > 0.05N_{\text{bkg}}$. Gray circles: sensitivity of the MAGIC-I single telescope with the Siegen (light gray, long dashed, Albert et al. (2008b)) and MUX readouts (dark gray, short dashed, Aleksić et al. (2012a)). Black triangles: stereo before the upgrade (Aleksić et al., 2012a). Squares: stereo after the upgrade: zenith angle below 30° (red, filled), 30 – 45° (blue, empty)

In order to find the optimal cut values in *Hadronness* and θ^2 in an unbiased way, we used an independent *training* sample of Crab Nebula data. The size of the *training* sample is similar to the size of the *test* sample from which the final sensitivity is computed. Different energy thresholds are achieved by varying a cut in the total number of photoelectrons of the images (for points < 300 GeV) or in the estimated energy of the events (above 300 GeV). For each energy threshold we perform a scan of cuts on the *training* subsample, and apply the best cuts to

the very poor signal-to-background ratio at low energies, leading to dramatic worsening of the sensitivity. Using optimized cuts one can recover some of the sensitivity lost at the lowest energies for mono observations.

¹Note that one of the main disadvantages of the mono observations was

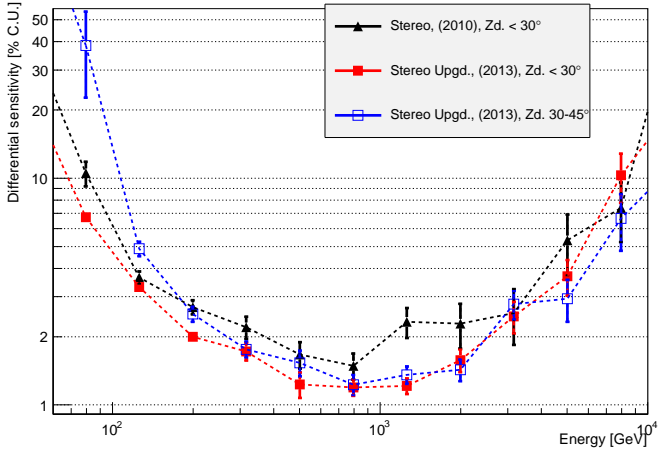


Figure 17: Differential (5 bins per decade in energy) sensitivity of the MAGIC Stereo system. We compute the flux of the source in a given energy range for which $N_{\text{excess}}/\sqrt{N_{\text{bkg}}} = 5$ with $N_{\text{excess}} > 10$, $N_{\text{excess}} > 0.05N_{\text{bkg}}$ after 50h of effective time.

the main sample obtaining the sensitivity value. The threshold itself is estimated as the peak of true energy distribution of MC events with a -2.6 spectral slope to which the same cuts were applied.

The integral sensitivity evaluated above is valid only for sources with a Crab Nebula-like spectrum. To assess the performance of the MAGIC telescopes for sources with an arbitrary spectral shape, we compute the differential sensitivity. Following the commonly used definition, we calculate the sensitivity in narrow bins of energy (5 bins per decade). The differential sensitivity is plotted for low and medium zenith angles in Fig. 17, and the values are summarized in Table 5 and 6 respectively.

The upgrade of the MAGIC-I camera and readout of the MAGIC telescopes has led to a significant improvement in sensitivity in the whole investigated energy range. The integral sensitivity reaches down to about 0.55% of C.U. around a few hundred GeV in 50h of observations. The improvement in the performance is especially evident at the lowest energies. In particular, in the energy bin 60-100 GeV, the differential sensitivity decreased from 10.5% C.U. to 6.7% C.U. reducing the needed observation time by a factor of 2.5. Observations at medium zenith angle have naturally higher energy threshold. Therefore the performance at the lowest energies is marred. Some of the sources, those with declination $> 58^\circ$, or $< -2^\circ$ can only be observed by MAGIC at medium or high zenith angles. Sources with declination between -2° and 58° , can be observed either at low zenith angles, or at medium zenith with a boost in sensitivity at TeV energies at the cost of a higher energy threshold.

The sensitivity of IACTs clearly depends on the observation time which can be spent observing a given source. In particular for transient sources, such as gamma-ray bursts or flares from Active Galactic Nuclei, it is not feasible to collect 50h of data within the duration of such an event. On the other hand, long, multi-year campaigns allow to gather of the order of hundred of hours (see e.g. ~ 140 h observations of M82

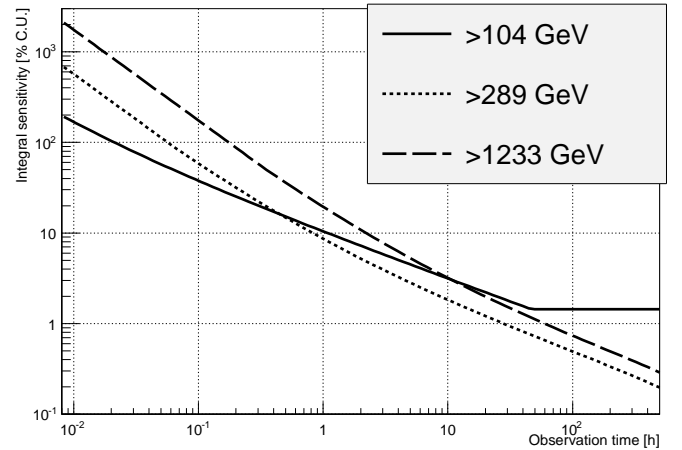


Figure 18: Dependence of the integral sensitivity of the MAGIC telescopes (computed according to $S_{\text{Li\&Ma,3Off}}$ prescription, see text for details) on the observation time, obtained with the low zenith Crab Nebula sample. Different line styles show different energy thresholds: > 104 GeV (solid), > 289 GeV (dotted), > 1233 GeV (dashed).

by VERITAS, Acciari et al. (2009) or ~ 160 h observations of Segue by MAGIC, Aleksić et al. (2014e)). In Fig. 18, using the γ and background rates from Table 3, we show how the sensitivity of the MAGIC telescopes depends on the observation time for different energy thresholds. For those exemplary calculations we use the Li&Ma definition of sensitivity with typical value of 3 Off positions for background estimation. In the medium range of observation times the sensitivity follows the usual $\propto 1/\sqrt{\text{time}}$ dependence. For very short observation times, especially for higher energies where the γ /hadron separation is very powerful, the limiting condition of at least 10 excess events leads to a dependence of $\propto 1/\text{time}$. On the other hand, for very long observations the sensitivity saturates at low energies. Note that the observation time at which the sensitivity saturates might be shifted by using stronger cuts, offering better γ to background rate, however at the price of increased threshold.

4.8. Off-axis performance

Most of the observations of the MAGIC telescopes are performed in the wobble mode with the source offset of 0.4° from the camera center. However, in the case of micro-scans of extended sources with sizes much larger than the PSF of the MAGIC telescopes, a γ -ray signal might be found at different distances from the camera center (see e.g. Aleksić et al., 2014c). Moreover, serendipitous sources (see e.g. detection of IC 310, Aleksić et al., 2010b) can occur in the FoV of MAGIC at an arbitrary angular offset from the pointing direction. Therefore, we study the performance of the MAGIC telescopes at different offsets from the center of the FoV with dedicated observations of the Crab Nebula at non-standard wobble offsets (see Table 1).

For easy comparison with the results presented in Aleksić et al. (2012a), we first compute the integral sensitivities as a function of the wobble offset at the same energy threshold of 290 GeV (see Fig. 19). The upgrade of the MAGIC tele-

$E_{\text{thresh.}}$ [GeV]	γ -rate [min ⁻¹]	bkg-rate [min ⁻¹]	$S_{N_{\text{ex}}/\sqrt{Nbkg}}$ [% C.U.]	$S_{\text{Li\&Ma,1Off}}$ [% C.U.]	$S_{\text{Li\&Ma,3Off}}$ [% C.U.]	$S_{\text{Li\&Ma,5Off}}$ [% C.U.]	$S_{N_{\text{ex}}/\sqrt{Nbkg}}$ [10 ⁻¹³ cm ⁻² s ⁻¹]
84	19.06 ± 0.21	8.73 ± 0.07	2.291 ± 0.025	2.29 ± 0.03	2.29 ± 0.03	2.29 ± 0.03	156.5 ± 1.7
86	18.80 ± 0.21	7.80 ± 0.06	2.073 ± 0.022	2.07 ± 0.03	2.07 ± 0.03	2.07 ± 0.03	137.1 ± 1.5
104	16.88 ± 0.19	4.88 ± 0.05	1.445 ± 0.015	1.714 ± 0.020	1.445 ± 0.022	1.445 ± 0.022	75.9 ± 0.8
146	6.17 ± 0.10	0.320 ± 0.013	0.837 ± 0.023	1.25 ± 0.03	1.004 ± 0.024	0.948 ± 0.022	28.6 ± 0.8
218	3.63 ± 0.07	0.070 ± 0.006	0.66 ± 0.03	1.06 ± 0.04	0.83 ± 0.03	0.78 ± 0.03	13.4 ± 0.7
289	2.94 ± 0.07	0.032 ± 0.004	0.56 ± 0.04	0.93 ± 0.05	0.72 ± 0.04	0.67 ± 0.04	7.6 ± 0.5
404	3.05 ± 0.07	0.030 ± 0.004	0.51 ± 0.04	0.87 ± 0.04	0.67 ± 0.04	0.63 ± 0.03	4.4 ± 0.3
523	2.51 ± 0.06	0.023 ± 0.003	0.55 ± 0.04	0.95 ± 0.05	0.72 ± 0.04	0.68 ± 0.05	3.2 ± 0.3
803	1.59 ± 0.05	0.0109 ± 0.0010	0.60 ± 0.03	1.12 ± 0.04	0.84 ± 0.03	0.78 ± 0.03	1.78 ± 0.10
1233	0.95 ± 0.04	0.0062 ± 0.0007	0.75 ± 0.05	1.53 ± 0.06	1.11 ± 0.05	1.02 ± 0.05	1.12 ± 0.08
1935	0.55 ± 0.03	0.0053 ± 0.0011	1.20 ± 0.15	2.50 ± 0.16	1.80 ± 0.13	1.66 ± 0.15	0.83 ± 0.10
2938	0.307 ± 0.022	0.0027 ± 0.0008	1.55 ± 0.24	3.7 ± 0.3	2.55 ± 0.25	2.32 ± 0.24	0.51 ± 0.08
4431	0.160 ± 0.016	0.0022 ± 0.0006	2.7 ± 0.5	6.6 ± 0.6	4.5 ± 0.5	4.1 ± 0.4	0.41 ± 0.07
6718	0.078 ± 0.011	0.0018 ± 0.0010	4.9 ± 1.6	12.9 ± 1.7	8.7 ± 1.5	7.8 ± 1.4	0.34 ± 0.11
8760	0.046 ± 0.008	0.0005 ± 0.0005	7.2 ± 1.3	16.9 ± 2.3	10.4 ± 1.7	9.0 ± 1.8	0.30 ± 0.05

Table 3: Integral sensitivity of the MAGIC telescopes obtained with the low zenith Crab Nebula data sample above a given energy threshold, $E_{\text{thresh.}}$. The sensitivity is calculated as $N_{\text{excess}} / \sqrt{Nbkg} = 5 (S_{N_{\text{ex}}/\sqrt{Nbkg}})$, or according to the 5σ significance obtained from Li & Ma (1983) (using 1, 3 or 5 background regions, $S_{\text{Li\&Ma,1Off}}$, $S_{\text{Li\&Ma,3Off}}$ and $S_{\text{Li\&Ma,5Off}}$). The sensitivity is computed for 50 h of observation time with the additional conditions $N_{\text{excess}} > 10$, $N_{\text{excess}} > 0.05Nbkg$ ($S_{N_{\text{ex}}/\sqrt{Nbkg,sys}}$). The γ -rate and bkg-rate columns show the rate of γ events from Crab Nebula and residual background respectively above the energy threshold of the sensitivity point.

scopes has improved the off-axis performance. For example, the sensitivity at offsets of $\sim 1^\circ$ has improved by $\sim 25\%$, which is more than the global 15% improvement seen at the usual offset of $\sim 0.4^\circ$. Interestingly there is not much difference in the γ rates associated with these sensitivity values. This suggests that most of the improvement in sensitivity comes from a better image reconstruction, possibly thanks to the higher pixelization of the new MAGIC-I camera, rather than from triggering more events due to larger trigger region.

4.9. Extended sources

Some of the sources might have an intrinsic extension. The sensitivity for detection such sources is spoiled for two reasons. First of all, the signal is diluted over a larger part of the sky. This forces us to loosen the angular θ^2 cut and hence accept more background. The new cut value can be roughly estimated to be

$$\theta_{\text{cut}} = \sqrt{\theta_0^2 + \theta_s^2}, \quad (2)$$

where θ_0 is a cut for a point-like source analysis, and θ_s is a characteristic source size. As the background events show a nearly flat distribution of $dN/d\theta^2$ such a cut will increase the background by θ_{cut}^2 . The looser θ^2 cut will affect the sensitivity in two ways. First of all, the sensitivity will be spoiled by a factor of $\sqrt{\theta_{\text{cut}}^2} = \theta_{\text{cut}}$ due to accepting more background events. Note however that the acceptance of the θ_{cut} cut for γ -rays can be larger than the acceptance of θ_0 cut for a point like source, as the γ -rays originating from the center of the source will still be accepted even if they are strongly misreconstructed.

For the further calculations let us assume $\theta_0 = 0.1$ (comparable to $\theta_{0,68}$ at the energies of a few hundred GeV, see Section 4.6) and a PSF shape as described by Fig. 15 (note that the

PSF is not much affected by the distance from the center of the camera, Aleksić et al. (2012a)) and a source with a flat surface brightness up to θ_s . In Fig. 20 we show the acceptance for background and γ events. We also compute a sensitivity “spoiling factor”, defined as the square root of the background acceptance divided by the γ acceptance and normalized to 1 for a point like source. As an example, let us assume a source with a radius of 0.5° . The optimal cut $\theta_s = 0.51$ computed according to Eq. 2 results in 26 times larger background than with cut $\theta_0 = 0.1$. This would correspond to ≈ 5 times worse sensitivity, however the cut contains $\approx 90\%$ of γ events, significantly larger than $\approx 70\%$ efficiency for a point like cut. Therefore the sensitivity is spoiled by a smaller factor, ≈ 4 .

A second effect which can spoil the sensitivity for extended sources is the loss of collection area for higher offsets from the camera center. For a source radius of e.g. 0.5° , the γ -rays can be observed up to an offset of 0.9° from the camera center. For such large offsets, the collection area is nearly a factor of 3 smaller than in the camera center. Using the γ -rates, which are proportional to the collection area, shown in Fig. 19 we can compute the average rate of γ rays for an arbitrary source profile. For this example of a source with constant surface density and a radius of 0.5° it turns out that the total average collection area is lower only by $\approx 20\%$ than for a point like source at the usual wobble offset of 0.4° . However, since a similar drop happens also for the background events, the net spoiling of the sensitivity due to this effect is only $\sim 10\%$.

Finally, we compute the radius Ξ of the MAGIC *effective field of view*. It is defined such that observations of an isotropic gamma-ray flux with a hypothetical instrument with a flat-top acceptance $R'(\xi) = R(0)$ for $\xi < \Xi$, and $R'(\xi) = 0$ for $\xi > \Xi$, would yield the same number of detected gamma

$E_{\text{thresh.}}$ [GeV]	γ -rate [min^{-1}]	bkg-rate [min^{-1}]	$S_{Nex/\sqrt{Nbkg}}$ [%C.U.]	$S_{\text{Li\&Ma,1Off}}$ [%C.U.]	$S_{\text{Li\&Ma,3Off}}$ [%C.U.]	$S_{\text{Li\&Ma,5Off}}$ [%C.U.]	$S_{Nex/\sqrt{Nbkg}}$ [$10^{-13} \text{ cm}^{-2} \text{ s}^{-1}$]
114	19.7 ± 0.4	7.68 ± 0.10	1.95 ± 0.03	1.95 ± 0.04	1.95 ± 0.04	1.95 ± 0.04	91.4 ± 1.6
119	19.3 ± 0.3	6.83 ± 0.10	1.76 ± 0.03	1.77 ± 0.03	1.76 ± 0.04	1.76 ± 0.04	78.4 ± 1.3
141	17.4 ± 0.3	4.23 ± 0.08	1.217 ± 0.021	1.55 ± 0.03	1.26 ± 0.03	1.22 ± 0.03	43.5 ± 0.8
210	5.60 ± 0.16	0.216 ± 0.017	0.76 ± 0.04	1.15 ± 0.05	0.92 ± 0.04	0.86 ± 0.04	16.1 ± 0.8
310	3.59 ± 0.12	0.070 ± 0.010	0.67 ± 0.05	1.07 ± 0.07	0.84 ± 0.06	0.79 ± 0.06	8.3 ± 0.7
401	3.19 ± 0.12	0.051 ± 0.008	0.65 ± 0.06	1.05 ± 0.08	0.82 ± 0.07	0.77 ± 0.06	5.6 ± 0.5
435	3.31 ± 0.12	0.053 ± 0.009	0.63 ± 0.06	1.03 ± 0.08	0.80 ± 0.07	0.75 ± 0.06	4.8 ± 0.4
546	2.98 ± 0.11	0.042 ± 0.008	0.63 ± 0.06	1.03 ± 0.08	0.80 ± 0.07	0.75 ± 0.06	3.4 ± 0.3
821	2.24 ± 0.10	0.0227 ± 0.0023	0.61 ± 0.04	1.06 ± 0.05	0.81 ± 0.05	0.76 ± 0.04	1.77 ± 0.12
1262	1.19 ± 0.07	0.0086 ± 0.0014	0.71 ± 0.07	1.38 ± 0.09	1.02 ± 0.08	0.94 ± 0.07	1.02 ± 0.10
1955	0.84 ± 0.06	0.0072 ± 0.0013	0.93 ± 0.11	1.84 ± 0.13	1.35 ± 0.11	1.24 ± 0.10	0.63 ± 0.07
2891	0.60 ± 0.05	0.013 ± 0.003	1.72 ± 0.24	3.2 ± 0.3	2.37 ± 0.24	2.2 ± 0.3	0.58 ± 0.08
4479	0.31 ± 0.04	0.0052 ± 0.0017	2.1 ± 0.4	4.5 ± 0.6	3.2 ± 0.5	3.0 ± 0.4	0.32 ± 0.07
7133	0.140 ± 0.024	0.0017 ± 0.0010	2.7 ± 0.9	7.1 ± 1.0	4.8 ± 0.9	4.3 ± 0.9	0.17 ± 0.06

Table 4: Integral sensitivity of the MAGIC telescopes obtained with the medium zenith ($30 - 45^\circ$) Crab Nebula data sample above a given energy threshold $E_{\text{thresh.}}$. Columns as in Table 3.

E_{min} [GeV]	E_{max} [GeV]	γ -rate [min^{-1}]	bkg-rate [min^{-1}]	$S_{Nex/\sqrt{Nbkg}}$ [%C.U.]	$S_{\text{Li\&Ma,1Off}}$ [%C.U.]	$S_{\text{Li\&Ma,3Off}}$ [%C.U.]	$S_{\text{Li\&Ma,5Off}}$ [%C.U.]	$S_{Nex/\sqrt{Nbkg}}$ [$10^{-12} \text{ cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$]
63	100	3.01 ± 0.13	4.06 ± 0.08	6.74 ± 0.23	8.8 ± 0.4	7.1 ± 0.3	6.8 ± 0.3	728 ± 25
100	158	4.29 ± 0.12	2.41 ± 0.06	3.31 ± 0.12	4.77 ± 0.14	3.87 ± 0.11	3.67 ± 0.10	137 ± 5
158	251	3.37 ± 0.08	0.54 ± 0.03	2.00 ± 0.08	2.95 ± 0.10	2.38 ± 0.08	2.25 ± 0.08	30.5 ± 1.3
251	398	1.36 ± 0.05	0.066 ± 0.010	1.72 ± 0.15	2.75 ± 0.20	2.16 ± 0.16	2.03 ± 0.15	9.3 ± 0.8
398	631	1.22 ± 0.04	0.027 ± 0.006	1.23 ± 0.16	2.10 ± 0.18	1.61 ± 0.18	1.51 ± 0.15	2.3 ± 0.3
631	1000	0.88 ± 0.04	0.0133 ± 0.0018	1.19 ± 0.10	2.18 ± 0.12	1.64 ± 0.09	1.53 ± 0.11	0.72 ± 0.06
1000	1585	0.58 ± 0.03	0.0059 ± 0.0007	1.21 ± 0.10	2.48 ± 0.11	1.80 ± 0.09	1.66 ± 0.09	0.230 ± 0.018
1585	2512	0.299 ± 0.021	0.0027 ± 0.0005	1.58 ± 0.18	3.75 ± 0.22	2.60 ± 0.19	2.36 ± 0.18	0.090 ± 0.010
2512	3981	0.166 ± 0.016	0.0020 ± 0.0005	2.5 ± 0.4	6.2 ± 0.5	4.3 ± 0.4	3.8 ± 0.4	0.041 ± 0.007
3981	6310	0.093 ± 0.012	0.0014 ± 0.0003	3.7 ± 0.7	10.2 ± 1.0	6.8 ± 0.7	6.1 ± 0.7	0.017 ± 0.003
6310	10000	0.060 ± 0.010	0.0046 ± 0.0015	10 ± 3	22 ± 3	16 ± 3	14.5 ± 2.3	0.013 ± 0.003

Table 5: Differential sensitivity of the MAGIC telescopes obtained with the low zenith observations of Crab Nebula data sample. The definitions of the sensitivities are as in Table 3. The γ -rate and bkg-rate columns show the rate of γ events from Crab Nebula and residual background respectively in the differential estimated energy bins.

E_{min} [GeV]	E_{max} [GeV]	γ -rate [min^{-1}]	bkg-rate [min^{-1}]	$S_{Nex/\sqrt{Nbkg}}$ [%C.U.]	$S_{\text{Li\&Ma,1Off}}$ [%C.U.]	$S_{\text{Li\&Ma,3Off}}$ [%C.U.]	$S_{\text{Li\&Ma,5Off}}$ [%C.U.]	$S_{Nex/\sqrt{Nbkg}}$ [$10^{-12} \text{ cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$]
63	100	0.40 ± 0.12	2.92 ± 0.11	39 ± 16	56 ± 16	45 ± 12	43 ± 11	4200 ± 1700
100	158	3.18 ± 0.16	2.89 ± 0.05	4.9 ± 0.4	7.0 ± 0.4	5.7 ± 0.3	5.4 ± 0.3	202 ± 15
158	251	2.67 ± 0.19	0.54 ± 0.04	2.52 ± 0.19	3.7 ± 0.3	2.99 ± 0.25	2.83 ± 0.22	38 ± 3
251	398	2.86 ± 0.13	0.305 ± 0.019	1.76 ± 0.14	2.64 ± 0.14	2.11 ± 0.11	2.00 ± 0.10	9.5 ± 0.8
398	631	1.76 ± 0.12	0.088 ± 0.006	1.53 ± 0.20	2.41 ± 0.16	1.90 ± 0.14	1.79 ± 0.13	2.8 ± 0.4
631	1000	1.44 ± 0.09	0.0376 ± 0.0021	1.23 ± 0.13	2.04 ± 0.12	1.58 ± 0.09	1.48 ± 0.10	0.74 ± 0.08
1000	1585	0.94 ± 0.08	0.0197 ± 0.0016	1.36 ± 0.12	2.38 ± 0.16	1.81 ± 0.13	1.69 ± 0.13	0.257 ± 0.023
1585	2512	0.67 ± 0.06	0.0111 ± 0.0015	1.43 ± 0.16	2.67 ± 0.22	2.00 ± 0.19	1.85 ± 0.18	0.082 ± 0.009
2512	3981	0.32 ± 0.05	0.0093 ± 0.0012	2.8 ± 0.4	5.3 ± 0.7	3.9 ± 0.6	3.7 ± 0.5	0.046 ± 0.007
3981	6310	0.20 ± 0.04	0.0042 ± 0.0017	2.9 ± 0.6	6.4 ± 1.2	4.6 ± 0.9	4.2 ± 0.9	0.014 ± 0.003
6310	10000	0.10 ± 0.03	0.00522 ± 0.00023	6.7 ± 1.9	14 ± 3	10 ± 3	9.2 ± 2.1	0.0083 ± 0.0023

Table 6: Differential sensitivity of the MAGIC telescopes obtained with the medium zenith ($30 - 45^\circ$) Crab Nebula data sample. Columns as in Table 5.

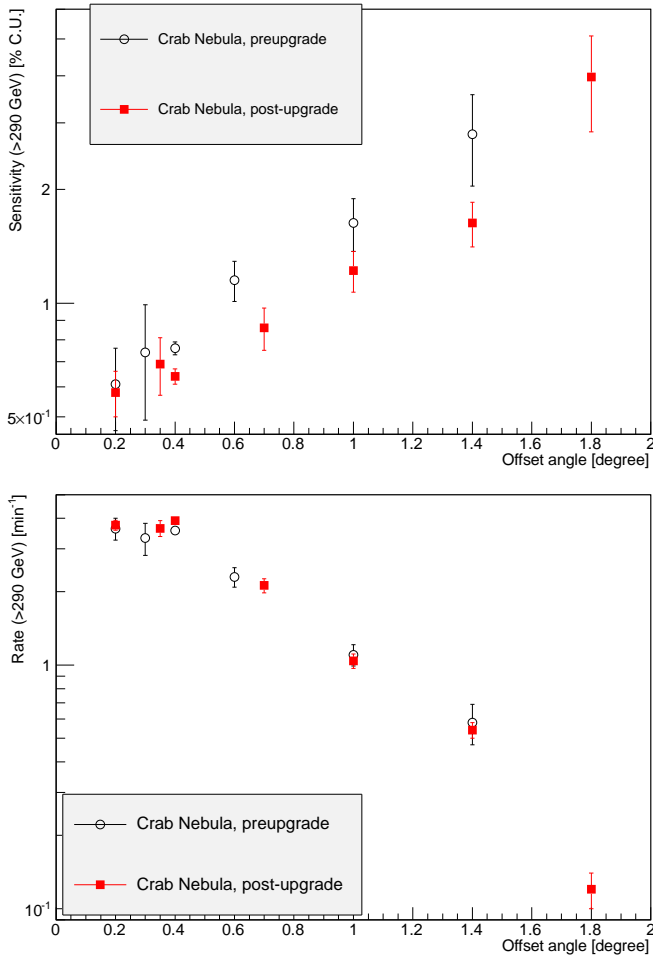


Figure 19: Top panel: integral sensitivity above 290 GeV for low zenith angle observations at different offsets, ξ , from the camera center. Bottom panel: corresponding γ -ray rates $R(\xi)$. Black empty circles: data from before the upgrade (Aleksić et al., 2012a), Red filled squares: current data (see Table 1)

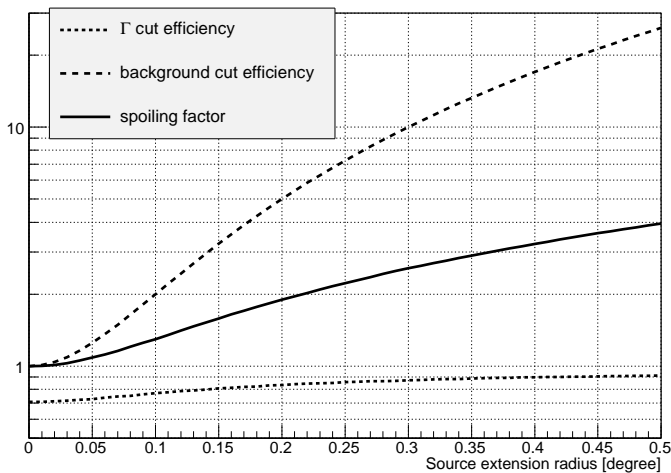


Figure 20: Dashed line: dependence of the amount of background integrated up to a cut determined by Eq. 2 as a function of the radius of the source, normalized to the background for a point like source. Dotted line: fraction of the total γ events contained within the cut. Solid line: sensitivity for an extended source divided by sensitivity for a point like source.

rays as with MAGIC, when no cuts on the arrival direction are applied. We can therefore obtain Ξ from the condition $\int_0^\Xi 2\pi\xi R(0) d\xi = \int_0^{1.8^\circ} 2\pi\xi R(\xi) d\xi$, where $R(\xi)$ is shown in bottom panel of Fig. 19, yielding $\Xi = 1^\circ$. We note, however, that standard observations of sources with an extension larger than 0.4° are technically difficult, as in that case the edge of the source would fall into the background estimation region. Nevertheless, the *effective field of view* is a useful quantity for non-standard observations of diffuse signals like, e.g. the cosmic electron flux (Borla Tridon, 2011; Aharonian et al., 2009).

5. Systematic uncertainties

The systematic errors of the IACT technique stem from many small individual factors which are only known with limited precision, and possibly change from one night to another. Most of those factors (e.g. uncertainties connected with the atmosphere, reflectivity of the mirrors) were not affected by the upgrade and hence the values reported in Aleksić et al. (2012a) are still valid for them. In this section we evaluate the component of the total systematic uncertainty of the MAGIC telescopes which changed for observations after the upgrade. We also estimate the total systematic uncertainty for various observation conditions.

5.1. Background subtraction

Dead pixels and stars in the field of view of the telescopes cause a small inhomogeneity in the distribution of the events in the camera plane. In addition, stereoscopy with just two telescopes produces a natural inhomogeneity, with the distribution of events being slightly dependent on the position of the second telescope. This effect was especially noticeable before the upgrade, due to the smaller trigger area of the old MAGIC-I camera. Both of these effects result in a slight rotational asymmetry of the camera acceptance. The effect is minimized by wobbling such that the source and background estimation positions in the camera are being swapped. Before the upgrade the systematic uncertainty of the background determination was $\lesssim 2\%$ (Aleksić et al., 2012a). Due to the larger trigger region this uncertainty has been reduced now to $\lesssim 1\%$.

Note that the effect of the background uncertainty depends on the signal to background ratio. In case of a strong source, where the signal \gtrsim background, it is negligible. However for a very weak source, with e.g. a signal to background ratio of $\sim 5\%$, the additional systematic error on the flux normalization just from the uncertainty of the background can amount up to $\sim 20\%$. Moreover, as the error will be energy dependent, it might lead to an additional uncertainty in the spectral index. Let us consider a hypothetical weak source with a spectrum reconstructed between E_{\min} and E_{\max} . The signal to background ratio of this source is SBR_{LE} and SBR_{HE} in the energy range $E_{\min} - \sqrt{E_{\min}E_{\max}}$ and $\sqrt{E_{\min}E_{\max}} - E_{\max}$ respectively. In this case we can roughly estimate the systematic uncertainty on the spectral index due to background inhomogeneity as:

$$\Delta\alpha_{SBR} = 2 \times \frac{\sqrt{(1\%/SBR_{LE})^2 + (1\%/SBR_{HE})^2}}{\ln(E_{\max}/E_{\min})}, \quad (3)$$

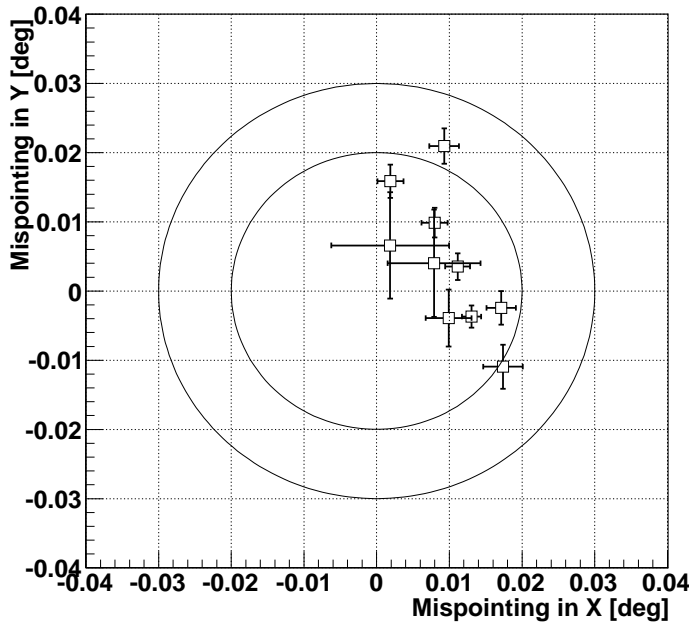


Figure 21: The difference in the reconstructed and nominal source position of the Crab Nebula during 10 nights of observations. The two circles show a distance of 0.02° and 0.03° .

where the 1% comes from the precision at which the background is estimated. Note that this formula was already used e.g. in Aleksić et al. (2012b, 2014b), however with a larger value of background uncertainty. For a strong source observed in a broad energy range, Eq. 3 gives a negligible number, e.g. for source observed between 0.08-6 TeV with a SBR of 25-60% $\Delta\alpha_{\text{SBR}} = 0.02$. However for a weak source, e.g. observed between 0.1 and 0.6 TeV with SBR = 6-15% we obtain $\Delta\alpha_{\text{SBR}} = 0.2$, increasing the total systematic error on the spectral index from 0.15 to 0.25.

5.2. Pointing accuracy

During the observations, the gravitational loads lead to a slight deformation of the telescopes structure and sagging of the camera. Most of the effect is corrected by the active mirror control and simultaneous observations of reference stars with CCD cameras placed at the center of the telescope's reflectors. However, a slight residual mispointing of the telescopes can affect the precision with which the position of a γ -ray source can be obtained. In order to evaluate this effect we analyse the Crab Nebula data night by night. For each night we construct a two dimensional distribution of the difference between the reconstructed and nominal source position in camera coordinates. After subtracting the background, the distribution is fitted with a two dimensional Gaussian to determine a possible systematic offset (see Fig. 21). Note that the 'Y' direction in the camera correspond to the Zenith axis of the telescope. We conclude that the systematic error on the reconstructed source position is $\lesssim 0.02^\circ$, comparable to value obtained in Aleksić et al. (2010a).

5.3. Energy scale

The absolute energy scale of IACTs is hard to determine. There are many systematics effects such as imprecise knowledge of the atmospheric transmission, mirror reflectivity, properties of the PMTs, etc. which affect it. If determined solely from the best knowledge of these parameters, it is expected to be accurate within 15-17% as shown in Aleksić et al. (2012a). The knowledge of the absolute energy scale of the MAGIC telescopes is validated by using inter-telescope calibration (see Section 4.3) and from the analysis of muon rings (Vacanti et al., 1994; Goebel et al., 2005). Both of those methods are burdened by their own systematic uncertainties. The inter-calibration of the telescopes using γ -ray events with similar *Impact* parameters improves greatly the relative size scale, however has a very weak handle on the absolute light scale of both telescopes. On the other hand, the muon analysis suffers from the fact that the light seen from muon events is emitted mostly up to the height of ~ 2 km above the telescopes, while the light from γ -ray events is generated mainly at the height of 10 km. This in addition makes the light spectrum of muon events shifted to lower wavelengths. Therefore the muon calibration has to be MC dependent.

A small miscalibration of the energy scale will affect the reconstructed spectrum in two ways, both having the highest impact at the lowest energies. For example, let us assume that MC simulations have a higher light scale, i.e. for a given energy the amount of light generated by a γ -ray shower at a given energy predicted by the MC simulations is larger than for real showers. In this case, real γ -ray showers will have a smaller *Size* parameter than predicted by the MC simulations. Thus some of them may not survive a data quality size cut, or not even trigger the telescopes lowering the real collection area w.r.t. the one predicted from MC simulations. This would artificially lower the reconstructed flux, especially at the low energies. On the other hand, the overestimation of the light shed into the cameras by the shower will introduce an unrecoverable bias in the energy estimation migrating events to lower energies. For the medium energies where the collection area is quite flat and the usual γ -ray spectra drop rapidly this will also artificially lower the reconstructed spectrum. For the lowest energies, below the analysis threshold, where the collection area drops very fast, more events will migrate in a given energy bin than escape from it, artificially increasing the reconstructed flux by a pile-up effect. It is not obvious to determine which of those effects will be dominant at a given energy. It will depend on the precise shape of the collection area, the spectrum of the source and the total miscalibration of the light scale. Therefore, the systematic error in the light scale can shift the spectrum at the lowest energies in both ways.

As the effect is mostly pronounced at the threshold we investigate the uncertainty of the light scale of the MAGIC telescopes by shifting the energy threshold in two different ways. We perform a set of full analyzes using MC simulations with the light scale artificially shifted by -25%, -10%, -5%, +5%, +10% and +25%. In the first scenario, we compare the reconstructed flux at a given energy for low and medium zenith angles (see Fig. 22). For the sake of comparison, we show in Fig. 22 the

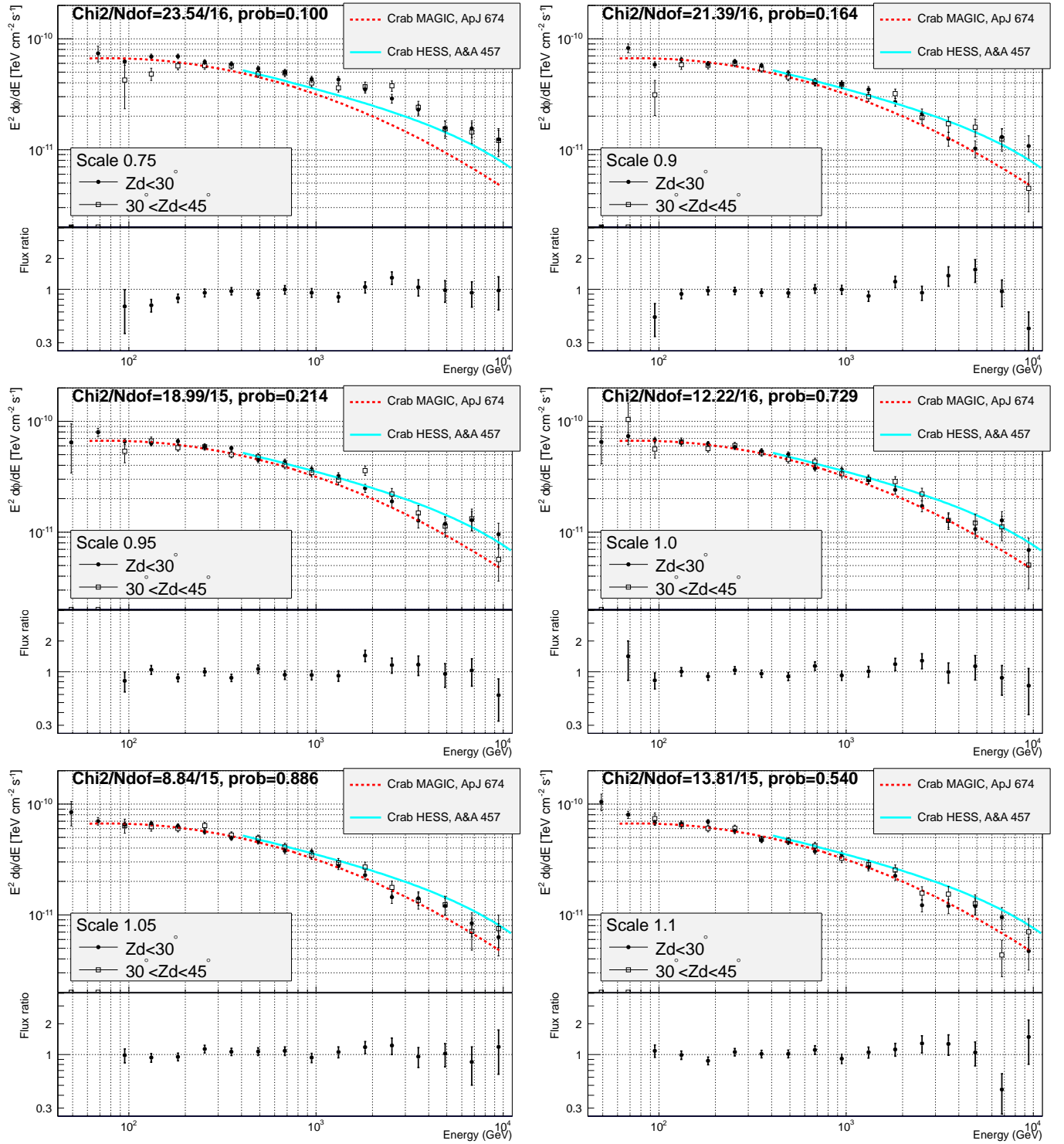


Figure 22: Spectra of the Crab Nebula obtained with the low (full circles) and medium (empty squares) zenith samples for MC with different energy scales (-25%, -10%, -5%, no scale, +5%, +10%, see header of legends). The bottom panel shows the ratio of the medium zenith flux to the low zenith flux. As a reference two historical Crab Nebula spectra (MAGIC, Albert et al. (2008b), and H.E.S.S. Aharonian et al. (2006)) are plotted with a dashed and a solid line.

two most spread apart historical Crab Nebula spectra observed by IACTs. For the low zenith spectrum, as the MC light scale is increased, an increasing pile-up effect shows up at low energies. However, as the observations at higher zenith do not reach such lowest energies, and at the energies of ~ 100 GeV, both spectra are consistent, such an overestimated light scale cannot be excluded a priori. On the other hand, when the MC light scale is decreased, there is a clear difference between the low and medium zenith samples at low energies. Moreover, even while the χ^2 value computed between the two spectra gives an acceptable probabilities of $\sim 16\%$ and $\sim 10\%$ for down scaling by 10% and 25% respectively, there is a clear structure in the flux ratio plot.

A second way to artificially increase the threshold of the MAGIC telescopes is to use a higher *Size* cut. In Fig. 23 we show the Crab Nebula spectra for a low zenith sample obtained for *Size* greater than 50, 100, 200 and 400 phe. As before, we perform a full analysis of the Crab Nebula data for each of the MC samples with artificially shifted light scale. We compute the ratio of the flux obtained with a given cut in the *Size* parameter to the reference cut of 50 phe. Note however that the plotted statistical uncertainties of the flux ratio are overestimated as the points between the spectra with different cuts in *Size* are strongly correlated (the same Crab Nebula sample was used to obtain each of them). In this case, the situation is opposite to the one with the energy threshold varying with the zenith angle of the observations. For the underestimated light scale the spectra at different size cuts are still consistent between each other, thus no constraint on the absolute light scale can be drawn. On the other hand, for the overestimated light scale, the flux ratio while going to lower energies has a characteristic V-shape (best visible with red curves in Fig. 23). Going to lower energies it first slightly drops (due to the direct effect on the collection area) and then sharply increases (consistent with a pile-up effect from the bias in energy estimation below the energy threshold). Combining both methods we validate that the systematic uncertainty of the absolute energy scale of the MAGIC telescopes is below 15%, similarly to the one expected in Aleksić et al. (2012a).

5.4. Systematic uncertainty in flux normalization and slope

Due to better camera and trigger homogeneity after the upgrade and also lack of dead time in the readout the total systematic uncertainty in the flux normalization is slightly lower than evaluated in Aleksić et al. (2012a). We checked the contribution of the mismatch in cut efficiencies to the systematic uncertainty and estimate it to be $< 12\%$ in the whole investigated energy range, compatible with the 10-15% uncertainties due to analysis and data/MC discrepancies reported in Aleksić et al. (2012a). The effect on the spectral indices was smaller than the statistical uncertainty obtained with this data sample. We estimate the uncertainty on the flux normalization to be 18% at low energies ($\lesssim 100$ GeV) and 11% in the energy range of a few hundred GeV. At the highest energies, $\gtrsim 1$ TeV, due to more pronounced MC/data mismatches the systematic uncertainty is a bit higher, namely 16%. The systematic uncertainty on the

reconstructed spectral slope of the sources is still ± 0.15 for a source with signal to background ratio of at least 25%.

5.5. Night to night systematic uncertainty

The total systematic error estimated in the previous section is a proper quantity to be used while comparing the MAGIC observations with the data of other instruments or with theoretical predictions. However, a significant fraction of the systematic error is nearly constant and will affect all the MAGIC data in the same way. On the other hand, part of the systematic uncertainty (e.g. atmospheric transmission at a given day or small changes in the optical PSF) will vary from one night to another resulting in slightly different estimations of the flux even from a steady source. In order to estimate this remaining, relative systematic error we use the method of Aleksić et al. (2012a). We divided our data into sub-samples and compute the flux for each of them. Then we compare the standard deviation, σ_F , of the distribution of the reconstructed fluxes with the typical uncertainty of individual points, δF .

We first divide the data according to data runs. Each data run is normally 20-min long, and correspond to one wobble pointing. With such time binning the integrated flux above 300 GeV is computed with a statistical precision of about 10% (see Fig.24). Three short data runs (with a much higher relative statistical error of the flux estimation, $> 20\%$) were removed from this study. Fitting the points with a constant we obtain $\chi^2/N_{dof} = 106/50$, corresponding to a fit probability of $6.7 \cdot 10^{-6}$. The spread of the individual points is $\sigma_F \sim 15\%$, suggesting a relative run-to-run systematic uncertainty of $\sim 11\%$. Note, that while so far there was no variability observed from the Crab Nebula, it is not excluded that a small intrinsic variability, below the accuracy of the current IACTs is present at the source itself. Therefore, the values derived with this method can be treated as conservative estimates of the run to run systematic uncertainty of the MAGIC telescopes. Adding such a systematic uncertainty of 11% in quadrature to the statistical uncertainties of the individual points (see the dashed error bars in Fig.24), and performing a fit with a constant, we obtain that the $\chi^2/N_{dof} = 46/50$, nominally corresponding to probability of 63%.

In the case of night-by-night binning δF is of the order of 5%. Interestingly, this light curve can be fitted with a constant flux with $\chi^2/N_{dof} = 16.5/9$, which corresponds to a barely acceptable probability of 5.7%. We conservatively conclude that the relative systematic errors are of the order of 11%. This value is similar to the values before the upgrade of the MAGIC telescopes Aleksić et al. (2012a) as well as to the one obtained by the H.E.S.S telescopes (Aharonian et al., 2006). It is plausible that most of this uncertainty is due to the atmospheric variations.

6. Conclusions

The upgrade of the readout and one of the cameras of the MAGIC telescopes have significantly improved their performance. The trigger threshold for low zenith observations is

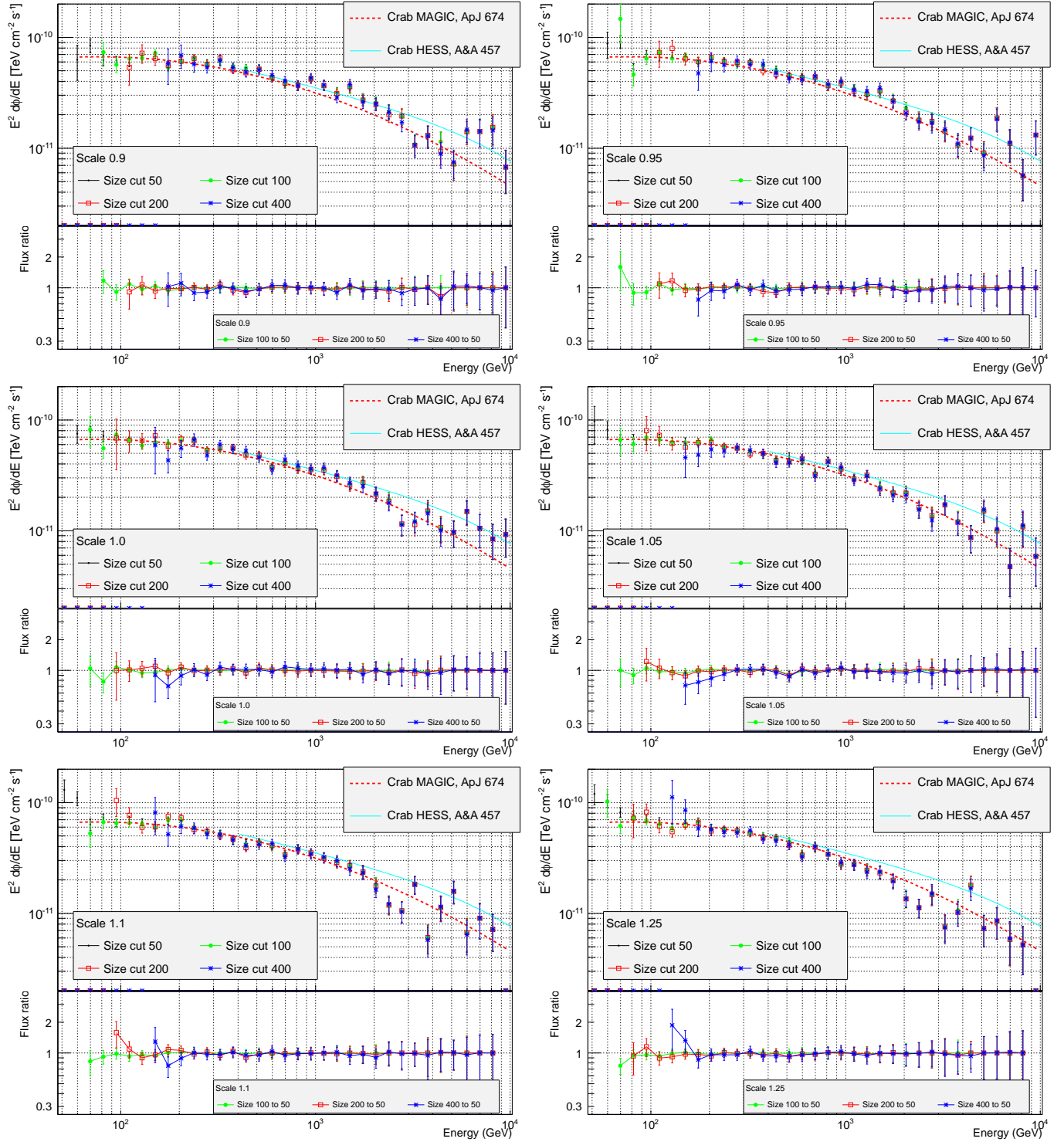


Figure 23: Spectra of the Crab Nebula obtained with the low zenith sample for different *Size* cuts and with scaled MC: -10%, -5%, no scale, +5%, +10% and +25% (see the legend header in different panels). The spectra are obtained with *Size* cuts of 50 phe (black dots), 100 phe (green full circles), 200 phe (red empty squares) and 400 phe (blue asterisks). At the bottom, smaller sub-panels with the ratio between a spectrum with a given *Size* cut and the reference one (> 50 phe) is shown. As a reference, two historical Crab Nebula spectra (MAGIC, Albert et al. (2008b), and H.E.S.S. Aharonian et al. (2006)) are plotted with a dashed and solid line.

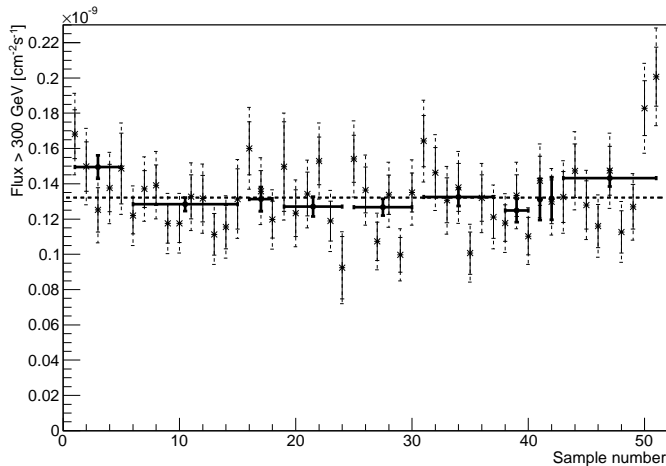


Figure 24: Integrated fluxes above 300 GeV of the Crab Nebula for different data runs (~ 20 min) as a function of the data sample number are shown with thin lines and asterisks. The thick lines show the corresponding night by night fluxes (vertical error bars show the data span). The second set of dashed error bars for the thin lines shows the statistical error added in quadrature with 11% systematic error. The dashed vertical line shows the mean flux obtained from the total data sample.

~ 50 GeV. With the 15h sample of Crab Nebula data, its spectrum could be reconstructed between 65 GeV and 13.5 TeV. Within systematic uncertainties it is consistent with previous measurements of the Crab Nebula with IACTs. The best performance of the MAGIC telescopes is achieved at medium energies, at a few hundred GeV. At those energies the images are sufficiently large to provide enough information for efficient reconstruction, while the rapidly falling power-law spectrum of the Crab Nebula still provides enough statistics. The energy resolution at these medium energies is as good as 16% with a negligible bias and the angular resolution is $\lesssim 0.07^\circ$. The sensitivity above 220 GeV is $(0.66 \pm 0.03)\%$ of C.U. for 50h of observations. At the lowest energies, below 100 GeV, the performance has improved drastically, reducing the needed observation time by a factor of 2.5. The larger trigger region and improved pixelization of the MAGIC-I camera have improved also the off-axis performance. A source with a Crab Nebula like spectrum, but 80 times weaker, can be detected at the offset of 1° within 50h of observations, making the MAGIC telescopes capable of efficient sky scans.

The performance of the MAGIC telescopes at medium zenith angles, $30 - 45^\circ$, is mostly similar to the one at low zenith angles. The higher threshold, however, significantly spoils all the performance parameters below ~ 200 GeV. For the highest energies, above a few TeV, better performance is achieved with observations at medium zenith angles.

We revised different sources of systematic uncertainties after the upgrade and studied in detail the uncertainty in the energy threshold. The larger trigger region has allowed us to lower the systematics connected with the background estimation by a factor of 2. From comparisons of reconstructed SEDs of the Crab Nebula for different energy thresholds we validated the systematic uncertainty in the energy scale to be below 15%. The sys-

tematic uncertainty on the flux normalization was estimated to be 11-18%, and on the spectral slope ± 0.15 . The part of the systematic uncertainty which can change from one observation run to another is estimated to be about 11%.

Thanks to the improvement in the performance achieved after the upgrade, the MAGIC telescopes have reached an unprecedented sensitivity. Since then, five new VHE γ -ray sources have already been discovered by MAGIC. Among them, 3C 58, is the least luminous pulsar wind nebula so far detected in the VHE γ -rays (Aleksić et al., 2014d).

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